



Publication Year	2018
Acceptance in OA	2020-10-19T08:06:23Z
Title	The WISSH quasars project. V. ALMA reveals the assembly of a giant galaxy around a $z = 4.4$ hyper-luminous QSO
Authors	Bischetti, M., PICONCELLI, Enrico, Feruglio, Chiara, Duras, F., BONGIORNO, ANGELA, Carniani, S., Marconi, Alessandro, Pappalardo, C., Schneider, R., Travascio, A., VALIANTE, ROSA, VIETRI, GIUSTINA, ZAPPACOSTA, Luca, Fiore, F.
Publisher's version (DOI)	10.1051/0004-6361/201833249
Handle	http://hdl.handle.net/20.500.12386/27872
Journal	ASTRONOMY & ASTROPHYSICS
Volume	617

The WISSH quasars project

V. ALMA reveals the assembly of a giant galaxy around a $z = 4.4$ hyper-luminous QSO[★]

M. Bischetti^{1,2}, E. Piconcelli¹, C. Feruglio³, F. Duras⁴, A. Bongiorno¹, S. Carniani^{5,6}, A. Marconi⁷, C. Pappalardo^{8,9}, R. Schneider¹⁰, A. Travascio^{1,10}, R. Valiante¹, G. Vietri^{11,12}, L. Zappacosta¹, and F. Fiore^{3,1}

¹ INAF - Osservatorio Astronomico di Roma, Via Frascati 33, 00078 Monte Porzio Catone (Roma), Italy
e-mail: manuela.bischetti@oa-roma.inaf.it

² Università degli Studi di Roma “Tor Vergata”, Via Orazio Raimondo 18, 00173 Roma, Italy

³ INAF - Osservatorio Astronomico di Trieste, via G.B. Tiepolo 11, 34143 Trieste, Italy

⁴ Dipartimento di Matematica e Fisica, Università degli Studi Roma Tre, via della Vasca Navale 84, 00146 Roma, Italy

⁵ Cavendish Laboratory, University of Cambridge, 19 J. J. Thomson Avenue, Cambridge CB3 0HE, UK

⁶ Kavli Institute for Cosmology, University of Cambridge, Madingley Road, Cambridge CB3 0HA, UK

⁷ INAF - Osservatorio Astrofisico di Arcetri, Largo E. Fermi 5, 50125 Firenze, Italy

⁸ Centro de Astronomia e Astrofísica da Universidade de Lisboa, OAL, Tapada da Ajuda, 1349-018 Lisboa, Portugal

⁹ Instituto de Astrofísica e Ciências do Espaço, Universidade de Lisboa, OAL, Tapada da Ajuda, 1349-018 Lisboa, Portugal

¹⁰ Università degli Studi di Roma “La Sapienza”, Piazzale Aldo Moro 5, 00185 Roma, Italy

¹¹ Excellence Cluster *Universe*, Technische Universität München, Boltzmannstr., 85748 Garching, Germany

¹² European Southern Observatory, Karl-Schwarzschild-Str. 2, 85748 Garching b. München, Germany

Received 17 April 2018 / Accepted 19 May 2018

ABSTRACT

We present an ALMA high-resolution ($0.18'' \times 0.21''$) observation of the $840\mu\text{m}$ continuum and [CII] $\lambda 157.74\mu\text{m}$ line emission in the WISE-SDSS selected hyper-luminous (WISSH) quasi-stellar object (QSO) J1015+0020, at $z \sim 4.4$. Our analysis reveals an exceptional overdensity of [CII]-emitting companions with a very small ($<150\text{ km s}^{-1}$) velocity shift with respect to the QSO redshift. We report the discovery of the closest companion observed so far in submillimetre observations of high- z QSOs. It is only 2.2 kpc distant and merging with J1015+0020, while two other [CII] emitters are found at 8 and 17 kpc. Two strong continuum emitters are also detected at <3.5 arcsec from the QSO. They are likely associated with the same overdense structure of J1015+0020, as they exceed by a factor of 100 the number of expected sources, considering the $\log(N) - \log(S)$ at $850\mu\text{m}$. The host galaxy of J1015+0020 shows a star formation rate (SFR) of about $100 M_{\odot} \text{ yr}^{-1}$, while the total SFR of the QSO and its companion galaxies is a factor of ~ 10 higher, indicating that substantial stellar mass assembly at early epochs may have taken place in the QSO satellites. For J1015+0020 we computed a black hole mass $M_{\text{BH}} \sim 6 \times 10^9 M_{\odot}$. As we resolve the [CII] emission of the QSO, we can compute a dynamical mass of $M_{\text{dyn}} \sim 4 \times 10^{10} M_{\odot}$. This translates into an extreme ratio $M_{\text{dyn}}/M_{\text{BH}} \sim 7$, i.e. two orders of magnitude smaller than what is typically observed in local galaxies. The total stellar mass of the QSO host galaxy plus the [CII] emitters in the ALMA field of view already exceeds $10^{11} M_{\odot}$ at $z \sim 4.4$. These sources will likely merge and develop into a giant galaxy of $\sim 1.3 \times 10^{12} M_{\odot}$. Under the assumption of constant $\dot{M}_{\text{acc}}$ or λ_{Edd} equal to the observed values, we find that the growth timescale of the host galaxy of J1015+0020 is comparable or even shorter than that inferred for the SMBH.

Key words. quasars: individual: SDSSJ101549.00+002020.03 – galaxies: high-redshift – galaxies: star formation – quasars: super-massive black holes – quasars: emission lines

1. Introduction

The most popular models of AGN-galaxy co-evolution (Di Matteo et al. 2005; Menci et al. 2008; Hopkins et al. 2008) include galaxy interactions (both major and minor mergers) and AGN-driven feedback (i.e. the injection of energy and entropy in the interstellar medium (ISM) through winds and shocks) among the major processes driving this phenomenon. These two processes are highly correlated, for example galaxy interactions may destabilise the gas and make it available for fuelling both star formation (SF) and nuclear accretion, giving rise to the growth of the central super-massive black hole (SMBH)

through luminous AGN phases. The AGN fraction increases in IR-luminous, star-forming sources (Nardini & Risaliti 2011; Rosario et al. 2013), and a correlation between the AGN luminosity (L_{Bol}) and the SF luminosity is observed for a wide range of L_{Bol} (Netzer 2009; Lutz et al. 2010). In turn, the AGN can power winds (Fiore et al. 2017) hampering further SF and nuclear gas accretion (as well as AGN-driven winds).

A linked growth of the SMBH and its host galaxy is observationally supported by the well-known local scaling relations between the SMBH mass (M_{BH}) and the physical properties of the host galaxy bulge (e.g. see Kormendy & Ho 2013), such as the dynamical mass (M_{dyn}) or the velocity dispersion. Theoretically, such relations can be shaped by merger events (see Alexander & Hickox 2012, and references therein),

[★] Based on data from ALMA cycle 4 program 2016.1.00718.S.

triggering at the same time bursts of nuclear and SF activity (Volonteri et al. 2015; Gabor et al. 2016; Anglés-Alcázar et al. 2017b). The $M_{\text{BH}}-M_{\text{dyn}}$ relation also indicates that the assembly of the giant galaxies can be probed by observing the QSOs with the most massive SMBHs shining at $z > 2$. This provides insights into key evolutionary phases not observable in the local universe and allows the investigation of the hotly debated role of mergers and AGN feedback. This field of research has been revolutionised by ALMA, thanks to its unprecedented sensitivity and broadband coverage because the QSO emission typically outshines the host galaxy at all wavelengths below few tens of μm . The most powerful observational tool to study the high- z QSOs host galaxies is the [CII] fine structure line at $157.75 \mu\text{m}$. [CII] is in fact the strongest line from the cool gas ($T < 10^4 \text{ K}$) and, given its low ionisation potential of 11.3 eV, traces both the neutral and ionised medium. It is also a tracer of SF activity (Maiolino et al. 2005; Carniani et al. 2013). Trakhtenbrot et al. (2017) found a wide variety of host galaxy properties of hyper-luminous QSOs at $z \sim 5$ in terms of possible SMBH fuelling mechanisms and SF activity, suggesting that galaxy–galaxy interactions may not be a necessary condition for either of the two processes.

[CII] has been investigated even in the most distant $z > 6$ QSOs. These objects typically reside in compact hosts where rotating disks are already in place and intense SF activity of tens to thousands of solar masses is ongoing (e.g. Wang et al. 2013, 2016; Ciccone et al. 2015; Díaz-Santos et al. 2016; Venemans et al. 2016, 2017a; Willott et al. 2017; Decarli et al. 2018). Recent studies of $z \sim 6$ QSOs reveal that they are powered by SMBH at the massive end of the black hole mass function (Jiang et al. 2007; De Rosa et al. 2011, 2014; Venemans et al. 2015; Bañados et al. 2016) and that their hosts are among the brightest and most massive galaxies at these redshifts. According to local relations (e.g. Jiang et al. 2011), these sources are therefore expected to assemble stellar masses typical of giant galaxies at $z = 0$. In the $M_{\text{BH}}-M_{\text{dyn}}$ plane, most of high- z QSOs lie above the local relation, as they are characterised by very low stellar-to-black hole mass ratios as low as ~ 10 . However, the number of $z > 4$ sources with available M_{BH} estimates from single epoch relations is very limited to date, while the bulk of the M_{BH} is still derived from the QSO L_{Bol} under the assumption of accretion at the Eddington limit. Furthermore, the majority of the high- z QSOs are still unresolved or only marginally resolved, thus affecting size and dynamical mass estimates with large uncertainty. Accordingly, high-resolution studies are of primary importance.

It is therefore crucial to study the SMBH and host galaxy growth at early epochs, i.e. $z \sim 2-4$ when both processes are maximised. Accordingly, we have undertaken the WISE-SDSS selected hyper-luminous (WISSH) QSOs project to study the most powerful AGN in the Universe, which happen to shine at these cosmic epochs (Bischetti et al. 2017). Similar to the $z \sim 6$ QSOs studied so far, WISSH have $L_{\text{Bol}} > 10^{47} \text{ erg s}^{-1}$ and are powered by accretion onto SMBH of $M_{\text{BH}} \sim 10^9-10^{10} M_{\odot}$ at rates close (or even higher than) the Eddington limit. Such huge luminosities at Eddington regimes are likely triggered by galaxy interactions (Menci et al. 2014; Valiante et al. 2014, 2016) and drive powerful winds that may affect the whole host galaxy. Therefore, these QSOs are ideal targets to shed light on the AGN-galaxy feeding and feedback cycle. To this purpose, we collected information about the AGN power and the multi-transition presence of nuclear and galaxy-scale winds from multiwavelength spectroscopy and photometry (Bischetti et al. 2017; Vietri et al. 2018). We also built up the far-infrared (FIR-) to-UV spectral energy distributions (SED) of 14 WISSH QSOs

with *Herschel* photometry (Duras et al. 2017) to derive the star formation rate (SFR) in their host galaxy.

From this WISSH-*Herschel* subsample, we selected the QSO J1015+0020 at $z = 4.4$ for a pilot ALMA observing programme aimed to characterise the host galaxy and environment properties of hyper-luminous QSOs. The high redshift and low declination guaranteed that this target would fit the ALMA band 7 well and be observed with good sensitivity. Specifically, we present here the results from a high-resolution ($0.18'' \times 0.21''$) ALMA observation of the $840 \mu\text{m}$ and [CII] $\lambda 157.74 \mu\text{m}$ line emission in J1015+0020. Throughout this paper, we assumed a ΛCDM cosmology with $H_0 = 67.3 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_{\Lambda} = 0.69$, and $\Omega_{\text{M}} = 0.31$ (Planck Collaboration XIII 2016).

2. ALMA observation and data analysis

In this work we present the ALMA Cycle 4 observation (project 2016.1.00718.S) of the WISSH QSO SDSS J1015+0020 (celestial coordinates RA 10:15:49.00, Dec +00:20:20.03). The observation was carried out on 5 March 2016 for a 0.6 hours on-source integration time in C36-5 configuration with a maximum projected baseline of 1396 m. We used the ALMA band 7 receiver and the frequency division mode of the ALMA correlator. This provided us with four spectral windows of 1.875 GHz width, for a total spectral coverage of 7.5 GHz, with a spectral resolution of 31.25 MHz. A first spectral window was centred at $\sim 352 \text{ GHz}$ to cover the expected observed frequency of the [CII] emission line, given the SDSS DR10 redshift $z_{\text{SDSS}} = 4.400$ (Pâris et al. 2014). A second window was put adjacent to it on the lower frequency side, in the case of a blueshift of the rest-frame UV lines with respect to the systemic emission of the source, i.e. the [CII] emission. The remaining two spectral windows were placed at higher frequencies to account for the continuum emission.

Visibilities were calibrated with the CASA 4.7.0 software (McMullin et al. 2007) in pipeline mode by applying the default phase, bandpass, and flux calibrators provided. Images were produced using the CASA task clean with natural weights, a 0.03 arcsec pixel size and a 30 km s^{-1} channel width, and a final beam size of $0.18'' \times 0.22''$. The ALMA field of view (FOV) of our observation, defined as the region in which the relative sensitivity is higher than 0.5, is a circular area with a radius of $\sim 8.5''$. The continuum map was obtained by averaging over all the spectral windows and excluding 1.2 GHz around the [CII] emission. The continuum emission in the spectral region of the [CII] line was first modelled by fitting the UV plane of the first two spectral windows by a first degree polynomial and then subtracted to the visibilities.

For all the sources detected with a signal-to-noise ratio higher than five in the ALMA FOV, continuum and [CII] flux densities were measured by fitting a 2D Gaussian model to the final map in the image plane (see Sect. 3.1). Furthermore, from spectral fitting of the [CII] emission we derived the parameters of the [CII] line profile (Sect. 3.2). For each source, we applied a spectral model including one or two Gaussian emission line components, based on the profile of the [CII] line.

3. Results from the ALMA observation

3.1. Continuum and [CII] emission maps

The WISSH quasar J1015+0020 is detected both in continuum and [CII] line emission. Furthermore, continuum and [CII] maps reveal multiple sources with an angular separation of less than 4 arcsec from the central QSO.

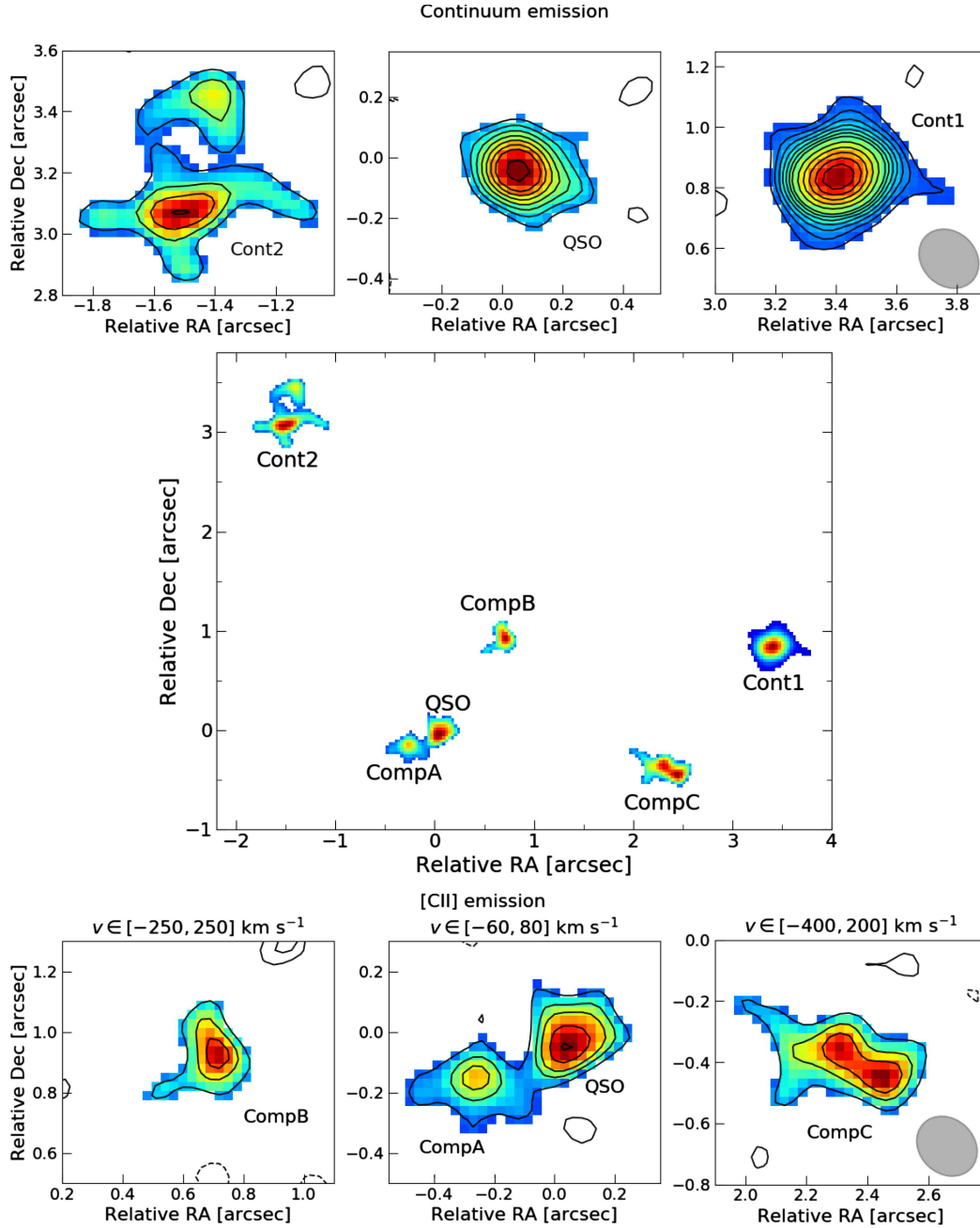


Fig. 1. Continuum and [CII] line emission maps of J1015+0020 and the other sources detected in the ALMA FOV. *Top panel:* continuum emission maps of the QSO, Cont1, and Cont2 sources. Contours range from 2σ to 8σ in steps of 1σ ($0.04 \text{ mJy beam}^{-1}$), while above 8σ are in steps of $0.16 \text{ mJy beam}^{-1}$. *Middle panel:* spatial distribution of all the detected sources, where coordinates are relative to the QSO location. *Bottom panel:* [CII] emission maps of the QSO and companion sources CompA, CompB, and CompC. Each map is integrated over the velocity range indicated above the panel. Contours are shown as in top panel with $\sigma = 0.060, 0.027, 0.070 \text{ Jy beam}^{-1} \text{ km s}^{-1}$ in the left, middle, and right panel, respectively. Dashed contours refer to -2σ . The ALMA beam is shown as a grey ellipse.

Figure 1 gives an insight into these crowded surroundings. Specifically, the top panels show the ALMA continuum map at $\sim 840 \mu\text{m}$, with a rms sensitivity of $0.04 \text{ mJy beam}^{-1}$, revealing cold dust emission from the host galaxy of J1015+0020 and two additional sources both at a distance of $\sim 3.5 \text{ arcsec}$ (Cont1 and Cont2 hereafter). The continuum emission from the QSO is detected at $\sim 9\sigma$ significance with a flux density of $f_{\text{cont}}^{\text{QSO}} \sim 600 \mu\text{Jy}$. The QSO continuum is resolved by the ALMA beam and has a deconvolved size of $0.16 \times 0.94 \text{ arcsec}$ (see Table 1). Cont1 and Cont2 are detected at $\sim 16\sigma$ and $\sim 6\sigma$ significance with a flux density ~ 1250 and $\sim 390 \mu\text{Jy}$, respectively. These sources are also

resolved, see Table 1. Cont1 and Cont2 are not detected in [CII] line emission and, therefore, we are not able to derive a spectral redshift for these sources.

The ALMA [CII] map also reveals multiple sources around J1015+0020. Figure 1 (bottom middle panel) shows the presence at a $\sim 6\sigma$ significance of a clearly distinct companion galaxy with an angular separation of only $0.33''$ (CompA hereafter), corresponding to a proper distance of $\sim 2.2 \text{ kpc}$ at the QSO redshift.

Moving to larger distance from the QSO, two additional [CII] emitters are detected at $1.2''$ (8.2 kpc) and $2.3''$ (16 kpc) with a significance of 6σ and 7σ , respectively. The angular

Table 1. Summary of continuum and [CII] line emission properties of J1015+0020 and the additional sources detected in the ALMA map.

(1) Source	J1015+0020	Cont1	Cont2	CompA	CompB	CompC
Continuum emission						
(2) f [μ Jy]	595 ± 64	1252 ± 76	388 ± 61	<120	<120	<125
(3) $a_{\max}$ ["]	0.16 ± 0.03	0.19 ± 0.02	0.28 ± 0.06	–	–	–
(4) $a_{\min}$ ["]	0.94 ± 0.03	0.10 ± 0.03	0.08 ± 0.06	–	–	–
[CII] emission						
(5) f [Jy km s ⁻¹]	0.47 ± 0.05	<0.15	<0.15	0.19 ± 0.03	0.35 ± 0.06	1.13 ± 0.16
(6) a_{maj} ["]	0.17 ± 0.03	–	–	0.28 ± 0.05	–	0.59 ± 0.19
(7) a_{min} ["]	0.09 ± 0.03	–	–	0.16 ± 0.04	–	0.19 ± 0.04
(8) FWHM [km s ⁻¹]	339 ± 38	–	–	60 ± 13	480 ± 45	409 ± 42
(9) $L_{\text{[CII]}}$ [$10^8 L_{\odot}$]	2.9 ± 0.3	<0.92	<0.92	1.2 ± 0.1	2.2 ± 0.4	7.0 ± 1.0
(10) Angular sep ["]	–	3.5	3.5	0.33	1.2	2.4

Notes. Rows give the following information: (1) source ID, (2) continuum flux, (3)–(4) major and minor deconvolved axes of the continuum emission, (5) integrated [CII] flux density, (6)–(7) major and minor deconvolved axes of the [CII] emission, (8) FWHM of the [CII] line profile, (9) [CII] luminosity, and (10) angular separation from J1015+0020. For [CII]-only detected sources, i.e. CompA, CompB, and CompC, 3σ upper limits on the continuum flux are given. For continuum-only detected sources, i.e. Cont1 and Cont2, we report 3σ upper limits on the integrated [CII] flux density, computed assuming a line width of 300 km s⁻¹ and taking into account the differential sensitivity within the FOV.

separations from J1015+0020 are 1.2" and 2.3", corresponding to ~ 8.2 kpc and 16 kpc, respectively. Hereafter, these sources are referred as CompB and CompC. All three [CII] emitters are not detected in continuum emission. The [CII] line emission of the QSO is detected at 8.7σ significance and has an integrated flux density of $f_{\text{[CII]}}^{\text{QSO}} = 0.47 \pm 0.05$ Jy km s⁻¹, derived from 2D-Gaussian fitting. CompA, CompB, and CompC have a [CII] flux density of ~ 0.19 , 0.35 , and 1.13 Jy km s⁻¹, respectively. The [CII] emission associated with J1015+0020 and CompA is marginally resolved by the ALMA beam (see Table 1). CompB is unresolved, while CompC has an elongated morphology which clearly rules out a point source nature.

3.2. Spectra and velocity maps

We extracted the [CII] spectrum of J1015+0020, CompA, CompB, and CompC from the ALMA cube in regions of 1.5–3 beam areas, according to the size of the source. The resulting continuum-subtracted spectra are shown in Fig. 2. The [CII] line profile of J1015+0020 is well reproduced by a single Gaussian component centred at 351.5 GHz, corresponding to a [CII]-based redshift $z_{\text{[CII]}} = 4.407$. This translates into a small velocity shift of about 400 km s⁻¹ between the [CII] emission and z_{SDSS} .

All the [CII] line-detected sources in the ALMA map have a very small velocity shift of ($\lesssim 150$ km s⁻¹) with respect to $z_{\text{[CII]}}$ and, therefore, we consider these as companions of J1015+0020. For each source, we calculate the velocity shift as the difference between $z_{\text{[CII]}}$ and the velocity that bisects the cumulative [CII] line flux. CompA, which is the closest companion, has an extremely narrow [CII] line profile with a $\text{FWHM}_{\text{[CII]}}^{\text{A}} = 60 \pm 13$ km s⁻¹ centred at the same redshift of the QSO. CompB and CompC are instead characterised by a boxy-double peaked [CII] line, with $\text{FWHM}_{\text{[CII]}}^{\text{B}} = 480 \pm 45$ km s⁻¹ and $\text{FWHM}_{\text{[CII]}}^{\text{C}} = 409 \pm 42$ km s⁻¹, respectively. These profiles may indicate multiple kinetic components or rotation. The velocity shift of CompB is also consistent with $z_{\text{[CII]}}$, while CompC shows a [CII] line characterised by a blueshift of 145 ± 30 km s⁻¹.

Figure 3 shows the velocity and velocity dispersion maps of J1015+0020 (top panel) and CompC (bottom panel). As for

the QSO, a small velocity gradient seems to be present in the north-south direction, from about -60 to $+40$ km s⁻¹, while the [CII] line of CompA on the QSO left is too narrow to show a gradient at our spectral resolution and appears as a monochromatic spot. The velocity dispersion map of J1015+0020 clearly shows an increase in the south direction and a peak dispersion of ~ 180 km s⁻¹ that is offset with respect to the peak of the [CII] emission, indicated by the black cross in Fig. 3. An increased central velocity dispersion is usually found at the AGN location (e.g. Trakhtenbrot et al. 2017), while we likely observe a perturbation of the gas. This can be interpreted in terms of (i) the very close presence of CompA merging with the QSO, (ii) an additional ongoing, disk scale merger, and (iii) a possible [CII] outflow component. Deeper observations with higher spatial resolution are needed to draw firm conclusions on the nature of this feature.

Two blobs with positive and negative velocities with respect to the QSO redshift are present in the velocity map of CompC. The velocity gradient is much larger, ranging from about -300 km s⁻¹ to $+150$ km s⁻¹. We may interpret this as a rotating disk seen at high inclination (see also Sect. 3.1) with a diameter of ~ 4 kpc. Another possibility is that CompC consists of two interacting sources, which are characterised by slightly blueshifted and redshifted [CII] emission with respect to the $z_{\text{[CII]}}$ of the QSO. This is in agreement with the velocity dispersion map, in which the red blob appears more perturbed than the blue blob. However, the quality of the data does not allow us to clearly discriminate between these hypotheses.

4. Overdensity around J1015+0020

4.1. Multiple companions

The ALMA observation has shed light on the crowded surroundings of J1015+0020. We detected three sources, in addition to the QSO, in [CII] line emission at an angular separation of $\lesssim 2.5''$ (see Sect. 3.1), corresponding to a proper distance of ~ 17 kpc at the observed redshift. This high number of sources is very surprising within the context of submillimetre observations of high- z QSOs, where a single companion at larger separation is usually observed (Trakhtenbrot et al. 2017). Among the [CII]

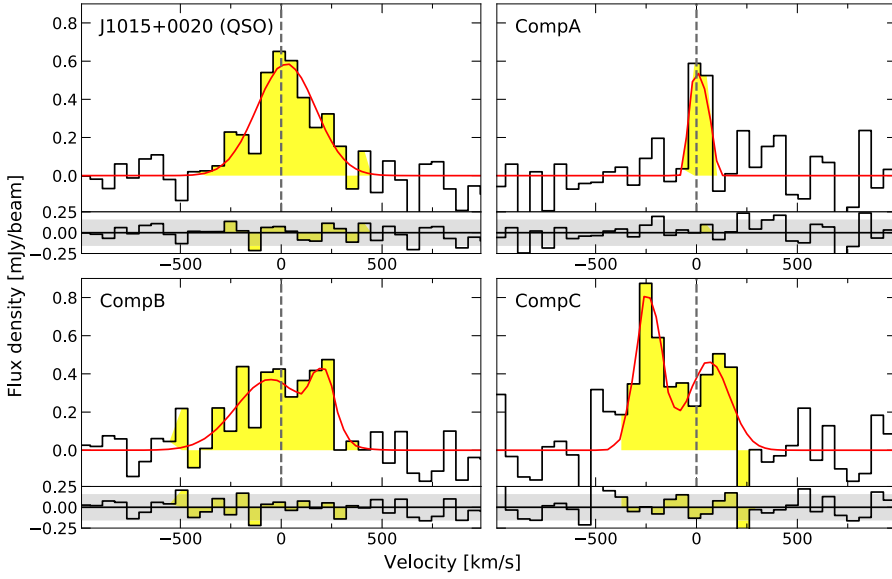


Fig. 2. ALMA continuum-subtracted spectra of the [CII] emission line of J1015+0020 and its companion [CII] emitters, extracted from an area of 1.5 to 2 beam areas, according to the source. The continuum-subtracted [CII] flux density is shown as a function of the relative velocity with respect to the QSO redshift $z_{\text{[CII]}}$, indicated by the dashed vertical line. The plotted channel width corresponds to 60 km s^{-1} . The red curve represents the best fit to the data of one or two Gaussians model.

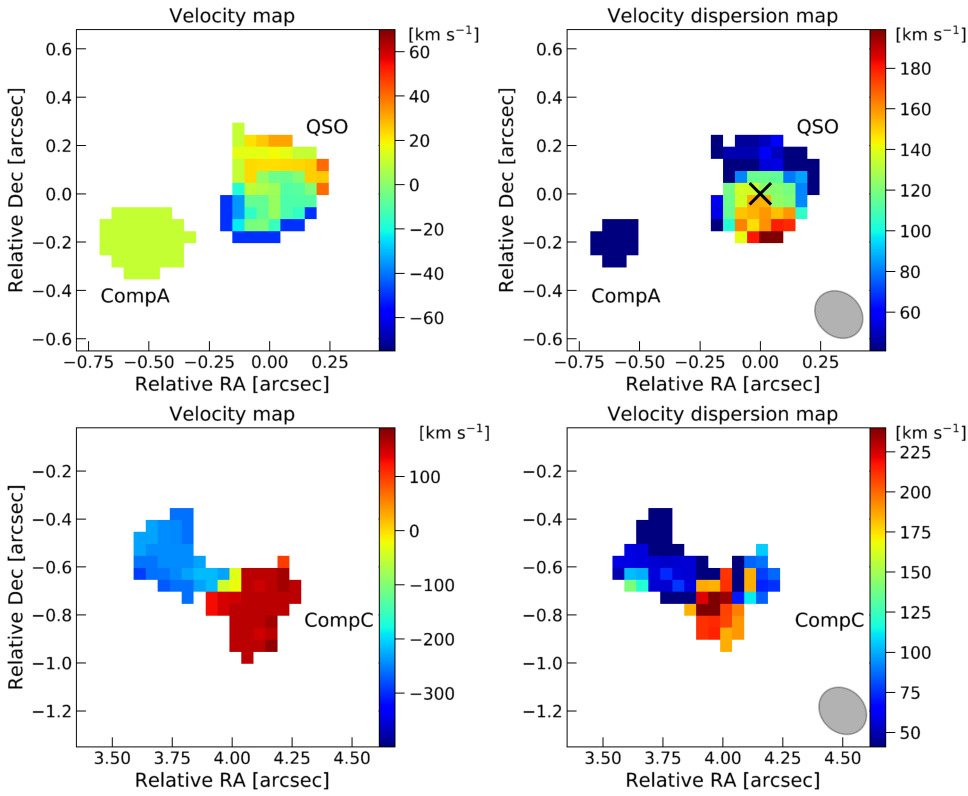


Fig. 3. Velocity and velocity dispersion maps of J1015+0020 (top panel) and CompC (bottom panel), corresponding to the emitting regions with a signal-to-noise ratio higher than 3 for J1015+0020 and higher than 2.5 for CompC. Colour bars indicate the velocity and velocity dispersion range of the maps. The ALMA beam is also shown in the right panels as a grey ellipse.

emitters in our high-resolution observation, we report the discovery of the closest companion observed so far around a high- z QSO. CompA is indeed located at $\sim 0.3''$, corresponding to only $\sim 2 \text{ kpc}$ from J1015+0020.

Furthermore, we detected two additional continuum emitting sources separated $\sim 3.5''$ from J1015+0020, which have comparable or even higher continuum flux density with respect to the hyper-luminous QSO. Cont1 and Cont2 lack [CII] line emission in the spectral band covered by our ALMA observation. The non-detection of [CII] emission from these two sources implies a comoving distance along the line of sight larger than 18.5 Mpc foreground and 25 Mpc background the QSO. In order to understand whether they are related to the QSO, we computed the expected number of field sources of any redshift

within a region of $4'' \times 4''$ around the central QSO. By using the $\log(N) - \log(S)$ at $850 \mu\text{m}$ derived by Simpson et al. (2015), we should expect ~ 0.02 field sources with a flux density $\geq 0.4 \text{ mJy}$, i.e. the value observed for Cont2, the faintest continuum source in the ALMA map. The expected number of sources in the same region, with a flux comparable or larger than that of Cont1 is even smaller, i.e. $\sim 3 \times 10^{-3}$. Indeed, assuming a Poissonian distribution with average number of successes equal to the expected number counts, the probability of having two detections is $\sim 2 \times 10^{-4}$. Given the limited sky coverage (about one square degree) of the Simpson et al. (2015) survey, the observed counts may not be representative of the whole sky. However, as the observed number of continuum sources around the QSO is a factor of 100 larger than the expected value, we conclude that

Cont1 and Cont2 are likely associated with the same overdensity traced by the [CII] emitters around J1015+0020. Recent deep submillimetre surveys carried out with ALMA (Carniani et al. 2015; Fujimoto et al. 2016; Aravena et al. 2016b) typically observed ~ 0.1 sources per ALMA pointing with similar flux to Cont2.

We did not detect any counterpart of the five companions in the Sloan Digital Sky Survey (SDSS), HST/ACS, UKIRT Infrared Deep Sky Survey (UKIDSS), and Wide-field Infrared Survey Explorer (WISE) images in correspondence of their coordinates. Stark et al. (2009) and Bouwens et al. (2015) measured typical densities of 0.01 galaxies with $\text{SFR} \sim 100 M_{\odot} \text{yr}^{-1}$ per ALMA band-7 pointing. Concerning number counts of [CII]-emitting galaxies at $z \sim 5$, Aravena et al. (2016a) predicted about 0.06 galaxies per pointing. In conclusion, it is very unlikely that the presence of three [CII]- and two continuum-emitting companions in the field surrounding J1015+0020 is due to a chance superposition of unrelated galaxies close to the line of sight. Accordingly, we are likely observing a very significant overdensity of star-forming galaxies around a powerful and massive QSO.

4.2. Star formation in and around the QSO host galaxy

The ALMA observation allows us to derive the $\sim 840 \mu\text{m}$ continuum flux of J1015+0020 and, therefore, extend the coverage of the SED presented in Duras et al. (2017) to larger wavelengths. Furthermore, ALMA has highlighted the presence of two continuum emitters (Cont1 and Cont2) at an angular separation $\sim 3.5''$ from J1015+0020 (see Sect. 3.1). Both these continuum emitters have a $840 \mu\text{m}$ continuum flux density comparable to that of the QSO. These sources cannot be resolved as distinct objects in the *Herschel* images, given the SPIRE point spread function ranging from $17.6''$ at $250 \mu\text{m}$ to $35.2''$ at $500 \mu\text{m}$. We therefore performed a new SED fit of J1015+0020 to remove the contamination of the continuum emitters and more accurately estimate the FIR luminosity (L_{FIR}) of the QSO. We added the ALMA data to the SDSS DR10 (Pâris et al. 2014), WISE (Wright et al. 2010), and *Herschel*/SPIRE (Pilbratt et al. 2010; Griffin et al. 2010) photometry, presented in Duras et al. (2017). We also included three additional near-IR photometric points from the UKIDSS large area survey (Warren et al. 2007a). In order to quantify the possible contribution of Cont1 and Cont2 to the SPIRE fluxes, we built the following SED:

- SED A, which includes photometric points from SDSS to ALMA $840 \mu\text{m}$. The latter takes into account the continuum emission $f_{\text{cont}}^{\text{total}}$ from the QSO, Cont1, and Cont2 (Sect. 3.1).
- SED B, which provides an estimate of the emission from the QSO by removing contamination from Cont1 and Cont2. Specifically, (i) we verified a negligible contribution from the two continuum emitters to the photometric points at wavelength $\leq 22 \mu\text{m}$; (ii) we rescaled the SPIRE fluxes by a factor $r_{\text{cont}} = 3.76$; the latter represents the ratio of $f_{\text{cont}}^{\text{total}}$ to the ALMA-based flux of the QSO $f_{\text{cont}}^{\text{QSO}}$ listed in Table 1; and (iii) we considered $f_{\text{cont}}^{\text{QSO}}$ as the flux density for the photometric point at $840 \mu\text{m}$.

The rest-frame SEDs A and B are shown in Fig. 4. Both are well modelled by the sum of an accretion disk plus torus emission component and a cold dust component in the far-IR. However, the SPIRE photometric points show an excess with respect to the best fit model in SED A. This feature is significantly reduced in SED B, indicating that the main contributors to this excess have been removed. Some residual contamination may still be present in particular at $500 \mu\text{m}$ because of

continuum emitters accounted in the SPIRE fluxes lying outside of the ALMA FOV, i.e. at an angular separation larger than $8.5''$, which translates to a distance ≥ 60 kpc at the QSO redshift. As expected, SED A provides a larger estimate of the FIR luminosity ($\log(L_{\text{FIR}}/\text{erg s}^{-1}) = 46.33 \pm 0.02$) than SED B ($\log(L_{\text{FIR}}/\text{erg s}^{-1}) = 45.69^{+0.14}_{-0.07}$). Accordingly, the SFR derived by following Kennicutt & Evans (2012) is reduced by a factor of 4, i.e. from about 940 to $220^{+68}_{-32} M_{\odot} \text{yr}^{-1}$. The difference with the SFR derived in Duras et al. (2017) is even higher, i.e. a factor of 6. This highlights the importance of sampling the FIR and submillimeter bands with high angular resolution in case of high L_{Bol} and high- z sources, for which a significantly enhanced merger rate is expected (Treister et al. 2012). Similar results have been indeed reported by Banerji et al. (2017) for $z \sim 2.5$ heavily-reddened QSOs, pointing out the need for ALMA observations to uncover these structures around luminous QSOs.

We measured a bolometric luminosity $L_{\text{Bol}} = (1.7 \pm 0.4) \times 10^{47} \text{erg s}^{-1}$ (the uncertainty is dominated by that on the QSO inclination). We therefore corrected the SFR derived from SED B according to Duras et al. (2017), who found that in hyper-luminous QSOs with $\log(L_{\text{Bol}}/\text{erg s}^{-1}) > 47.0$ the AGN contributes to about 50% of the total FIR luminosity. Accordingly, the resulting SFR of J1015+0020 is $\sim 100 M_{\odot} \text{yr}^{-1}$. Throughout the paper, we use quantities derived by SED B as representative of the physical properties of J1015+0020 and its host galaxy.

Our ALMA observation reveals that a significant percentage of the $840 \mu\text{m}$ flux is not associated with the host galaxy of J1015+0020, but instead arises from the surrounding continuum emitters, which likely belong to the QSO overdensity (Sect. 4). We roughly characterised the spectral shape of Cont1, which is the strongest continuum emitting source in our ALMA map (see Fig. 1). Cont1 is detected only by ALMA, but SDSS to WISE non-detections can be used to derive upper limits on its flux in all these bands. Specifically, we fitted the star-forming M82 and the starburst Arp220 galaxy templates (Polletta et al. 2007) to the Cont1 photometry, requiring that they match the ALMA point and leaving the redshift free to vary from $z = 0$ to 5. A M82-like template does not match with the upper limits at any redshift in this range. A $\sim 10 M_{\odot} \text{yr}^{-1}$ star-forming galaxy would be in fact visible in the UKIDSS and WISE bands. On the contrary, a more extreme starburst, i.e. an Arp220-like template, is compatible with the SED of Cont1. This suggests that Cont1 is undergoing intense star formation activity of hundreds of $M_{\odot} \text{yr}^{-1}$, in agreement with the FIR excess (corresponding to $\sim 700 M_{\odot} \text{yr}^{-1}$) observed in SED A associated with the *Herschel* photometry. Such a value is derived under the assumption that Cont1 and Cont2 sources belong to the same structure of J1015+0020, which is supported by the very unlikely possibility to have two field sources with such a close angular separation (see Sect. 4).

We can estimate the star formation activity of the [CII] emitters revealed around J1015+0020 with the relation from Sargsyan et al. (2014),

$$\text{SFR}(M_{\odot} \text{yr}^{-1}) = 1.0 \times 10^{-7} L_{[\text{CII}]} / L_{\odot}, \quad (1)$$

which gives SFR within a 50% uncertainty for low redshift star-forming galaxies. For CompA, CompB, and CompC we derive [CII] luminosities of $\sim 1.2 \times 10^8 L_{\odot}$, $2.2 \times 10^8 L_{\odot}$, and $7.0 \times 10^8 L_{\odot}$, respectively (see Table 1). According to Eq. (1), these values translate to $\text{SFR}^{\text{A}} \sim 12 M_{\odot} \text{yr}^{-1}$, $\text{SFR}^{\text{B}} \sim 22 M_{\odot} \text{yr}^{-1}$, and $\text{SFR}^{\text{C}} \sim 70 M_{\odot} \text{yr}^{-1}$. We note that, in case of J1015+0020, Eq. (1) would correspond to a SFR of only $30 M_{\odot} \text{yr}^{-1}$, indicating a discrepancy with the L_{FIR} -based value. This is likely due to

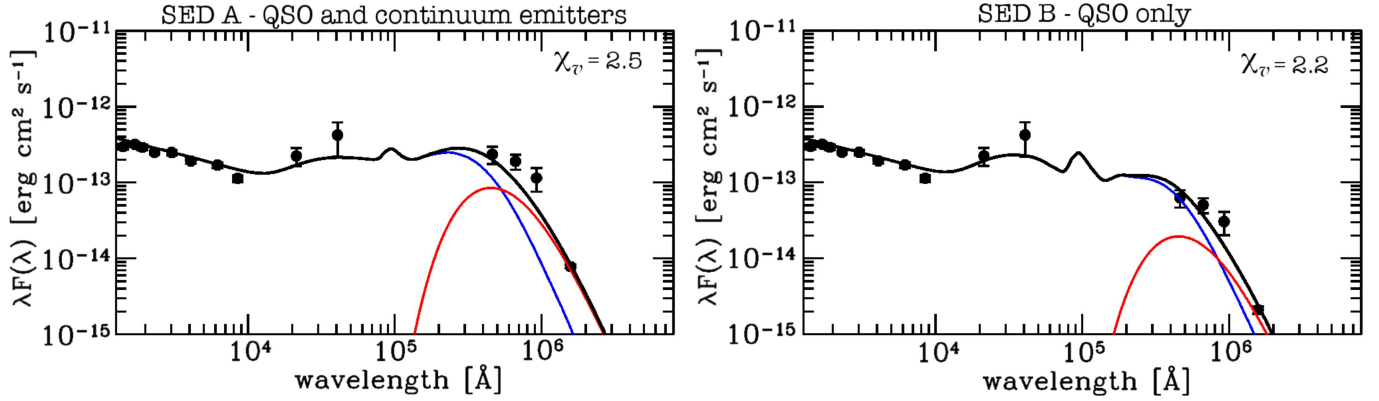


Fig. 4. Rest-frame SED of J1015+0020 before (left panel) and after (right panel) removing the contamination to the FIR fluxes of the continuum emitters in the ALMA FOV. Black circles indicate the rest-frame photometric points. The black curve represents the total best fit model, while blue and red curves refer to the accretion disk plus torus and cold dust emission, respectively.

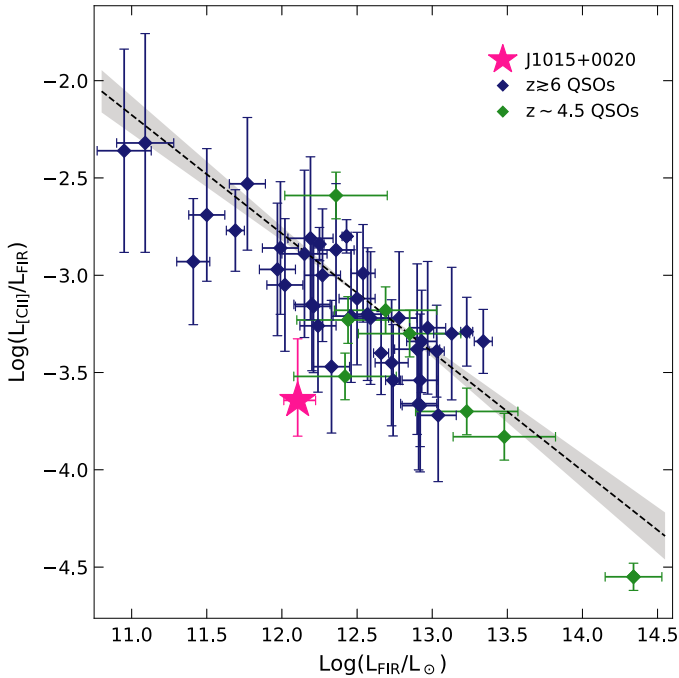


Fig. 5. Ratio of $L_{\text{CII}}/L_{\text{FIR}}$ as a function of L_{FIR} for J1015+0020, compared to high- z QSOs from literature (see Sect. 4 for details) with relative errors. For sources whose uncertainty on L_{FIR} was not available, we assume the average value of the sample. J1015+0020 is indicated by the magenta star, while blue(green) symbols refer to $z \geq 6$ ($z \sim 4.5$) QSOs. The best fit relation obtained from orthogonal linear regression and its 1σ error are shown by the dashed line and shaded area.

the huge radiative power of the AGN in hyper-luminous sources, whose ionising effect reduces the [CII] emission.

Using the [CII] flux derived in Sect. 3.1 and Eq. (1) in Solomon & Vanden Bout (2005), we compute the [CII] luminosity of J1015+0020, $L_{\text{CII}} = (2.9 \pm 0.3) \times 10^8 L_{\odot}$. This implies a $\log(L_{\text{CII}}/L_{\text{FIR}}) = -3.64^{+0.31}_{-0.18}$, which is among the lowest values found for high- z QSO so far. This is shown in Fig. 5, where J1015+0020 (magenta star) is compared to a sample of $z \sim 4\text{--}7$ QSOs from literature. Specifically, we collected L_{CII} and L_{FIR} for 42 sources from the works of Wang et al. (2013, 2016), Venemans et al. (2016, 2017a), Willott et al. (2013, 2015, 2017), Kimball et al. (2015), Díaz-Santos et al. (2016) and Decarli et al. (2017, 2018), with relative uncertainties. Whether

these are not available, we assume the average uncertainty within the sample. The $L_{\text{CII}}/L_{\text{FIR}}$ ratio for $z \sim 4.5\text{--}7$ QSOs span about 1.5 dex and, although with large scatter, a negative trend is evident. We obtain the relation $\log(L_{\text{CII}}/L_{\text{FIR}}) = \alpha \log(L_{\text{FIR}}/L_{\odot}) + \beta$, with $\alpha = -0.61 \pm 0.06$ and $\beta = 4.5 \pm 0.8$ by fitting the data with an orthogonal linear regression accounting for errors on both axes. The slope α is slightly steeper than the value of -0.53 derived in Willott et al. (2017) for $z \sim 6$ QSOs. This difference is due to the addition, in our sample, of the hyper-luminous sources from Trakhtenbrot et al. (2017) and Kimball et al. (2015), populating the high L_{FIR} tail of the sample.

5. SMBH versus host galaxy properties

5.1. SMBH and dynamical masses

We performed a spectral analysis of the rest-frame UV region around the CIV emission line in the SDSS DR10 (Pâris et al. 2014) spectrum of J1015+0020 with the aim of estimating the SMBH mass based on the width of the CIV line profile and luminosity at 1350 Å.

Black hole masses based on CIV can be affected by a factor of a few to ten uncertainties (Shen & Liu 2012). However, in case of J1015+0020 the profile of this line indicates a small deviation from a symmetric profile associated with Keplerian velocity. The velocity shift Δv_{CIV} of the CIV line with respect to the [CII]-based redshift of the QSO is indeed moderate ($\sim 1000 \text{ km s}^{-1}$) for this AGN luminosity regime, once compared to other WISSH QSOs. Vietri et al. (2018) measured the CIV velocity shift with respect to the systemic $H\beta$ emission for a sample of 18 WISSH QSOs, finding values up to 8000 km s^{-1} with an average $\Delta v_{\text{CIV}} \sim 3000 \text{ km s}^{-1}$, which is in agreement with the large shifts typically observed in high-luminosity QSOs (e.g. Sulentic et al. 2017; Hamann et al. 2017). Therefore, we are likely observing the CIV emitting region at large inclination angle, as also indicated by the ALMA observation (see below in this section). According to Vietri et al. (2018), higher line-of-sight inclinations correspond to smaller distortions of the CIV line profile, while low inclinations are associated with very broad, asymmetric profiles due to outflowing gas.

Specifically, we fit the spectral region between 1300 Å and 1700 Å with a model consisting of a power law to reproduce the continuum emission: one Gaussian component to account for the BLR emission of CIV and a second Gaussian component to fit possible CIV wings associated with outflowing gas (see Fig. 6).

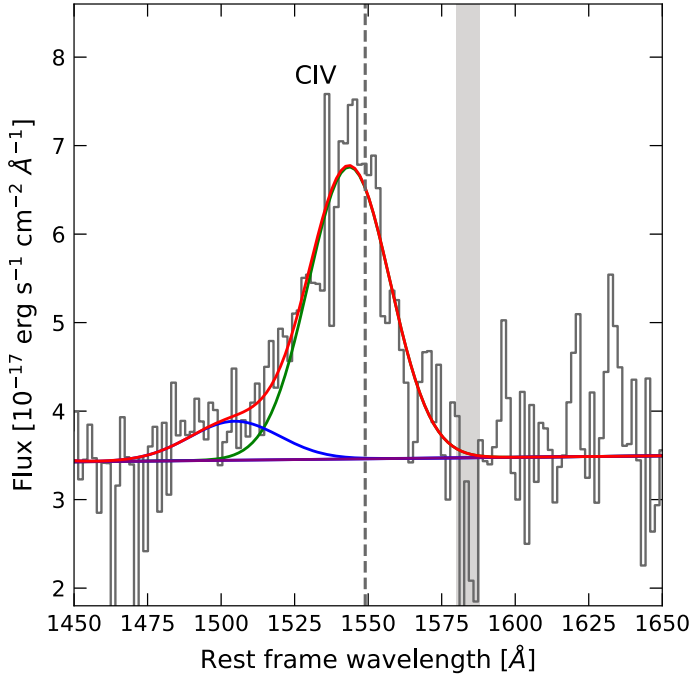


Fig. 6. Rest-frame UV spectrum of J1015+0020, corresponding to the CIV spectral region. The total best fit model is shown in red, while the green(blue) curve refers to the CIV core(wing) emission. Continuum emission is shown in purple. The vertical dashed line corresponds to the z_{CII} of the QSO and the grey shaded region indicates the spectral region excluded from the fit because of telluric features.

We find that the CIV profile is best fitted by the combination of two Gaussians with dominant contribution from the BLR component. The best fit values are $FWHM_{\text{CIV}}^{\text{BLR}} = 6330 \pm 270 \text{ km s}^{-1}$ and $\lambda L_{\lambda 1350\text{\AA}} = (5.0 \pm 0.8) \times 10^{46} \text{ erg s}^{-1}$. These quantities are used to derive M_{BH} according to the single epoch relation from Vestergaard & Peterson (2006) and considering the empirical correction proposed by Coatman et al. (2017) for high-luminosity QSOs to take into account the blueshift $\Delta v_{\text{CIV}} = 1050 \pm 310 \text{ km s}^{-1}$ affecting the CIV line profile as follows:

$$\log\left(\frac{M_{\text{BH}}}{M_{\odot}}\right) = 6.71 + 0.53 \log\left[\frac{\lambda L_{\lambda 1350\text{\AA}}}{10^{44} \text{ erg s}^{-1}}\right] + 2 \log\left(\frac{FWHM_{\text{CIV}}}{1000 \text{ km s}^{-1}}\right) - 2 \log\left[\alpha \left(\frac{\Delta v_{\text{CIV}}}{1000 \text{ km s}^{-1}}\right) + \beta\right], \quad (2)$$

where $\alpha \sim 0.4$ and $\beta \sim 0.6$. The resulting black hole mass of J1015+0020 is $5.7_{-2.1}^{+3.4} \times 10^9 M_{\odot}$, where the uncertainties are dominated by the intrinsic 0.2 dex scatter in the Δv_{CIV} -corrected relation from Coatman et al. (2017). This leads to an Eddington ratio of $\lambda_{\text{Edd}} = 0.23_{-0.09}^{+0.14}$.

The measurements of the FWHM and size of the [CII] emission can be used to estimate the dynamical masses of both the QSO host galaxy and CompA, CompB, and CompC. Under the assumption that the ISM is mainly distributed in an inclined, rotating disk, the dynamical mass can be expressed as

$$M_{\text{dyn}}/M_{\odot} = 9.8 \times 10^8 \left(\frac{D_{\text{[CII]}}}{\text{kpc}}\right) \left(\frac{FWHM_{\text{[CII]}}}{100 \text{ km s}^{-1}} \frac{1}{\sin(i)}\right)^2, \quad (3)$$

where $D_{\text{[CII]}}$ is the deconvolved major axis of the [CII]-emitting region, computed as 1.5 times the deconvolved major axis (Wang et al. 2013), and i is the inclination angle between the

Table 2. Inclination (in units of deg) and dynamical mass (in units of $M_{\odot}$) of J1015+0020 and the [CII] emitters detected in the ALMA map.

Source	J1015+0020	CompA	CompB	CompC
i [deg]	54 ± 12	55 ± 7	–	72 ± 3
$\log(M_{\text{dyn}}/M_{\odot})$	10.6 ± 0.3	9.4 ± 0.3	>10.7	11.2 ± 0.2
$\log(M_{\text{BH}}/M_{\odot})$	9.8 ± 0.2	–	–	–

Notes. For CompB, which is only marginally resolved on one axis in our observation, we assume as size of the [CII] emitting region $1.5'' \times 0.18''$, where $0.18''$ is the minor axis of the ALMA beam. The black hole mass of the QSO (see Sect. 5.1) is also listed.

line of sight and the polar axis of the disk. The FWHM of the [CII] line is related to the circular velocity in the disk by the relation $v_{\text{circ}} = 0.75 \times FWHM_{\text{[CII]}} / \sin(i)$ (Wang et al. 2013). In case of a resolved source the inclination of the disk can be estimated from the ratio of semi-minor to semi-major axes as $i = \arcsin(a_{\text{min}}/a_{\text{maj}})$. The resulting inclination values for J1015+0020 and the [CII] emitters resolved by the ALMA beam are listed in Table 2. We also report the statistical uncertainty associated with i derived from the 2D Gaussian fit of the ALMA data. Nonetheless, we point out that inclination estimates of marginally resolved sources (such as J1015+0020 and CompA), detected at moderate signal-to-noise ratio, can be altered by non-circular beam shapes. In our ALMA observation we might also be losing the more extended, low surface brightness [CII] emission and, thus, underestimating $D_{\text{[CII]}}$. Moreover, the deconvolved size of the source and the semi-axes ratio are estimated from a 2D-Gaussian profile that might not well reproduce the surface-brightness distribution; further discussion on these issues can be found in Trakhtenbrot et al. (2017) and references therein. Deeper, higher resolution observations are therefore needed to reach a firm conclusion concerning the inclination of the detected sources and, therefore, their dynamical masses. Keeping in mind all these limitations, we use Eq. (3) to estimate the dynamical mass of J1015+0020, i.e. $\log(M_{\text{dyn}}/M_{\odot}) = 10.6 \pm 0.3$. We note that assuming the galaxy to be supported by velocity dispersion would lead to a smaller M_{dyn} by a factor of 3.

J1015+0020 is characterised by an extreme ratio of about 1 : 7 of M_{BH} with respect to M_{dyn} , as shown in Fig. 7. Our QSO is compared to the high- z QSOs with single epoch M_{BH} estimates from the sample introduced in Sect. 4. We also plot the relative uncertainties. Whether these are not available, we consider a 0.3 dex error on the (MgII-based) M_{BH} estimates, while we propagate the inclination error to the uncertainty on M_{dyn} . For unresolved sources, we plot the $M_{\text{dyn}} \sin^2(i)$ value as lower limit on the true dynamical mass. The $M_{\text{BH}}-M_{\text{dyn}}$ relation from Jiang et al. (2011), derived from local galaxies in a wide range of $M_{\text{dyn}} \sim 10^9-10^{12} M_{\odot}$, is also shown for comparison. According to the local relation, the $M_{\text{BH}}-M_{\text{dyn}}$ ratio at the observed black hole mass should be about 1 : 600, translating into a host galaxy dynamical mass of $\sim 4 \times 10^{12} M_{\odot}$, suggesting that we are observing the cradle of what would be a giant galaxy at $z = 0$.

As we also resolve CompA and CompC in our ALMA observation, we are able to estimate their dynamical masses $\log(M_{\text{dyn}}^{\text{A}}/M_{\odot}) = 9.4 \pm 0.3$ and $\log(M_{\text{dyn}}^{\text{C}}/M_{\odot}) = 11.2 \pm 0.2$ (see Table 2). CompB is instead only marginally resolved on one axis in our observation and an assumption on $D_{\text{[CII]}}$ is therefore necessary. We use $D_{\text{[CII]}} = 1.5 \times 0.18''$ (where $0.18''$ is the minor axis of the ALMA beam) and Eq. (3) to derive a lower limit on its dynamical mass of $M_{\text{dyn}}^{\text{B}} \sin^2(i) = 5.4 \times 10^{10} M_{\odot}$. Given their

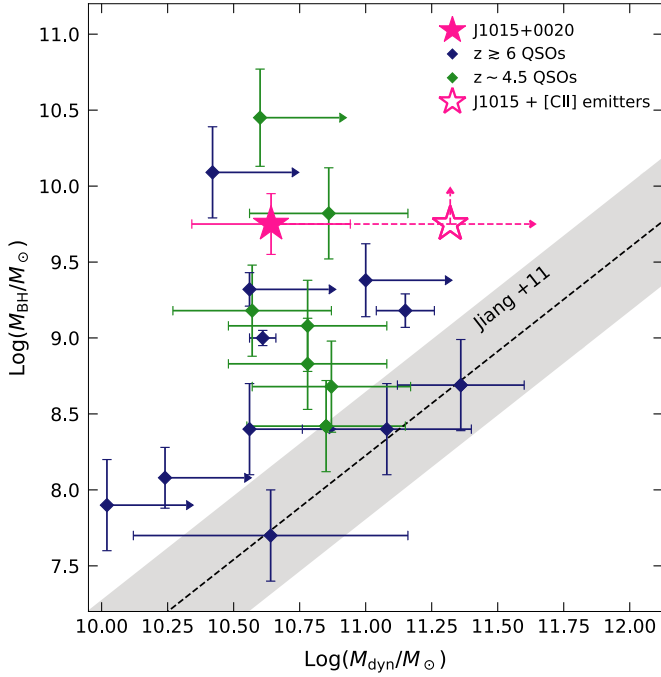


Fig. 7. Black hole mass as a function of dynamical mass of J1015+0020, compared with a sample of high- z , luminous, and hyper-luminous QSOs from literature. Specifically, the magenta star refers to our target, while diamonds refer to $z \sim 4.8$ – 7.1 QSOs observed in [CII] with ALMA (see details in text). M_{dyn} are computed according to Eq. (3), while M_{BH} are single epoch estimates. The best fit $M_{\text{BH}}-M_{\text{dyn}}$ relation from Jiang et al. (2011) is also indicated by the dashed line with the relative 0.42 dex intrinsic scatter (shaded region).

small distance from J1015+0020 all these companions are likely going to merge and eventually build up the mass of the QSO host galaxy. By combining the dynamical masses of the QSO and [CII] emitters we obtain a large value of $M_{\text{dyn}}^{\text{Tot}} \sim 2.2 \times 10^{11} M_{\odot}$ already at $z = 4.4$, as shown in Fig. 7.

5.2. SMBH and host galaxy growth

In the previous sections we have discussed the presence of multiple companions at very close distance ($\lesssim 23$ kpc) from J1015+0020, which likely contributes to the final mass of the QSO host galaxy. By comparing the mass accretion rate of the SMBH with the SFR, we can in principle understand how high- z QSOs have grown to reach their location in Fig. 7, under the assumption that we observe them when most of the black hole and galaxy mass is being assembled.

The value $\dot{M}_{\text{acc}}$ can be derived as $\dot{M}_{\text{acc}} = L_{\text{Bol}}/\epsilon c^2$, once assumed a standard accretion efficiency $\epsilon = 0.1$. The resulting value $\dot{M}_{\text{acc}} = 30 \pm 7 M_{\odot} \text{ yr}^{-1}$ for J1015+0020 is shown in Fig. 8a as a function of the SFR, compared to the 42 high- z QSOs sample introduced in Sect. 5.1. We also include five WISSH QSOs with available M_{BH} and L_{Bol} estimates presented in Bischetti et al. (2017), Duras et al. (2017), Vietri et al. (2018). If unavailable in literature, we compute L_{Bol} by using the bolometric correction from $\lambda L_{1450 \text{ \AA}}$ of Runnoe et al. (2012), assuming as uncertainty the 0.1 dex scatter found for this correction. We also compute the SFR from the $L_{\text{FIR}}(8\text{--}1000 \mu\text{m})$ according to the relation reported in Kennicutt & Evans (2012), with an associated scatter of 0.2 dex, and dividing the SFR by a factor of two in case of QSOs with $L_{\text{Bol}} > 10^{47} \text{ erg s}^{-1}$, as suggested by Duras et al. (2017).

From the relation of Fiore et al. (2017) between mass outflow rate and AGN bolometric luminosity $\log(\dot{M}_{\text{out}}/M_{\odot} \text{ yr}^{-1}) = 0.76 \times \log(L_{\text{Bol}}/\text{erg s}^{-1}) - 32$, we can define the locus of points in Fig. 8a with unitary mass loading factor $\eta = \dot{M}_{\text{out}}/\text{SFR}$, which translates into a $\dot{M}_{\text{acc}} = 1.32 \times \log(\text{SFR}) - 3.64$. J1015+0020 and most of the high- z QSOs clearly lie above this line, suggesting that they are potentially able to develop massive molecular outflows affecting the growth of their host galaxies. Future deep ALMA observations of the CO emission in these objects will be able to confirm this prediction.

From the non-detection of the five QSO companions in the WISE bands, we can compute an upper limit on the L_{Bol} of a possible AGN contribution in these sources. Specifically, by using Mrk231 and NGC6420 templates (Polletta et al. 2007; Fiore et al. 2008), we derive a value of $\log(L_{\text{Bol}}/\text{erg s}^{-1}) \lesssim 43.84$, corresponding to a $\dot{M}_{\text{acc}} \lesssim 0.01 M_{\odot} \text{ yr}^{-1}$. In the companion galaxies, where the AGN effect is absent or very low, it is possible to have star formation activity at a comparable or even higher SFR (orange and red stars in Fig. 8a, see also Sect. 4.2) with respect to the QSO host galaxy itself. This suggests that a significant percentage of stellar mass may be assembled in the QSO satellites and then contribute at later times to the QSO host galaxy mass by mergers.

We compare the SMBH growth and stellar mass assembly timescales in J1015+0020. A basic estimate of the present M_{BH} growth timescale can be derived as the ratio $t_{\text{acc}} = M_{\text{BH}}/\dot{M}_{\text{acc}}$, if a constant mass accretion rate equal to the observed value is assumed. Following the same approach, i.e. by assuming a constant SFR, one can estimate the present stellar mass growth timescale as $t_{\text{SFR}} = M_{\text{dyn}}/\text{SFR}$. Few combined [CII] and CO observations of high- z QSOs are available so far (Wang et al. 2013; Venemans et al. 2017b), providing a wide range of molecular gas fractions contributing to the total dynamical mass, from $\sim 10\%$ to 80% , if a negligible dark matter content in the inner regions of the galaxy is assumed (Genzel et al. 2017). Therefore, the resulting t_{SFR} should be considered as upper limits of the real stellar mass assembly timescales. Figure 8b shows t_{acc} as a function of t_{SFR} for J1015+0020 and the same high- z QSOs plotted in Fig. 7. The value t_{SFR} ranges from about 30 Myr to 1 Gyr, while the black holes have reached their current mass in ~ 1 – 300 Myr. The two growth timescales appear comparable, bearing in mind the large uncertainties affecting these measurements, as most sources lie close to the 1:1 relation. This suggests that the QSOs of Fig. 7 are moving in the $M_{\text{BH}}-M_{\text{dyn}}$ plane in parallel to the local relation. However, in most of the $z \sim 4.5$ QSOs the $t_{\text{acc}}/t_{\text{SFR}}$ ratio is < 1 . This is likely due to the extreme accretion rates $\lambda_{\text{Edd}} \sim 0.8$ – 5 measured in these sources.

An alternative approach consists of estimating the exponential growth timescale of a black hole accreting at constant Eddington rate $\lambda_{\text{Edd}} = L_{\text{Bol}}/L_{\text{Edd}}$, where $L_{\text{Edd}}/\text{erg s}^{-1} = 1.28 \times 10^{38} M_{\text{BH}}/M_{\odot}$ is the Eddington luminosity. Specifically, following Volonteri & Rees (2005):

$$t_{\text{Edd}} = \tau_{\text{acc}} \frac{\epsilon}{1 - \epsilon} \lambda_{\text{Edd}} \times \ln(M_{\text{BH}}/M_{\text{BH}}^0)$$

where $\tau_{\text{acc}} \sim 0.45$ Gyr is the characteristic accretion timescale (Shapiro 2005) and $M_{\text{BH}}^0 \sim 10^3 M_{\odot}$ is the initial mass of the black hole seed (see Fig. 8c). We thus calculate the exponential stellar mass growth timescale of the host galaxy by assuming constant $\tau = \text{SFR}/M_{\text{dyn}}$ ratio, equal to the observed value as follows:

$$t_{\tau} = \frac{1}{\tau} \left[\ln(M_{\text{dyn}}) - \ln(M_{\text{dyn}}^0) \right],$$

where M_{dyn}^0 is the initial dynamical mass. Specifically, we compute $M_{\text{dyn}}^0 = 9.5 \times 10^8 M_{\odot}$ from the minimum mass of a

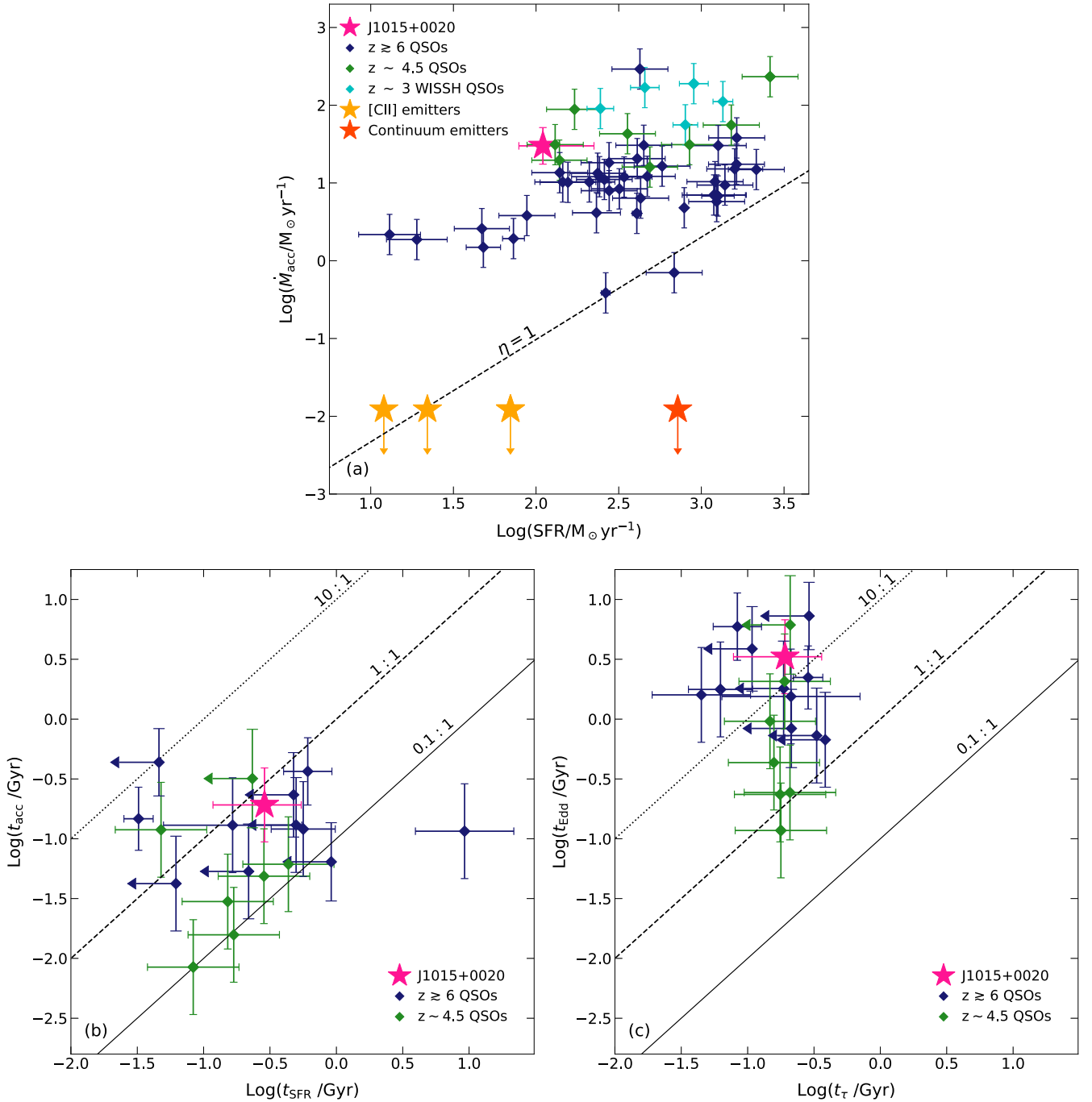


Fig. 8. *Panel a:* black hole accretion rate as a function of the SFR for J1015+0020 (magenta star) compared to the sample of $z \gtrsim 4.5$ QSOs described in Sect. 4 and five WISSH QSOs from Bischetti et al. (2017), Vietri et al. (2018), Duras et al. (2017). Orange (red) stars refer to the [CII](continuum) emitters detected in the ALMA map. The locus of points with unitary mass loading factor is also indicated by the dashed line. *Panel b:* present BH vs. host galaxy growth timescales, derived assuming constant $\dot{M}_{\text{acc}}$ and SFR equal to the observed values. The dotted, dashed, and solid lines indicate a 10:1, 1:1, and 0.1:1 growth timescale, respectively. *Panel c:* exponential BH vs. host galaxy growth timescales, derived assuming constant λ_{Edd} and $\text{SFR}/M_{\text{dyn}}$. Lines as in panel b.

star-forming dark matter halo (Finlator et al. 2011) by rescaling for the cosmological baryon fraction Ω_{M} , since for the majority of the sources (including J1015+0020) we are not able to distinguish between gas and stellar mass. We find that t_{Edd} is on average a factor of 10 larger than t_{acc} (see Fig. 8c), ranging from 300 Myr to 6 Gyr. This translates into a typical ratio of exponential black hole mass growth to exponential stellar mass growth of 10:1, as t_{τ} is comparable to t_{SFR} . According to this scenario, the host galaxies of high- z QSOs have grown in a shorter timescale than their central SMBHs. A shorter BH growth timescale can be obtained

by either assuming a more massive seed (e.g. $10^5 M_{\odot}$) or a super-Eddington accretion regime onto stellar mass seeds ($\sim 100 M_{\odot}$; e.g. Volonteri 2010; Valiante et al. 2017).

According to this scenario, the host galaxies of high- z QSOs are growing faster than their central SMBHs. Finally, we note that adopting a $\tau = 2.4 \text{ Gyr}^{-1}$, typical of $z = 6-7$ galaxies (Stark et al. 2009; González et al. 2010), would shift most QSOs towards larger stellar t_{τ} , i.e. closer to the 1:1 relation. This indicates that the bulk of the high- z QSOs sample considered here are undergoing a peculiar phase of intense SF activity.

6. Conclusions

In this work, we report on the analysis of the ALMA high-resolution $0.18'' \times 0.21''$ observation of the $840\mu\text{m}$ continuum and [CII] $\lambda 157.74\mu\text{m}$ line emission of the WISSH quasar J1015+0020 and its surrounding (8.5 arcsec radius) field. This data allows us to characterise the host galaxy and environment properties of this hyper-luminous QSO at $z \sim 4.4$. Our main findings can be summarised as follows:

- The ALMA observation reveals an exceptional overdensity of [CII]-emitting companions with a very small ($<150\text{ km s}^{-1}$) velocity shift with respect to the QSO redshift. Specifically, we report the discovery of the closest companion observed so far in submillimeter observations of high- z QSOs. This companion is only 2.2 kpc distance and merging with the QSO, as indicated by the increased velocity dispersion in the host of J1015+0020 offset from the AGN location. The other two [CII] emitters are located at 8 and 17 kpc.
- We also detected two continuum emitters (Cont1 and Cont2 in Table 1) within an angular separation of less than 3.5 arcsec, which are characterised by a $840\mu\text{m}$ continuum flux density comparable to that of the hyper-luminous QSO. These sources do not show line emission in the ALMA band but are likely physically associated with J1015+0020. In fact, they exceed by a factor of 100 the number of expected sources according to the $850\mu\text{m}$ $\log(N)-\log(S)$ (see Sect. 4). Comparing the five companions detected around J1015+0020 with typical number density of galaxies from deep submillimeter surveys, clearly indicates the presence of a significant overdensity of star-forming galaxies around this $z \sim 4.4$ QSO. We are observing the early phase of the formation of a giant, massive galaxy, assembled by merging of the ALMA-detected companions with a $\sim 10^{10} M_{\odot}$ SMBH at its centre.
- We are able to accurately build up the SED of the emission from the QSO, by quantifying and removing the contribution to the $250\text{--}840\mu\text{m}$ fluxes from Cont1 and Cont2 (see Sect. 4.2). The QSO host galaxy has a SFR of about $100 M_{\odot} \text{ yr}^{-1}$, while the bulk of the SF activity takes place in Cont1 and Cont2, for which we derive a SFR $\sim 700 M_{\odot} \text{ yr}^{-1}$. The [CII] emitters contribute to additional $100 M_{\odot} \text{ yr}^{-1}$. A significant percentage of the stellar mass assembly at earlier epochs may have therefore taken place in the companion galaxies, more than in the QSO host galaxy itself (e.g. Anglés-Alcázar et al. 2017a).
- For J1015+0020 we measure a SMBH mass of $\sim 6 \times 10^9 M_{\odot}$ using a single epoch relation based on the CIV emission line profile, which is a reliable tracer of the mass since it does not exhibit a strong blueshifted wing associated with non-virial motions (Coatman et al. 2017; Vietri et al. 2018, see Sect. 7). From the [CII] line profile and emitting region, we also compute the dynamical mass of the QSO host galaxy, $M_{\text{dyn}} \sim 4 \times 10^{10} M_{\odot}$. This translates into an extreme $M_{\text{dyn}}/M_{\text{BH}}$ ratio of ~ 7 , which is a factor of 100 smaller than what typically observed in local galaxies. According to the local relation from Jiang et al. (2011), such a SMBH mass should be hosted in giant galaxy with a stellar mass of $\sim 1.3 \times 10^{12} M_{\odot}$. Remarkably, the total stellar mass of QSO plus [CII]-emitting companions already exceeds $10^{11} M_{\odot}$ at $z \sim 4.4$.
- J1015+0020, as most of the 47 QSOs at $z \gtrsim 4.5$ from the sample described in Sect. 5.1, is potentially able to drive a massive molecular outflow affecting the SFR in the host galaxy, according to the relation from Fiore et al. (2017) between

$\dot{M}_{\text{out}}$ and L_{Bol} (see Sect. 5.2). Dedicated ALMA observations of the CO emission in these objects are needed to further investigate this prediction. Figure 8b and c compare the SMBH versus galaxy growth timescales of high- z QSOs by assuming a constant ($\dot{M}_{\text{acc}}$, SFR) and (λ_{Edd} , τ), respectively. We find that the present growth rate of the host galaxy is comparable to that inferred for the SMBH (Fig. 8b), while Fig. 8c suggests that the time necessary to reach the observed $\dot{M}_{\text{dyn}}$ is shorter than the time required to the SMBH to accrete the observed mass.

Acknowledgements. This paper makes use of the following ALMA data: ADS/JAO.ALMA#2016.1.00718.S. ALMA is a partnership of ESO (representing its member states), NSF (USA), and NINS (Japan), together with NRC (Canada), MOST and ASIAA (Taiwan), and KASI (Republic of Korea), in cooperation with the Republic of Chile. The Joint ALMA Observatory is operated by ESO, AUI/NRAO, and NAOJ. Funding for SDSS-III has been provided by the Alfred P. Sloan Foundation, the Participating Institutions, the National Science Foundation, and the U.S. Department of Energy Office of Science. The SDSS-III web site is <http://www.sdss3.org/>. SDSS-III is managed by the Astrophysical Research Consortium for the Participating Institutions of the SDSS-III Collaboration including the University of Arizona, the Brazilian Participation Group, Brookhaven National Laboratory, Carnegie Mellon University, University of Florida, the French Participation Group, the German Participation Group, Harvard University, the Instituto de Astrofísica de Canarias, the Michigan State/Notre Dame/JINA Participation Group, Johns Hopkins University, Lawrence Berkeley National Laboratory, Max Planck Institute for Astrophysics, Max Planck Institute for Extraterrestrial Physics, New Mexico State University, New York University, Ohio State University, Pennsylvania State University, University of Portsmouth, Princeton University, the Spanish Participation Group, University of Tokyo, University of Utah, Vanderbilt University, University of Virginia, University of Washington, and Yale University. The UKIDSS project is defined in Lawrence et al. (2007). UKIDSS uses the UKIRT Wide Field Camera (WFCAM; Casali et al. 2007) and a photometric system described in Hewett et al. (2006). The pipeline processing and science archive are described in Irwin et al. (in prep.) and Hambly et al. (2008). We have used data from the 2nd data release, which is described in detail in Warren et al. (2007b). We thank E. Merlin for the assistance with the HST photometry and M. Ginolfi for the useful discussion on the detection of sources in interferometric images. This research project was supported by the DFG Cluster of Excellence “Origin and Structure of the Universe” (www.universe-cluster.de). C.P. acknowledges support from the Science and Technology Foundation (FCT, Portugal) through the Post-doctoral Fellowship SFR/BDP/90559/2012, PEst-OE/FIS/UI2751/2014, and PTDC/FIS-AST/2194/2012. S.C. acknowledges financial support from the Science and Technology Facilities Council (STFC). LZ acknowledges financial support under ASI/INAF contract I/037/12/0. CF acknowledges support from the European Union Horizon 2020 research and innovation programme under the Marie Skłodowska-Curie grant agreement No 664931.

References

- Alexander, D. M., & Hickox, R. C. 2012, *New Astron. Rev.*, **56**, 93
- Anglés-Alcázar, D., Faucher-Giguère, C.-A., Kereš, D., et al. 2017a, *MNRAS*, **470**, 4698
- Anglés-Alcázar, D., Faucher-Giguère, C.-A., Quataert, E., et al. 2017b, *MNRAS*, **472**, L109
- Aravena, M., Decarli, R., Walter, F., et al. 2016a, *ApJ*, **833**, 71
- Aravena, M., Decarli, R., Walter, F., et al. 2016b, *ApJ*, **833**, 68
- Bañados, E., Venemans, B. P., Decarli, R., et al. 2016, *ApJS*, **227**, 11
- Banerji, M., Carilli, C. L., Jones, G., et al. 2017, *MNRAS*, **465**, 4390
- Bischetti, M., Piconcelli, E., Vietri, G., et al. 2017, *A&A*, **598**, A122
- Bouwens, R. J., Illingworth, G. D., Oesch, P. A., et al. 2015, *ApJ*, **803**, 34
- Carniani, S., Marconi, A., Biggs, A., et al. 2013, *A&A*, **559**, A29
- Carniani, S., Maiolino, R., De Zotti, G., et al. 2015, *A&A*, **584**, A78
- Casali, M., Adamson, A., Alves de Oliveira, C., et al. 2007, *A&A*, **467**, 777
- Cicone, C., Maiolino, R., Gallerani, S., et al. 2015, *A&A*, **574**, A14
- Coatman, L., Hewett, P. C., Banerji, M., et al. 2017, *MNRAS*, **465**, 2120
- Decarli, R., Walter, F., Venemans, B. P., et al. 2017, *Nature*, **545**, 457
- Decarli, R., Walter, F., Venemans, B. P., et al. 2018, *ApJ*, **854**, 97
- De Rosa, G., Decarli, R., Walter, F., et al. 2011, *ApJ*, **739**, 56
- De Rosa, G., Venemans, B. P., Decarli, R., et al. 2014, *ApJ*, **790**, 145
- Díaz-Santos, T., Assef, R. J., Blain, A. W., et al. 2016, *ApJ*, **816**, L6
- Di Matteo, T., Springel, V., & Hernquist, L. 2005, *Nature*, **433**, 604

- Duras, F., Bongiorno, A., Piconcelli, E., et al. 2017, *A&A*, **604**, A67
- Finlator, K., Oppenheimer, B. D., & Davé, R. 2011, *MNRAS*, **410**, 1703
- Fiore, F., Feruglio, C., Shankar, F., et al. 2017, *A&A*, **601**, A143
- Fiore, F., Grazian, A., Santini, P., et al. 2008, *ApJ*, **672**, 94
- Fujimoto, S., Ouchi, M., Ono, Y., et al. 2016, *ApJS*, **222**, 1
- Gabor, J. M., Capelo, P. R., Volonteri, M., et al. 2016, *A&A*, **592**, A62
- Genzel, R., Schreiber, N. M. F., Übler, H., et al. 2017, *Nature*, **543**, 397
- González, V., Labbé, I., Bouwens, R. J., et al. 2010, *ApJ*, **713**, 115
- Griffin, M. J., Abergel, A., Abreu, A., et al. 2010, *A&A*, **518**, L3
- Hamann, F., Zakamska, N. L., Ross, N., et al. 2017, *MNRAS*, **464**, 3431
- Hambly, N. C., Collins, R. S., Cross, N. J. G., et al. 2008, *MNRAS*, **384**, 637
- Hewett, P. C., Warren, S. J., Leggett, S. K., & Hodgkin, S. T. 2006, *MNRAS*, **367**, 454
- Hopkins, P. F., Hernquist, L., Cox, T. J., & Kereš, D. 2008, *ApJS*, **175**, 356
- Jiang, L., Fan, X., Vestergaard, M., et al. 2007, *AJ*, **134**, 1150
- Jiang, Y.-F., Greene, J. E., & Ho, L. C. 2011, *ApJ*, **737**, L45
- Kennicutt, R. C., & Evans, N. J. 2012, *ARA&A*, **50**, 531
- Kimball, A. E., Lacy, M., Lonsdale, C. J., & Macquart, J.-P. 2015, *MNRAS*, **452**, 88
- Kormendy, J., & Ho, L. C. 2013, *ARA&A*, **51**, 511
- Lawrence, A., Warren, S. J., Almaini, O., et al. 2007, *MNRAS*, **379**, 1599
- Lutz, D., Mainieri, V., Rafferty, D., et al. 2010, *ApJ*, **712**, 1287
- Maiolino, R., Cox, P., Caselli, P., et al. 2005, *A&A*, **440**, L51
- McMullin, J. P., Waters, B., Schiebel, D., Young, W., & Golap, K. 2007, *ASP Conf. Ser.*, **376**, 127
- Menci, N., Fiore, F., Puccetti, S., & Cavaliere, A. 2008, *ApJ*, **686**, 219
- Menci, N., Gatti, M., Fiore, F., & Lamastra, A. 2014, *A&A*, **569**, A37
- Nardini, E., & Risaliti, G. 2011, *MNRAS*, **415**, 619
- Netzer, H. 2009, *MNRAS*, **399**, 1907
- Pâris, I., Petitjean, P., Aubourg, É., et al. 2014, *A&A*, **563**, A54
- Pilbratt, G. L., Riedinger, J. R., Passvogel, T., et al. 2010, *A&A*, **518**, L1
- Planck Collaboration XIII. 2016, *A&A*, **594**, A13
- Polletta, M., Tajer, M., Maraschi, L., et al. 2007, *ApJ*, **663**, 81
- Rosario, D. J., Santini, P., Lutz, D., et al. 2013, *ApJ*, **771**, 63
- Runnoe, J. C., Brotherton, M. S., & Shang, Z. 2012, *MNRAS*, **422**, 478
- Sargsyan, L., Samsonyan, A., Leboutteiller, V., et al. 2014, *ApJ*, **790**, 15
- Shapiro, S. L. 2005, *ApJ*, **620**, 59
- Shen, Y., & Liu, X. 2012, *ApJ*, **753**, 125
- Simpson, J. M., Smail, I., Swinbank, A. M., et al. 2015, *ApJ*, **807**, 128
- Solomon, P. M., & Vanden Bout, P. A. 2005, *ARA&A*, **43**, 677
- Stark, D. P., Ellis, R. S., Bunker, A., et al. 2009, *ApJ*, **697**, 1493
- Sulentic, J. W., del Olmo, A., Marziani, P., et al. 2017, *A&A*, **608**, A122
- Trakhtenbrot, B., Lira, P., Netzer, H., et al. 2017, *ApJ*, **836**, 8
- Treister, E., Schawinski, K., Urry, C. M., & Simmons, B. D. 2012, *ApJ*, **758**, L39
- Valiante, R., Schneider, R., Salvadori, S., & Gallerani, S. 2014, *MNRAS*, **444**, 2442
- Valiante, R., Schneider, R., & Volonteri, M. 2016, *PASA*, **33**, e032
- Valiante, R., Agarwal, B., Habouzit, M., & Pezzulli, E. 2017, *PASA*, **34**, e031
- Venemans, B. P., Bañados, E., Decarli, R., et al. 2015, *ApJ*, **801**, L11
- Venemans, B. P., Walter, F., Zschaechner, L., et al. 2016, *ApJ*, **816**, 37
- Venemans, B. P., Walter, F., Decarli, R., et al. 2017a, *ApJ*, **837**, 146
- Venemans, B. P., Walter, F., Decarli, R., et al. 2017b, *ApJ*, **845**, 154
- Vestergaard, M., & Peterson, B. M. 2006, *ApJ*, **641**, 689
- Vietri, G., Piconcelli, E., Bischetti, M., et al. 2018, *A&A*, **617**, A81
- Volonteri, M. 2010, *A&ARv*, **18**, 279
- Volonteri, M., & Rees, M. J. 2005, *ApJ*, **633**, 624
- Volonteri, M., Capelo, P. R., Netzer, H., et al. 2015, *MNRAS*, **449**, 1470
- Wang, R., Wagg, J., Carilli, C. L., et al. 2013, *ApJ*, **773**, 44
- Wang, R., Wu, X.-B., Neri, R., et al. 2016, *ApJ*, **830**, 53
- Warren, S. J., Hambly, N. C., Dye, S., et al. 2007a, *MNRAS*, **375**, 213
- Warren, S. J., Cross, N. J. G., Dye, S., et al. 2007b, *ArXiv e-prints* [[astro-ph/0703037](#)]
- Willott, C. J., Omont, A., & Bergeron, J. 2013, *ApJ*, **770**, 13
- Willott, C. J., Bergeron, J., & Omont, A. 2015, *ApJ*, **801**, 123
- Willott, C. J., Bergeron, J., & Omont, A. 2017, *ApJ*, **850**, 108
- Wright, E. L., Eisenhardt, P. R. M., Mainzer, A. K., et al. 2010, *AJ*, **140**, 1868