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On the size of the CO-depletion radius in the IRDC G351.77–0.51

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ABSTRACT

An estimate of the degree of CO-depletion (f_D) provides information on the physical conditions occurring in the innermost and densest regions of molecular clouds. A key parameter in these studies is the size of the depletion radius, i.e. the radius within which the C-bearing species, and in particular CO, are largely frozen on to dust grains. A strong depletion state (i.e. $f_D > 10$, as assumed in our models) is highly favoured in the innermost regions of dark clouds, where the temperature is < 20 K and the number density of molecular hydrogen exceeds a few $\times 10^4$ cm $^{-3}$. In this work, we estimate the size of the depleted region by studying the Infrared Dark Cloud (IRDC) G351.77–0.51. Continuum observations performed with the *Herschel Space Observatory* and the *Large APEX BOlometer CAmera*, together with APEX C 18 O and C 17 O $J = 2 \rightarrow 1$ line observations, allowed us to recover the large-scale beam- and line-of-sight-averaged depletion map of the cloud. We built a simple model to investigate the depletion in the inner regions of the clumps in the filament and the filament itself. The model suggests that the depletion radius ranges from 0.02 to 0.15 pc, comparable with the typical filament width (i.e. ~ 0.1 pc). At these radii, the number density of H $_2$ reaches values between 0.2 and 5.5×10^5 cm $^{-3}$. These results provide information on the approximate spatial scales on which different chemical processes operate in high-mass star-forming regions and also suggest caution when using CO for kinematical studies in IRDCs.

Key words: astrochemistry – molecular processes – stars: formation – ISM: molecules – Galaxy: abundances – submillimetre: ISM.

1 INTRODUCTION

In the last two decades, evidence of CO-depletion has been widely found in numerous sources (e.g. Caselli et al. 1999; Kramer et al. 1999; Bergin et al. 2002; Fontani et al. 2012; Wiles et al. 2016), with important consequences for the chemistry, such as an increase in the fraction of deuterated molecules (Bacmann et al. 2003). Cold and dense regions within molecular clouds with temperatures, T , < 20 K (e.g. Caselli et al. 2008) and number densities, $n(\text{H}_2)$, $> \text{a few} \times 10^4$ cm $^{-3}$ (e.g. Crapsi et al. 2005; see also Bergin & Tafalla 2007, and references therein) provide the ideal conditions that favour depletion of heavy elements on to the surface of dust

grains. Here, chemical species can contribute not only in the gas-phase chemistry, but also in surface reactions by being frozen on to the dust grains. This chemical dichotomy can lead to the formation of complex molecules, modifying the chemical composition over a significant volume of a cloud (e.g. Herbst & van Dishoeck 2009).

How much of the CO is depleted on to the surface of dust grains is usually characterized by the depletion factor (e.g. Caselli et al. 1999; Fontani et al. 2012), defined as the ratio between the ‘expected’ CO abundance with respect to H $_2$ ($\chi_{\text{CO}}^{\text{E}}$) and the observed one ($\chi_{\text{CO}}^{\text{O}}$):

$$f_D = \frac{\chi_{\text{CO}}^{\text{E}}}{\chi_{\text{CO}}^{\text{O}}} \quad (1)$$

The main CO isotopologue (i.e. $^{12}\text{C}^{16}\text{O}$) is virtually always optically thick, and thus its intensity is not proportional to the CO column

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density. One of the less abundant CO isotopologues (i.e. C^{18}O , as in this work) can be used to obtain a much more accurate estimate of f_{D} .

In general, in dense and cold environments, such as massive clumps in an early stage of evolution, the estimate of the depletion degree provides important information such as the increase of the efficiency of chemical reactions that occur on the dust grain surface, favoured by the high concentration of the depleted chemical species.

High-mass star-forming regions are potentially more affected by large-scale depletion: the high volume densities of H_2 both in the clumps and in the surrounding intra-clump regions make these sources more prone to high levels of molecular depletion (e.g. Giannetti et al. 2014) because the time-scale after which depletion becomes important (see e.g. Caselli et al. 1999, 2002) decreases with increasing volume density. In fact, the depletion time-scale (τ_{dep}) is inversely proportional to the absorption rate, $\kappa_{\text{abs}} = \sigma \langle v \rangle n_{\text{g}} S$ [s^{-1}], of the freezing-out species, where σ is the dust grain cross-section, $\langle v \rangle$ is the mean velocity of the Maxwellian distribution of gaseous particles, n_{g} is the dust grain number density, and S is the sticking coefficient (see also Bergin & Tafalla 2007, for more details).

In different samples of young high-mass star-forming regions, the observed depletion factors vary between ≈ 1 , in the case of complete absence of depletion, and a few tens (e.g. in Thomas & Fuller 2008; Fontani et al. 2012; Feng et al. 2016). The largest values reported are those in Fontani et al. (2012) where f_{D} can reach values of up to 50–80, a factor of ~ 10 larger than those observed in samples of low-mass clouds (e.g. Bacmann et al. 2002, 2003; Ceccarelli et al. 2007; Christie et al. 2012). Giannetti et al. (2014) noted that the method used by Fontani et al. (2012) to calculate $N(\text{H}_2)$ yielded values ~ 2.7 times larger than that they themselves used to obtain this quantity, mainly due to the different dust absorption coefficient, κ , assumed [i.e. an absorption coefficient at $870 \mu\text{m}$, $\kappa_{870\mu\text{m}} = 1.8 \text{ cm}^2 \text{ g}^{-1}$ assumed by Giannetti et al. (2014), while $\kappa_{870\mu\text{m}} \sim 0.77 \text{ cm}^2 \text{ g}^{-1}$ assumed by Fontani et al. (2012) and derived here following Beuther et al. (2005) and Hildebrand (1983)].

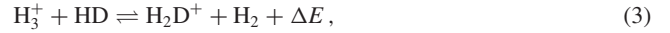
Rescaling the depletion factors of these sources under the same assumptions, we note that the typical values of f_{D} vary from a minimum of 1 to a maximum of ~ 10 , except for some particular cases. However, there are studies, which were made with much higher resolution, that reported even larger depletion factors: this is the case for the high-mass core in the Infrared Dark Cloud (IRDC) G28.34+0.06 where, at a spatial resolution of $\sim 10^{-2}$ pc (using a source distance of 4.2 kpc; Urquhart et al. 2018), f_{D} reaches values of 10^2 – 10^3 , the highest values of f_{D} found to date (Zhang et al. 2009).

Observing depletion factors ranging from 1 to 10 means that along the line of sight there will be regions in which the CO is completely in gas phase (i.e. not depleted), and other regions in which the depletion factor reaches values larger than 10. The knowledge of the size of the region within which most of the CO is largely frozen on to dust grains (the depletion radius) provides information on the approximate spatial scales on which different chemical processes operate in high-mass star-forming regions. At these scales, for example, it is reasonable to think that derivation of H_2 column densities from CO lines (see e.g. Bolatto, Wolfire & Leroy 2013, and reference therein) and the studies of the gas dynamics using the carbon monoxide isotopologues lines are strongly affected by depletion process, as not all of the CO present is directly observable. Furthermore, the disappearance of CO from the gas phase also favours the deuteration process (e.g. Roberts & Millar 2000; Bacmann et al. 2003). This is due to the following reaction that becomes increasingly inefficient, while the amount of frozen

CO increases:



In a high CO-depletion regime, H_3^+ is destroyed less efficiently by reaction (2), remaining available for other reactions. Deuterium enrichment in the gas phase is enhanced by the formation of H_2D^+ through the reaction:



which proceeds from left to right, unless there is a substantial fraction of ortho- H_2 (e.g. Gerlich, Herbst & Roueff 2002), the only case in which reaction (3) is endothermic and naturally proceeds in the direction in which the observed deuterated fraction decreases. For this reason, for a higher degree of CO-depletion, the formation of H_2D^+ by reaction (3) is more efficient due to the greater amount of H_3^+ available.

Single-cell and one-dimensional numerical calculations (e.g. Sipilä, Caselli & Harju 2013; Kong et al. 2015) are fast enough to include detailed treatment of the depletion process together with large comprehensive networks to follow the chemistry evolution in these environments.

However, when we move to more accurate three-dimensional simulations, strong assumptions are needed to limit the computational time and allow the study of the dynamical effects on a smaller set of chemical reactions. Körtgen et al. (2017) and Körtgen et al. (2018), for example, analysed the dependence of the deuteration process on the dynamical evolution, exploring how the chemical initial conditions influence the process. They performed 3D magnetohydrodynamical (MHD) simulations fully coupled for the first time with a chemical network that describes deuterium chemistry under the assumption of complete depletion (i.e. $f_{\text{D}} = \infty$) on to a spatial scale of ~ 0.1 pc from the collapse centre. However, the real extent of the depletion radius is still unknown, and this uncertainty might alter the theoretical models that describe the chemical evolution of star-forming regions.

One way to shed light on these issues is by mapping high-mass star-forming regions on both large and small scales with the goal to determine the depletion factor fluctuations over a broad range of densities and temperatures.

In this paper, we selected the nearest and most massive filament found in the $870 \mu\text{m}$ APEX Telescope Large AREA Survey of the Galaxy (ATLASGAL) [see Schuller et al. (2009) and Li et al. (2016)], the IRDC G351.77–0.51, presented in Section 2. In Section 3, we present the dust temperature map obtained by fitting the spectral energy distribution (SED) using continuum observations of the *Herschel* Infrared Galactic Plane Survey (Hi-GAL; Molinari et al. 2010). In addition, we compute the column density maps of H_2 and C^{18}O . Combining the H_2 and C^{18}O column density maps allows us to produce the large-scale depletion map of the entire source. We evaluate variations of depletion over the whole structure with the aim to understand which scales are affected by depletion the most. By assuming a volume density distribution profile for the H_2 and a step-function to describe the abundance profile of C^{18}O with respect to H_2 , we estimate the size of the depletion radius (Section 4). In Section 5, we finally explore how much the estimates of the depletion radius are affected by the model's assumptions.

2 SOURCE AND DATA: IRDC G351.77–0.51

In Fig. 1, we show a continuum image at $870 \mu\text{m}$ of G351.77–0.51 from the ATLASGAL survey (Li et al. 2016), in which the clumps

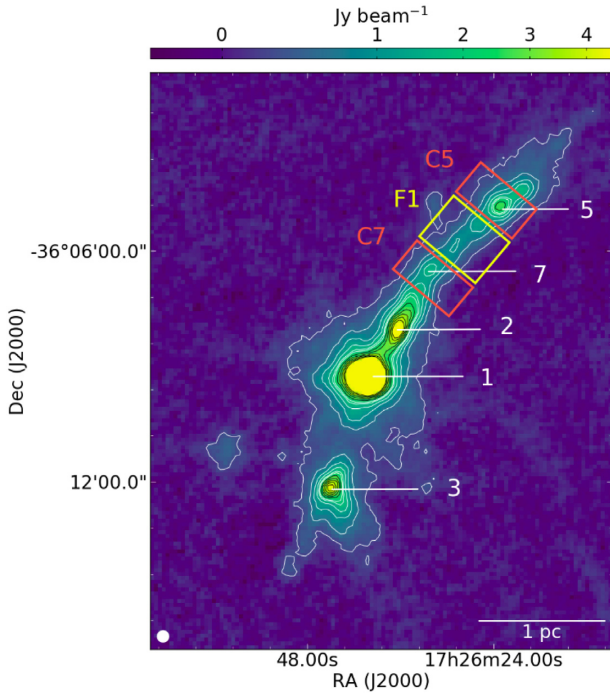


Figure 1. Large APEX BOlometer CAmera (LABOCA) map of the 870 μm dust continuum emission from the IRDC G351.77–0.51. The white contours are from 5σ (0.2 Jy beam^{-1}) to 2 Jy beam^{-1} while the black contours are from 2.4 to 4 Jy beam^{-1} in steps of 10σ . The white labels indicate some of the dust clumps identified by Leurini et al. (2011b). The orange and yellow boxes (i.e. regions C5, C7, and F1) are linked to the models discussed in Section 4.

(labels 1, 2, 3, 5, and 7) appear well pronounced along the filamentary structure of the star-forming region. Among the twelve original clumps defined in Leurini et al. (2011b), we selected only the five clumps for which data of both C^{17}O and C^{18}O are available (see Section 3.2.1).

G351.77–0.51 is the most massive filament in the ATLASGAL survey within 1 kpc from us. Leurini et al. (2019) estimate $M \sim 2600 M_{\odot}$ [twice that listed by Li et al. (2016), who used different values for dust temperature and opacity]. Following the evolutionary sequence of massive clumps defined by König et al. (2017) – originally outlined by Giannetti et al. (2014) and Csengeri et al. (2016) – the main clump of G351.77–0.51 was classified as an infrared-bright (IRB) source,¹ and revealed hot-core features (see Giannetti et al. 2017a). Leurini et al. (2011a) and Leurini et al. (2019) studied the velocity field of the molecular gas component. They discussed the velocity gradients by considering whether they might be due to rotation, or outflow(s) around clump-1, or indicative of multiple velocity components detected in several C^{18}O spectra.

In this paper, we assume the same nomenclature as Leurini et al. (2019) to identify the structures that compose the complex network of filaments of G351.77–0.51: below, we will refer to the central filament as ‘main body’ or ‘main ridge’. It is well visible in Fig. 1 as the elongated structure that harbours the five clumps,

identified by the white labels. The gas that constitutes the main body is cold and chemically young, in which Giannetti et al. (2019) find high abundances of $\text{o-H}_2\text{D}^+$ and N_2D^+ and an age $\lesssim 10^5$ yr for the clump-7 (see that paper for more details). The northern part of the main body appears in absorption against the mid-IR background of our Galaxy up to $70 \mu\text{m}$ and in emission in the (sub)millimetre range (Faúndez et al. 2004). The analysis of the $J=2 \rightarrow 1$ C^{18}O molecular line (see Section 3.2) confirmed the presence of less dense molecular sub-filaments, linked to the main body: we refer to them as ‘branches’, following the naming convention of Schisano et al. (2014).

2.1 C^{18}O map

The $\text{C}^{18}\text{O } J=2 \rightarrow 1$ observations were carried out with the *Atacama Pathfinder Experiment* 12m submillimetre telescope (APEX) between 2014 August and November. The observations were centred at 218.5 GHz, with a velocity resolution of 0.1 km s^{-1} . The whole map covers an approximate total area of 234 arcmin^2 . We refer to Leurini et al. (2019), for a more detailed description of the data set.

The root mean square (rms) noise was calculated iteratively in each pixel because it is not uniform in the map, due to the different exposure times. The first estimate of rms noise was obtained from the unmasked spectra, then, any emission higher than 3σ was masked and the rms was recomputed. The cycle continues until the difference between the rms noise of two consecutive iterations (i.e. $\sigma_{i+1} - \sigma_i$) is equal to 10^{-4} K . We estimate a typical final noise of 0.30 K (between 0.15 and 0.45 K) per velocity channel.

3 RESULTS

3.1 Temperature and H_2 column density maps

Calculating C^{18}O column densities requires the gas excitation temperature, $T_{\text{ex}}^{\text{C}^{18}\text{O}}$. We derived $T_{\text{ex}}^{\text{C}^{18}\text{O}}$ from the dust temperature, T_{dust} , map, which can be obtained from the *Herschel* data through pixel-by-pixel SED fits. This allows one to determine also the H_2 column density map from the emission between 160 and $500 \mu\text{m}$.

In this work, we adopt the dust temperature map presented in Leurini et al. (2019). To estimate the dust temperature of the whole filamentary structure of G351.77–0.51, these authors used the *Herschel* (Pilbratt et al. 2010) Infrared Galactic Plane Survey (Molinari et al. 2010) images at 500 , 350 , and $250 \mu\text{m}$ from the *Spectral and Photometric Imaging Receiver* (SPIRE; Griffin et al. 2010) and at $160 \mu\text{m}$ from the *Photodetector Array Camera & Spectrometer* (PACS; Poglitsch et al. 2010). The authors adopted a model of two emission components (Schisano et al., in preparation), splitting the fluxes observed by *Herschel* in each pixel into the filament and background contributions, as discussed by Peretto et al. (2010). Then, they fitted the filament emission pixel-by-pixel with a grey body model deriving the dust temperature and the H_2 column density. They assumed a dust opacity law $\kappa_0(\nu/\nu_0)^\beta$ with $\beta = 2$, $\kappa_0 = 0.1 \text{ cm}^2 \text{ g}^{-1}$ at $\nu_0 = 1250 \text{ GHz}$ (Hildebrand 1983). This prescription assumes a gas-to-dust ratio equal to 100. A detailed description of the procedure is given in Leurini et al. (2019).

Their maps, shown in Fig. 2, have a resolution of 36 arcsec , i.e. the coarsest resolution in the *Herschel* bands, that is comparable to the resolution of our C^{18}O data (29 arcsec). In most of the

¹This class of object shows a flux larger than 2.6 Jy at $21\text{--}24 \mu\text{m}$ and no compact radio emission at $4\text{--}8 \text{ GHz}$ within 10 arcsec of the ATLASGAL peak.

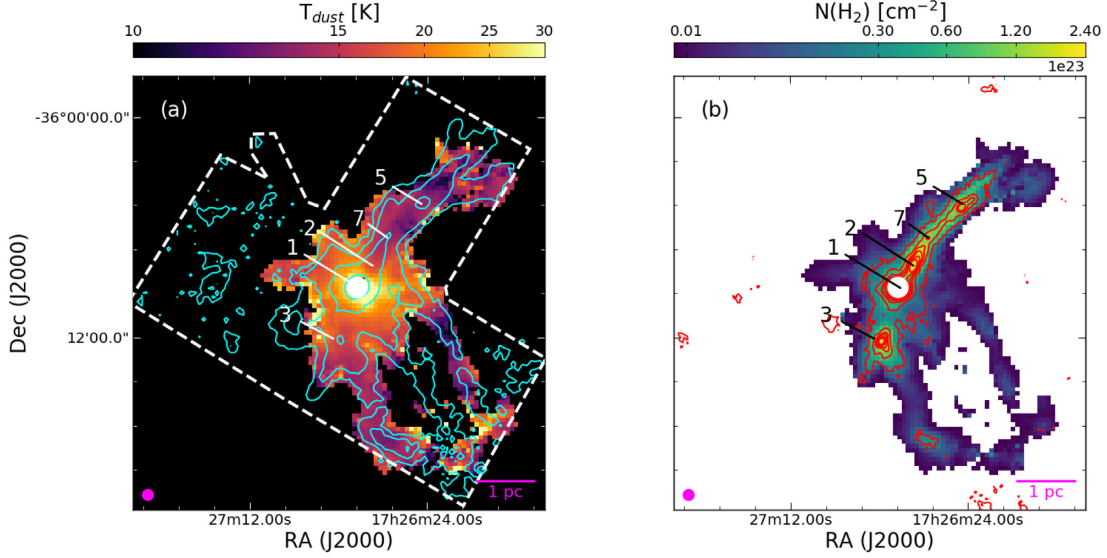


Figure 2. (a) Dust temperature map adapted from Leurini et al. (2019), generated by a pixel-by-pixel SED fitting of the 160–500 μm continuum fluxes of the Hi-Gal Survey. The white dashed lines mark the region observed in C^{18}O with APEX (Section 2.1). The cyan contours are defined on the integrated intensity map of the APEX C^{18}O $J=2 \rightarrow 1$ line in Fig. 3(a) at 3σ , 9σ , 27σ , and 81σ contours, where $3\sigma = 0.9 \text{ K km s}^{-1}$; (b) H_2 column density map adapted from the pixel-by-pixel SED fitting from Leurini et al. (2019), scaled to a gas mass-to-dust ratio equal to 120 (see the text). The red contours are the same of the LABOCA map of the 870 μm dust continuum emission shown in Fig. 1. In both the panels, we masked the saturated hot-core region in the *Herschel* data (white circle).

map, the dust temperature [panel (a) of Fig. 2] ranges between 10 and 30 K, while along the main body and in the region to the south of clump-3 in Fig. 1 typically $T_{\text{dust}} \lesssim 15 \text{ K}$ are found.

$T_{\text{ex}}^{\text{C}^{18}\text{O}}$ was derived from T_{dust} by applying the empirical relation defined by Giannetti et al. (2017a), who suggest a relation $T_{\text{ex}}^{\text{C}^{18}\text{O}} = 1.46 \times T_{\text{dust}} - 12.9$ with an intrinsic scatter of 6.7 K. However, this relation is only valid for dust temperatures between ~ 10 and $\sim 45 \text{ K}$, while at low T_{dust} it may underestimate $T_{\text{ex}}^{\text{C}^{18}\text{O}}$. Our data are concentrated in this low-temperature regime, where $T_{\text{dust}} \lesssim 9 \text{ K}$ would be translated into negative $T_{\text{ex}}^{\text{C}^{18}\text{O}}$ following the relation mentioned earlier. Therefore, we decided to use the flattest curve allowed by the uncertainties at 2σ of the original fit to limit this issue: $T_{\text{ex}}^{\text{C}^{18}\text{O}} = T_{\text{dust}} - 2.6$ and imposing a lower limit of $T_{\text{ex}} = 10.8 \text{ K}$, equal to the separation between the levels of the $J=2 \rightarrow 1$ transition of the C^{18}O line.

$T_{\text{ex}}^{\text{C}^{18}\text{O}}$ typically ranges between ~ 11 and 28 K with peaks of up to 40 K . Most of the pixels along the main ridge show a temperature of $12\text{--}13 \text{ K}$. At these temperatures, CO is efficiently removed from the gas phase and frozen on to dust grains for densities exceeding few $\times 10^4 \text{ cm}^{-3}$ (see Giannetti et al. 2014). This suggests that the depletion factor in these regions will have the highest values of the entire IRDC (as confirmed by the analysis of the depletion map – Section 4.1).

We modified the H_2 column density map of Leurini et al. (2019) by using a different value of gas mass-to-dust ratio, γ , equal to 120 as derived by Giannetti et al. (2017b), by assuming a galactocentric distance $D_{\text{GL}} = 7.4 \text{ kpc}$ (Leurini et al. 2019). The resulting H_2 column density map is shown in Fig. 2(b). The saturated region in the *Herschel* data, corresponding to the hot-core clump, was masked out and excluded from our analysis. In the coldest regions of Fig. 2(a), the molecular hydrogen reaches a column density of $2.4 \times 10^{23} \text{ cm}^{-2}$, decreasing to $4.2 \times 10^{21} \text{ cm}^{-2}$ in the regions in which the dust is warmer. Fig. 2(b) shows that the high-column density material (i.e. $\gtrsim 1 \times 10^{23} \text{ cm}^{-2}$) is distributed in a

single filamentary feature and in clump-like structures, the so-called ‘main-body’ in Leurini et al. (2019).

3.2 Column densities of C^{18}O

Under the assumption of local thermal equilibrium (LTE), we derived the C^{18}O column density, $N(\text{C}^{18}\text{O})$, from the integrated line intensity of the $(2-1)$ transition. Following Kramer & Winnewisser (1991), the general formula for $N(\text{C}^{18}\text{O})$ is

$$N(\text{C}^{18}\text{O}) = \frac{C_\tau}{\eta_c} \frac{3h}{8\pi^3 \mu^2} \frac{Z}{2} e^{\frac{E_{\text{low}}}{k_B T_{\text{ex}}}} \left[1 - e^{-\frac{h\nu_{J,J-1}}{k_B T_{\text{ex}}}} \right]^{-1} \times [J(T_{\text{ex}}, \nu_{\text{C}^{18}\text{O}}) - J(T_{\text{bg}}, \nu_{\text{C}^{18}\text{O}})]^{-1} \int T_{\text{MB}} d\nu$$

$$= C_\tau f(T_{\text{ex}}) \int T_{\text{MB}} d\nu, \quad (4)$$

where C_τ is the optical depth correction factor defined as $\tau/[1 - \exp(-\tau)]$ and the τ is the optical depth of the $J=2 \rightarrow 1$ transition of C^{18}O (see Section 3.2.1); h and k_B are the Planck and Boltzmann constants, respectively; η_c is the filling factor; μ is the dipole moment of the molecule; Z is the partition function; E_{low} is the energy of the lower level of the transition and $\nu_{J,J-1}$ is the frequency of the $J \rightarrow J-1$ transition of the considered molecule (in this case, the C^{18}O $J=2 \rightarrow 1$, equal to $\sim 219.5 \text{ GHz}$); $J(T_{\text{ex}}, \nu) = (h\nu/k_B)(\exp(h\nu/k_B T) - 1)^{-1}$; $T_{\text{bg}} \cong 2.7 \text{ K}$ is the background temperature; T_{MB} is the main beam temperature, and its integral over the velocity range covered by the C^{18}O line is shown in Fig. 3(a). In the last line of equation (4), $f(T_{\text{ex}})$ incorporates all the constants and the terms depending on T_{ex} . We further considered possible saturation effects of the continuum *Herschel* maps, in the hot-core region (i.e. clump-1 in Fig. 1).

In the following paragraphs, we discuss the steps that allowed us to derive the C_τ map (Fig. 3b) and the final C^{18}O column density map (Fig. 3c), necessary to produce the depletion map discussed in Section 4.1.

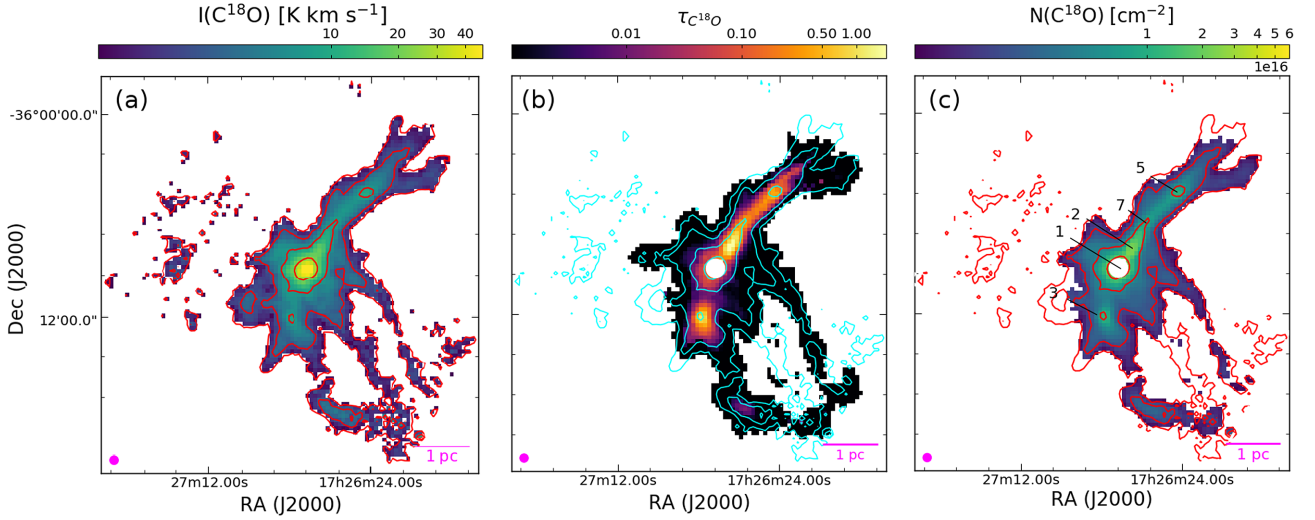


Figure 3. (a) Integrated intensity distribution of the $\text{C}^{18}\text{O } J=2 \rightarrow 1$ molecular line transition detected with APEX in units of K km s^{-1} . (b) Map of the opacity from the $N(\text{H}_2)$ map, following the procedure described in Section 3.2.1. The final τ - $N(\text{H}_2)$ relation used is: $\log_{10}(\tau_{\text{C}^{18}\text{O}}) = 2.5 \log_{10}(N(\text{H}_2)) - 57.4$. (c) C^{18}O column density map obtained from the integrated intensity distribution, shown in Fig. 3(a), using equation (4). The red and cyan contours are the 3σ , 9σ , 27σ , and 81σ contours of the integrated flux density distribution in panel (a), where $3\sigma = 0.9 \text{ K km s}^{-1}$.

3.2.1 Opacity correction

We estimate the optical depth of $\text{C}^{18}\text{O } J=2 \rightarrow 1$ in the clumps by means of the detection equation (see Hofner et al. 2000, for more details). We use the C^{18}O and $\text{C}^{17}\text{O } J=2 \rightarrow 1$ APEX observations presented in Leurini et al. (2011b), assuming the relative abundance, ϕ , equal to 4 as found by Wouterloot et al. (2008).² Both transitions were observed in seven single-pointing observations, centred at the coordinates of the clumps. We did not consider the data of clumps 9 and 10 (defined in Leurini et al. 2011b), because they are not part of the source, nor the data of the hot core (i.e. clump-1), due to the saturation of the continuum observations (especially at 250 micron).

The ratio between the peaks of the two CO isotopologue line intensities in each clump is then equal to

$$R_{\text{C}^{18}\text{O}, \text{C}^{17}\text{O}} = \frac{T_{\text{MB}, \text{C}^{18}\text{O}}}{T_{\text{MB}, \text{C}^{17}\text{O}}} = \frac{\eta_{\text{C}^{18}\text{O}} [J(T_{\text{ex}, \text{C}^{18}\text{O}}, \nu_{\text{C}^{18}\text{O}}) - J(T_{\text{bg}}, \nu_{\text{C}^{18}\text{O}})] (1 - e^{-\tau_{\text{C}^{18}\text{O}}})}{\eta_{\text{C}^{17}\text{O}} [J(T_{\text{ex}, \text{C}^{17}\text{O}}, \nu_{\text{C}^{17}\text{O}}) - J(T_{\text{bg}}, \nu_{\text{C}^{17}\text{O}})] (1 - e^{-\tau_{\text{C}^{17}\text{O}}})}, \quad (5)$$

where the considered transition of C^{18}O and C^{17}O is the $J=2 \rightarrow 1$, at frequencies $\nu_{\text{C}^{18}\text{O}} = 219.5 \text{ GHz}$ and $\nu_{\text{C}^{17}\text{O}} = 224.7 \text{ GHz}$, respectively; τ is the optical depth of the two considered CO isotopologues, with $\tau_{\text{C}^{18}\text{O}} = \phi \tau_{\text{C}^{17}\text{O}}$.

In equation (5), we assumed the $J=2 \rightarrow 1$ transition of the two CO isotopologues, under the same excitation conditions (i.e. $T_{\text{ex}, \text{C}^{17}\text{O}} = T_{\text{ex}, \text{C}^{18}\text{O}}$; e.g. Martin & Barrett 1978; Ladd 2004), are tracing the same volume of gas. This assumption has the consequence that the filling factor, η , is the same for both the transitions and therefore it can be eliminated from equation (5).

We fitted a linear relation between $\log_{10}(\tau)$ and $\log_{10}[N(\text{H}_2)]$,³ and then computed C_τ to correct $N(\text{C}^{18}\text{O})$ following the schematic

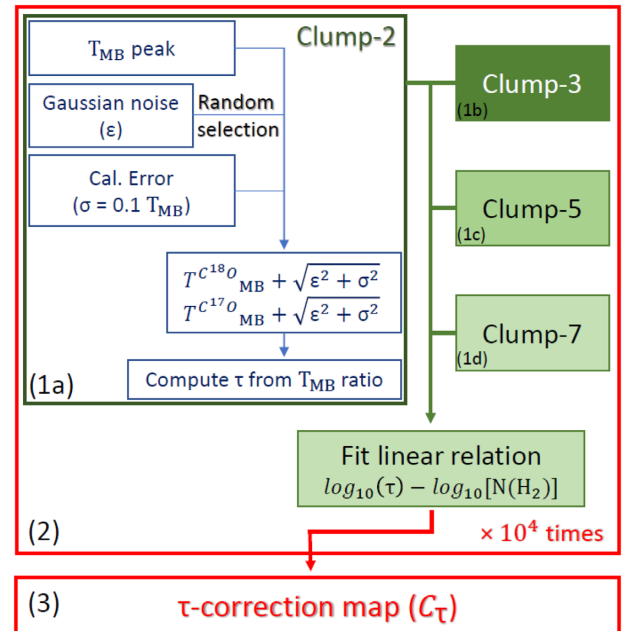


Figure 4. Schematic diagram showing how the correction C_τ was calculated starting from the C^{18}O and C^{17}O single-pointing observation of clumps 2, 3, 5 and 7 in Leurini et al. (2011b).

diagram in Fig. 4. The best linear fit was computed with a least-squares regression by extracting the C^{17}O and C^{18}O peak fluxes 10^4 times (step 2 in Fig. 4). The basic assumption – see step 1a – was that the rms (ϵ) is normally distributed.⁴ For each clump, the τ value was calculated following equation (5), including the contribution of the rms, and assuming a 10 percent calibration uncertainty (σ), summed in quadrature to ϵ (i.e. $T_{\text{MB}, \text{C}^{18}\text{O}} = T_{\text{MB}}^{\text{C}^{18}\text{O}} + \sqrt{\epsilon^2 + \sigma^2}$ and the same for $T_{\text{MB}, \text{C}^{17}\text{O}}$). We note that when applying this method

² Assuming a galactocentric distance of the source equal to $D_{\text{GL}} = 7.4 \text{ kpc}$ as in Leurini et al. (2019).

³ We used a τ - $N(\text{H}_2)$ relation and not a τ - $N(\text{C}^{18}\text{O})$ one because $N(\text{H}_2)$ is not affected by opacity (i.e. κ_ν correction already applied in Section 3.1) and the first relation has less scatter than the second.

⁴ This step was performed using the `numpy.random.normal` function of NUMPY (Oliphant 2006) v1.14.

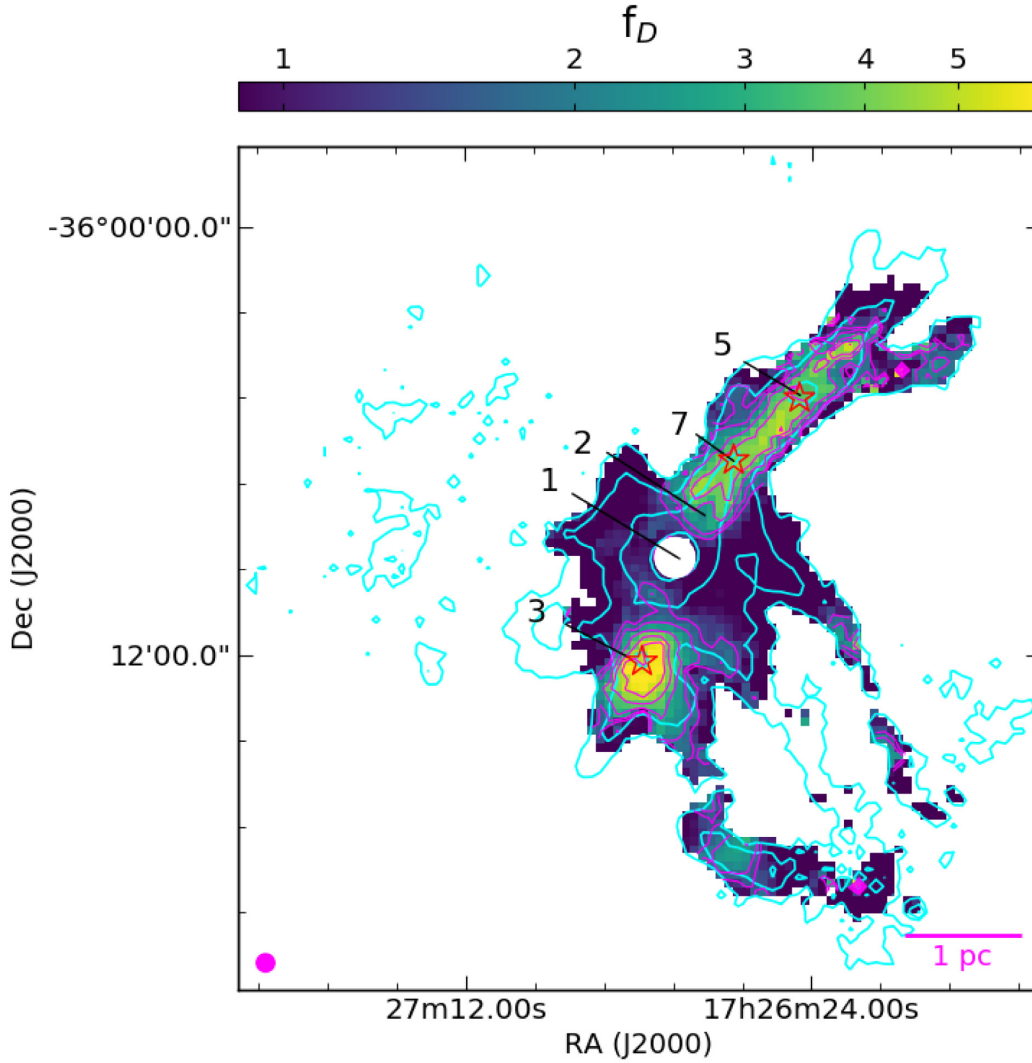


Figure 5. Depletion factor (f_D) map obtained by taking the ratio between the expected and observed $N(\text{C}^{18}\text{O})$. We assumed a canonical abundance of 2.1×10^{-7} of $N(\text{C}^{18}\text{O})$ relative to $N(\text{H}_2)$. The cyan contours are the same as in Fig. 3, while the ones in magenta are defined at $f_D = 1.5, 2, 3, 4$ and 5 . The red stars in clump-5 and 7 indicate locations of the clumps reported in Leurini et al. (2019), while the one in clump-3 is set on the coordinates of the candidate H II region reported in Anderson et al. (2014).

to select different values of T_{MB} in the C^{18}O and C^{17}O spectra, the errors do not change significantly if the distributions are built by 10^3 elements or more. Applying this procedure to each clump – steps 1 (a–d) – we derived the linear fit $\log_{10}(\tau) - \log_{10}[N(\text{H}_2)]$. We repeated this procedure 10^4 times and we generated a cube of CO column density maps, one for each solution $\log_{10}(\tau) - \log_{10}[N(\text{H}_2)]$ found. At the end of this procedure, a distribution of 10^4 values of $N(\text{C}^{18}\text{O})$ has been associated with each pixel, used to compute the average value of $N(\text{C}^{18}\text{O})$ and its relative error bar in each pixel. We finally imposed an upper limit for the correction at $\tau = 2$. This value would only be achieved in the saturated hot-core region and thus no pixel included in our analysis is affected by this condition. We then generated the opacity map from the $N(\text{H}_2)$ map. The results of this procedure are shown in Fig. 3(b), together with the C^{18}O integrated intensity map – Fig. 3(a). We note that the estimated τ values range between ~ 0.1 and ~ 1.9 along the main body of G351.77–0.51.

The final C^{18}O column density map was derived by including the opacity correction using the best-fitting final $\log_{10}(\tau) - \log_{10}[N(\text{H}_2)]$ relation – i.e. $\log_{10}(\tau_{\text{C}^{18}\text{O}}) = 2.5\log_{10}[N(\text{H}_2)] - 57.4$. The opacity-

corrected column density map of C^{18}O is shown in Fig. 3(c). Our correction has increased the C^{18}O column densities by up to a factor of 2.3, and remains almost constant in the branches. Over the whole structure, the column densities of C^{18}O range between 1 and $6 \times 10^{16} \text{ cm}^{-2}$.

To summarize, the C^{18}O column density map was generated under the assumption of LTE following equation (4). Possible saturation in the continuum data and opacity effects have been considered. The column density map was then used to evaluate the depletion map, as discussed in the next section.

4 DISCUSSION

4.1 The large-scale depletion map in G351.77–0.51

The final CO-depletion factor map – Fig. 5 – was generated by taking the ratio between the expected and the observed CO emission, using an abundance of C^{18}O with respect to H_2 (i.e. $\chi_{\text{C}^{18}\text{O}}^{\text{E}}$) equal to 2.1×10^{-7} (see Giannetti et al. 2017b).

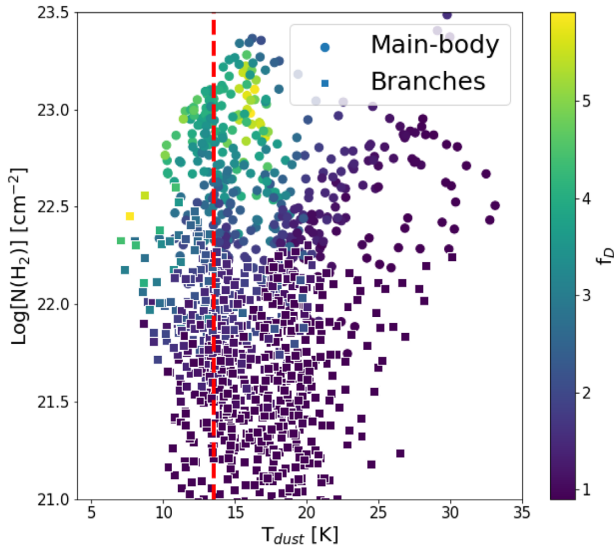


Figure 6. Pixel-by-pixel scatter plot of the whole structure detected in Fig. 5. The circles represent the pixels of the main body, while the squares indicate those of the branches. The red-dashed line represents the 10.8 K lower limit imposed on the $T_{\text{ex}}^{\text{C}^{18}\text{O}}$ (as discussed in Section 3.1) that corresponds to $T_{\text{dust}} \sim 13.4$ K.

As visible in Fig. 5, in almost half of the map the depletion factor is < 1.5 , highlighting the absence of strong CO-depletion around the core and along the branches in the southern directions with respect to the main ridge. This effect can have two different causes for the region of the central clump and for the branches, respectively: In the first case (i.e. the surroundings of clump-1 in Fig. 1), the absence of depletion can be linked to the intense star formation activity, demonstrated in previous papers (e.g. Leurini et al. 2011b; König et al. 2017; Giannetti et al. 2017a and Leurini et al. 2019). The increase in temperature induced by the forming stars is able to completely desorb the ice mantles around dust grains in which the CO molecules are locked. This effect lowers the observed depletion factor, until it reaches unity. Following equation (2) in Schuller et al. (2009), we computed $N(\text{H}_2)$ in the hot-core (saturated) region from the ATLASGAL peak flux at $870 \mu\text{m}$, using $\kappa_{870\mu\text{m}} = 0.76 \text{ cm}^2 \text{ g}^{-1}$ consistent with the dust opacity law discussed in Section 3.1. To obtain a depletion of 1 here, consistent with the surroundings, the dust temperature should be 80 K.

To evaluate the effects of $N(\text{H}_2)$ and T_{dust} on f_D , Fig. 6 was obtained by the pixel-by-pixel combination of Figs 2(b), (a) and 5. Each point represents a pixel shared between the three maps, where dots and squares are used to distinguish the pixels of the main body from those of the branches, respectively.

Instead, in the branches (square markers in Fig. 6) we notice that depletion reaches values ~ 3.5 . This result suggests that even in these structures the depletion process can start to occur. On the other hand, where f_D is close to 1, the lower density disfavours a high degree of depletion (e.g. see Caselli et al. 1999). Furthermore, we should consider that the observed depletion factor is averaged along the line of sight and in the beam.

Along the main ridge and in the surroundings of clump-3 (Fig. 1), the depletion factor ranges between 1.5 and 6 and reaches its maximum in clump-3. Both regions appear in absorption at $8 \mu\text{m}$ in the Leurini et al. (2011b) maps, showing H_2 column densities of a few 10^{22} cm^{-2} . In particular, it may appear counter-intuitive that we observe the highest depletion factor of the whole structure in

clump-3, as an H II region has been identified in Wide-field Infrared Survey Explorer (WISE) at a distance of only ≈ 8 arcsec from the centre of clump-3 (Anderson et al. 2014). For these reasons, clump-3 should show similar depletion conditions to those in clump-1. However, such a high depletion factor suggests dense and cold gas close to the H II region contained in the clump. Within the cloud, the degree of depletion could be maintained if self-shielding would attenuate the effects of the radiation field of the H II region. These ideas are supported by the analysis of the 8 and $24 \mu\text{m}$ maps shown in Leurini et al. (2019). At the location of the clump-3, a region slightly offset with respect to the bright spot associated with the H II region is clearly visible in absorption at both wavelengths.

In clumps 5 and 7 (i.e. regions C5 and C7 in Fig. 1), along the main ridge of G351.77–0.51, the average depletion factors are $f_{D,C7} = 3.4^{+0.4}_{-0.5}$ (in clump-7; peaking at 4.5) and $f_{D,C5} = 3.1^{+0.5}_{-0.6}$ (in clump-5; peaking at 4.3), respectively (see Section 4.2.1 for more details about error estimation). Compared to the average values of the samples mentioned in Section 1, our values are slightly lower. However, we should consider that the observed depletion factors are affected by many factors such as the opacity correction applied and the $\text{C}^{18}\text{O}/\text{H}_2$ abundance ratio, which can vary up to a factor of 2.5 (e.g. Ripple et al. 2013).

Along the main ridge, the depletion factors reach values of ~ 6 . This is comparable to what is found in Hernandez et al. (2011), where the authors studied depletion factor variations along the filamentary structure of IRDC G035.30-00.33 with IRAM 30m telescope observations. Along the filament, they estimate a depletion factor of up to 5. Of course, also in this case the considerations made earlier hold, but if we consider the different opacity corrections applied and $\text{C}^{18}\text{O}/\text{H}_2$ relative abundance assumed, we note that the difference between our results and those of Hernandez et al. (2011) is not larger than $\Delta f_D \sim 1$.

To summarize, the final depletion map of G351.77–0.51 – Fig. 5 – shows widespread CO-depletion in the main body, as well as at various locations in the branches. Comparing our results with those of Hernandez et al. (2011) in G035.30-00.33, we note that in both cases the phenomenon of depletion affects not only the densest regions of the clumps, but also the filamentary structures that surround them. This result suggests that CO-depletion in high-mass star-forming regions affects both small and large scales.

4.2 Depletion modelling

High densities and low gas temperatures favour CO-depletion. Given this, it is reasonable to think that $\text{C}^{18}\text{O}/\text{H}_2$ is not constant within a cloud. This quantity varies as a function of location, following the volume density and gas temperature distributions/profiles.

The regions in which the depletion degree is higher are those where the dust surface chemistry becomes more efficient, due to the high concentration of frozen chemical species on the dust surface. To understand how the efficiency of the various types of chemical reactions changes, we need to understand how the depletion degree varies within the dark molecular clouds.

In order to reproduce the average depletion factor observed in G351.77–0.51, we built a simple 1D model describing the distributions of C^{18}O and H_2 . We focus our attention on the main ridge, identifying three distinct regions: clumps 5 and 7, and the filamentary region between them (i.e. regions C5, C7, and F1 in Fig. 1, respectively).

The model assumes that both profiles have radial symmetry with respect to the centre of the ridge, i.e. the ‘spine’. R_{flat} is the distance relative to the spine within which the density profiles remain roughly

flat [i.e. $n(\text{H}_2, R_{\text{flat}}) = 0.5n(\text{H}_{2,\text{spine}})$ if $p = 2$]. We normalize the volume density profiles of the model with respect to R_{flat} , while $n(\text{H}_{2,\text{spine}})$ is the central volume density of H_2 .

Our H_2 volume density profile is described by

$$n(\text{H}_2) = n(\text{H}_{2,\text{spine}}) \left[1 + \left(\frac{R}{R_{\text{flat}}} \right)^\alpha \right]^{-p/2}, \quad (6)$$

up to a maximum distance $R_{\text{max}} \sim 0.2$ pc, two times larger than the filament width estimated in Leurini et al. (2019), that encompasses the entire filamentary structure.

For the clumps (i.e. regions C5 and C7), we assume $p = 2$ and equation (6) takes the functional form described by Tafalla et al. (2002). In this case, the free parameters of the model are α , $n(\text{H}_{2,\text{spine}})$, and R_{flat} . Starting from R_{flat} , both the distributions of C^{18}O and H_2 scale as a power law with the same index: $\alpha_{\text{C5}} = 1.9$ and $\alpha_{\text{C7}} = 1.8$, for the clump-5 and clump-7, respectively. The solutions for α were found by exploring the parameter space defined by the results of Beuther et al. (2002), i.e. $1.1 < \alpha < 2.1$. The best-fitting models return the values of $n(\text{H}_{2,\text{spine}})$ and R_{flat} equal to $8 \times 10^6 \text{ cm}^{-3}$ and 5.5×10^{-3} pc in the case of clump-5 and $5.8 \times 10^7 \text{ cm}^{-3}$ and 2×10^{-3} pc for the clump-7.

In the case of the main body (i.e. region F1), $\alpha = 2$, the volume density profile of H_2 is described by a Plummer-like profile (see Plummer 1911, for a detailed discussion). For this model, the free parameters are p and $n(\text{H}_{2,\text{spine}})$, while R_{flat} corresponds to the thermal Jeans length, λ_J , for an isothermal filament in hydrostatic equilibrium, calculated at $R=0$:

$$\lambda_J = \frac{c_s^2}{G \mu m_{\text{H}} N(\text{H}_2)_{R=0}}, \quad (7)$$

where c_s^2 is the isothermal sound speed for the mean temperature in the region F1 (equal to 13.8 K), G is the gravitational constant in units of $[\text{cm}^3 \text{ g}^{-1} \text{ s}^{-2}]$, μ is the mean molecular weight of interstellar gas ($\mu = 2.3$), m_{H} is the hydrogen mass in grams, and $N(\text{H}_2)_{R=0}$ is the averaged H_2 column density calculated at $R=0$. Applying equation (7) to region F1, we obtained $R_{\text{flat}} \sim 0.008$ pc, while the value of p index was found equal to 1.9 after exploring the range of values reported in Arzoumanian et al. (2011). We note that an R_{flat} variation of 30 per cent does not significantly change the $N(\text{H}_2)$ profile.

To simulate depletion effects, we introduced the depletion radius, R_{dep} , that is the distance from the spine where the degree of depletion drastically changes, following equations:

$$\left. \begin{aligned} R \leq R_{\text{dep}} \quad f_D &= (10, \infty); \quad \chi_{\text{C}^{18}\text{O}}^{\text{E}} = \frac{\chi_{\text{C}^{18}\text{O}}^{\text{E}}}{f_D} \\ R > R_{\text{dep}} \quad f_D &= 1; \quad \chi_{\text{C}^{18}\text{O}}^{\text{E}} = \chi_{\text{C}^{18}\text{O}}^{\text{E}} \end{aligned} \right\}, \quad (8)$$

where the canonical abundance for C^{18}O ($\chi_{\text{C}^{18}\text{O}}^{\text{E}} = 2.1 \times 10^{-7}$) was set by using the findings of Giannetti et al. (2017b) and references therein. The lower limit of $f_D = 10$ for $R < R_{\text{dep}}$ is motivated by the observational constraints for high-mass clumps, which in most cases show depletion values between 1 and 10 (e.g. Thomas & Fuller 2008; Fontani et al. 2012; Giannetti et al. 2014, as discussed in Section 1). Instead, the choice of extending the model until reaching the theoretical condition of full depletion (i.e. $f_D = \infty$) is done to study the effect that such a drastic variation of f_D has on the size of R_{dep} .

In Fig. 7, we present a sketch of the model: The profile for $n(\text{H}_2)$ is indicated in blue, that for $\chi_{\text{C}^{18}\text{O}}$ is shown in green (bottom panel), while the profiles for $n(\text{C}^{18}\text{O})$ appear in orange and red (two values of f_D are considered). Radial profiles are then convolved

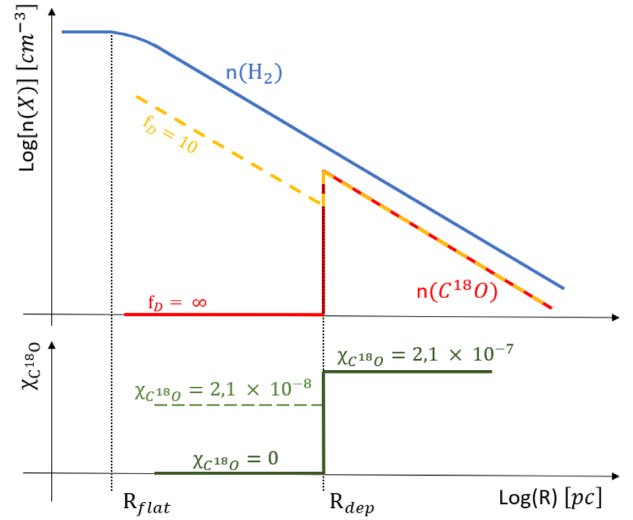


Figure 7. Schematic representation of the model applied to simulate the depletion effect inside/outside the depletion radius. In blue, we show the number density profile of H_2 ; in green, the step-function reproducing the variation of the degree of depletion inside/outside R_{dep} . The C^{18}O profiles are obtained by the product of the other two.

with the *Herschel* beam at $\lambda = 500 \mu\text{m}$ [i.e. full width at half-maximum (FWHM) = 36 arcsec]. For each profile, we calculate the set of values that best represent the observed data in each considered region (i.e. C5, C7, and F1 of Fig. 1): We generate a set of profiles of $n(\text{H}_2)$ by exploring the parameter space defined by the free parameters of each modelled region, and we select the curve that maximizes the total probability of the fit [i.e. $P_{\text{TOT}} = \prod p(r_i)$, where $p(r_i)$ are the probability density values corresponding to the error bars in Fig. 8, and calculated at the various distances r_i from the spine]. The same was done for the profile of $n(\text{C}^{18}\text{O})$, selecting the best profile from a set of synthetic profiles generated by assuming different sizes of R_{dep} .

In the next section, we present the results of the model applied to the clumps and the filament region, providing an estimate for both R_{dep} and $n(\text{H}_2)$ calculated at R_{dep} .

4.2.1 Estimate of R_{dep}

In Figs 8(a) and (b), we show the best-fitting mean-radial column density profiles of $N(\text{H}_2)$ and $N(\text{C}^{18}\text{O})$ of clump-5, respectively. In blue, we show the column density profiles before the convolution while in orange the convolved profiles. The black-dashed line in the central panels is the convolved undepleted profile of $N(\text{C}^{18}\text{O})$, assuming a constant $\chi_{\text{C}^{18}\text{O}} = 2.1 \times 10^{-7}$ for all R . The observed data are plotted as red points. In Fig. 8(c), we report (in green) the depletion factor as a function of the distance from the projected centre in clump-5 assuming an abundance of C^{18}O for $R < R_{\text{dep}}$ equal to 2.1×10^{-8} , while in magenta the same profile for $\chi_{\text{C}^{18}\text{O}} = 0$. Figs 8(d), (e), and (f) are the same for the clump-7, respectively. In both cases, we estimate R_{dep} by varying the abundance $\chi_{\text{C}^{18}\text{O}}$ for $R < R_{\text{dep}}$ considering the two limiting cases, i.e. with $f_D = 10$ and ∞ .

All error bars have been defined taking the range between the 5th and 95th percentiles. In the case of $N(\text{H}_2)$, the uncertainties are generated by computing the pixel-to-pixel standard deviation of the H_2 column density along the direction of the spine, and

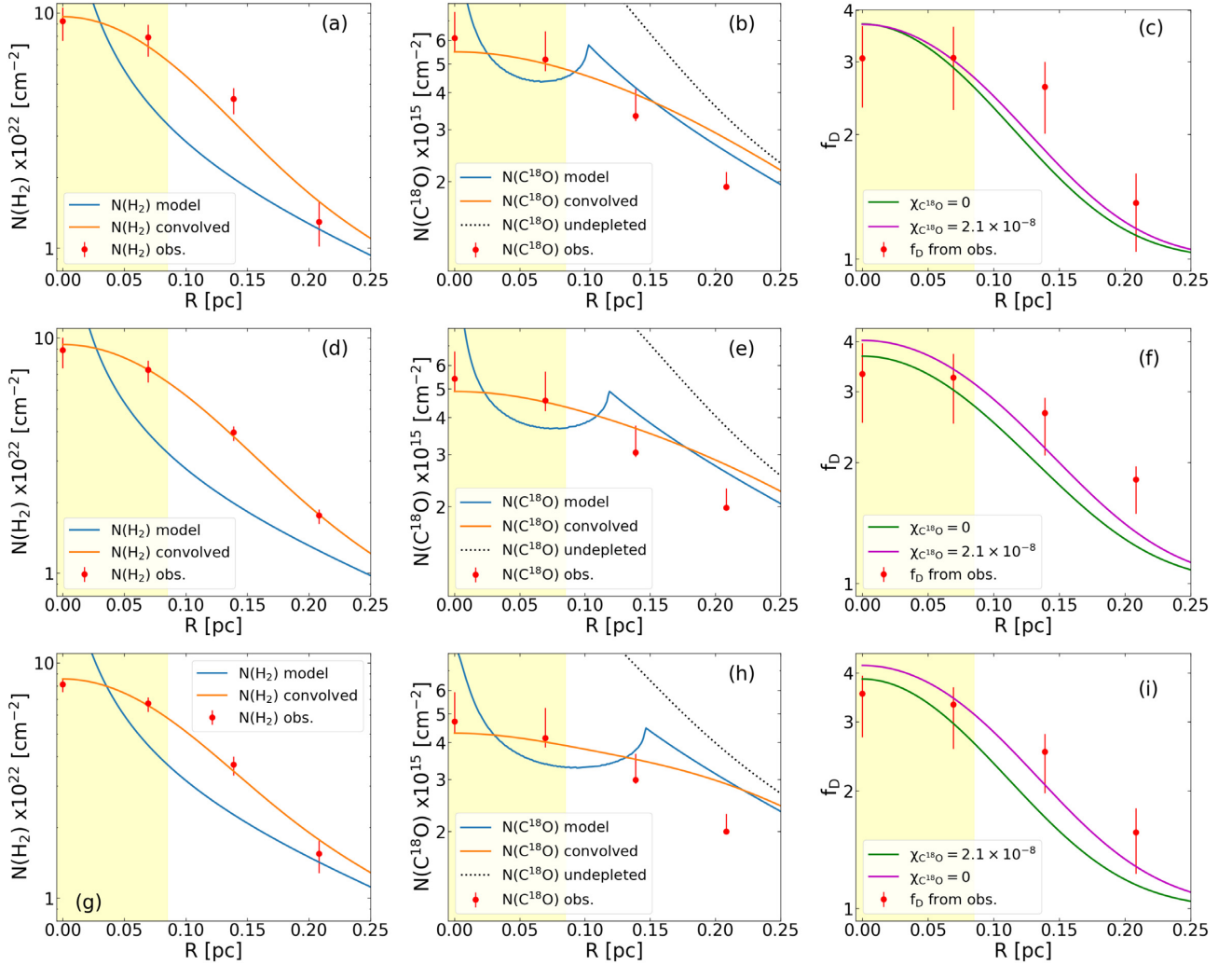


Figure 8. From top: Column density profiles of H_2 and C^{18}O in panels (a) and (b), respectively, while panel (c) shows the los- and beam-averaged depletion factor profile from region C5 in Fig. 1. Panels (d), (e), and (f) are the same as (a), (b), and (c), respectively, for region C7, while panels (g), (h), and (i) refer to the region F1. The blue profiles in panels (a), (d), and (g) are the synthetic column density profiles obtained by the integration of the number density distribution; the solid-blue lines in panels (b), (e), and (h) are the same for C^{18}O assuming $\chi_{\text{C}^{18}\text{O}} = 2.1 \times 10^{-8}$ for $R < R_{\text{dep}}$, and the estimated position of R_{dep} corresponds to the cusp. The orange profiles in panels (a), (b), (d), (e), (g), and (h) are the convolution results of the blue profiles and the *Herschel* beam at $500 \mu\text{m}$, while the black-dotted line in panels (b), (e), and (h) indicates the convolved undepleted profile of $N(\text{C}^{18}\text{O})$ assuming a constant $\chi_{\text{C}^{18}\text{O}} = 2.1 \times 10^{-7}$. The green and magenta profiles in panels (c), (f), and (i) show the depletion factor profiles obtained by assuming different abundances for $R < R_{\text{dep}}$ (see equation 8). The red points are the observed values for each quantity plotted. The yellow-shaded area shows the dimensions of the *Herschel* half width at half-maximum (HWHM) at $\lambda = 500 \mu\text{m}$ (i.e. FWHM=36 arcsec; 0.17 pc at the source distance). All profiles are plotted as a function of the radial distance from the spine.

assuming a Gaussian distribution of this quantity around the mean value of $N(\text{H}_2)$. The errors associated with $N(\text{C}^{18}\text{O})$ are computed propagating the uncertainty on the line integrated intensity and on the optical depth correction (as discussed in Section 3.2.1), using a Monte Carlo approach. Finally, the uncertainties on f_D are estimated computing the abundance, randomly sampling the probability distributions of $N(\text{H}_2)$ and $N(\text{C}^{18}\text{O})$.

In clump-5, the estimated $R_{\text{dep}, i}$, ranges between a minimum of ~ 0.07 pc and a maximum of ~ 0.10 pc. In clump-7 instead, which shows a larger average depletion factor compared to clump-5, R_{dep} ranges between ~ 0.07 and 0.12 pc (for $f_D = \infty$ and 10, respectively). Comparing Figs 8(c) and (f), both synthetic profiles are able to well reproduce the observed ones up to a distance of ~ 0.15 pc. In the case of total depletion regime (i.e. $f_D = \infty$), at R_{dep} we estimate a volume density $n(\text{H}_2) = 5.7 \times 10^4$ and 5.3×10^4

cm^{-3} , for clump-5 and 7, respectively. Our results are comparable with those found in low-mass prestellar cores (i.e. Caselli et al. 2002; Ford & Shirley 2011) and in high-mass clumps in early stages of evolution (i.e. Giannetti et al. 2014). Furthermore, it is important to note that both values should be considered as upper limits due to the approximation in equation (8) for all the models discussed. For the case in which $f_D = 10$, in clump-5 at a distance equal to $R_{\text{dep}} = 0.10$ pc, the model provides a volume density of $n(\text{H}_2) = 3.1 \times 10^4 \text{ cm}^{-3}$. Similarly, in the case of clump-7, $R_{\text{dep}} = 0.12$ pc is reached at a density of $n(\text{H}_2) = 2.2 \times 10^4 \text{ cm}^{-3}$.

In Figs 8(g) and (h), we show the mean-radial column density profiles of $N(\text{H}_2)$ and $N(\text{C}^{18}\text{O})$, respectively, in the region between the two clumps (i.e. F1 region in Fig. 1), while the depletion profile is shown in panel (i). The results are in agreement with those estimated for clumps 5 and 7, showing a $0.08 \text{ pc} < R_{\text{dep}} < 0.15 \text{ pc}$ (comparable

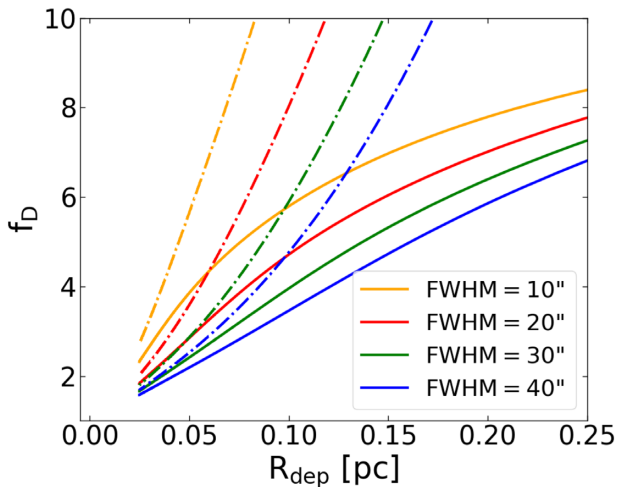


Figure 9. f_D variation as a function of R_{dep} for the best-fitting model applied to clump-5. The full lines represent the f_D – R_{dep} relation assuming $f_D = 10$ for $R < R_{\text{dep}}$, while the dash-dotted lines are the ones for $f_D = \infty$. The blue profile in Fig. 8(a) was convolved at different angular resolutions in order to predict how the same relation changes with different observational conditions. Colours are linked to the different FWHMs used for the convolution.

with the typical filament width ~ 0.1 pc; e.g. Arzoumanian et al. (2011), which corresponds to densities of $4.2 \times 10^4 \text{ cm}^{-3} > n(\text{H}_2) > 2.3 \times 10^4 \text{ cm}^{-3}$ for $f_D = \infty$ and 10, respectively. The summary of the R_{dep} estimates is shown in Table 1 (Section 5).

Once the model presented in Section 4.2 has been calibrated with the data, it allows one to estimate how the size of R_{dep} can vary according to the obtained f_D . An example of these predictions is shown in Fig. 9. Starting from the best-fitting synthetic profile obtained for clump-5 (i.e. the blue profile in Fig. 8a), we generated a family of $N(\text{C}^{18}\text{O})$ profiles, convolved with different beams (i.e. FWHM between 10 and 40 arcsec). The value of f_D reported in Fig. 9 is calculated at $R = 0$ according to equation (1), by varying the dimensions of R_{dep} from 0 up to the maximum size of the model radius for both $f_D = (10, \infty)$ at $R < R_{\text{dep}}$ (i.e. full and dash-dotted lines, respectively). We stress that this type of prediction is strongly dependent on the model assumptions (i.e. the assumed $\text{C}^{18}\text{O}/\text{H}_2$ abundance) and valid only for sources for which clump-5 is representative.

Nevertheless, a variation in R_{dep} is notable and it changes by varying the value of f_D for $R < R_{\text{dep}}$: at FWHM = 10 arcsec for example, a factor of 1.5 in f_D (from 4 to 6) corresponds to a factor of 1.6 in the R_{dep} size if we assume $f_D = \infty$ for $R < R_{\text{dep}}$, while for $f_D = 10$ it corresponds to a factor of 2.2. These variations are slightly lowered if FWHM increases: a factor of 1.4 and 1.6 for $f_D = (\infty, 10)$ at $R < R_{\text{dep}}$, respectively, for FWHM = 40 arcsec.

4.2.2 Comparison with 3D models

As a first step, we compared the estimated values of R_{dep} with the results obtained by Körtgen et al. (2017). These authors imposed a condition of total depletion on the whole collapsing region of their 3D MHD simulations, to model the deuteration process in prestellar cores. The initial core radius range is $0.08 < R_c < 0.17$ pc, depending on the simulation set-up. This assumption is necessary to reduce the computational costs of the 3D simulations but, as discussed in Section 1, the outcome can strongly deviate from reality when we move to larger scales.

For G351.77–0.51, the largest estimated depletion radius is $\lesssim 0.15$ pc, comparable with the initial core size assumed by Körtgen et al. (2017).

We note also that, imposing the condition of total depletion on a scale of 0.17 pc, Körtgen et al. (2017) find a deuterium fraction of H_2D^+ reaching a value of 10^{-3} (100 times the canonical values reported in Oliveira et al. 2003) on scales comparable with our size of R_{dep} , after $\sim 4 \times 10^4$ yr (see their fig. 3). This result is in agreement with the ages suggested by the depletion time-scale estimated in our models following Caselli et al. (1999):

$$\tau_{\text{dep}} \sim \frac{10^9}{S n(\text{H}_2)} \text{ yr}, \quad (9)$$

assuming the estimated volume densities at R_{dep} and a sticking coefficient $S = 1$.

A similar picture is reported by Körtgen et al. (2018), who show that the observed high deuterium fraction in prestellar cores can be readily reproduced in simulations of turbulent magnetized filaments. After $\sim 10^4$ yr, the central region of the filament shows a deuterium fraction 100 times higher than the canonical value, and has a radius of ~ 0.1 pc (see their fig. 1). The dimensions of the regions showing a high deuterium fraction can be compared with our estimates of R_{dep} since the two processes are connected, as shown in Section 1.

The comparison of our findings with both small- and large-scale simulations suggests that the total-depletion assumption could provide reliable results within reasonable computational times.

5 THE INFLUENCE OF A DIFFERENT ABUNDANCE PROFILE ON R_{DEP}

It is important to stress that the model discussed in this paper presents a number of limitations:

(i) A first caveat is represented by how we obtain the final values from observations (i.e. red dots in Fig. 8), with which our models were then optimized. The main assumption we made was the canonical $\text{C}^{18}\text{O}/\text{H}_2$ abundance, assumed equal to 2.1×10^{-7} , which actually can vary by up to a factor of 2.5 (Section 4.1). While knowing these limits, we note that the results shown in the final f_D map of G351.77–0.51 (i.e. Fig. 5) are in good agreement with what was found by Hernandez et al. (2011) and Hernandez et al. (2012) in the IRDC G035.30–00.33, the only other case in which a filamentary-scale depletion factor map has been shown to date;

(ii) A second limitation comes from the H_2 volume density profile we assumed in Section 4.2 to reproduce the data, which ignores all the sub-structure of clumps. Although this assumption is strong, it seems reasonable because the best-fitting $N(\text{H}_2)$ profiles shown in Fig. 8 – panels (a), (d), and (g) – are in agreement with the data points;

(iii) The third and strongest assumption we made is the shape of the $\chi_{\text{C}^{18}\text{O}}$ profile to describe the C^{18}O variation with respect to H_2 within the main ridge of G351.77–0.51. In Fig. 10, the step-function profile assumed in Section 4.2 is shown as an orange-dashed line. All the results of Figs 8 and 9, discussed in Sections 4.2.1 and 4.2.2, assume this abundance profile.

We therefore examined how much the estimates of the R_{dep} size, reported in the previous paragraphs, depend on the abundance profile assumed. For these tests, we studied the three $\chi_{\text{C}^{18}\text{O}}$ profiles shown in Fig. 10. In addition to the already discussed step-function, we tested two profiles describing a smoother variation of $\chi_{\text{C}^{18}\text{O}}$ as a function of R . In the first test, the $\text{C}^{18}\text{O}/\text{H}_2$ abundance was

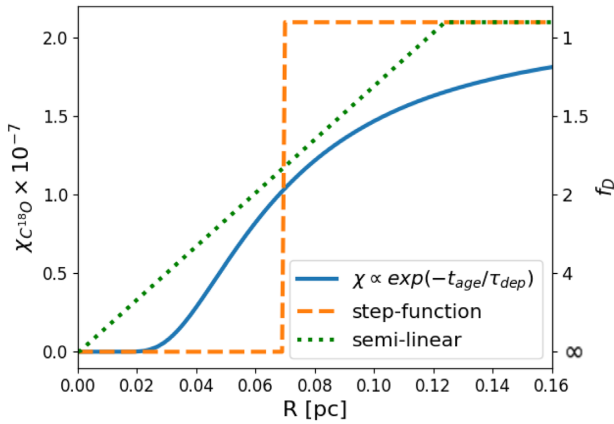


Figure 10. Sketch of the abundance profiles assumed to describe the $\chi_{\text{C}^{18}\text{O}}$ variation in our tests. The orange-dashed profile shows the step-function described in Section 4.2. In this plot, only the position of the step-function discontinuity has a physical meaning, equal to the estimated size of R_{dep} in the C5 model with $\chi_{\text{C}^{18}\text{O}} = 0$ for $R < R_{\text{dep}}$; the green-dotted profile describes a more gradual variation of $\chi_{\text{C}^{18}\text{O}}$ for $R < R_{\text{drop}}$ as a linear profile between canonical $\text{C}^{18}\text{O}/\text{H}_2$ abundance and the value of $\chi_{\text{C}^{18}\text{O}}$ at the centre of the model; the blue profile reproduces the trend $\chi_{\text{C}^{18}\text{O}} \propto \exp(-t_{\text{age}}/\tau_{\text{dep}})$, assuming only the CO absorption as dominant process (see Section 5).

assumed to be constant up to a certain distance, called R_{drop} , at which the abundance starts to linearly decrease (the green-dotted profile in Fig. 10). In the second test, we considered the physical case in which $\chi_{\text{C}^{18}\text{O}} \propto \exp(-t_{\text{age}}/\tau_{\text{dep}})$, where t_{age} is the age of the source, as expected integrating the evolution of $n(\text{C}^{18}\text{O})$ over time, $dn(\text{C}^{18}\text{O})/dt$, only considering depletion [i.e. $dn(\text{C}^{18}\text{O})/dt = -\kappa_{\text{ads}}n(\text{C}^{18}\text{O})$], and ignoring desorption]. The latter case is shown as a blue profile in Fig. 10, which reaches zero in the innermost regions of the model (i.e. in this example, for $R \lesssim 0.02$ pc), due to the high volume densities of H_2 (see equation 9).

Since we no longer have a strong discontinuity in the new abundance profiles, we have to redefine the concept of R_{dep} . In the new tests, R_{dep} will correspond to the distance at which $\chi_{\text{C}^{18}\text{O}} = 2.1 \times 10^{-8}$ (i.e. the distance at which the 90 per cent of the CO molecules are depleted), and its new estimates are shown in Table 1.

We note that the exponential profile describes the physically more realistic case, whereas the other two are simpler approximations, easier to model. Therefore, we compare the results of the exponential case with those described by the other two.

As visible in Table 1, the new R_{dep} values estimated from the model with the exponential profile are within a factor of about 2–3 of those found from the semilinear and step-function models, respectively. These results suggest that, although the assumption of a step-function to describe the $\chi_{\text{C}^{18}\text{O}}$ profile may seem too simplistic, it actually succeeds reasonably well in reproducing the results predicted by a physically more realistic profile. Furthermore, the step-function and semilinear profiles represent two extreme (and opposites) cases of a drastic and a constant change in the $\chi_{\text{C}^{18}\text{O}}$ profile, respectively, while the exponential profile can be interpreted as a physically more realistic situation. Because of this, even if the relative 3σ level errors on R_{dep} , calculated by probability distributions, are always less than 15 per cent of the values reported in Table 1, the uncertainty of a factor of 2–3 on the estimates of R_{dep} from the exponential profile seems more realistic.

Finally, it is worth discussing the predictions of the depletion time-scales in the three models. In clump-5, for example, the R_{dep}

Table 1. Summary of the estimated R_{dep} assuming the profiles in Fig. 10 and discussed in Section 5. For all the estimates of R_{dep} shown in the table, we estimate the relative 3σ level error to be less than 15 per cent.

Regions Profiles	C5 [pc]	C7 [pc]	F1 [pc]
Step-function ($f_{\text{D}} = 10$; $R < R_{\text{dep}}$)	0.10	0.12	0.15
Step-function ($f_{\text{D}} = \infty$; $R < R_{\text{dep}}$)	0.07	0.07	0.08
Exponential	0.04	0.04	0.05
Semilinear	0.02	0.02	0.03

of the model with the semilinear profile provides $n(\text{H}_2) = 5.5 \times 10^5 \text{ cm}^{-3}$. With this result, we estimated $\tau_{\text{dep}} = 1.8 \times 10^3 \text{ yr}$, smaller by a factor of about 10–20 than those obtained using the model with the step-function profile, for which $\tau_{\text{dep}} = 1.8 \times 10^4 \text{ yr}$ ($f_{\text{D}} = \infty$ for $R < R_{\text{dep}}$) or $3.2 \times 10^4 \text{ yr}$ ($f_{\text{D}} = 10$), respectively. This large discrepancy can be explained once again by noting that these tests represent the two extreme cases with respect to the exponential profile and therefore define the lower and upper limits for the estimate of τ_{dep} .

For the exponential model, the best-fitting t_{age} is $2.1 \times 10^4 \text{ yr}$, while at R_{dep} , $n(\text{H}_2) = 1.7 \times 10^5 \text{ cm}^{-3}$, which corresponds to a $\tau_{\text{dep}} = 5.8 \times 10^3 \text{ yr}$. As mentioned above, for this model, only the adsorption process of CO is included, neglecting desorption at $T \gtrsim 20 \text{ K}$. Nevertheless, the estimated t_{age} and the corresponding lower limit of τ_{dep} suggest that the contribution of the adsorption process is negligible at these scales. Furthermore, both ages seem in good agreement with the values of τ_{dep} provided by the step-function profiles, suggesting again that this is a reasonable approximation to describe the profile of $\chi_{\text{C}^{18}\text{O}}$.

6 CONCLUSIONS

In order to add a new piece of information to the intriguing puzzle of the depletion phenomenon and its consequences on the chemical evolution in star-forming regions, in the first part of this work we presented and discussed the depletion factor derived for the whole observed structure of G351.77–0.51. In the second part, we reported the results of the estimated depletion radius, as obtained from a simple toy model described in Section 4.2, testing different hypotheses (Section 5).

G351.77–0.51, with its proximity and its filamentary structure, seems to be the perfect candidate for this type of study. We use *Herschel* and Large APEX Bolometer Camera (LABOCA) continuum data, together with APEX $J(2-1)$ line observations of C^{18}O , to derive the f_{D} map (Fig. 5). Along the main body and close to clump-3, the observed depletion factor reaches values as large as ~ 6 , whereas in lower density higher temperature structures f_{D} is close to 1, as expected.

We constructed a simple model that could reproduce the observed f_{D} in two clumps along the main ridge and along the filament itself.

As results of this work, we conclude the following:

(i) In many regions of the spine and the branches of G351.77–0.51, the depletion factor reaches values larger than ~ 2.5 . This highlights that even in the less prominent structures, as the branches, the depletion of CO can start to occur, altering the chemistry of the interstellar medium and making it more difficult to study the gas dynamics and to estimate the mass of cold molecular gas (using CO);

(ii) We find that CO-depletion in high-mass star-forming regions affects not only the densest regions of the clumps, but also the filamentary structures that surround them;

(iii) The model assumed to estimate the size of the depletion radius suggests that it ranges between 0.15 and 0.08 pc, by changing the depletion degree from 10 to the full depletion state. These estimates agree with the full depletion conditions used in the 3D simulations of Körtgen et al. (2017) and Körtgen et al. (2018). This result highlights that such assumptions are not so far from the limits constrained by the observations. We also show that under certain assumptions, it is possible to estimate the size of the depleted region from f_D (see Fig. 9);

(iv) We verified that by changing the shape of the profile assumed to describe the $C^{18}O/H_2$ variations inside the filament and clumps, the estimates of the size of R_{dep} do not change more than a factor of 3. This difference was interpreted as the final uncertainty associated with the new estimates of R_{dep} from the physically more realistic case of the exponential profile, where $R_{\text{dep}} \sim 0.04\text{--}0.05$ pc;

(v) At $R = R_{\text{dep}}$, the model shows number densities of H_2 between 0.2 and $5.5 \times 10^5 \text{ cm}^{-3}$. Following Caselli et al. (1999), we estimated the characteristic depletion time-scale for the clump-5, the clump-7, and the filament region included between them (i.e. $\tau_{\text{dep}} \sim 5$ and 0.2×10^4 yr, respectively). It is interesting to note that at similar ages, in Körtgen et al. (2017) and Körtgen et al. (2018) the simulated deuterium degree also suggests a regime of high depletion on scales that are consistent with the R_{dep} estimated by our model.

Observations at higher resolutions (even in the case of G351.77–0.51) are necessary to put stronger constraints on the models. Nevertheless, that our results are comparable with the modelling presented by Körtgen et al. (2017) and Körtgen et al. (2018) suggests that even if on a large scale it is necessary to include a detailed description of the depletion processes, on the sub-parsec scales we can exploit the conditions of total depletion used in the simulations. This assumption does not seem to strongly alter the expected results but has the enormous advantage of a considerable decrease of computational costs of the 3D simulations. It would however be necessary in the future to explore the effect of properly modelling the depletion in 3D simulations, to confirm the findings reported in this work.

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REFERENCES

- Anderson L. D., Bania T. M., Balser D. S., Cunningham V., Wenger T. V., Johnstone B. M., Armentrout W. P., 2014, *ApJS*, 212, 1
- Arzoumanian D. et al., 2011, *A&A*, 529, L6
- Astropy Collaboration, 2013, *A&A*, 558, A33
- Astropy Collaboration, 2018, *AJ*, 156, 123
- Bacmann A., Lefloch B., Ceccarelli C., Castets A., Steinacker J., Loinard L., 2002, *A&A*, 389, L6
- Bacmann A., Lefloch B., Ceccarelli C., Steinacker J., Castets A., Loinard L., 2003, *ApJ*, 585, L55
- Bergin E. A., Tafalla M., 2007, *ARA&A*, 45, 339
- Bergin E. A., Alves J., Huard T., Lada C. J., 2002, *ApJ*, 570, L101
- Beuther H., Schilke P., Menten K. M., Motte F., Sridharan T. K., Wyrowski F., 2002, *ApJ*, 566, 945
- Beuther H., Schilke P., Menten K. M., Motte F., Sridharan T. K., Wyrowski F., 2005, *ApJ*, 633, 535
- Bolatto A. D., Wolfire M., Leroy A. K., 2013, *ARA&A*, 51, 207
- Caselli P., Walmsley C. M., Tafalla M., Dore L., Myers P. C., 1999, *ApJ*, 523, L165
- Caselli P., Walmsley C. M., Zucconi A., Tafalla M., Dore L., Myers P. C., 2002, *ApJ*, 565, 344
- Caselli P., Vastel C., Ceccarelli C., van der Tak F. F. S., Crapsi A., Bacmann A., 2008, *A&A*, 492, 703
- Ceccarelli C., Caselli P., Herbst E., Tielens A. G. G. M., Caux E., 2007, in Reipurth B., Jewitt D., Keil K., eds, *Protostars and Planets V*. Univ. Arizona Press, Tucson, p. 47
- Christie H. et al., 2012, *MNRAS*, 422, 968
- Crapsi A., Caselli P., Walmsley C. M., Myers P. C., Tafalla M., Lee C. W., Bourke T. L., 2005, *ApJ*, 619, 379
- Csengeri T. et al., 2016, *A&A*, 586, A149
- Faúndez S., Bronfman L., Garay G., Chini R., Nyman L.-Å., May J., 2004, *A&A*, 426, 97
- Feng S., Beuther H., Zhang Q., Henning T., Linz H., Ragan S., Smith R., 2016, *A&A*, 592, A21
- Fontani F., Giannetti A., Beltrán M. T., Dodson R., Rioja M., Brand J., Caselli P., Cesaroni R., 2012, *MNRAS*, 423, 2342
- Ford A. B., Shirley Y. L., 2011, *ApJ*, 728, 144
- Gerlich D., Herbst E., Roueff E., 2002, *Planet. Space Sci.*, 50, 1275
- Giannetti A. et al., 2014, *A&A*, 570, A65
- Giannetti A., Leurini S., Wyrowski F., Urquhart J., Csengeri T., Menten K. M., König C., Güsten R., 2017a, *A&A*, 603, A33
- Giannetti A. et al., 2017b, *A&A*, 606, L12
- Giannetti A. et al., 2019, *A&A*, 621, L7
- Griffin M. J. et al., 2010, *A&A*, 518, L3
- Herbst E., van Dishoeck E. F., 2009, *ARA&A*, 47, 427
- Hernandez A. K., Tan J. C., Caselli P., Butler M. J., Jiménez-Serra I., Fontani F., Barnes P., 2011, *ApJ*, 738, 11
- Hernandez A. K., Tan J. C., Kainulainen J., Caselli P., Butler M. J., Jiménez-Serra I., Fontani F., 2012, *ApJ*, 756, L13
- Hildebrand R. H., 1983, *Q. J. R. Astron. Soc.*, 24, 267
- Hofner P., Wyrowski F., Walmsley C. M., Churchwell E., 2000, *ApJ*, 536, 393
- Hunter J. D., 2007, *Comput. Sci. Eng.*, 9, 90
- Kong S., Caselli P., Tan J. C., Wakelam V., Sipilä O., 2015, *ApJ*, 804, 98
- König C. et al., 2017, *A&A*, 599, A139
- Körtgen B., Bovino S., Schleicher D. R. G., Giannetti A., Banerjee R., 2017, *MNRAS*, 469, 2602
- Körtgen B., Bovino S., Schleicher D. R. G., Stutz A., Banerjee R., Giannetti A., Leurini S., 2018, *MNRAS*, 478, 95
- Kramer C., Winnewisser G., 1991, *A&AS*, 89, 421
- Kramer C., Alves J., Lada C. J., Lada E. A., Sievers A., Ungerechts H., Walmsley C. M., 1999, *A&A*, 342, 257
- Ladd E. F., 2004, *ApJ*, 610, 320

- Leurini S., Codella C., Zapata L., Beltrán M. T., Schilke P., Cesaroni R., 2011a, *A&A*, 530, A12
- Leurini S., Pillai T., Stanke T., Wyrowski F., Testi L., Schuller F., Menten K. M., Thorwirth S., 2011b, *A&A*, 533, A85
- Leurini S. et al., 2019, *A&A*, 621, A130
- Li G.-X., Urquhart J. S., Leurini S., Csengeri T., Wyrowski F., Menten K. M., Schuller F., 2016, *A&A*, 591, A5
- McKinney W., 2010, in van der Walt S., Millman J., eds, Proc. 9th Python Sci. Conf. (SCIPY 2010), Data Structures for Statistical Computing in Python. p. 51, Available at: <http://conference.scipy.org/proceedings/scipy2010/pdfs/mckinney.pdf>
- Martin R. N., Barrett A. H., 1978, *ApJS*, 36, 1
- Molinari S. et al., 2010, *PASP*, 122, 314
- Oliphant T., 2006, NumPy: A Guide to NumPy, Trelgol Publishing, USA, Available at: <https://archive.org/details/NumPyBook/page/n25>
- Oliveira C. M., Hébrard G., Howk J. C., Kruk J. W., Chayer P., Moos H. W., 2003, *ApJ*, 587, 235
- Peretto N. et al., 2010, *A&A*, 518, L98
- Pilbratt G. L. et al., 2010, *A&A*, 518, L1
- Plummer H. C., 1911, *MNRAS*, 71, 460
- Poglitsch A. et al., 2010, *A&A*, 518, L2
- Ripple F., Heyer M. H., Gutermuth R., Snell R. L., Brunt C. M., 2013, *MNRAS*, 431, 1296
- Roberts H., Millar T. J., 2000, *A&A*, 361, 388
- Schisano E. et al., 2014, *ApJ*, 791, 27
- Schuller F. et al., 2009, *A&A*, 504, 415
- Sipilä O., Caselli P., Harju J., 2013, *A&A*, 554, A92
- Tafalla M., Myers P. C., Caselli P., Walmsley C. M., Comito C., 2002, *ApJ*, 569, 815
- Thomas H. S., Fuller G. A., 2008, *A&A*, 479, 751
- Urquhart J. S. et al., 2018, *MNRAS*, 473, 1059
- Wiles B., Lo N., Redman M. P., Cunningham M. R., Jones P. A., Burton M. G., Bronfman L., 2016, *MNRAS*, 458, 3429
- Wouterloot J. G. A., Henkel C., Brand J., Davis G. R., 2008, *A&A*, 487, 237
- Zhang Q., Wang Y., Pillai T., Rathborne J., 2009, *ApJ*, 696, 268

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