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Authors	Pizzuti, Lorenzo, Saltas, Ippocratis D., Casas, Santiago, Amendola, Luca, BIVIANO, ANDREA
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Future constraints on the gravitational slip with the mass profiles of galaxy clusters

Lorenzo Pizzuti,¹★ Ippocratis D. Saltas,² Santiago Casas,³ Luca Amendola⁴ and Andrea Biviano⁵

¹*Dipartimento di Fisica, Sezione di Astronomia, Università di Trieste, Via Tiepolo 11, I-34143 Trieste, Italy*

²*CEICO, Institute of Physics of the Czech Academy of Sciences, Na Slovance 2, CZ-182 21 Praha 8, Czech Republic*

³*AIM, CEA, CNRS, Université Paris-Saclay, Université Paris Diderot, Sorbonne Paris Cité, F-91191 Gif-sur-Yvette, France*

⁴*Institut für Theoretische Physik, Universität Heidelberg, Philosophenweg 16, D-69120 Heidelberg, Germany*

⁵*INAF-Osservatorio Astronomico di Trieste, via G. B. Tiepolo 11, I-34143 Trieste, Italy*

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ABSTRACT

The gravitational slip parameter, defined as the ratio between the two scalar potentials appearing in the linearly perturbed Friedmann–Lemaître–Robertson Walker metric, is an important discriminator between large classes of gravity theories at cosmological and astrophysical scales. In this work we use a combination of simulated information of galaxy cluster mass profiles, inferred by strong + weak lensing analyses and by the study of the dynamics of the cluster member galaxies, to reconstruct the gravitational slip parameter η and predict the accuracy with which it can be constrained with current and future galaxy cluster surveys. Performing a full-likelihood statistical analysis, we show that galaxy cluster observations can constrain η down to percent level already with a few tens of clusters. We discuss the significance of possible systematics, and show that the clusters masses and numbers of galaxy members used to reconstruct the dynamics mass profile have a mild effect on the predicted constraints.

Key words: dark energy – cosmology: theory.

1 INTRODUCTION

The dark energy problem has challenged our understanding of gravity at large scales in the Universe. To explain it, a variety of different theoretical scenarios have been proposed based on the introduction of new degrees of freedom beyond the standard paradigm of general relativity (GR) (for a review see e.g. Amendola & Tsujikawa 2010; Clifton et al. 2012). The recent impressive measurement of the speed of gravitational waves provided the most stringent test on viable modified gravity theories so far, with direct implications for cosmology and the large-scale structure of the Universe. However, understanding the nature of the late-time acceleration of the Universe calls for further, complementary observational information, which could shed light on the viability of the surviving theory space.

From a phenomenological point of view, modifications of GR leave two key imprints at the level of large-scale structures in the Universe. In particular, they affect the effective strength of gravity, impacting the clustering of matter at large scales, while they may also impact the way gravity interacts with light, leaving

a characteristic imprint on weak lensing observables. Both effects, however, could be equally due to new gravitational degree(s) of freedom that do not modify gravity, such as minimally coupled scalar–tensor models (quintessence, k-essence), so that they cannot be used to detect genuine departures from GR. A crucial parameter in this regard is the gravitational slip η defined as the ratio between the effective gravitational coupling of light to that of matter (to be defined later on). In GR and in models where the new degree of freedom is minimally coupled to gravity one has $\eta = 1$, while an $\eta \neq 1$ would be the smoking gun for a departure from GR (see e.g. Amendola et al. 2013a) at late times, at least as long as dark matter can be described as a perfect fluid.¹ Crucially, η can be reconstructed in a model-independent manner by combining galaxy kinematics and weak lensing observations (Amendola et al. 2013b; Motta et al. 2013), while it is directly linked to the properties of gravitational wave propagation at any scale (Saltas et al. 2014; Amendola et al. 2017; Sawicki et al. 2017). Therefore, its reconstruction from observations is key in testing departures from GR and/or

¹We notice here that free streaming neutrinos can also be a source of gravitational slip. However, at late times in the cosmological evolution their contribution can be neglected.

★ E-mail: lorenzo.pizzuti92@hotmail.it

distinguishing amongst large classes of gravity models at both cosmological and astrophysical scales.

Galaxy clusters are ideal laboratories in this regard, as they can provide us with simultaneous information about the local gravitational potentials sourcing the dynamics of their member galaxies as well as the lensing of light, which can be combined to reconstruct η . In this sense, tests of gravity based on η can be understood as a consistency check between the cluster's mass profile as inferred from kinematical and lensing observations, respectively. Current and future surveys observing at different, complementary bands, such as *Euclid*, Dark Energy Survey (DES), *XMM-Newton*, or eROSITA² will provide a large amount of data for hundreds of galaxy clusters, allowing for accurate joint lensing and dynamics mass profile determinations of some dozens of clusters. This work adds to previous efforts to constrain gravity with galaxy clusters in various theoretical set-ups, e.g. in Sakstein et al. (2016), Wilcox et al. (2016), Wilcox et al. (2015), Terukina et al. (2014), and it is complementary to conceptually similar reconstructions of η using redshift-space distortions, galaxy–galaxy lensing and the background expansion (Pinho, Casas & Amendola 2018).

The goal of this is two fold: On the one hand, we aim at predicting the potential statistical accuracy of future constraints on η from galaxy cluster kinematics and lensing, while at the same time understanding their dependence on key factors such as the number of member galaxies and the mass of the clusters in the sample. Our analysis is based on the dynamical mass profile reconstructions, performed with the *MAMPOSST* method of Mamon, Biviano & Boué (2013), of a set of simulated spherically symmetric isolated clusters. This approach will give us enough flexibility to investigate how different cluster masses and densities of tracers affect our constraints. For the complementary lensing information, we will simulate the likelihood distribution relying on the lensing mass profile inferred by Umetsu et al. (2012), Umetsu et al. (2016), Caminha et al. (2016) on the massive, relaxed galaxy cluster MACS 1206, which belongs to a sample of 20 X-ray selected objects analysed within the CLASH/CLASH-VLT collaborations (Postman et al. 2012; Rosati et al. 2014).

It should be noted that theories modifying gravity typically introduce a characteristic time and scale dependence on the gravitational slip parameter $\eta = \eta(t, k)$, which is often captured through appropriate parametrizations. Unlike previous constraints of gravity (Sakstein et al. 2016; Wilcox et al. 2016, 2015; Terukina et al. 2014) based on particular gravity set-ups, in our analysis we will make no assumption on the particular functional dependence of η on time and scale, the only assumption being that of a Navarro–Frenk–White parametrization for the cluster mass profile. In this regard, our constraints will be sufficiently general – any departures from the expectation $\eta = 1$ is incorporated in a difference between the mass profile parameters inferred by dynamics and lensing probes, which can then be translated into constraints on the parameter space of particular theories of gravity.

We show that already with ~ 15 clusters we can get at 1σ an accuracy of ~ 5 per cent for a scale independent and ~ 10 per cent for a scale-dependent η . We investigate the effect of varying the number of tracers (i.e. member galaxies) in the dynamics analysis, the dependence of constraints on the cluster mass, as well as the uncertainties in the simulated lensing probability distribution.

We structure the paper as follows. Section 2 lays down the main assumptions and definitions of our analysis, while in Section 3 we introduce the essential theoretical framework. A detailed exposition of all our simulations is presented in Section 4, and finally, the results of the statistical analysis are explained in Section 5. The main results are summarized in the last section of the paper, while a summary of the predicted constraints can be found in Table 2.

2 ESSENTIAL ASSUMPTIONS AND DEFINITIONS

An essential step towards forecasting the potential of future galaxy–cluster surveys to constrain the gravitational slip parameter is having a handle upon a measurement of clusters dynamical and lensing masses, as we explain in more detail in Section 3. In this regard, we will combine information from a set of reliable simulated data based on current and future observations. The analysis will inevitably rely on certain fundamental simplifications and assumptions about the structure and dynamics of the clusters and the underlying cosmology, which is important to lay out before we proceed. They can be summarized as follows:

(i) The clusters will be modelled as spherical inhomogeneities in a complete equilibrium, that is, in thermal and hydrostatic equilibrium. In other words, the clusters will be assumed to be relaxed, with their mass distribution and gravitational potentials being time-independent. Although from an observational viewpoint it might be challenging to identify truly relaxed clusters, this simplification will allow us to avoid the introduction of systematics and/or degeneracies that we would otherwise encounter when dealing with a non-relaxed cluster. We comment on the issue of departures from relaxation in Section 3. Additional assumptions related to the anisotropy of the cluster's velocity field are also explained in Section 4.1.

(ii) The density distribution of the cluster will be dominated by dark matter, which we will model by a Navarro–Frenk–White (NFW) profile. The member galaxies will be assumed to be collisionless and non-interacting point particles moving along geodesics and populated according to the same profile. The validity of this assumption is ensured in the case of a relaxed cluster, where galaxy–galaxy encounters are negligible and the dynamics of member galaxies traces the total gravitational potential, whereas it is expected to break down as soon as non-linear galactic interactions become important. In this case the Jeans equation is no more valid and cannot be used to reconstruct the cluster mass profile.

(iii) The background cosmology will be that of the Λ CDM model, while large-scale structures will be sourced by linear scalar fluctuations Ψ and Φ appearing in the linearly perturbed Friedmann–Robertson–Walker metric. In spherical coordinate, when the conformal Newtonian gauge choice (see Mukhanov 2001) is adopted, the perturbed metric reads

$$ds^2 = -(1 + \Psi(t, \mathbf{x}))dt^2 + a(t)^2(1 + \Phi(t, \mathbf{x}))d\mathbf{x}^2; \quad (1)$$

Φ and Ψ are generally called Bardeen potentials. Under this convention, we define the gravitational slip parameter as $\eta \equiv \Phi/\Psi$. The choice of a Λ CDM background guarantees agreement with observations of the Universe's background history (e.g. from supernovae). A modification of gravity will then introduce new degrees of freedom and effects which change the evolution of large-scale perturbations already at the linear level (see e.g. Amendola & Tsujikawa 2010.)

(iv) The functional form of the gravitational slip parameter $\eta = \eta(t, k)$ as a function of time and scale depends on the particular

²<http://www.darkenergysurvey.org/>, <http://www.mpe.mpg.de/heg/www/Projects/erosita/index.php>, <http://sci.esa.int/xmm-newton/43133-slew-survey-and-catalogue/>. For a review on future surveys see Sartoris et al. (2012).

gravity action under consideration. Here, we shall make no a priori assumption about the actual gravity model, proceeding in a sufficiently general manner considering two distinct cases: A case where η is assumed to be scale independent throughout our clusters sample, and another one where it depends on the cluster's spatial scale through the characteristic scale r_s of the NFW profile function. These two cases encompass the predictions of scalar–tensor theories at different (subhorizon) scales, as explained in Section 3.

3 GRAVITATIONAL SLIP AS A CONSISTENCY CONDITION ON CLUSTER DYNAMICAL AND LENSING MASS

Testing the gravitational slip at the galaxy cluster level can be understood as a consistency check between the lensing and dynamical mass of the cluster, M_{dyn} and M_{lens} , respectively. In GR, $M_{\text{dyn}} = M_{\text{lens}}$, which is no longer true as soon as gravity is modified.³ To understand why this is so, let us assume a cluster to be in a state of dynamical equilibrium with the total gravitational potential Ψ . If interactions and dissipative effects are negligible, the cluster member galaxies act like collisionless tracers of the gravitational potential and their dynamics are governed by the Jeans equation:

$$\frac{\partial(v\sigma_r^2)}{\partial t} + 2\beta(r)\frac{v\sigma_r^2}{r} = -v(r)\frac{\partial\Psi}{\partial r}, \quad (2)$$

with $v(r)$ the number density of tracers, σ_r^2 the velocity dispersion along the radial direction and $\beta \equiv 1 - (\sigma_\theta^2 + \sigma_\phi^2)/2\sigma_r^2$ the velocity anisotropy profile. The gravitational potential Ψ is connected to the matter density distribution through the Poisson equation:

$$\nabla^2\Psi = 4\pi G\rho_m, \quad (3)$$

where ρ_m stands for the total mass density of the cluster at a given radius r . The total density is assumed to be dominated by the dark matter component, i.e. $\rho_m = \rho_{\text{DM}} + \rho_{\text{gas}} + \rho_{\text{gal}} \simeq \rho_{\text{DM}}$, as baryons provide a negligible contribution. The total mass of the (spherical) cluster with total radius R is given by the mass conservation equation:

$$M_{\text{tot}}(R) = 4\pi \int_0^R r^2 \rho_m(r) dr. \quad (4)$$

In order to model the dark matter density we will follow a parametric approach through the use of an NFW profile Navarro, Frenk & White (1997), which describes the distribution of collisionless cold dark matter (CDM) in spherical cluster-sized haloes in an equilibrium configuration, defined through

$$\rho_{\text{NFW}}(r) = \frac{\rho_s}{(r/r_s)(1+r/r_s)^2}, \quad (5)$$

with ρ_s and r_s two free parameters to be fit from observations. The NFW profile is in very good agreement with N -body simulations. In particular, r_s is a characteristic scale radius, corresponding to the distance from the centre of the cluster where $d(\ln \rho)/d \ln r = -2$, while ρ_s is defined as

$$\rho_s = \rho_c(z) \frac{\Delta_{\text{vir}}}{3} \frac{c^3}{\log(1+c) - c/(1+c)}. \quad (6)$$

Here, $\rho_c(z) = 3H^2(z)/8\pi G$ is the critical density of the Universe at redshift z , and $c \equiv r_{\text{vir}}/r_s$ is the so-called *concentration* parameter,

³Our definition of modified gravity is that of Amendola et al. (2017), defined as any theory modifying the dynamics of the genuine degrees of freedom of GR, that is the two polarizations of the spin-two field.

while r_{vir} denotes the radius enclosing a mean overdensity Δ_{vir} times the critical density of the Universe. In this paper, we will use $\Delta_{\text{vir}} = 200$ (and so, $r_{\text{vir}} = r_{200}$), which is close to the density at virialization predicted by the spherical collapse model in an EdS universe $\Delta_{\text{sc}} = 178$.

By inserting equations (6) into (4) we obtain the halo mass as a function of radius as:

$$M(r) = M_{200} \frac{\ln(1+r/r_s) - (r/r_s)(1+r/r_s)^{-1}}{\ln(1+c) - c/(1+c)}, \quad (7)$$

with $M(r_{200}) = M_{200}$ the total mass enclosed in a sphere of radius r_{200} . Despite its relative simplicity, the NFW profile has been widely adopted in the literature to describe the mass profile of gravitationally bound structures, as it has been shown to provide a generally good fit to observational data (e.g. Biviano et al. 2013; Umetsu et al. 2016). The choice of a parametric description of the density perturbations has the noticeable advantage to simplify the extension of the analysis to specific modified gravity theories, where departures from GR can be easily incorporated in additional parameters in the model, as we will discuss in more detail later on. Now, assuming spherical symmetry the Poisson equation can be integrated to yield

$$\Psi(r) = G \int_0^r \frac{ds}{s^2} M_{\text{dyn}}(s). \quad (8)$$

The above equation provides a definition for a cluster dynamical mass, that is, the mass one would infer through galaxy kinematics measurements, related to the Newtonian potential Ψ . We can derive a similar definition for the lensing mass considering that, due to the conformal structure of covariant Maxwell's equations, light ray propagation is affected by both time–time and space–space component of the metric. In this regard, we define the lensing or Weyl potential as (see e.g. Pratten et al. 2016) $\Phi_{\text{lens}} \equiv (\Phi + \Psi)/2$, which satisfies the Poisson-type equation

$$\nabla^2(\Phi + \Psi) = 8\pi G\rho_{\text{lens}}. \quad (9)$$

Under the assumption of spherical symmetry we have

$$M_{\text{lens}} = \frac{r^2}{2G} \frac{d}{dr}(\Phi + \Psi). \quad (10)$$

Equations (8) and (9) allow us to solve for the potential Φ ,

$$\Phi(r) = G \int_0^r \frac{ds}{s^2} [2M_{\text{lens}}(s) - M_{\text{dyn}}(s)]. \quad (11)$$

Combining the expressions for the potentials Φ and Ψ we can compute the ratio Φ/Ψ , to arrive to a relation for the gravitational slip parameter η in terms of the lensing and dynamical mass of the cluster as,

$$\eta(r) = \frac{\int_0^r \frac{ds}{s^2} [2M_{\text{lens}}(s) - M_{\text{dyn}}(s)]}{\int_0^r \frac{ds}{s^2} M_{\text{dyn}}(s)}. \quad (12)$$

Equation (12) will be our starting point in predicting the constraints on η at the cluster level. Notice that, in GR it is $M_{\text{dyn}} = M_{\text{lens}}$ and $\eta = 1$.

In the usual description of η in general classes of gravity models, one usually works in momentum (Fourier) space, where the scale dependence is controlled through the Fourier wavenumber k . For example, for the popular class of the general Horndeski scalar–tensor theories in the sub-horizon quasi-static regime η acquires the form

$$\eta = h_2 \times \frac{1 + h_4 k^2}{1 + h_5 k^2}, \quad (13)$$

with the functions $h_i = h_i(t)$ defined in De Felice, Kobayashi & Tsujikawa (2011). For the particular case of Brans–Dicke/ $f(R)$ gravity they depend on the scalaron Compton wavelength. However, since our simulations, and hence our analysis, will be in real space it is essential to discuss how our parametrization relates to the Fourier space analyses. To get the real space η one would have to express the linearized equations for Φ and $\Phi + \Psi$ in real space, assuming a parametrization for the density profile. In general, we expect the resulting expression for generic theories of gravity to acquire the following schematic form

$$\eta(r, t) = f_1(t)(1 + f_2(r, t)), \quad (14)$$

with GR corresponding to the limit $f_1 \rightarrow 1, f_2 \rightarrow 0$. A scale-independent η is sourced only by f_1 , which is in principle time-dependent – in our analysis, we will confine ourselves to a fixed redshift, which enforces η to be a constant for this case. Without the assumption of a model-dependent parametrization our approach will be rather phenomenological, yet sufficiently general to encompass any gravity theory. In particular, in view of equation (12) and the assumption of an NFW profile for the matter distribution, it is easy to see that η will be sourced through a difference of the inferred parameters r_{200}, r_s from lensing and dynamics respectively, i.e

$$\eta(r; r_{200}^{\text{dyn}}, r_s^{\text{dyn}}) \neq \eta(r; r_{200}^{\text{lens}}, r_s^{\text{lens}}), \quad (15)$$

unless $(r_{200}^{\text{dyn}}, r_s^{\text{dyn}}) = (r_{200}^{\text{lens}}, r_s^{\text{lens}})$. From an observational viewpoint, the existence of some new degree of freedom sourcing $\eta \neq 1$, will manifest itself as a tension between the inferred NFW parameters from lensing and dynamics, respectively, when trying to reconstruct the mass distribution of the cluster. In this regard, the strategy in our analysis will be as follows: the synthetic clusters will be constructed with different initial conditions implying different distribution for the inferred NFW profiles from dynamics, $(r_{200}^{\text{dyn}}, r_s^{\text{dyn}})$. The resulting distributions from the dynamics will be in principle different from those from lensing, hence sourcing gravitational slip under equation (12). The derived constraints on η can then be mapped to specific theories of gravity. Below, we explain the simulations and statistical analysis in detail, before we dwell upon the predicted constraints.

4 GALAXY CLUSTER SIMULATIONS AND STATISTICAL ANALYSIS

Predicting the constraints on η from current or future galaxy cluster observations requires a sample of galaxy clusters along with predictions for the dynamical and lensing mass of each member of the sample according to equation (12). The *principal aim of our analysis* will be to estimate the number of reliable lensing and dynamics mass reconstructions, which will be available from upcoming imaging and spectroscopic surveys, needed in order to constrain η at the *percent level* of statistical accuracy. For this purpose, we will neglect the impact of systematic effects in mass profile determinations, e.g. deviation from dynamical relaxation and spherical symmetry, or errors introduced by non-suitable parametrizations of the density profile. Therefore, we will account only for the statistical power of the presented method – in particular, we will consider reliable probability distributions in the mass profile parameter space as would be obtained from the analysis of high-quality imaging and spectroscopic data. An extensive study of the impact of systematics in constraining modified gravity models with cluster mass profiles will be performed in a forthcoming work (Pizzuti et al., in preparation).

Table 1. A summary with the characteristics of the synthetic clusters sample used in our analysis ordered according to increasing mass. The second column shows the parameters r_{200} for each cluster, sampled with a fixed step size of 0.15 Mpc, while the third column the corresponding r_s parameter derived according to the empirical concentration–mass relation (18). The fourth column shows the scale radius r_v , obtained by fitting the projected number density profile from the phase space with a projected NFW model, which is additionally required in the *MAMPOSSt* procedure. The fifth column denotes the number of tracers within a sphere of radius $r = r_{200}$, computed according to equation (17).

Cluster ID	r_{200} (Mpc)	r_s (Mpc)	r_v (Mpc)	$N_g(r_{200})$
1	0.750	0.185	0.169	300
2	0.900	0.094	0.097	324
3	1.050	0.177	0.178	343
4	1.200	0.133	0.144	384
5	1.350	0.231	0.219	437
6	1.500	0.258	0.256	500
7	1.650	0.227	0.252	570
8	1.800	0.360	0.387	648
9	1.950	0.511	0.466	732
10	2.100	0.515	0.473	823
11	2.250	0.539	0.503	920
12	2.400	0.614	0.641	1024
13	2.550	0.848	0.807	1133
14	2.700	0.993	1.005	1249
15	2.850	1.112	1.117	1371

The analysis carried out in Pizzuti et al. (2016) on the X-ray selected CLASH cluster MACS 1206 has shown that accurate dynamics and lensing mass profile reconstructions allow to constrain $\eta(r)$ down to 30 per cent statistical uncertainties with a single galaxy cluster when the systematic effects are under control. It should be noted that the requirement of accurate lensing mass measurements restricts the availability of accessible data. Such measurements, for example, are outside the reach of upcoming surveys such as *Euclid* or LSST.

In reconstructing a cluster mass profile from the dynamics of the member galaxies we proceed in *two steps* as follows: we first generate a synthetic sample of 15 equilibrium configurations of isolated, self-gravitating systems populated by collisionless point-like tracers according to a spherical NFW profile. Each synthetic cluster is constructed with a different characteristic mass, corresponding to different values of the virial radius r_{200} in the NFW profile, covering two orders of magnitude in cluster masses (see Table 1). As for the scale radius r_s , we assigned the value of this parameter according to the mass–concentration relation of Merten et al. (2015), as discussed in Section 4.1. An interesting outcome of our analysis, as we will see in Section 5, is that the constraints do not show an appreciable dependence on halo masses and concentrations. The velocity dispersion of the tracers is assigned following the solution of the Jeans equation, as explained in Section 4.1. The particular number of 15 synthetic haloes is chosen as a fair compromise between computing time and a statistically representative sample of cluster masses and concentrations range. The second step, discussed in Section 4.2, consists in inputting positions and velocities of the halo’s particles into the *MAMPOSSt* procedure of Mamon et al. (2013), to perform a maximum-likelihood fit of the cluster mass profile based on the tracers of the velocity field. For the lensing part we simulate the probability distribution $P_L(r_s, r_{200})$ relying on the results of the analyses of Umetsu et al. (2012), Caminha et al. (2016), and Umetsu et al. (2016), which derived the strong + weak lensing mass profile for the galaxy cluster MACS 1206. The profile

is parametrized as an NFW model with ~ 6 per cent and 30 per cent uncertainties on the mass profile parameters r_{200} and r_s , respectively. Moreover, it was found that the joint probability distribution of the NFW parameters is almost Gaussian with a non-zero covariance between r_s and r_{200} .

We thus use a bivariate Gaussian distribution,

$$P_L(r_s, r_{200}) = \frac{1}{2\pi\sigma_{r_s}\sigma_{r_{200}}\sqrt{1-\rho^2}} \exp \left\{ - \left[\frac{(r_s - \bar{r}_s)^2}{\sigma_{r_s}^2} + \frac{(r_{200} - \bar{r}_{200})^2}{\sigma_{r_{200}}^2} - \frac{2\rho(r_s - \bar{r}_s)(r_{200} - \bar{r}_{200})}{\sigma_{r_s}\sigma_{r_{200}}} \right] \frac{1}{2(1-\rho^2)} \right\}, \quad (16)$$

assumed to be centred on the best-fitting values of the NFW parameters $\bar{r}_s, \bar{r}_{200}$ given by the dynamics reconstruction, as we will explain in Section 5. Here ρ indicates the correlation. In general, the virial radius r_{200} , which is related to the total cluster mass, can be constrained much better than the characteristic radius of the halo mass profile, expressed in terms of r_s . As show in table 2 of Umetsu et al. (2016), typical uncertainties on the scale radius given by a lensing probe are of the order of ~ 30 – 40 per cent, while r_{200} can be recovered up to ~ 5 – 10 per cent. For our reference analysis we consider $\sigma_{r_{200}} = 0.07 \times r_{200}$ and $\sigma_{r_s} = 0.30 \times r_s$, while for the correlation in $P_L(r_s, r_{200})$ we adopt the value $\rho = 0.67$, found by fitting a bivariate Gaussian on the posterior distribution of Umetsu et al. (2012). It is worth to point out that the accuracy of lensing and dynamics cluster mass profile reconstructions expected from future surveys will be much lower than the results of the CLASH and CLASH-VLT collaborations. In Section 5, we will further discuss the robustness of our results when varying both correlation and variances in the lensing distribution and the number of tracers used in the dynamical fit.

4.1 Synthetic galaxy clusters

Here we discuss the procedure for the generation of the simulated clusters and we list the main properties of the 15 produced synthetic haloes we will use in our analysis. Under the assumption of dynamical relaxation, member galaxies in clusters act as collisionless tracers of the total gravitational potential and their dynamics is governed by the Jeans equation. In order to avoid systematic effects induced by lack of equilibrium in the mass profile determination, we generate each halo as an isolated system of collisionless particles assuming spherical symmetry, dynamical relaxation, and 3D-Gaussian distribution of the velocities of the tracers (galaxies). Given a Λ CDM background with $H_0 = 70 \text{ km Mpc}^{-1} \text{ s}^{-1}$, $\Omega_m = 0.3$, and $\Omega_\Lambda = 0.7$,⁴ we populate each halo at redshift $z = 0$ according to an NFW density profile, equation (7), up to ~ 9 virial radii.

It is natural to expect that the accuracy of the cluster's reconstructed dynamical mass will depend on the total number of its galaxy members within r_{200} . For this purpose, we perform our analysis for two different choices in this regard. In the first case, all 15 synthetic clusters are generated with the same number of tracers equal to 1000, an optimistic, although not unrealistic choice, given the expectation of the number of cluster galaxies with accurate measured spectroscopic redshift from present and future surveys. In the second case we assigned the number of tracers assuming that it is dependent on the cluster mass, requiring a minimum of

300 particles within r_{200} for the least massive halo in the sample. In particular, for the total number of tracers within r_{200} of the i -th cluster $N_g^{(i)}(r_{200})$ we use

$$N_g^{(i)}(r_{200}) = \begin{cases} 300 & i = 1, \\ \left(M_{200}^{(i)} / M_{200}^{(1)} \right) \times \frac{N_c}{i} & i > 1. \end{cases} \quad (17)$$

We choose $N_c = 375$ in order to guarantee a maximum number of tracers less than 1400 for the most massive cluster. It is worth to notice that the choice of equation (17) is totally arbitrary and still an optimistic simplification, since the number of spectroscopically confirmed member galaxies in clusters strongly depends on several factors, such as observational time, completeness of the sample (see e.g. Biviano et al. 2013) and contamination from field galaxies. The resulting $N_g(r_{200})$ is displayed in the fifth column of Table 1.

The end result of the sample consists of 15 objects with masses spanning a range from 5.1×10^{13} to $2.9 \times 10^{15} M_\odot$. For each halo, we determine the scale radius r_s according to the c - M relation (extrapolated at $z = 0$) of Merten et al. (2015),⁵ which analysed 19 of the X-ray selected clusters of the CLASH sample at $z \sim 0.19$ – 0.89 to find

$$c(M_{200}, z) = A \times \left(\frac{1.37}{1+z} \right)^B \times \left(\frac{M_{200}}{8 \times 10^{14} M_\odot h^{-1}} \right)^K, \quad (18)$$

with

$$A = 3.66 \pm 0.16, \quad B = -0.14 \pm 0.52, \quad K = -0.32 \pm 0.18. \quad (19)$$

We have randomly sampled the value of the parameters A , B , and K for every halo assuming that they are Gaussian-distributed around the central value and standard deviation of (19). The characteristic end-values for each synthetic cluster are summarized in Table 1.

In order to better explore the dependency of the derived constraints on the density of tracers used in the dynamics analysis, we further consider an additional synthetic cluster with $r_{200} = 1.96 \text{ Mpc}$, $r_s = 0.27 \text{ Mpc}$ generated six times by varying the number of particles $N_g(r_{200})$ from 50 to 2000.

To each particle at a radial distance r from the centre we assign a velocity whose components in spherical coordinates are Gaussian-distributed with a squared dispersion $\sigma_r^2(r) = \{\sigma_r^2(r), \sigma_\theta^2(r), \sigma_\phi^2(r)\}$, given by the solution of the (spherical) Jeans equation, equation (2). Note that the velocity anisotropy profile $\beta(r)$ is in principle unknown, and one of the major sources of systematics in the reconstruction of cluster mass profile with the analysis of the dynamics of member galaxies. Different methods have been developed in the literature to reconstruct $\beta(r)$, e.g. through fitting suitable parametrized profiles, or non-parametric techniques based on inversions of the Jeans equation (e.g. Binney & Mamon 1982; Host et al. 2009), however, the latter generally requires additional information and assumptions. In our case, we will use a fixed parametric form for $\beta(r)$, namely the Tiert model ‘T’ of Tiert et al. (2007),

$$\beta(r)_T = \beta_\infty \times \frac{r}{r + r_c}, \quad (20)$$

where the normalization of β_∞ identifies the value of the anisotropy at large radii, while r_c is a characteristic scale radius. The Tiert profile is a generalized version of the Mamon and Lokas profile (Mamon & Lokas 2005), which has been shown to provide a good fit to the average cluster-sized haloes anisotropy profile over a set

⁴Note that the background cosmology enters only in the definition of M_{200} for each cluster.

⁵See also Duffy et al. (2008).

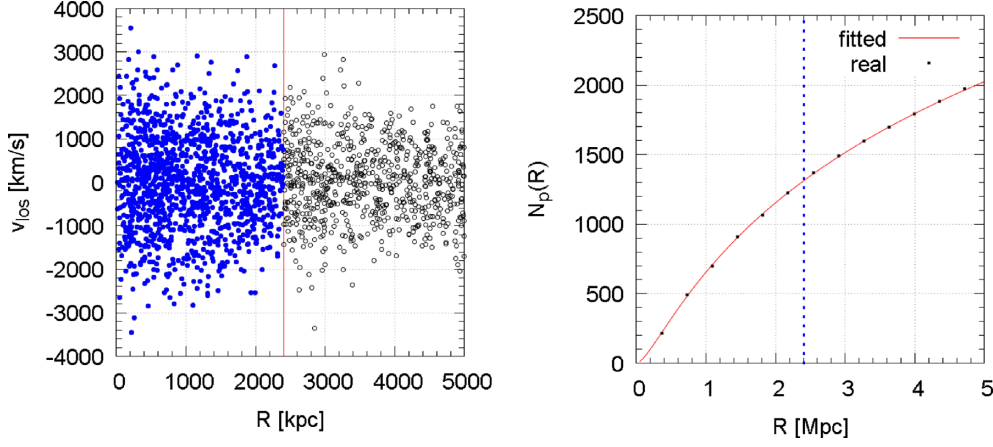


Figure 1. *Left:* Projected phase space ($R, v_z = v_{\text{los}}$) for one of the synthetic clusters (ID 12) generated with $N(r_{200}) = 1024$. Each point indicates a cluster member; the red vertical line denotes the virial radius r_{200} . With blue points we mark the tracers used in the dynamical mass profile reconstruction through the *MAMPOSSt* procedure. *Right:* Cumulative projected number density profile of the tracers as a function of projected radius R (black points). The red solid curve shows the best-fitting projected NFW profile with scale radius r_v indicated in Table 1; vertical blue dotted line corresponds to $R = r_{200}$.

of cosmological simulations (see e.g. Mamon, Biviano & Murante 2010). Here, we will assume that r_c coincides with the scale radius of the mass profile r_s , and we will also fix $\beta_\infty = 0.5$ to generate our sample. With the above, the solution for the radial dispersion σ_r^2 is given by Mamon & Łokas (2005)

$$\sigma_r^2(r) = \frac{1}{v(r)} \int_0^\infty \exp \left[2 \int_r^s \beta(t) \frac{dt}{t} \right] v(s) \frac{d(\Psi(s))}{ds} ds, \quad (21)$$

which is connected to the angular one through

$$\sigma_\theta^2(r) = [1 - \beta(r)] \sigma_r^2(r). \quad (22)$$

Since we are assuming spherical symmetry we impose $\sigma_\phi^2 \equiv \sigma_\theta^2$.

4.2 Reconstructing the clusters dynamical mass: the *MAMPOSSt* method

Armed with the previously constructed synthetic clusters, each with a given (r_s, r_{200}) and a velocity distribution for its tracers, our next step is the reconstruction of what would be the cluster's mass profile from the kinematical data.

To this purpose, we employ the *MAMPOSSt* (*Modeling Anisotropy and Mass Profiles of Observed Spherical Systems*) procedure developed by Mamon et al. (2013), which provides a numerical framework for deriving the inferred cluster mass profiles from the analysis of the dynamics of the member galaxies. The code performs a maximum-likelihood fit to the distribution of the positions and velocity field of the galaxies in projected phase space (R, v_z), where R is the projected radius of each tracer from the cluster center and v_z the velocity along the line of sight (los), in the cluster rest frame. As an example, in the left-hand panel of Fig. 1 we show the projected phase space of the 12th synthetic cluster in the sample, generated with 1024 particles within r_{200} . The radius $R = r_{200}$ is identified by the red vertical line, while the blue points indicate all the cluster members lying in projection within r_{200} .

Given a parametric form of the mass profile $M(r)$ and of the velocity anisotropy profile $\beta(r) = 1 - \sigma_\theta^2/\sigma_r^2$, the code solves the spherical Jeans equation (2) to derive the probability density of finding an object at position (R, v_z) in the projected phase space as

$$q(R, v_z) = \frac{2\pi R g(R, v_z)}{N_p(R_{\text{max}}) - N_p(R_{\text{min}})}. \quad (23)$$

Here $N_p(R)$ stands for the predicted cumulative number of objects at projected radius R , and $g(R, v_z)$ is the surface probability density of observed objects, which in the case of a 3D-Gaussian velocity distribution takes the form

$$g(R, v_z) = \sqrt{\frac{2}{\pi}} \int_R^\infty \frac{rv(r)}{\sqrt{r^2 - R^2}} \frac{dr}{\sigma_r(r) \sqrt{1 - \beta(r) R^2/r^2}} \times \exp \left[-\frac{v_z^2}{2(1 - \beta(r) R^2/r^2) \sigma_r^2(r)} \right]. \quad (24)$$

Given equation (23), the likelihood is simply obtained by the sum over all the tracers in the projected phase space,

$$-\log \mathcal{L} = -\sum_{i=1}^n \log q(R_i, v_{z,i} | \vec{\theta}), \quad (25)$$

where as usual, $\vec{\theta}$ stands for the model-parameters vector. *MAMPOSSt* additionally requires a parametric model of the projected number density profile of the tracers (see Mamon et al. 2013), with a characteristic scale radius r_v . Since we are working with collisionless particles, by construction the number density profile $v(r)$ scales exactly as the NFW mass profile $\rho(r)$ (i.e. $r_v \equiv r_s$), however, the projected number density profile, needed to compute the probability of equation (23), is obtained in *MAMPOSSt* by integrating the three-dimensional profile along the line of sight, assuming that it extends to infinity. This leads to a value of r_v which can be slightly different from r_s and for this reason we fit the projected number density profile from the phase space of each synthetic cluster and use the best-fitting values of r_v as the input for the *MAMPOSSt* analysis (see the fourth column of Table 1). This is the so-called ‘Split’ running mode of *MAMPOSSt*, widely used in the literature (e.g. Biviano et al. 2013; Balestra et al. 2016). This is also more reliable from an observational point of view: in a real cluster it is not guaranteed that the distribution of the tracers (i.e. the member galaxies) scales in the same as the total density profile (see e.g. Biviano & Salucci 2006; Budzynski et al. 2012).

5 RESULTS

We now proceed with the results of our analysis aimed at constraining the gravitational slip η from the sample of simulated synthetic

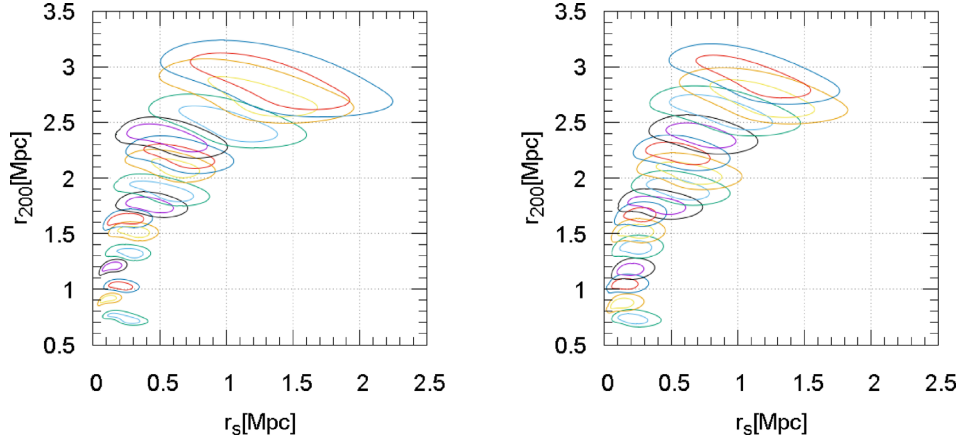


Figure 2. 1σ and 2σ contours from the *MAMPOSSt* analysis of the sample of synthetic clusters in a sequence of increasing input r_{200} (cf. Table 1). The *left* plot shows the case where a fixed number of tracers are considered within r_{200} , while the *right* plot refers to the case where the number of tracers is proportional to the cluster mass (see Section 4.1).

clusters described in Section 4. In particular, we will apply the *MAMPOSSt* procedure to each halo in order to simultaneously constrain the mass profile parameters (r_s , r_{200}) and the velocity anisotropy parameter β_∞ , considering all the particles with a projected radius $R \in [0.05 \text{ Mpc}, r_{200}]$ in the fit. This choice has been made to be consistent with real observations, although the range on which we perform the fit could be arbitrary in this ideal case. Indeed, for an observed cluster we cannot assume the validity of the Jeans equation beyond r_{200} , while at very small radii the dynamics is dominated by the presence of the brightest central galaxy (BCG), which is the dominant galaxy that sits at the bottom of the potential well of relaxed galaxy clusters (see e.g. Biviano et al. 2013). The final likelihood distribution of dynamics $\mathcal{L}^{\text{dyn}} = \mathcal{L}^{\text{dyn}}(r_{200}^{\text{dyn}}, r_s^{\text{dyn}})$ is then obtained by marginalizing over β_∞ . In Fig. 2, we show the resulting 1σ and 2σ iso-probability contours for all the synthetic clusters obtained for what will be our *reference case* with $N_g(r_{200}) = 1000$ (left plot), and the case with a variable number of tracers within r_{200} (right plot). It is worth to point out that the best-fitting values $\bar{r}_s$, $\bar{r}_{200}$ given by the *MAMPOSSt* analysis are not expected to coincide with the true values listed in Table 1, even if these are always included within the 1σ region. This fact is not surprising, since each synthetic projected phase space represents only one possible statistical realization of a halo with given mass, concentration, and anisotropy (we stress again that velocities of the tracers are randomly generated over a Gaussian distribution with given $\tilde{\sigma}_r^2(r)$). For the purposes of our analysis, we assume as fiducial model the GR expectation $\eta = 1$, which is obtained by imposing that the lensing probability distribution $\mathcal{L}^{\text{lens}} = P_L(r_{200}^{\text{lens}}, r_s^{\text{lens}})$ peaks at the best-fitting values $\bar{r}_s$, $\bar{r}_{200}$ given by the dynamics, through an appropriate rescaling which preserves the relative errors $(\sigma(r_s)/r_s)^{\text{lens}}$ and $(\sigma(r_{200})/r_{200})^{\text{lens}}$.

5.1 Case A: scale-independent η

The first part of the analysis considers η to be constant at all scales. This implies that in equation (12) the scale radius of the NFW profile should be the same as inferred by a lensing or dynamics probe, i.e. $r_s^{\text{lens}} = r_s^{\text{dyn}} \equiv r_s$. Thus, the expression of η as a function of the mass profiles M_{lens} , M_{dyn} reduces to a constant ratio where the only relevant parameters are the virial radii

$$r_{200}^{\text{lens}}, r_{200}^{\text{dyn}},$$

$$\eta = 2 \times \frac{\ln(1 + c^{\text{dyn}}) - c^{\text{dyn}}/(1 + c^{\text{dyn}})}{\ln(1 + c^{\text{lens}}) - c^{\text{lens}}/(1 + c^{\text{lens}})} \left(\frac{r_{200}^{\text{lens}}}{r_{200}^{\text{dyn}}} \right)^3 - 1, \quad (26)$$

where $c^{\text{lens/dyn}} = r_{200}^{\text{lens/dyn}}/r_s$ are the lensing and dynamics concentrations, respectively. Following the procedure of Pizzuti et al. (2016), each of the 15 pairs of joint dynamics-lensing likelihoods are Monte Carlo sampled for 10 000 sets of values $(r_s, r_{200}^{\text{dyn}}, r_{200}^{\text{lens}})$ to obtain the corresponding distribution for η according to equation (26). The analysis is iterated for the case of synthetic clusters produced with a fixed and a variable number of tracers, respectively (15 clusters for each case), as explained in Section 4.1.

The results are shown in Figs 3 and 4. In particular, Fig. 3 shows the 1σ and 2σ constraints on η from each individual cluster, while Fig. 4 shows the combined constraints from many clusters. An important feature shown in Fig. 3 is that the constraints on η are practically insensitive to the choice of the number of member galaxies used within the *MAMPOSSt* fit, masses, and concentration of the cluster, in agreement with the discussion of Mamon et al. (2013). From an observational viewpoint, this is particularly encouraging, since it broadens the range of clusters that may be used from current or future surveys. Our results are in agreement with Mamon et al. (2013) which used simulated DM haloes to test the performance of the *MAMPOSSt* code and found that there is only a modest reduction in the uncertainties in mass profile parameters, when the number of tracers increases from 100 to 5000. Furthermore, from Fig. 4 it can be seen that the error on η obtained by combining together the distributions of many clusters (N_{clusters}) scales approximately as the expected theoretical behaviour $\sim (N_{\text{clusters}})^{-1/2}$ (black dashed line in all panels). We predict that, for our reference analysis (top left panel) the *combined analysis of 15 clusters* lead to a constraint

$$\eta = 1.00 \pm 0.05 \text{ at } 1\sigma. \quad (27)$$

When varying the number of particles in the dynamics fit (top right-hand panel) the scaling relation mildly deviates from $(N_{\text{clusters}})^{-1/2}$, as a consequence of the slightly larger uncertainties on η from the analysis of the smallest haloes. The effect is however averaged out when increasing N_{clusters} and the scaling is almost identical to the

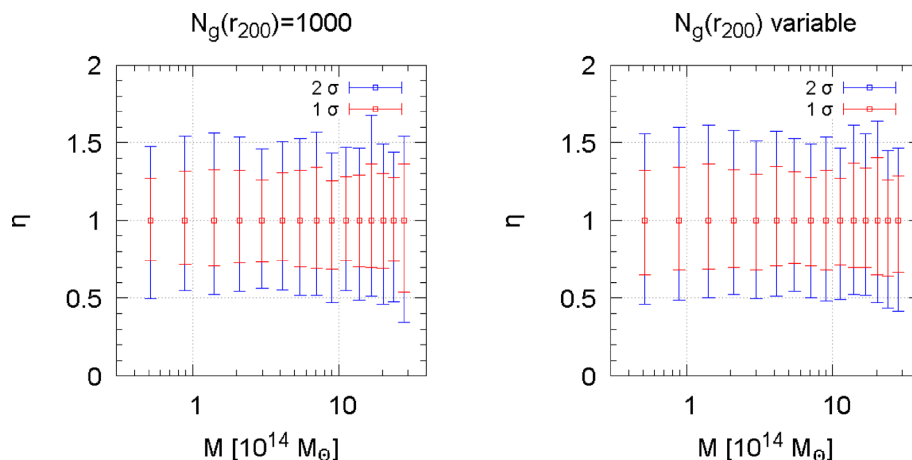


Figure 3. 1σ and 2σ constraints on the gravitational slip parameter η from the combination of galaxy cluster dynamics (simulated) and weak lensing data (MACS1206) for the sample of 15 clusters considered in this work. The *left* plot corresponds to cluster simulations with a fixed number of tracers ($N_g = 1000$), while the *right* plot to the case where the number of tracers is proportional to the cluster mass (see Section 4.1). It can be seen that, the constraints are practically insensitive to the values of mass and concentration of the cluster, as well as to the variation of the number of tracers according to cluster total mass.

reference case for a number tracers larger than 500. We will further discuss this point in the next section.

In order to investigate the effect on our predictions induced by variations in the lensing information [still relying on the simplified assumption of Gaussianity for the lensing distribution $P_L(r_s, r_{200})$] we repeat the analysis doubling the uncertainties in r_{200} and changing the value of the correlation ρ in the lensing information. We have considered three alternatives for ρ : a totally uncorrelated distribution ($\rho = 0$), a small correlation $\rho = 0.15$, and $\rho = 0.99$. The results are shown in the lower panels of Fig. 4 (for the correlation we plot only the case $\rho = 0.15$). For completeness, it is worth noticing that we have also considered a change in the relative uncertainties of the scale radius $\sigma(r_s)/r_s$ from $0.3 \rightarrow 0.45$, but the variation has only marginal effects on the final distributions. The marginalized constraints on r_s obtained by the *MAMPOSSt* fit on our synthetic sample are on average of the same order of the one used in the lensing distribution for the reference analysis $\sigma(r_s)/r_s \sim 0.3$. Therefore, modifying the lensing information on r_s in practice does not affect the final predictions for η . We nevertheless emphasize that mild effects can be enhanced or suppressed depending on the adopted value of the correlation ρ . All the constraints obtained in this case are listed in the second column of Table 2.

5.2 The effect of the density of member galaxies

Realistic surveys cover a wide range of clusters, each with a different number of tracers depending on the details of its evolution and environment. In this regard, an important question regards the dependence of constraints on the density of tracers within the cluster. Here, we will provide an answer considering the case of a scale-independent η , and we will show that the effect is practically negligible above a given density of tracers. For this purpose, we generated six realizations of a new cluster, characterized by fixed $r_{200} = 1.96$ Mpc, $r_s = 0.27$ Mpc, and a Tired velocity anisotropy profile with $\beta_\infty = 0.5$, increasing the number of particles enclosed within r_{200} in each realization, considering the cases $N_g(r_{200}) = 50, 100, 200, 500, 1000, 2000$. Notice that we chose to keep the lensing probability distribution identical to the reference case.

In Fig. 6, we show the 1σ and 2σ error bars (left-hand panel), along with the symmetrized 1σ and 2σ errors (right-hand panel) of

a scale-independent η as a function of $N_g(r_{200})$. It appears that above ~ 100 tracers within r_{200} , the goodness of the resulting bounds on η increases only marginally, passing from 27 per cent uncertainties at 1σ for $N(r_{200}) = 100$, to 21 per cent for $N(r_{200}) = 2000$; the effect is almost irrelevant for $N(r_{200}) > 500$.

Extrapolating the result of $\eta(N_g = 100)$ with the naive theoretically expected scaling $\sim N_{\text{cluster}}^{-1/2}$, we obtain

$$\eta(N_g = 100) = 1.00 \pm 0.07(1\sigma)_{-0.11}^{+0.13}(2\sigma) \quad 15 \text{ clusters}, \quad (28)$$

$$\eta(N_g = 100) = 1.00 \pm 0.03(1\sigma)_{-0.04}^{+0.05}(2\sigma) \quad 100 \text{ clusters}. \quad (29)$$

This shows that very stringent constraints on the gravitational slip can be obtained with a reasonably small number of clusters for which a hundred measured spectroscopic redshifts are available, together with reliable lensing mass profile reconstructions. Nevertheless, it is important to notice that the goodness of the expected constraints is strongly dependent on the assumed CLASH-like lensing probability distribution, which is based on high-quality data of a restricted sample of selected objects. Thus, realistic constraints obtainable with upcoming data sets will rely on the analysis of only few dozen clusters for which such accurate measurements will be available; however, our results show that this will be sufficient to bring the uncertainties on η down to less than 10 per cent.

Moreover, it is worth pointing out that the analysis carried out in this paper relies on the assumption that the best-fitting values of r_s, r_{200} from the simulated lensing distribution always coincides with the *MAMPOSSt* best fit in order to guarantee that the resulting distribution of η is centred over the GR expectation $\eta = 1$. However, one should take into account that, even in this ideal situation where all the systematics are neglected, the smaller is the number of tracers in the *MAMPOSSt* analysis, the larger is the difference between the true value of the mass profile parameters and the one fitted by *MAMPOSSt* (the so called bias parameter, see e.g. Mamon et al. 2013). In Fig. 7 we plot the error bars on η obtained by centering the lensing distribution on the true values $r_{200} = 1.96$ Mpc, $r_s = 0.27$ Mpc instead of on the *MAMPOSSt* best-fitting values. As expected, the bias in η tends to be larger for smaller number of tracers, even if GR is always included within 1σ .

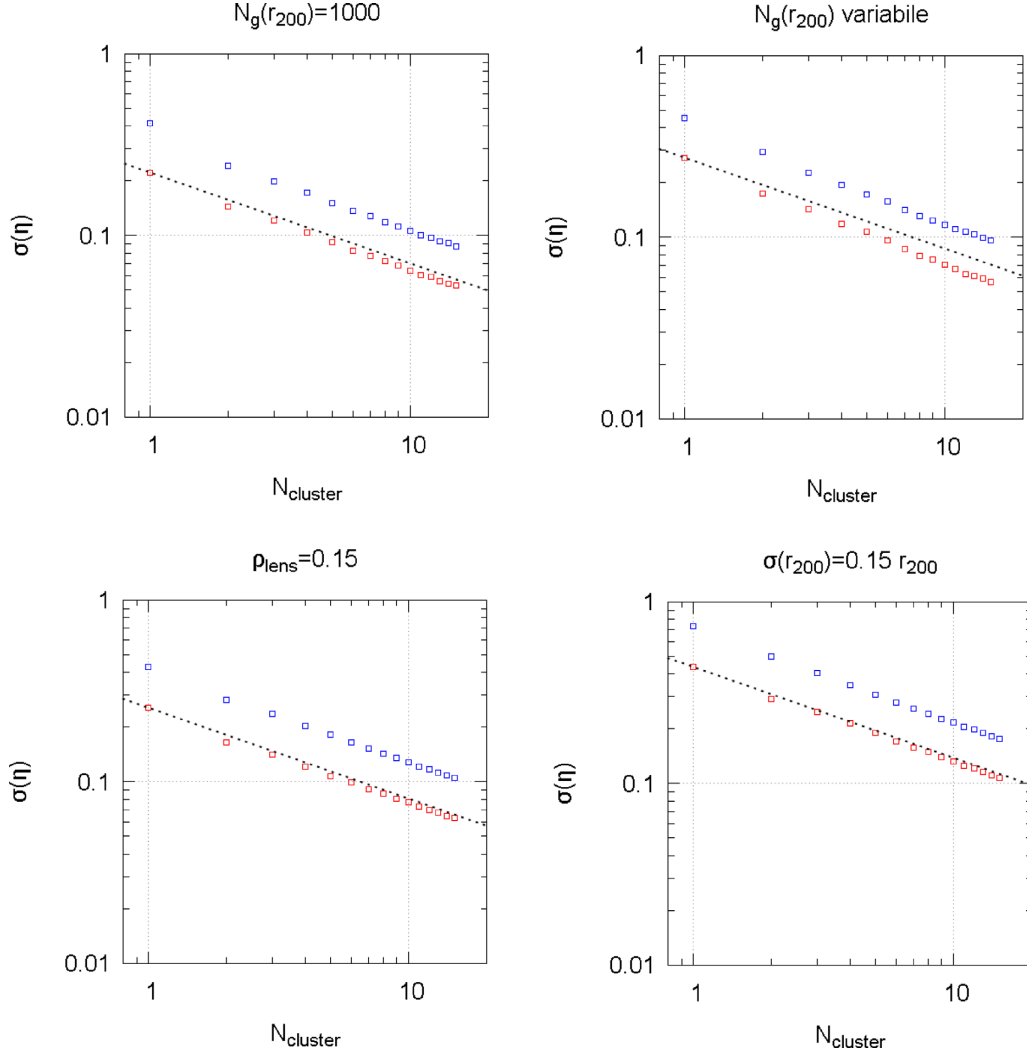


Figure 4. 1σ (red, lower squares) and 2σ (blue, upper squares) constraints on the gravitational slip parameter η from the combination of galaxy cluster dynamics (simulated) and weak lensing data (MACS1206) as described in the text. The *upper left* plot corresponds to cluster simulations with a fixed number of tracers ($N_g = 1000$), while the *upper right* one to the case where the tracers number is proportional to the cluster mass (see Section 4.1). The *bottom left* plot shows the constraints obtained when changing the correlation of the lensing distribution from $\rho = 0.67$ to $\rho = 0.15$; finally, the *bottom right* plot is for the case when the standard deviation of r_{200} is doubled. In all cases the fiducial model is GR, $\eta = 1$. The black-dotted curve denotes the theoretically expected scaling of the 1σ errors, $\sim(N_{\text{cluster}})^{-1/2}$. Comparison of the two upper plots shows that the predicted errors appear to be marginally sensitive to the choice of the number of tracers in the cluster; in particular, in the case of variable number of tracers in the dynamical fit, the theoretical relation $\sim(N_{\text{cluster}})^{-1/2}$ is recovered for $N_{\text{gal}} \gtrsim 500$ (see Section 5.2).

Table 2. Constraints from the combined analysis of multiple number of clusters for scale-independent (third column) and scale-dependent (fourth column) gravitational slip η . In the latter case, the error is evaluated at the reference scale $r = 1.5$ Mpc. In each row refers to the constraints from the combined 15 clusters-likelihood. When varying the lensing parameters (correlation ρ and relative errors, respectively), the number of tracers in the dynamics fit is fixed to $N_g = 1000$ for all clusters.

Case	$\eta = \text{constant}$	$\eta = \eta(r = 1.5\text{Mpc})$
Reference	1.00 ± 0.05 (1σ) ± 0.09 (2σ)	$1.00^{+0.10}_{-0.09}$ (1σ) ± 0.21 (2σ)
Variable N_g	1.00 ± 0.05 (1σ) ± 0.09 (2σ)	1.00 ± 0.12 (1σ) $^{+0.25}_{-0.24}$ (2σ)
$\rho = 0.15$	1.00 ± 0.06 (1σ), ± 0.11 (2σ)	1.00 ± 0.13 (1σ), ± 0.27 (2σ)
$\rho = 0$	$1.00^{+0.06}_{-0.05}$ (1σ), ± 0.10 (2σ)	1.00 ± 0.12 (1σ), ± 0.26 (2σ)
$\rho = 0.99$	1.00 ± 0.03 (1σ), ± 0.11 (2σ)	1.00 ± 0.08 (1σ), ± 0.19 (2σ)
$\sigma(r_{200}) = 0.15 \times r_{200}$	$1.00^{+0.11}_{-0.10}$ (1σ) $^{+0.18}_{-0.17}$ (2σ)	1.00 ± 0.15 (1σ) ± 0.30 (2σ)
$\sigma(r_s) = 0.45 \times r_s$	1.00 ± 0.05 (1σ) ± 0.09 (2σ)	1.00 ± 0.15 (1σ) ± 0.31 (2σ)

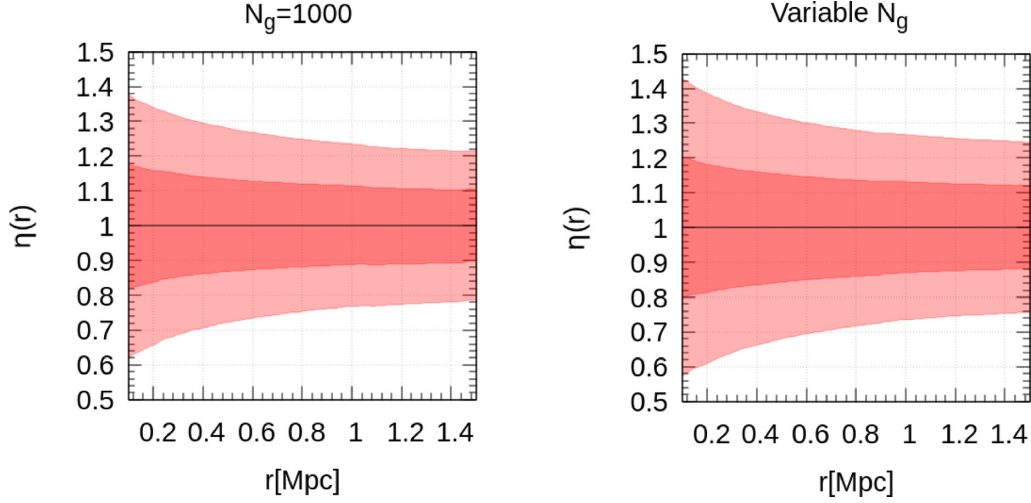


Figure 5. Constraints on a scale-dependent η as a function of the radius from the centre of the cluster, accounting for the sample of 15 synthetic clusters, as explained in the text. *Left:* constant number of member galaxies ($N_g = 1000$) throughout our synthetic cluster sample. *Right:* variable number of member galaxies proportional to the cluster mass (see Section 4.1). Darker and lighter shaded regions correspond to 1σ and 2σ bands, respectively.

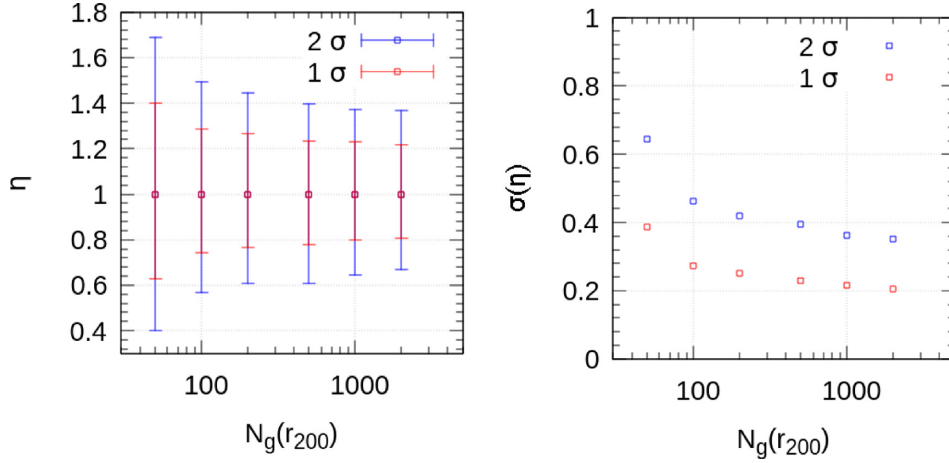


Figure 6. Constraints on a scale-independent η as a function of the number of tracers considered within r_{200} of a single cluster, as explained in Section 5.2. *Left-hand panel:* 1σ and 2σ errorbars. *Right-hand panel:* Scaling of the 1σ and 2σ symmetrized errors with the tracers number.

5.3 Case B: scale-dependent η

We now consider $\eta \equiv \eta(r)$ to be scale dependent, i.e. the scale radius of the lensing mass profile (r_s^{lens}), and that of the dynamics mass profile (r_s^{dyn}) are independent parameters for each cluster. As shown in Pizzuti et al. (2016), the constraints on η depend on the distance from the cluster centre, with the scaling of the uncertainties with r determined by the actual values of r_s^{lens} , r_s^{dyn} . In order to combine the distributions of η obtained for each single cluster we considered 100 equally spaced radial bins in the range $[0.1, 1.5]$ Mpc. The choice of the lower bound is dictated by consistency with real observations – indeed, as already explained earlier, the assumptions of the Jeans analysis cannot be met at very small radii where the dynamics is expected to be dominated by dissipative effects. As for the upper bounds, for an observed cluster the analysis cannot generally be performed beyond the virial radius $r > r_{200}$, where lack of dynamical equilibrium as well as contamination from the large-scale structure introduces systematic effects in both lensing and dynamics mass profile reconstructions. We considered the value $r_{\text{max}} = 1.5$ Mpc as a fair compromise that accounts for the different

scales of the mass profiles. In each radial bin we Monte Carlo sampled the joint probability distribution

$$\log \mathcal{L}(r_s^{\text{lens}}, r_{200}^{\text{lens}}, r_s^{\text{dyn}}, r_{200}^{\text{dyn}}) =$$

$$\log \mathcal{L}_{\text{dyn}}(r_s^{\text{dyn}}, r_{200}^{\text{dyn}}) + \log P_L(r_s^{\text{lens}}, r_{200}^{\text{lens}}), \quad (30)$$

for 10 000 trials, and combined, through multiplication, the resulting likelihoods for $\eta(r)$ to get the final probability distribution $P(\eta, r)$. In the left and right plot of Fig. 5 we show the results for the reference analysis of $N_g = 1000$, and for the case of variable number of tracers in the *MAMPOSSt* fit, respectively. In each panel the coloured bands refers to one (inner red shaded area) and two (outer red shaded region) σ bands for the 15 cluster-combined distributions. The constraints for all the sub-cases analysed in the scale-dependent configuration, evaluated at the radius $r = 1.5$ Mpc, are summarized in the third column of Table 2. As now $r_s^{\text{dyn}} \neq r_s^{\text{lens}}$, we consider also the variation of the uncertainties in the lensing scale radius r_s^{lens} from 30 per cent to 45 per cent which produces an observable effect on the bounds on η .

6 SUMMARY AND DISCUSSION

The gravitational slip parameter $\eta \equiv \Phi/\Psi$ is of paramount importance for tests of gravity at cosmological and astrophysical scales. While a possible future detection of $\eta \neq 1$ would be the smoking gun for a departure from the standard gravity paradigm of GR at late-times in the cosmological evolution, a non-detection would disfavour large families of modified gravity scenarios such as the popular $f(R)$ scalar–tensor theories. Here, we elaborated on the potential of galaxy cluster observations to constrain η , through a joint analysis of dynamics and lensing information. As we discussed in detail in Section 4.1, our ground for the analysis of dynamics of cluster galaxies was based on a sample of synthetic (simulated) clusters, whose dynamical mass profile were subsequently derived by means of the *MAMPOSSt* numerical framework, in turn to be combined with a lensing probability distribution simulated on the basis of the available information from the analysis of the cluster MACS 1206.

It is important to remind that our reconstruction of η did not invoke the assumption of a particular gravity models or parametrization. In our view, $\eta \neq 1$ is sourced through a difference in the respective values of the NFW parameters (r_{200} , r_s) when inferred from dynamics and lensing, respectively. This phenomenology inclined view is general and allows to translate the derived constraints on η to constraints on the theory space of specific models. Nevertheless, we inevitably had to rely on assumptions such as virialization, spherical symmetry (see below) and an NFW parametrization for the matter density. In addition, the analysis was performed at a fixed redshift $z = 0$, which prevented us to capture any possible time-dependence of η . As concerns the lensing distribution, it is worth to notice that we have performed our analysis under the simplified assumption of Gaussianity for the lensing probability distribution $P_L(r_s, r_{200})$ (even if well-motivated by the reference observational data of MACS 1206). The effect of a variation of the lensing information such as the relative errors and correlation was investigated to find that it has a sizable effect on the derived constraints as can be seen from Table 2.

We can summarize our *main results* as follows:

(1) For a scale-independent η , and assuming a fiducial with $\eta = 1$, we can reach an accuracy of 5 per cent (9 per cent) at 1σ (2σ) with 15 clusters. The explicit results are also summarized in Table 2, and Fig. 5.

(2) In a similar manner, for the more realistic case of a scale-dependent $\eta = \eta(r)$, when evaluated at the reference scale $r = 1.5$ Mpc and accounting for 15 clusters we find an accuracy of 10 per cent at 1σ and 21 per cent at 2σ . Of course, it should come as no surprise that the constraints are less optimistic compared to the simpler scale-independent case.

(3) Optimization of future data analyses in this context requires an understanding of the dependence of the constraints on factors such as the clusters masses and the number of galaxy members in the clusters. Our investigation revealed that the cluster mass (parametrized through r_{200} in the NFW profile) has a rather mild effect on η . This is best seen from the constraints based on the individual clusters as a function of the cluster mass, as it is shown in Fig. 3. According to the procedure explained in Section 4.1 and shown in Table 2, varying the number of tracers in the clusters according to the halo total mass has either none or a very mild quantitative effect on the constraints. To better understand this behaviour, we have further investigated the effect of increasing the density of tracers used in the dynamics fit for one single cluster. The results of this exercise show that after $N_g(r_{200}) \sim 10^2$ the effect

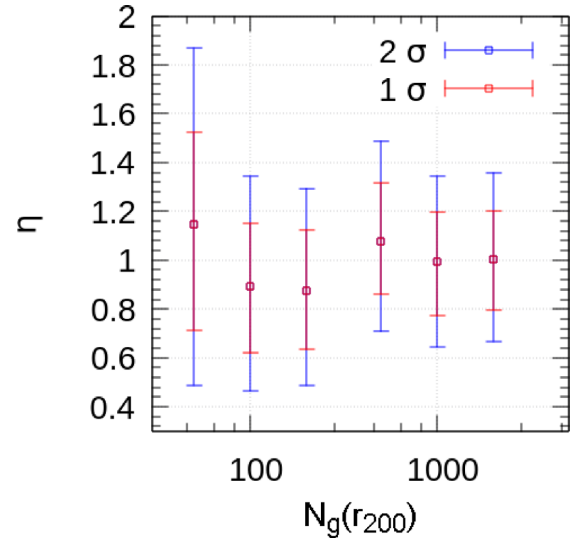


Figure 7. Similar to Fig. 6, but with the simulated lensing distribution now centred on the true values of the cluster mass profile parameters. Notice that, the fiducial is not centred around $\eta = 1$ anymore.

on the constraints for a scale-independent η is moderate, becoming irrelevant for $N_g(r_{200}) \gtrsim 500$. (See Figs 6 and 7.)

This paper has been devoted to the study of the constraining power of galaxy cluster mass profiles as a probe for modified gravity. It is worth stressing out again that the analysis carried out addressed the problem only from a statistical point of view, neglecting the effect of possible systematics. In principle, several issues can introduce biases in the determination of the gravitational slip – for example, deviations from spherical symmetry and dynamical relaxation, which are the main assumptions of our method, can be sources of differences between the mass profile inferred from lensing and dynamics of the member galaxies, respectively, giving rise to spurious detections of gravitational slip, $\eta \neq 1$. Moreover, as already mentioned in Section 4.1, the velocity anisotropy profile $\beta(r)$ is generally an unknown in analyses of cluster dynamics. In the parametric approach implemented in our version of the *MAMPOSSt* code, the choice of the anisotropy model introduces additional systematics that should be taken into account. A similar statement can be made concerning the parametrization of the mass profile, as discussed in Pizzuti et al. (2016). An additional source of systematics not considered here is the presence of interlopers in the dynamical analysis, namely galaxies that are projected in the cluster area but are not dynamically bound to the cluster.

The large amount of data for several hundred clusters that will be provided by current and future imaging and spectroscopic surveys, can allow for determinations of the gravitational slip parameter with uncertainties of few percents. Even if the fraction of clusters for which CLASH-like data will be available is expected to be relatively small, the analysis of ~ 20 objects can bring the errors on η down to less than 10 per cent. This naturally calls for an accurate calibration of the impact of systematics and understanding of the relevant statistics. Such constraints are complementary to other cosmological probes such as the cosmic microwave background, gravitational waves and electromagnetic counterpart detections, cluster number counts and redshift-space distortions, allowing to test gravity at cosmological scales down to an unprecedented level of accuracy.

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