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# The distribution of dust in edge-on galaxies: I. The global structure

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## ABSTRACT

In this first paper in a series we present a study of the global dust emission distribution in nearby edge-on spiral galaxies. Our sample consists of 16 angularly large and 13 less spatially resolved galaxies selected from the DustPedia sample. To explore the dust emission distribution, we exploit the *Herschel* photometry in the range 100–500  $\mu\text{m}$ . We employ Sérsic and 3D disc models to fit the observed 2D profiles of the galaxies. Both approaches give similar results. Our analysis unequivocally states the case for the presence of extraplanar dust in between 6 and 10 large galaxies. The results reveal that both the disc scale length and height increase as a function of wavelength between 100 and 500  $\mu\text{m}$ . The dust disc scale height positively correlates with the dust disc scale length, similar to what is observed for the stellar discs. We also find correlations between the scale lengths and scale heights in the near- and far-infrared which suggest that the stellar discs and their dust counterparts are tightly connected. Furthermore, the intrinsic flattening of the dust disc is inversely proportional to the maximum rotation velocity and the dust mass of the galaxy: more massive spiral galaxies host, on average, relatively thinner dust discs. Also, there is a tendency for the dust-to-stellar scale height ratio to decrease with the dust mass and rotation velocity. We conclude that low-mass spiral galaxies host a diffuse, puffed-up dust disc with a thickness similar to that of the stellar disc.

**Key words:** galaxies: ISM – galaxies: photometry – galaxies: structure.

## 1 INTRODUCTION

Galaxies represent complex gravitationally bound systems, the main constituents of which are enigmatic dark matter, stars, gas, and dust. The subject of interstellar dust, which manifests itself by the absorption of ultraviolet (UV) and optical photons, light scattering, and thermal emission at far-infrared/submillimetre (FIR/submm) wavelengths due to the absorbed radiation, has been extensively studied for decades (see e.g. reviews by Savage & Mathis 1979; Stein & Soifer 1983; Draine 2003, 2011; Galliano, Galametz & Jones 2018). The optical manifestation of dust in the form of a sharp dust lane, observed in many edge-on galaxies, serves as clear evidence that, while only contributing minorly to the total mass of the interstellar medium (ISM), it plays a crucial role in the radiative transfer (RT) in galaxies, and represents a tremendous obstacle for studying the interior of galaxies in the UV and optical domains. Moreover, the physical role of dust in the ISM is of great importance. Dust serves as a catalyst for the transformation of atomic to molecular hydrogen (Gould & Salpeter 1963; Goldsmith, Li & Krčo 2007) and

participates in the cooling and heating of the ISM (Draine 1978; Bakes & Tielens 1994; Tielens 2005). Observations and physical laboratories help us infer the characteristics of dust which are then linked to physical processes in the ISM at work. This builds up our understanding of how dust forms, evolves, and is destroyed in the context of the formation and evolution of galaxies.

The overall spatial distribution of dust in galaxies can be inferred in several ways. One way is by fitting optical images of edge-on galaxies with a prominent dust lane using RT simulations which allows one to retrieve the parameters of the dust disc (Xilouris et al. 1997, 1998, 1999; Bianchi 2007; De Geyter et al. 2013, 2014; Mosenkov et al. 2016, 2018). In RT modelling, it is generally assumed that dust, similar to stars, follows an exponential distribution both in the radial and vertical directions. The fact that the global radial distribution of dust in galaxies is essentially exponential has been confirmed in a number of studies where 1D, azimuthally averaged profiles for individual galaxies (Alton et al. 1998a; Muñoz-Mateos et al. 2009; Hunt et al. 2015; Casasola et al. 2017) and averaged profiles for a sample of galaxies (Smith et al. 2016) have been analysed. However, recently Mosenkov et al. (2019, hereafter M19) found that for 320 disc galaxies observed with the *Herschel* Space Observatory (Pilbratt

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et al. 2010) and fitted with a 2D Sérsic (Sérsic 1963; Sérsic 1968) profile, the pure exponential decline of the radial dust profile is not entirely correct – the radial distribution of cold dust emission has, on average, a Sérsic index (which characterizes the shape of the model surface brightness profile) of  $0.7 \pm 0.4$  at 250  $\mu\text{m}$ . They concluded that a substantial fraction of dust discs can be well-fitted with a Gaussian-like profile.

As to the vertical distribution of dust in galaxies, an exponential law is also commonly used (see e.g. Verstappen et al. 2013), but observational evidence for employing this law instead of, for example, a  $\text{sech}^2$ -law is missing, mainly due to the poor resolution of contemporary FIR observations. Moreover, there is not even a consensus on which law to use for describing the stellar vertical distribution in galaxies (see Comerón et al. 2011b, *axis* review by van der Kruit & Freeman 2011, and a recent study on the Milky Way by Mosenkov et al. 2021).

Apart from our Galaxy, which is viewed exactly edge-on from the Earth, dozens of nearby edge-on galaxies offer the best opportunity for exploring both the radial and vertical distributions of stars, gas, and dust in great detail. The natural and most common approach for studying the dust distribution in edge-on galaxies is by exploiting and analysing FIR/submm observations of the thermal emission of dust (Alton et al. 1998b, 2000b; Verstappen et al. 2013; Hughes et al. 2014; Bocchio et al. 2016a; Yoon et al. 2021). However, other wavelength ranges are also widely used for studying the dust structure in galaxies: UV (mainly through absorption and photon scattering off dust grains; Hodges-Kluck & Bregman 2014; Seon et al. 2014; Shinn & Seon 2015; Baes & Viaene 2016; Hodges-Kluck, Cafmeyer & Bregman 2016; Jo et al. 2018; Shinn 2019), optical (through absorbing dust structures viewed against the background stellar light of the galaxy; Howk & Savage 1997, 1999, 2000; Dalcanton, Yoachim & Bernstein 2004; Rossa et al. 2004; Thompson, Howk & Savage 2004; Howk 2005; Masters et al. 2010; Holwerda et al. 2012b, 2019; Narayan, Dettmar & Saha 2020), mid-infrared (MIR; Irwin & Madden 2006; Burgdorf, Ashby & Williams 2007; Irwin et al. 2007), as well as near-infrared (NIR; Montgomery & Clemens 2014), and optical polarimetry (Seon 2018). According to these studies, there is now ample evidence for the presence of extraplanar (i.e. at high vertical distances) clouds and filaments of dust in disc galaxies.

The high-latitude clouds of dust in galaxies can be produced by several possible mechanisms. First, since thick stellar discs and haloes are ubiquitous in disc galaxies (see e.g. Yoachim & Dalcanton 2006; Comerón, Salo & Knapen 2018), it is reasonable to suggest that thick dust discs or haloes are related to their stellar counterparts through the population of AGB stars present in these thick components (Howk 2012). Secondly, the dust in the thin disc could be ejected by hydrodynamic or magnetohydrodynamic flows, hit by interstellar shocks or galactic fountains or chimneys (Shapiro & Field 1976; Bregman 1980; Houck & Bregman 1990; de Avillez, Berry & Kahn 1998) and entrained together with gas until it is finally expelled out above the mid-plane, forming a thick dust disc or halo, or even expelled further out in the intergalactic medium (Zaritsky 1994; Thompson et al. 2004; Ménard et al. 2010). This scenario suggests that small grains are partly destroyed, which is supported both observationally and theoretically (Borkowski et al. 2006; Williams et al. 2006; Micelotta, Jones & Tielens 2010; Bocchio, Jones & Slavin 2014; Bocchio et al. 2016a). Other mechanisms include flows driven by magnetic field instabilities (Parker 1966, 1992), slow global winds (e.g. cosmic ray driven winds; Popescu et al. 2000a), and radiation pressure (Ferrara et al. 1991; Franco et al. 1991; Ferrara 1993; Davies et al. 1998).

Motivated by the study of M19, who considered 2D *Herschel* profiles for a large sample of DustPedia (Davies et al. 2017) galaxies, oriented at arbitrary angles, we present a detailed analysis of the global structure of the cold dust emission with a focus on the vertical dust structure in nearby, well-resolved edge-on galaxies. As such, this is the first systematic study for probing the 3D distribution of dust emission and exploring the general scaling relations between the stellar and dust structural parameters. Given the superior angular resolution and sensitivity of the *Herschel* observations, we are able to carry out a spatially resolved analysis for the off-plane dust structure. Bocchio et al. (2016a) has explored the vertical extent of dust emission around NGC 0891 and revealed the presence of a second, thick dust disc component. In the present paper one of our goals is to expand their study and demonstrate the existence of extraplanar FIR emission in other nearby, well-resolved edge-on galaxies.

For this purpose we use two alternate models for describing the dust emission profiles in galaxies: Sérsic and 3D exponential disc. We also study the dependence of the retrieved parameters on wavelength in the FIR domain, compare the dust structure with the stellar distribution, and examine several general scaling relations of galaxies which involve the parameters of the dust structure. Finally, we link the obtained results with the above proposed scenarios for the presence of extraplanar dust in galaxies.

In this article we study the distribution of the FIR/submm emission, which is not the same as the mass-density distribution of dust. The monochromatic luminosity density distribution can be approximated by a modified blackbody (MBB; see e.g. section 3 in Utomo et al. 2019) model written as

$$I_\nu(\lambda) = \kappa(\lambda_0) \Sigma_d \left( \frac{\lambda_0}{\lambda} \right)^\beta B_\nu(T_d, \lambda), \quad (1)$$

where  $\kappa(\lambda_0)$  is the dust absorption cross-section per unit mass (or opacity) at a reference wavelength  $\lambda_0$ ,  $\Sigma_d$  is the dust mass surface density,  $\beta$  is the dust emissivity index (which characterizes the composition of dust grains and is typically  $\lesssim 2.0$ , see Galliano et al. 2018 and references therein), and  $B_\nu(T_d, \lambda)$  is the Planck function for the temperature  $T_d$  at wavelength  $\lambda$ . As one can see, apart from the dust mass surface density, the luminosity density at the given wavelength also depends on the dust temperature, which, in its turn, relates to the interstellar radiation field (ISRF), and optical properties of dust – each of them can be a function of the coordinates inside the galaxy. Smith et al. (2016) fitted MBB spectral-energy distributions (SEDs) to combined radial profiles of 117 galaxies at 100–500  $\mu\text{m}$  and showed that the dust temperature gradually decreases from  $\sim 25$  K in the galaxy centre to  $\sim 15$  K at the optical radius (see their fig. 4). The behaviour of the emissivity index  $\beta$  is uncertain: in some galaxies it increases with radius, whereas in others it decreases (Hunt et al. 2015). The dust mass absorption coefficient is poorly constrained, and it is highly uncertain how it varies with radius, albeit Clark et al. (2019) found a significant variation in  $\kappa(500)$  (measured at 500  $\mu\text{m}$ ) for M 74 (by a factor of 2.3) and M 83 (by a factor of 5.3). If we assume that  $\beta$  and  $\kappa(\lambda_0)$  do not vary significantly inside the galaxy (usually they are fixed in SED modelling), the constructed mass density profiles of cold dust in galaxies generally follow the FIR/submm luminosity density distribution (although with a different scale length depending on wavelength, see e.g. Smith et al. 2016; Casasola et al. 2017). In this paper by the dust distribution we mean the FIR surface brightness distribution, although, in fact, these terms should not be used interchangeably.

The outline of this paper is as follows. In Section 2, we present the sample of nearby edge-on galaxies with available *Herschel*

observations. In Section 3, we describe how we prepare the data for a subsequent analysis presented in Section 4. The results of our structural decomposition are provided in Section 5. In Section 6, we present several new galaxy scaling relations based on the obtained results and summarize our findings in Section 7.

## 2 THE DATA AND SAMPLE

In our study, we exploit the DustPedia imagery (including the error maps) in the five *Herschel* bands (where applicable, see below): PACS 100, PACS 160, SPIRE 250, SPIRE 350, and SPIRE 500  $\mu\text{m}$ . In these bands, the pixel scales are as follows: 3, 4, 6, 8, and 12 arcsec, respectively. The point spread function (PSF) FWHMs for these wavebands are 8 arcsec (PACS 100), 12 arcsec (PACS 160), 18 arcsec (SPIRE 250), 24 arcsec (SPIRE 350), and 36 arcsec (SPIRE 500). The PACS 100 waveband has thus the best resolution among the *Herschel* bands that we consider in this study. The dedicated reduction of the *Herschel* maps is presented in Clark et al. (2018).

To select edge-on galaxies with available *Herschel* observations, we use the DustPedia (Davies et al. 2017) galaxy sample and data base<sup>1</sup> (Clark et al. 2018). From the whole DustPedia sample, we pre-selected 50 edge-on galaxies based on the galaxy inclination angle  $i \geq 85^\circ$  estimated in M19. From this preliminary sample we rejected galaxies in which the disc is (i) obviously less inclined so that the spirals are well-seen (e.g. NGC 4359, NGC 7090), and/or (ii) the galaxy has a deformed, irregular shape (e.g. NGC 4747). Also, we should take into account that the poor resolution of the *Herschel* telescope constrains our ability to recover the structural parameters of the galaxy. Therefore, we removed galaxies with a radial extent, determined at S/N = 3, of less than 2 arcmin in the PACS 100 band and/or with a vertical extent of a few pixels only (e.g. NGC 4289, NGC 4710, NGC 4758).

After this careful selection, we created a final sample of 29 genuine edge-on galaxies with an inclination angle  $i \gtrsim 85^\circ$ , which is either taken from the literature (mostly based on RT modelling or kinematical analysis) or roughly estimated by positioning the dust lane with respect to the galaxy mid-plane (see Mosenkov et al. 2015). In the latter case or when the dust lane is barely visible or warped (e.g. in NGC 3628), the uncertainty on  $i$  is set to a conservative value of  $1.5^\circ$ . For NGC 4244, the value of its inclination angle in the literature varies from  $82^\circ$  (Comerón et al. 2011a) to  $88^\circ$  (Holwerda et al. 2012a). We decided to use the latter estimate as it is retrieved from RT modelling, and, due to the large difference in the estimated inclination angles, assign to it a large error of  $2^\circ$ . The uncertainty on the galaxy inclination is taken into account when we compute the uncertainties of the model parameters of our photometric fitting described in Section 4.

We split by hand our final sample into two subsamples roughly based on their angular size D25: the main sample ( $D25 \gtrsim 4$  arcmin) and the additional sample ( $D25 \lesssim 4$  arcmin). The main sample comprises 16 large galaxies, most of which have RT models in the literature (see Table 1). The high spatial resolution of these nearby galaxies makes them particularly well-suited for studying the structure of their dust component in several PACS and SPIRE wavebands. The additional sample has 13 galaxies smaller in angular size and none of them have RT models. As these galaxies have a lower spatial resolution, we only consider the PACS 100 band for them. This supplementary sample supports the results of our study, especially for low-mass galaxies (see below). In Section 5, we examine if our

selected galaxies are resolved in the vertical direction based on the retrieved parameters of our modelling.

We list the general properties of the galaxies under study in Table 1 and display their composite optical Sloan Digital Sky Survey (SDSS; York et al. 2000; Ahumada et al. 2020) *gri* (or Digitized Sky Survey; Reid et al. 1991; Lasker et al. 1996) and composite *Herschel* 100, 160, 250  $\mu\text{m}$  images in Fig. 1. As one can see from the table, the galaxies in our sample are nearby, have a large angular extent, and are mostly of type Sb–Sc, but span fairly wide ranges of stellar masses  $9 \lesssim \log M_*/M_\odot \lesssim 11$ , dust masses ( $6 \lesssim \log M_{\text{dust}}/M_\odot \lesssim 8$ ), and maximum rotation velocities ( $70 \text{ km s}^{-1} \lesssim v_{\text{rot}} \lesssim 300 \text{ km s}^{-1}$ ). Most galaxies demonstrate well-defined optical dust lanes, although in low-mass galaxies (ESO 373-008, NGC 4183, NGC 4244, NGC 5023), the dust distribution looks more diffuse and less structured. This is compatible with the results of Dalcanton et al. (2004) and Holwerda et al. (2019) who found that dust lanes start to appear in galaxies with  $M_* \gtrsim 10^9 M_\odot$ . In NGC 3628 and NGC 4631, the stellar (and dust) discs are strongly disturbed due to interaction with close neighbours, which is manifested in various conspicuous tidal features in deep optical images (see Mosenkov et al. 2020 and references therein).

To trace the old stellar population in the galaxies (which makes up the bulk of the stellar mass), we use the 3.4  $\mu\text{m}$  (W1) imaging available thanks to the *WISE* space observatory (Wright et al. 2010). The pixel size is 1.375 arcsec and FWHM = 6.1 arcsec. The reason why we use *WISE* data instead of, for example, the *Spitzer* Space Telescope (Werner et al. 2004) with a much better resolution is that (i) all galaxies in our sample are imaged by *WISE* and (ii) the poor *WISE* resolution is more compatible with the *Herschel* resolution as fine structural details in the large nearby galaxies are smeared out and their global stellar structure (decomposed into a single disc and a bulge) can be easily compared to the dust structure traced by *Herschel*. We note that the emission in the *WISE* W1 band can be biased by hot dust from massive star-forming regions or dust around active nuclei and emission from the 3.3  $\mu\text{m}$  polycyclic aromatic hydrocarbon (PAH) feature (Meidt et al. 2012). As shown by Querejeta et al. (2015) using *Spitzer* 3.6  $\mu\text{m}$  and 4.5  $\mu\text{m}$  data, the contamination by dust can be as much as 20–30 per cent for Sa–Sc galaxies which constitute our sample. The distribution of this contaminating emission in edge-on galaxies is concentrated along a very thin disc the vertical scale of which is typically smaller than the *WISE* resolution and much smaller than the thickness of the general stellar disc – see stellar and non-stellar maps provided by the *Spitzer* Survey of Stellar Structure in Galaxies (S<sup>4</sup>G; Sheth et al. 2010; Muñoz-Mateos et al. 2013) collaboration in Querejeta et al. (2015)<sup>2</sup> for edge-on galaxies including some of the galaxies from our sample. In principle, this can introduce a bias in tracing the stellar distribution and make its vertical profile thinner than it is in reality. However, given that the contamination of the dust emission does not usually exceed 30 per cent of the total galaxy flux at 3.6  $\mu\text{m}$ , the influence of this contamination should not significantly change the results of our profile fitting of the stellar distribution described in Section 4.

## 3 DATA PREPARATION

Below we describe our preparation of the *Herschel* data. For *WISE* 3.4  $\mu\text{m}$ , we use already prepared data from M19. As this study is focused on the global structure of dust in edge-on spiral galaxies,

<sup>1</sup><http://dustpedia.astro.noa.gr/>

<sup>2</sup><https://irsa.ipac.caltech.edu/data/SPITZER/S4G/galaxies/>

**Table 1.** The general properties of edge-on galaxies from the main (upper part) and additional (bottom part) samples.

| Galaxy      | RA<br>(J2000) | Dec.<br>(J2000) | D<br>(Mpc)<br>(1) | D25<br>(arcmin)<br>(2) | $\log M_*$<br>( $M_\odot$ )<br>(3) | Type<br>(4) | $i$<br>(deg)<br>(5) | Ref.<br>(6)     | $v_{\text{rot}}$<br>( $\text{km s}^{-1}$ )<br>(7) | $\log M_{\text{dust}}$<br>( $M_\odot$ )<br>(8) |
|-------------|---------------|-----------------|-------------------|------------------------|------------------------------------|-------------|---------------------|-----------------|---------------------------------------------------|------------------------------------------------|
| ESO 209-009 | 07:58:15      | −49:51:15       | 12.59             | 6.49                   | 10.12                              | SBc         | $87.1 \pm 1.5$      | this study*     | $141.4 \pm 3.6$                                   | $7.21 \pm 0.05$                                |
| ESO 373-008 | 09:33:22      | −33:02:01       | 9.68              | 3.7                    | 9.56                               | Sc          | $89.5 \pm 1.5$      | this study*     | $93.7 \pm 2.8$                                    | $6.7 \pm 0.06$                                 |
| IC 2531     | 09:59:56      | −29:37:01       | 36.65             | 6.53                   | 10.83                              | Sc          | $89.5 \pm 0.2$      | 1,2,3           | $228.3 \pm 3.2$                                   | $7.75 \pm 0.06$                                |
| NGC 0891    | 02:22:33      | + 42:20:57      | 9.86              | 13.03                  | 10.82                              | Sb          | $89.7 \pm 0.3$      | 1,4,5           | $212.1 \pm 5.0$                                   | $7.7 \pm 0.05$                                 |
| NGC 3628    | 11:20:17      | + 13:35:23      | 10.81             | 11.04                  | 10.82                              | SBb         | $88.4 \pm 1.5$      | 5               | $212.1 \pm 3.4$                                   | $7.61 \pm 0.05$                                |
| NGC 4013    | 11:58:31      | + 43:56:48      | 18.79             | 4.88                   | 10.73                              | SABb        | $89.9 \pm 0.1$      | 1,6,7           | $181.7 \pm 4.3$                                   | $7.36 \pm 0.04$                                |
| NGC 4183    | 12:13:17      | + 43:41:55      | 16.67             | 4.29                   | 9.68                               | Sc          | $85.5 \pm 0.1$      | 8               | $103.7 \pm 2.0$                                   | $6.88 \pm 0.06$                                |
| NGC 4217    | 12:15:51      | + 47:05:30      | 18.54             | 5.53                   | 10.77                              | Sb          | $87.9 \pm 0.1$      | 6,7             | $187.6 \pm 4.1$                                   | $7.51 \pm 0.04$                                |
| NGC 4244    | 12:17:30      | + 37:48:26      | 4.35              | 16.22                  | 9.2                                | Sc          | $88 \pm 2.0$        | 9*, 10*, 11, 12 | $89.0 \pm 2.0$                                    | $6.38 \pm 0.07$                                |
| NGC 4302    | 12:21:42      | + 14:35:54      | 14.32             | 6                      | 10.37                              | Sc          | $89.6 \pm 0.1$      | 6               | $167.8 \pm 3.1$                                   | $7.2 \pm 0.05$                                 |
| NGC 4437    | 12:32:46      | + 00:06:54      | 8.39              | 9.02                   | 9.97                               | Sc          | $86 \pm 1.5$        | 13*             | $139.5 \pm 2.6$                                   | $7.11 \pm 0.05$                                |
| NGC 4631    | 12:42:08      | + 32:32:29      | 7.35              | 14.45                  | 10.42                              | SBcd        | $86 \pm 1.5$        | 14*             | $138.5 \pm 3.3$                                   | $7.28 \pm 0.06$                                |
| NGC 5023    | 13:12:13      | + 44:02:28      | 6.58              | 6.22                   | 8.81                               | Sc          | $87 \pm 1.5$        | 15              | $80.3 \pm 1.8$                                    | $6.02 \pm 0.08$                                |
| NGC 5529    | 14:15:34      | + 36:13:36      | 44.46             | 5.79                   | 11.06                              | SABc        | $87.1 \pm 0.1$      | 1,6,7           | $282.4 \pm 5.2$                                   | $7.97 \pm 0.07$                                |
| NGC 5746    | 14:44:56      | + 01:57:18      | 27.04             | 7.24                   | 11.1                               | Sb          | $86.8 \pm 0.1$      | 6               | $311.1 \pm 7.8$                                   | $7.85 \pm 0.06$                                |
| NGC 5907    | 15:15:54      | + 56:19:44      | 17.22             | 11.3                   | 10.87                              | SABc        | $86.5 \pm 0.3$      | 16*, 1, 17, 7   | $226.6 \pm 4.3$                                   | $7.85 \pm 0.05$                                |
| IC 0610     | 10:26:28      | + 20:13:41      | 29.24             | 1.89                   | 10.26                              | Sbc         | $90 \pm 1.5$        | this study*     | $138.1 \pm 4.1$                                   | $6.83 \pm 0.06$                                |
| IC 2233     | 08:13:59      | + 45:44:32      | 14                | 2.88                   | 9.21                               | SBc         | $88.5 \pm 1.5$      | 18*             | $77.1 \pm 1.8$                                    | $5.91 \pm 0.16$                                |
| NGC 3454    | 10:54:30      | + 17:20:39      | 20.61             | 2.41                   | 9.7                                | SBc         | $85 \pm 1.5$        | this study*     | $92.7 \pm 2.5$                                    | $6.69 \pm 0.08$                                |
| NGC 3501    | 11:02:47      | + 17:59:22      | 22.59             | 4.31                   | 10.2                               | Sc          | $87 \pm 1.5$        | this study*     | $138.8 \pm 3.0$                                   | $7.05 \pm 0.06$                                |
| NGC 3592    | 11:14:27      | + 17:15:36      | 20.51             | 2.13                   | 9.53                               | Sc          | $86 \pm 1.5$        | this study*     | $78.9 \pm 1.9$                                    | $6.45 \pm 0.18$                                |
| NGC 4222    | 12:16:23      | + 13:18:25      | 19.77             | 2.83                   | 9.8                                | Scd         | $90 \pm 1.5$        | this study*     | $99.8 \pm 2.2$                                    | $6.89 \pm 0.08$                                |
| NGC 4330    | 12:23:17      | + 11:22:05      | 18.71             | 2.29                   | 9.91                               | Sc          | $84 \pm 1.5$        | 19*             | $112.3 \pm 3.3$                                   | $6.75 \pm 0.06$                                |
| NGC 4607    | 12:41:12      | + 11:53:12      | 19.95             | 2.86                   | 10.03                              | SBbc        | $90 \pm 1.5$        | this study*     | $99.1 \pm 3.4$                                    | $6.84 \pm 0.06$                                |
| UGC 07321   | 12:17:34      | + 22:32:25      | 23.33             | 4.82                   | 9.61                               | Scd         | $90 \pm 1.5$        | this study*     | $93.9 \pm 2.3$                                    | $6.97 \pm 0.12$                                |
| UGC 07387   | 12:20:17      | + 04:12:06      | 38.19             | 1.99                   | 9.98                               | Scd         | $87 \pm 1.5$        | this study*     | $116.8 \pm 3.3$                                   | $6.91 \pm 0.1$                                 |
| UGC 07522   | 12:25:58      | + 03:25:49      | 35.48             | 2.62                   | 10.36                              | Sc          | $90 \pm 1.5$        | this study*     | $143.4 \pm 3.7$                                   | $7.3 \pm 0.07$                                 |
| UGC 07982   | 12:49:50      | + 02:51:05      | 19.23             | 2.91                   | 9.74                               | Sc          | $87 \pm 1.5$        | this study*     | $98.4 \pm 2.8$                                    | $6.63 \pm 0.1$                                 |
| UGC 09242   | 14:25:21      | + 39:32:22      | 24.32             | 4.21                   | 9.21                               | Scd         | $90 \pm 1.5$        | this study*     | $78.7 \pm 2.5$                                    | $6.15 \pm 0.19$                                |

*Note.* Columns: (1) Distance taken from Clark et al. (2018) as ‘best’ (preferred) distance estimate, (2) optical diameter taken from HyperLeda<sup>3</sup> (Makarov et al. 2014) based on their  $\log d_{25}$  parameter, (3) stellar mass from M19, (4) morphological type from HyperLeda, (5) weighted average inclination angle and its weighted average uncertainty based on the references listed in column Ref., (6) references for inclination angle mainly estimated from RT modelling or using the approach proposed in Mosenkov et al. (2015) (\* denotes non-RT methods), (7) apparent maximum rotation velocity of gas and its error taken from HyperLeda, (8) total dust mass from Nersesian et al. (2019) derived from SED modelling using a modified black-body (MBB) model.

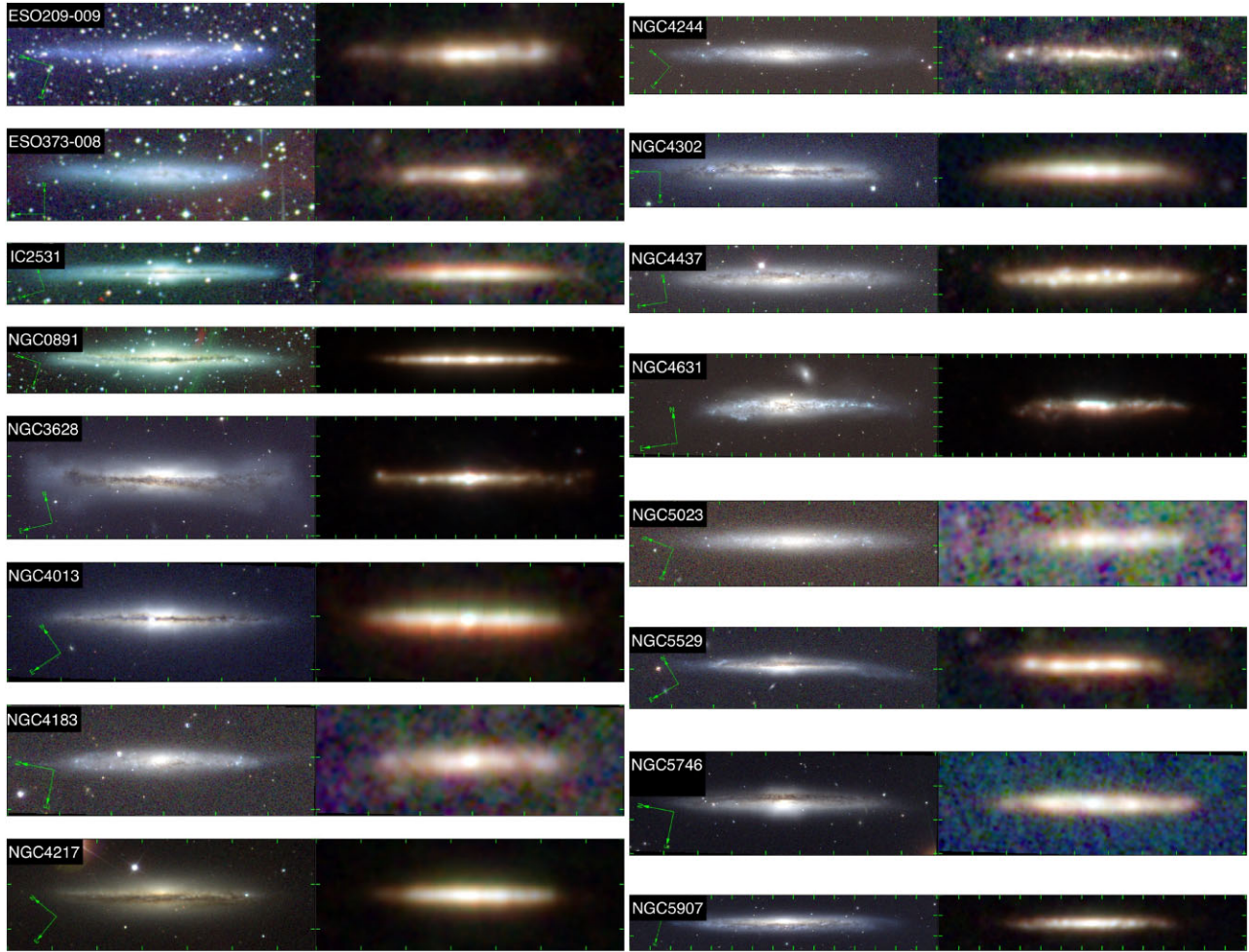
References: (1) Xilouris et al. (1999). (2) Mosenkov et al. (2016). (3) Peters et al. (2017). (4) Xilouris et al. (1998). (5) Shinn & Seon (2015). (6) Bianchi (2007). (7) Mosenkov et al. (2018). (8) Shinn (2018). (9) Olling (1996). (10) Comerón et al. (2011a). (11) MacLachlan et al. (2011). (12) Holwerda et al. (2012a). (13) Mosenkov et al. (2020). (14) Dumke, Krause & Wielebinski (2004). (15) Alton et al. (2000a). (16) van der Kruit & Searle (1982). (17) Just, Möllenhoff & Borch (2006). (18) Matthews & Uson (2008). (19) Fossati et al. (2018).

especially on detecting extraplanar dust emission, we should pay special attention to the preparation of the *Herschel* data. The *WISE* data are only used for tracing the global stellar structure of the galaxies as compared to their dust counterpart, with no emphasis on the thick stellar disc. Thus, we are satisfied with the preparation of the *WISE* data which has been done in M19.

For each galaxy image, we carefully examined the background emission to remove extrinsic sources. To mask out all objects, including the target galaxy, we employed SExtractor (Bertin & Arnouts 1996) with the parameters DETECT\_THRESH = 2.0 and DETECT\_MINAREA = 3. The size of the created mask was enlarged by a factor of 1.5, as determined empirically, to totally cover all scattered light. Careful masking of the galaxy image was crucial at this point to completely exclude all emission from any sources except for the background. Therefore, we manually revisited the created masks to ensure that all detected objects had been completely masked and no scattered light from them was present. Then, we iteratively fitted the background using polynomials of varying degrees starting from zero (a flat background). We examined the background in the resultant background-subtracted image to make a decision about

which polynomial of the background to use (typically, the second degree was sufficient to fit the background around the galaxy). Some images were contaminated by Galactic cirrus, which structure cannot be fitted using a single polynomial. In such cases, we partitioned the frame into several large-scale areas, each fitted with a separate polynomial. The typical size of such structures in our images was much larger than the area occupied by the detected sources, and, thus, the fitting for these parts using a sufficient number of pixels was possible. This procedure allowed us to clean the frames from the undesired large-scale background in our images.

All the *Herschel* galaxy images were rebinned to the PACS 100 image, so that they all would have the same pixel scale (for comparison purposes) and a unified mask (to take into account the same number of pixels while fitting, see below). All galaxy images were rotated by an average angle estimated by M19, who fitted the galaxy profiles with a single Sérsic model with a free position angle (see Section 4). Thus, in our final images the galactic discs are aligned with the horizontal direction. This was done to easily extract the cumulative (integrated) galaxy profiles along the major (parallel to the mid-plane) and minor (perpendicular to the mid-plane) axes



**Figure 1.** In each panel, left-hand images: composite RGB snapshots for the main sample galaxies of the *g*-, *r*-, and *i*-passband frames from the SDSS or Digitized Sky Survey (POSS2/UKSTU Red, Blue and IR); right-hand images: composite RGB snapshots based on the *Herschel* 100, 160, and 250  $\mu\text{m}$  images. The green ticks on the axes showcase a 1 arcmin size.

in our subsequent analysis. After the rotation, we equally cropped the images for each galaxy, so that the final frames in all five bands would have the same size.

The rebinning, rotation, and cropping procedures were also applied to the error maps which contain information about the weight for each pixel as a measure of the photometric quality.

Finally, for each galaxy we created a unified mask image consisting of a sum of the individual masks in each band which we initially created for the background subtraction (excluding, of course, the target galaxy itself from this mask). Despite the presence of some contaminating objects which appear in one filter and not in others, the unified mask allows us to create cumulative radial and vertical profiles which take into account the same pixels from the galaxy in all bands used. This unified mask is exploited in our further photometric fitting of the galaxy images so that we exclude the possibility of the dependence of our models on the masks used in different bands.

Knowing the extended point spread function (PSF) is of crucial importance for accurate fitting of galaxy profiles, especially when exploring the vertical surface brightness distribution in edge-on galaxies (see e.g. a comprehensive analysis in Sandin 2014). As shown by Sandin (2014) for the *Hubble Space Telescope*, the size of the PSF model should be ‘at least 1.5 times as extended as the vertical distance of the edge-on galaxy’. In our study, we exploit deep PSFs

for the PACS instrument from Bocchio, Bianchi & Abergel (2016b), combined from Vesta and Mars observations, and for the SPIRE instrument based on Neptune observations which were used for the flux calibration of this photometer (Bendo et al. 2013). These PSF images have a large dynamical range ( $\sim 10^6$ ) and radius ( $\sim 8$  arcmin), which are sufficient for the purposes of this study – for all galaxies in our sample, the size of the PSF is larger than the maximal vertical/radial distance in our final galaxy frames, meaning the whole range of vertical (and radial) distances is covered by the PSF. The PSF images have been rotated (to align the direction of the spacecraft’s Z-axis with its direction on the galaxy images) and rebinned to  $3 \text{ arcsec pix}^{-1}$  to match the respective final galaxy images. The PSF image for the *WISE* 3.4  $\mu\text{m}$  data was taken from the Explanatory Supplement to the *WISE* All-Sky Data Release Products.<sup>3</sup> Along with the PSF, we use the term ‘line’ spread function (LSF) as a 1D beam produced by averaging the 2D PSF along one (vertical or radial) direction (see Verstaappen et al. 2013; Bocchio et al. 2016a).

All the steps in data preparation described above were carried out using the IMAN Python package.<sup>4</sup>

<sup>3</sup><https://wise2.ipac.caltech.edu/docs/release/allsky/expsup/>

<sup>4</sup>[https://bitbucket.org/mosenkov/iman\\_new/src/master/](https://bitbucket.org/mosenkov/iman_new/src/master/)

#### 4 THE FITTING METHOD

Using a single Sérsic model, M19 performed 2D fitting of cold dust emission profiles for 320 DustPedia galaxies, including all objects from our sample. Here, we modify the modelling approach which was applied in M19 in several aspects.

First, in contrast to M19, who used automatic image preparation for galaxy fitting, the background subtraction and masking in this study were revised to accurately account for all possible sources that could potentially contaminate the galaxy emission. This is of high importance for studying the off-plane emission in edge-on galaxies which can be easily confused (and misinterpreted) with an undersubtracted sky background or scattered light from other objects in the image.

Secondly, for galaxies with bright FIR emission clumps or other unresolved components located in the plane of the galaxy, these features are fitted using a Gaussian or Sérsic function. This is done to minimize the influence of such compact sources on the overall galaxy profile instead of masking them out.

Thirdly, we use two alternative models in our study: a simple Sérsic model and a 3D broken exponential disc (see the description of these profiles in Sections 4.1 and 4.2, respectively). Using such a complex, non-single exponential profile to describe the FIR emission in spiral galaxies is based on the M19 modelling results: the dust emission profiles often exhibit a deficit of emission in the centre (a plateau, or a core) with respect to the general exponential fitting function (also noted by Muñoz-Mateos et al. 2009). As shown by M19, the average Sérsic index for 250  $\mu\text{m}$  galaxy profiles is  $0.7 \pm 0.4$ , and half of the sample galaxies they studied have  $n_{250} < 0.7$ . This indicates that the dust discs are not purely exponential and are often better described by a Gaussian function. They also concluded that the dust emission deficiency often occurs within 0.3–0.4 optical radii. In Section 5, we will demonstrate that this is also the case for many edge-on galaxies in our sample.

Fourthly, in this study we employ the IMFIT code (Erwin 2015) instead of GALFIT (Peng et al. 2002, 2010) and GALFITM (Bamford et al. 2011; Häußler et al. 2013), which were used in M19. This is done because the latter two codes do not include 3D fitting models. In Section 5, we compare the results of our single Sérsic modelling with those from M19 to assess the consistency of the two fitted models with each other.

Below, we briefly describe the two alternate functions (Sérsic and 3D disc models) for our photometric analysis. In our fitting, we use the Levenberg–Marquardt algorithm to retrieve the best-fitting parameters by minimizing the  $\chi^2$  value and convolve the models using the PSF images described in Section 3. Weight maps for each galaxy image are taken into account while modelling. For each galaxy image in each band, we do not keep fixed the centre or the position angle of the component used, as the astrometric calibration in the DustPedia images is not ideal and, for example, the bright centre of the galaxy can be slightly shifted (by 1–2 pixels at maximum) in different bands. For galaxies with a complex FIR profile with bright clumps and/or a bright central source (a bulge, a bar, or an AGN), we add other components to our model which can be described by Sérsic, Gaussian, or PSF profiles.

We do not repeat fitting of the *WISE* 3.4  $\mu\text{m}$  images with a single Sérsic model since it has been carried out in M19 for the whole DustPedia sample. However, we perform a bulge-disc decomposition for all edge-on galaxies in our sample to retrieve the parameters of their stellar discs and compare them with those of the dust emission profiles.

##### 4.1 Sérsic model

The Sérsic function (Sérsic 1963; Sérsic 1968) is characterized by the following major-axis intensity profile (see also Baes & Ciotti 2019):

$$I(r) = I_e \exp \left\{ -b_n \left[ \left( \frac{r}{r_e} \right)^{1/n} - 1 \right] \right\}, \quad (2)$$

where  $r$  is the radius,  $I_e$  is the surface brightness at the effective (half-light) radius  $r_e$ , and the Sérsic index  $n$  controls the shape of the intensity profile. The function  $b_n$  is calculated via the polynomial approximation found by Ciotti & Bertin (1999) (when  $n > 0.36$ ) and MacArthur, Courteau & Holtzman (2003) (when  $n \leq 0.36$ ).

For galaxies from the main sample, we start fitting with a 250  $\mu\text{m}$  galaxy image. We chose this band as it has an average resolution among the five *Herschel* bands considered and negligible contamination of the warm dust emission as compared to the PACS bands with higher resolution. The derived  $n$  parameter is then kept fixed during the fitting of the galaxy images in the remaining bands. This approach is justified because, according to M19 (see their fig. 6), the Sérsic index of the dust emission in the 100–500  $\mu\text{m}$  wavebands varies within 10–15 per cent of the average value without any systematics with wavelength, and, thus, can be taken independent (constant) of wavelength.

For galaxies from the additional sample, the fitting, of course, is done in only one PACS 100 waveband.

Since for some sample galaxies the inclination angle is slightly different from  $90^\circ$ , we correct the apparent flattening  $q$  (another free parameter in our model) for the inclination to compute the intrinsic flattening  $q_0$  using the Hubble formula (Hubble 1926):

$$q_0^{\text{seraic}} = \frac{\sqrt{q^2 - \cos^2 i}}{\sin i}. \quad (3)$$

The uncertainty of this parameter includes the error on  $q$  from IMFIT and the inclination uncertainty given in Table 1.

##### 4.2 3D disc model

The function we describe below represents a 3D luminosity density model for an axisymmetric disc, the vertical  $Z$ -axis of which is inclined by the angle  $i$  to the line of sight. The radial profile in this model is a broken exponential function (Freeman 1970) and the vertical profile has a simple exponential decline. The 3D luminosity density model  $j(R, z)$  at radius  $R$  from the central axis and at height  $z$  from the galaxy mid-plane is given by

$$j(R, z) = j_{\text{rad}}(R) e^{-|z|/h_z}, \quad (4)$$

where  $h_z$  is the vertical scale height of the disc.

The radial distribution in the disc plane is described by

$$j_{\text{rad}}(R) = S j_0 e^{-\frac{R}{h_{\text{in}}}} \left[ 1 + e^{\alpha(R - R_b)} \right]^{\frac{1}{\alpha} \left( \frac{1}{h_{\text{in}}} - \frac{1}{h_{\text{out}}} \right)}, \quad (5)$$

where  $j_0$  is the central luminosity density,  $h_{\text{in}}$  is the exponential scale length of the inner disc region before the break at the radius  $R_b$ , and  $h_{\text{out}}$  is the exponential scale length of the outer disc region beyond the break radius (see Erwin, Pohlen & Beckman 2008). The parameter  $\alpha$  controls the abruptness of the transition between the inner and outer regions of the disc. In our fitting, we constrain this parameter between 0.1 (steady change of the profile) and 50 (abrupt break). The scaling factor  $S$  is given by

$$S = \left( 1 + e^{-\alpha R_b} \right)^{-\frac{1}{\alpha} \left( \frac{1}{h_{\text{in}}} - \frac{1}{h_{\text{out}}} \right)}. \quad (6)$$

If the 3D luminosity density distribution has no break along the radial profile, then it can be given by

$$j(R, z) = J_0 e^{-R/h_R - |z|/h_z}. \quad (7)$$

As noted above, the inner galaxy region often demonstrates a deficit of FIR emission which means that  $h_{\text{in}}$  should have a very large value. For all such profiles, we fix  $h_{\text{in}} = 1000$  pix, for simplicity. Also, for all galaxies with a broken exponential profile, we do not fit the break radius  $R_b$ , but manually set it (and keep it fixed) by analysing cumulative radial profiles of the galaxy in all wavebands with the sufficient resolution in the vertical direction. Further on in the paper, we define  $h_R = h_{\text{in}}$  if  $h_{\text{in}}$  is of the same order as  $h_{\text{out}}$  (i.e. if the intensity within the break radius is not constant) and  $h_R = h_{\text{out}}$  otherwise.

In our fitting, the model inclination  $i$  is kept fixed at the galaxy inclination listed in Table 1. This is done to reduce the number of free parameters and to avoid possible degeneracies in our modelling. For calculating the uncertainties of the fit parameters and to take into account the effect of inclination (for very thin discs, even a very little change of inclination can be dramatic for the observed profile, see e.g. Verstappen et al. 2013), we run 50 sets of models with different (but fixed during fitting) inclination angles. The values of the inclination are drawn from a normal distribution with the mean and standard deviation taken from Table 1.

For *WISE* 3.4  $\mu\text{m}$ , we only consider a simple exponential 3D disc without a break as almost all galaxies in our sample do not show strong breaks in the *WISE* radial profiles. This is likely due to the poor resolution of the *WISE* data coupled with the fact that the stellar populations do not show a flat profile in the inner part of the galaxy which is observed in many dust emission profiles, as shown in M19.

As in the case of the Sérsic modelling, we start our fitting with the SPIRE 250 band and then use the output parameters for this band as the first guess parameters for the other *Herschel* bands. For the additional sample, we only fit the images at 100  $\mu\text{m}$ .

To find the intrinsic flattening  $q_{0,\lambda}^{\text{disc}}$  for both the broken exponential and simple 3D disc models, we determine the apparent flattening for an outer isophote of 0.003 Jy pix $^{-1}$  of the non-convolved model rotated to 90°. This isophote level was estimated empirically for all galaxies in our sample. It has a sufficient S/N ratio  $>3$  in all wavebands used and is outside of the break radius if it exists for all the galaxies in the sample.

## 5 RESULTS

In this section, we present the results of our fitting and compare them with the literature. We consider dependencies of the retrieved parameters on wavelength and explore different scaling relations including correlations with the general galaxy observables.

The results of our fitting for the Sérsic and 3D disc models are provided in Table A1, with the respective errors if available. Cumulative radial and vertical profiles of the sample galaxies with the superimposed models in the PACS 100 waveband are shown in Fig. A1.<sup>5</sup> These profiles were created by averaging the intensities in the four quadrants created if we dissect the galaxy along the major and minor axes. Then, we find an average for all the pixels perpendicular to the mid-plane for the given radius (cumulative radial profile) or parallel to the mid-plane for the given height (cumulative vertical profile). This approach is used to significantly increase the S/N ratio

in the resultant profile and to smooth out bright non-axisymmetric features. In our plots, we also add the LSF to demonstrate how the resolution can affect the observed and model profiles.

### 5.1 Notes on the fit results

For the whole sample (if we exclude outliers with extraplanar dust emission, see Section 5.3), we find a reduced  $\chi^2 = 1.04 \pm 0.45$  for the 3D disc models and  $0.97 \pm 0.44$  for the Sérsic models in all *Herschel* bands under study. These similar values of the reduced  $\chi^2$  suggest that both models are almost equally good in describing the 2D galaxy profiles. Therefore, in Section 6 we will use the results for both models to strengthen our conclusions on galaxy scaling relations.

According to the results of our modelling of the *Herschel* images, not all galaxies in our main and additional samples have reliable values of the apparent flattening and, consequently, the disc scale height in the wavebands we selected for our analysis. The apparent width of the emission profile in the vertical direction serves as a good criterion for examining the reliability of the fitting: it should be larger than the PSF HWHM = FWHM/2 in this specific waveband,  $r_{e,\lambda} \cdot q_\lambda > 0.8 \text{ HWHM}_\lambda$  (see fig. 7 and corresponding text in Gadotti 2009). Indeed, for some galaxies in Fig. A1, the FWHM of the vertical cumulative profile is closely or almost completely aligned with the PSF profile so that the shape of the inner part of the vertical galaxy profile is virtually described by the PSF shape. We mark such galaxies in Table A1 by \*. In our main and additional samples, 13 and 8 galaxies, respectively, are vertically resolved in PACS 100. Seven galaxies from our main sample appeared to be vertically resolved in the three PACS 100–SPIRE 250 bands, and only four galaxies are vertically resolved in all five PACS 100–SPIRE 500 wavebands.

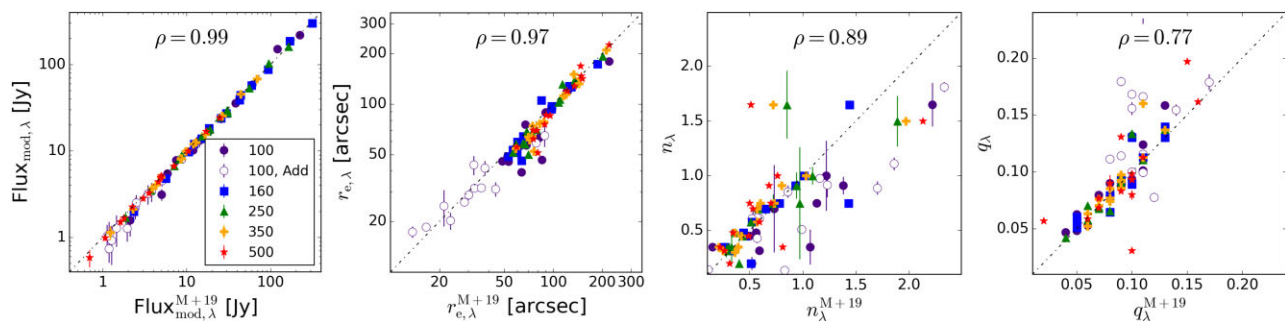
We note that the galaxy can have a reliable estimate of the exponential scale length and a poorly constrained estimate of the scale height due to the insufficient *Herschel* resolution in this waveband. We do not consider such galaxies in our statistics (marked by \* in Table A1) and scaling relations in Section 6 where the disc scale height  $h_{z,\lambda}$ , the intrinsic disc flattening  $q_{0,\lambda}^{\text{disc}}$ , and the Sérsic apparent flattening  $q_\lambda$  or intrinsic flattening  $q_{0,\lambda}^{\text{Sérsic}}$  are present.

### 5.2 Validity of the results

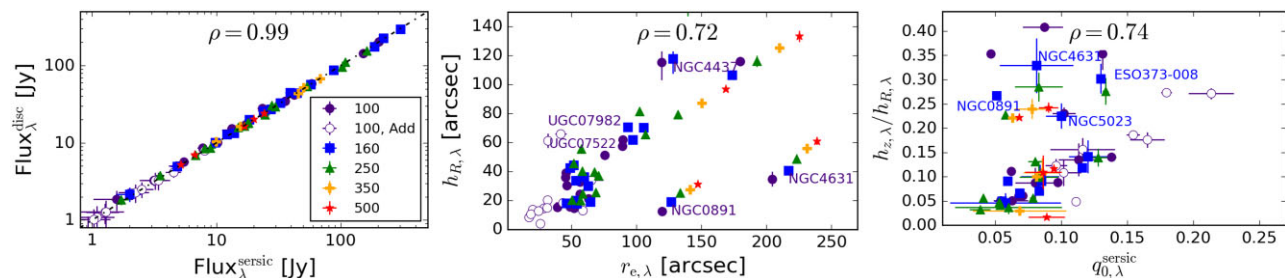
In Fig. 2, we compare the results of our Sérsic fitting with those presented in M19 for the same edge-on galaxies. As one can see, the results agree fairly well. The fluxes are in almost perfect agreement (the Pearson correlation coefficient  $\rho = 0.99$ ), with slightly lower values for smaller galaxies as compared to the M19 results. For the effective radius, the agreement is excellent ( $\rho = 0.97$ ). The Sérsic index is the most sensitive parameter in our fitting, but even for it, a strong correlation ( $\rho = 0.89$ ) with a few outliers is seen. These outliers have a larger Sérsic index  $n_\lambda^{\text{M}+19}$  in M19 because in the present fitting, the models of these galaxies consist of several components (Sérsic dust disc + dust clumps, a bright point-like source at the centre, a bulge, or a bar etc. – see Section 5.3) instead of a single Sérsic profile in M19. As to the apparent flattening, most galaxies have consistent values in both sets of Sérsic models, with several outliers being preferentially among galaxies from the additional sample. This discrepancy is likely related to the poorer spatial resolution of these galaxies: their vertical profiles are essentially described by the PSF (see Section 5.1). Overall, we conclude that our fits generally agree with those in M19.

In Fig. 3, we validate our fitting results for the 3D disc models as compared to the Sérsic models. In the left-hand panel we compare the

<sup>5</sup>Profiles in all wavebands used and for the Sérsic model are provided in the online material.



**Figure 2.** Comparison of the best-fitting parameters for the models obtained in M19 (X-axis) and those obtained in this work (Y-axis) for the Sérsic model. The results for the different bands are depicted by different symbols and colours. The dot-dashed lines show one-to-one relations.



**Figure 3.** Left-hand panel: Comparison of the 3D disc and Sérsic model fluxes within apertures with the semimajor and semiminor axes from Clark et al. (2018). Middle: Comparison of the 3D disc scale length and the effective radius of the Sérsic model. Right-hand panel: Comparison of the disc scale length-to-scale height ratio versus the intrinsic flattening for the Sérsic model. The dot-dashed line shows a one-to-one relation.

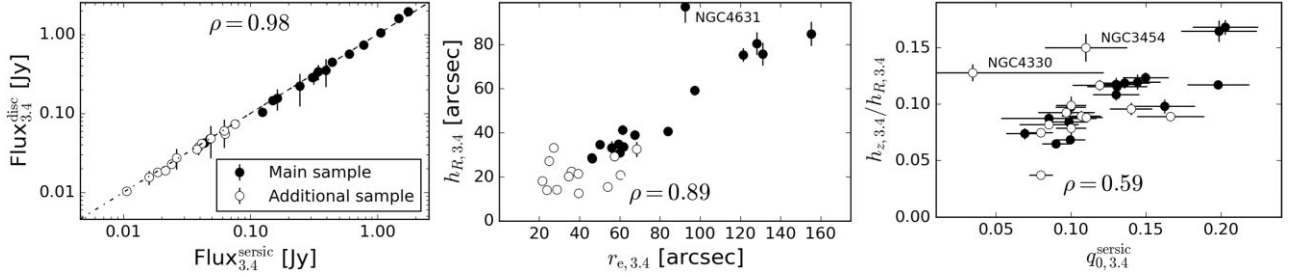
modelled fluxes within the apertures provided in Clark et al. (2018). As one can see, all data points follow a one-to-one relation. In the middle plot of Fig. 3, we display how the disc scale length and the effective radius of the Sérsic model (obtained in this study) agree with each other. Despite five outliers, a tight correlation between these parameters is apparent ( $\rho = 0.72$ ). NGC 0891, NGC 4437, and NGC 4631 are well-resolved nearby galaxies with a complex structure for which a broken exponential profile with two radial scale lengths is more accurate than a simple Sérsic function. UGC 07522 and UGC 07982 are, on the contrary, from the additional sample and have rather large disc scale lengths with respect to their effective radii. Finally, in the right-hand panel of Fig. 3, we show the correlation between the intrinsic flattening for the 3D disc and Sérsic models. Again, NGC 0891 and NGC 4631 deviate from the main trend ( $\rho = 0.74$ ), probably due to the presence of a thick dust disc and an unresolved superthin dust disc in NGC 0891 (Bocchio et al. 2016a). In general, we can conclude that both models agree fairly well with few exceptions.

Fig. 4 presents the comparison of our 3D disc models and Sérsic models from M19 for the *WISE* 3.4  $\mu\text{m}$  data. As can be seen, the fluxes and geometrical parameters of the models are in good agreement, except for several outliers. Both NGC 3454 and NGC 4330 have puffed-up stellar discs with boxy isophotes which may explain why these galaxies do not follow the general trend between the intrinsic disc and Sérsic flattenings. Also, 10 galaxies from our sample are in common with a sample of 175 edge-on galaxies from Mosenkov, Sotnikova & Reshetnikov (2010) who decomposed them into a bulge and disc using the Two Micron All Sky Survey (2MASS; Skrutskie et al. 2006) *JHK<sub>s</sub>*-band images. Our comparison with their study shows an excellent agreement for the disc scale length and scale height, with  $\rho = 0.99$  and  $\rho = 0.98$ , respectively.

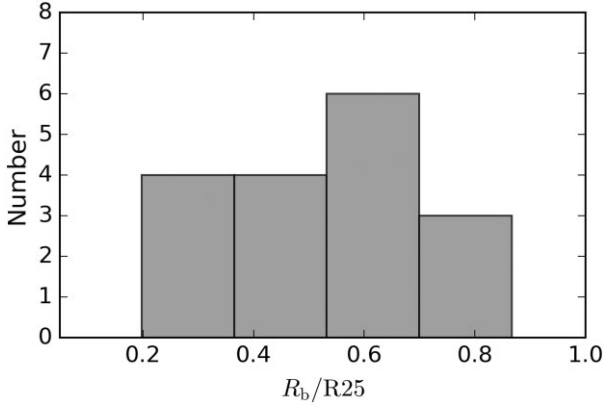
### 5.3 Cumulative intensity profiles

In this subsection, we describe cumulative radial and vertical intensity profiles with the 3D disc model shown in Fig. A1 and in the online material.

Seventeen galaxies out of 29 demonstrate a noticeable break in their radial profiles which we included in our broken exponential disc models for these galaxies. M19 studied the distribution of such characteristic (break) radii for a large sample of DustPedia galaxies and found a bimodal distribution with one peak centred at  $\sim(0.3\text{--}0.4) R_{25}$  and another at  $\sim 0.6 R_{25}$ , where  $R_{25}$  is half of the optical diameter  $D_{25}$ . They attributed the inner (shorter) break radius to the region where the dust emission deficit is observed and the outer (longer) break radius to the break on the corresponding downbending (Type II, see e.g. Erwin et al. 2008) stellar profile. The number of galaxies in our sample is too low to study the statistics of break radii: they uniformly span from  $\sim 0.2$  to  $\sim 0.8 R_{25}$  (see Fig. 5). However, as shown in M19, the shortest break radii denote the region where the dust emission deficit should be observed (depression in the dust emission profile with a very large inner scale length  $h_{R,\text{in}}$  – e.g. for IC 2531, NGC 4437, NGC 5023). This FIR emission deficit reflects the depression in the dust-mass density profile which, in its turn, generally follows the distribution of total (H I and H<sub>2</sub>) gas (see fig. A2 in Casasola et al. 2017 where they compare the radial distribution of dust, stars, and gas in 18 face-on spiral galaxies). FIR surface brightness profiles without a depression in the centre are mostly found in galaxies with enhanced star formation in the central region which can be related to bar instability, formation of a pseudo-bulge, or AGN activity (see section 5.4.4 in M19). By comparing the profiles of the same galaxy in different *Herschel* bands, we can notice that this enhanced emission peak is usually very bright at 100  $\mu\text{m}$  (where the contribution of the emission from stochastically heated warm dust is already small but not negligible) and gets lower



**Figure 4.** Comparison of the 3D disc model obtained in this study with the Sérsic model from M19 for 3.4  $\mu\text{m}$  data: the model fluxes (left-hand panel), the disc scale length versus the Sérsic effective radius (middle), and the disc relative thickness versus the intrinsic flattening for the Sérsic model (right-hand panel). The dot-dashed line shows a one-to-one relation.



**Figure 5.** Distribution by the break radius normalized by the optical radius  $R25 = D25/2$  (listed in Table 1) for 17 galaxies in the main and additional samples which were fitted with a broken exponential disc model.

and even disappears at longer wavelengths – see, for example, the radial profiles of NGC 3628, NGC 4013, and NGC 4631 in different FIR/submm wavebands. This dimming of the central source with increased wavelength stems from the fact that the contribution of the warm dust emission to the total emission decreases with wavelength and, thus, only the global cold dust disc is essentially traced in the SPIRE wavebands.

The longer break radii beyond  $\sim 0.5 R25$  (e.g. in NGC 4217, NGC 4631, NGC 5907) characterize a smooth transition between the inner profile with a larger scale length and the outer region with a smaller scale length (all broken exponential disc models in our sample have downbending profiles). As strongly supported by M19 (see their fig. 13), the dust and stellar radial profiles correlate with each other very well and are both observed in the radial range  $\sim (0.4 - 0.8) R25$ . For our sample, we also compared the radial profiles at 100 and 3.4  $\mu\text{m}$  and found a good consistency between them, except for the cases where a central depression of dust emission is observed.

Among 16 galaxies in the main sample, 10 galaxies have additional components apart from the dust disc (bright clumps, bright central point sources, bulges, and bars). For example, NGC 0891 has a bright Gaussian source at the centre that is probably related to the bulge or a bar. NGC 3628 is modelled with a bright point source at the centre and a Sérsic bar. The NGC 4244 model has five Gaussians to fit the brightest clumps of dust emission. NGC 4631 hosts a bright Sérsic bar, and NGC 4437 has nine point-like sources which describe bright emission clumps along the galaxy mid-plane. All these features are seen in our cumulative profiles. Galaxies in the additional sample

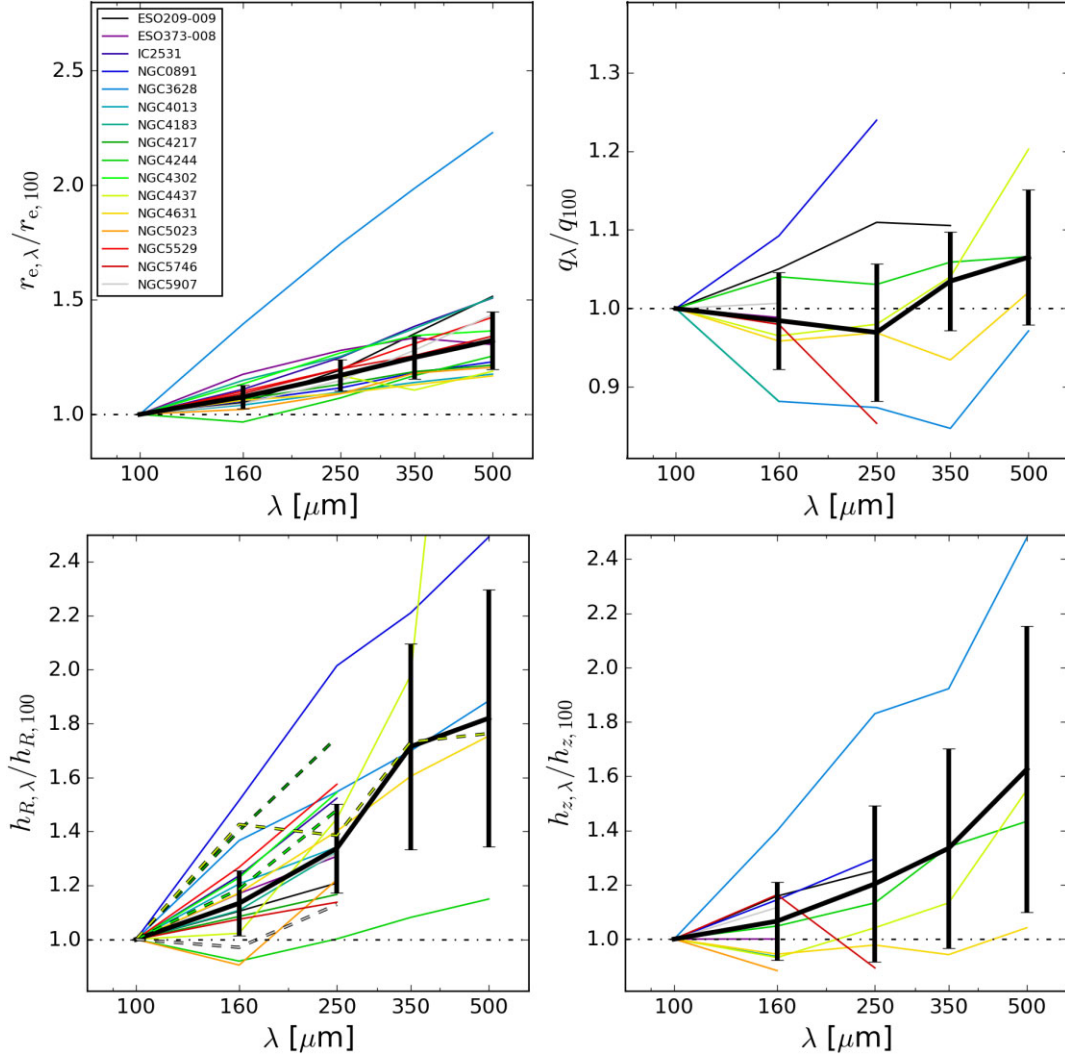
exhibit smoother profiles due to the worse spatial resolution and thus are only fitted with a single disc component.

The vertical profiles for six galaxies (ESO 209-009, NGC 0891, NGC 3628, NGC 4013, NGC 4437, and NGC 4631) show the presence of another (thicker) dust emission component in all or almost all wavebands. This component clearly dominates over the model at larger vertical distances above the galaxy mid-plane, usually starting at  $z = 1-2$  kpc. For these galaxies, the average  $\chi^2 = 19.8 \pm 10.1$  compared to  $\chi^2 = 1.04 \pm 0.45$  for the whole sample without these outliers. This proves that another dust component is required to better describe the observed profiles of these six galaxies. Four galaxies from the main sample (NGC 4244, ESO 373-008, NGC 5529, and NGC 5746) show bumps in their cumulative vertical profiles above the model profile which can point to the existence of another vertically extended component. Six galaxies from the additional sample (IC 2233, NGC 3592, UGC 07321, UGC 07387, UGC 07522, and UGC 09242) also show an excess of FIR emission compared to our models at large vertical distances. We note that this extraplanar FIR emission, which typically starts to dominate at heights 1–2 kpc above the mid-plane, cannot be explained by the scattered light of the PSF since we accounted for this effect while modelling. Also, variations in the galaxy inclination angle as described in Section 4 cannot explain the observed excess of dust emission (the red spread along the model profile in Fig. A1 shows a one- $\sigma$  spread of the model profiles for varied inclination angles, as described in Section 4.2). This indicates that the observed extraplanar emission is real and produced by the dust at high distances above the mid-plane.

#### 5.4 Dependence of the structural parameters on wavelength

In Fig. 6, we showcase how the structural parameters change with wavelength for both the Sérsic and 3D disc models (see also Table 2 where we provide a quantitative description of the gradients). As one can see, the effective radius, on average, increases up to 30 per cent from PACS 100 to SPIRE 500 which is lower (probably because of the edge-on orientation of the galaxies in this study) than what was found in M19 for the whole DustPedia sample:  $r_{e,160}/r_{e,100} = 1.10 \pm 0.27$ ,  $r_{e,250}/r_{e,100} = 1.23 \pm 0.29$ ,  $r_{e,350}/r_{e,100} = 1.38 \pm 0.34$ ,  $r_{e,500}/r_{e,100} = 1.57 \pm 0.41$ . NGC 3628 is an obvious outlier and shows an increase of the effective radius by a factor of 2.2 which is probably related to the peculiar structure of its dust (and stellar) disc perturbed by the interaction with a neighbour.

A similar positive gradient is found for the disc scale length in Fig. 6, bottom left-hand panel. NGC 891, NGC 3628, and NGC 4437, some of the best spatially resolved galaxies in our sample, exhibit a sharp increase of the disc scale length ( $h_{R,500}/h_{R,100} \gtrsim 1.8-2$ ) with wavelength. The dashed lines depict the change of the scale



**Figure 6.** Dependence of the Sérsic and 3D disc model parameters on wavelength for each galaxy (coloured lines) and an average among all these galaxies for each band (black thick line, with error bars representing standard deviations). The dashed coloured lines correspond to the outer scale length  $h_{\text{out}}$  for four galaxies (NGC 4217, NGC 4302, NGC 4437, and NGC 5907) where no deficit in the emission profile is observed and  $h_{\text{in}}$  is comparable to  $h_{\text{out}}$ .

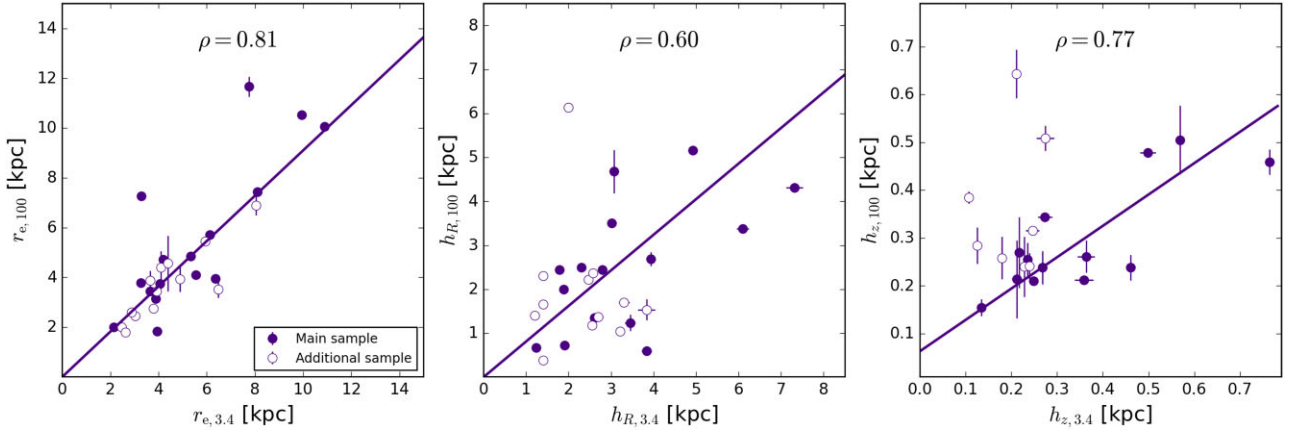
**Table 2.** Distributions of the fit parameters (mean value and standard deviation) depending on wavelength.

| Band      | $r_{e,\lambda}/r_{e,100}$ | $q_{\lambda}/q_{100}$ | $h_{R,\lambda}/h_{R,100}$ | $h_{z,\lambda}/h_{z,100}$ |
|-----------|---------------------------|-----------------------|---------------------------|---------------------------|
| PACS 160  | $1.08 \pm 0.05$           | $0.98 \pm 0.06$       | $1.13 \pm 0.12$           | $1.07 \pm 0.14$           |
| SPIRE 250 | $1.17 \pm 0.07$           | $0.97 \pm 0.09$       | $1.34 \pm 0.16$           | $1.20 \pm 0.29$           |
| SPIRE 350 | $1.25 \pm 0.09$           | $1.03 \pm 0.06$       | $1.71 \pm 0.38$           | $1.33 \pm 0.37$           |
| SPIRE 500 | $1.32 \pm 0.12$           | $1.07 \pm 0.09$       | $1.82 \pm 0.48$           | $1.63 \pm 0.53$           |

length for the outer part of the disc ( $h_{\text{out}}$ ) after the break. As one can see, the outer disc scale length also increases in the same way as the inner scale length. The agreement with the literature on the radial scale length is also good. For 18 face-on spiral galaxies, Casasola et al. (2017) found the following ratios of the scale lengths at different wavelengths (we computed them using their table 7):  $h_{R,160}/h_{R,100} = 1.10 \pm 0.01$ ,  $h_{R,250}/h_{R,100} = 1.30 \pm 0.03$ ,  $h_{R,350}/h_{R,100} = 1.45 \pm 0.05$ , and  $h_{R,500}/h_{R,100} = 1.60 \pm 0.01$ . Smith et al. (2016) found  $h_{R,500}/h_{R,250} = 1.07$  (see their table 1) for a combined radial profile of 45 large spiral galaxies versus 1.08 in this study.

As discussed in Alton et al. (1998a), Casasola et al. (2017), and M19, the observed steady increase of the effective radius and the scale length with wavelength in the FIR domain is related to the dust heating (i.e. the gradient of the cold-dust temperature) by the diffuse ISRF which is gradually decreasing with radius. Also, the dust-mass surface density distribution profiles are flatter (i.e. have a larger scale length) than the stellar mass density distribution (see table 7 in Casasola et al. 2017). This effect was also demonstrated by Bianchi (2008) using a RT model of NGC 0891 (see his fig. 10).

As shown in the upper right-hand panel of Fig. 6, the apparent flattening of the emission profiles does not change significantly with wavelength (typically, within 10 per cent, i.e. within the uncertainty of this parameter), although the statistics here are rather poor, especially at larger wavelengths. This suggests that the scale height of the dust emission profile should increase with wavelength, similar to what we observe for the radial extension. This is confirmed by the dependence of the disc scale height on wavelength from our 3D disc modelling (see the bottom right-hand panel of Fig. 6). Although the number of galaxies with the resolved vertical emission profile at larger FIR wavelengths is low, most galaxies in PACS 160 and



**Figure 7.** Dependencies between the structural parameters at 3.4 and 100  $\mu\text{m}$ : for the effective radii (left-hand panel), for the radial scale length (middle panel), and for the disc scale heights (right-hand panel). The thick lines depict the corresponding linear regression lines.

SPIRE 250 clearly demonstrate a gradual increase of the dust disc thickness with wavelength. For the three largest galaxies, the scale height grows by more than 40 per cent from 100 to 500  $\mu\text{m}$ . Again, NGC 3628 with the very disturbed stellar and dust discs is an obvious outlier and demonstrates a substantial increase of the dust disc scale height with wavelength by a factor of 2.5. Bocchio et al. (2016a) studied the vertical profile of NGC 0891 in the NIR, MIR, and FIR and found the presence of thin and thick dust discs in this galaxy. They revealed that the scale height of the thin disc demonstrates a dramatic gradient through the FIR: it changes by a factor of two from 100 to 250  $\mu\text{m}$ . The increase of the scale height for our single disc model of NGC 0891 is less prominent (30 per cent) but is comparable to the change of the scale height of their thick disc.

The observed gradient of the scale height of the sample galaxies through the FIR could indicate that the dust temperature noticeably drops within its scale height, which naturally leads to an increase of the relative FIR emission with height at longer wavelengths. Another explanation could be the influence of the emission from point-like sources (dust clumps, ends of the spiral arms positioned along the line of sight, rings) which emit at relatively warm temperatures and thus become less pronounced at longer wavelengths (plus the poorer resolution smears them out). This can increase the scale height at larger wavelengths and create the observed gradient. Bocchio et al. (2016a) also noted for NGC 0891 that contamination from an unresolved super thin disc, which represents the collective emission from dust heated by nearby hot stars in star-forming regions, may also produce this effect.

## 6 SCALING RELATIONS

One of the aims of this paper is exploring how the retrieved parameters of the dust disc correlate with each other, the stellar parameters, and the general galaxy parameters. Below we consider the results only for the PACS 100 band in which we obtained models for all galaxies from the main and additional samples. However, we ensured that the scaling relations presented below are essentially valid in the other *Herschel* wavebands.

General properties of galaxies (such as colour, mass, luminosity, size, rotational velocity, velocity dispersion, etc.) are not distributed randomly, but are mutually related and form specific scaling relations (see e.g. Mosenkov et al. 2010; Mosenkov, Sotnikova & Reshetnikov 2014; Mosenkov et al. 2015 with some caveats on their use and physical meaning). These relations provide important constraints on

the proposed scenarios of galaxy formation and evolution: any theory and numerical and cosmological hydrodynamical simulations should be able to reproduce the observed scaling relations (see e.g. Camps et al. 2016; Trčka et al. 2020; Camps et al. 2022). Below we present galaxy scaling relations which characterize the connection between the global dust and stellar structure in spiral galaxies.

### 6.1 Dust versus stellar structural parameters

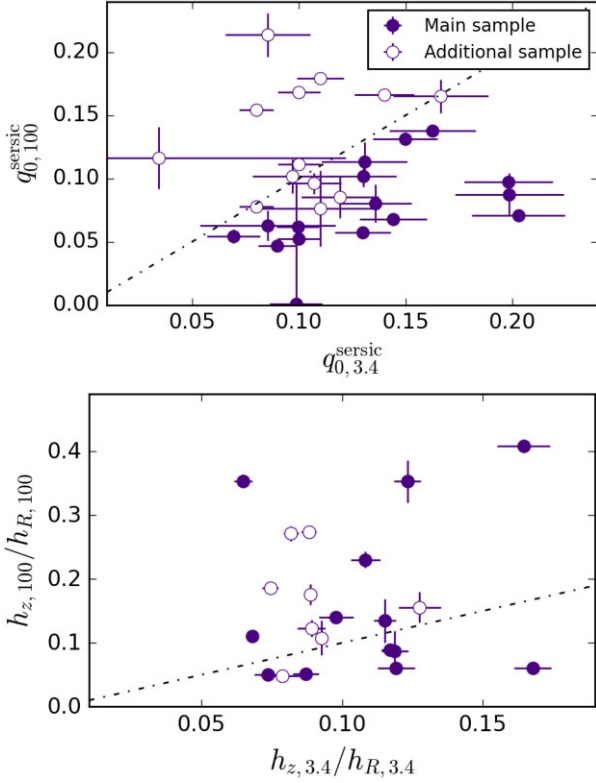
In Fig. 7 we show how the dust and stellar radial and vertical scales correlate with each other. We can see that the effective radii of our Sérsic models at 100 and 3.4  $\mu\text{m}$  agree very well ( $\rho = 0.81$ ) which has been noted in M19 for DustPedia galaxies with arbitrary inclination angles. For our edge-on galaxies, we obtain  $\langle r_{e,100}/r_{e,3.4} \rangle = 0.91 \pm 0.21$  versus  $0.90 \pm 0.27$  from M19.

Similarly, the disc scale lengths at 3.4 and 100  $\mu\text{m}$  are correlating with  $\rho = 0.60$ : more extended stellar discs host more extended dust discs. This has been noted earlier by Casasola et al. (2017) who found  $\langle h_{R,100}/h_{R,3.6} \rangle = 0.83 \pm 0.05$  compared to our  $\langle h_{R,100}/h_{R,3.4} \rangle = 0.81 \pm 0.31$ .

The main ambition of this study is to explore the vertical structure of the dust disc. In Fig. 7, right-hand plot, one can see how the scale heights at 3.4 and 100  $\mu\text{m}$  compare with each other. The relation between the vertical scales appears to be even tighter ( $\rho = 0.77$ ) than for the disc scale lengths: the thicker the stellar disc, the thicker the dust disc, on average. We can see that the scale height at 100  $\mu\text{m}$  is similar to the scale height at 3.4  $\mu\text{m}$ :  $\langle h_{z,100}/h_{z,3.4} \rangle = 0.90 \pm 0.24$ . However, there are some obvious outliers, mostly low-mass galaxies from the additional sample, which do not follow this relation. The reason for this will be seen in Section 6.3.

Interestingly, the intrinsic flattening of the dust disc does not correlate with the intrinsic flattening of the stellar disc, based on both the Sérsic and 3D disc models (see Fig. 8). This indicates that the relative thickness of the dust disc can be rather different from the relative thickness of the stellar disc (see also Mosenkov et al. 2015 who considered the dependence of the disc relative thickness on *Hubble* stage and bulge-to-total luminosity ratio). Also, note that for some dust discs the relative thickness  $h_{z,100}/h_{R,100}$  is exceptionally large and reaches 0.3–0.4 versus  $h_{z,3.4}/h_{R,3.4} = 0.10$ –0.15 for the stellar disc.

Finally, in Fig. 9 we can see the dependence of the disc scale height on the radial scale length for both the 3.4 and 100  $\mu\text{m}$  models. This is a typical scaling relation between the geometrical scales of spiral



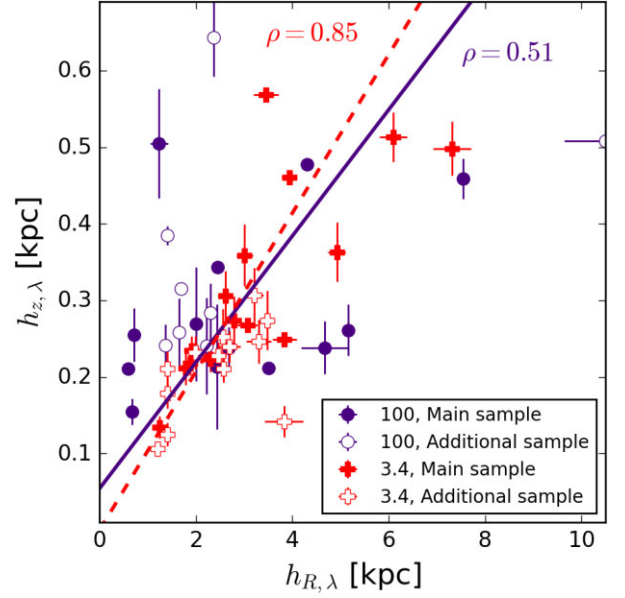
**Figure 8.** Correlations between the intrinsic flattening at 3.4 and 100  $\mu\text{m}$  for the Sérsic model (upper) and 3D disc model (bottom). The dot-dashed line shows a one-to-one relation.

galaxies: the larger the galaxy in the radial direction, the thicker, on average, it should be. This scaling effect is well-seen for the stellar discs in Mosenkov et al. (2010) (see their fig. 10). Interestingly, both the dust and stellar discs from our sample follow similar relations linking the scale length and scale height ( $\rho = 0.51$ , with slope  $k = 0.098$  and  $\rho = 0.85$ , with slope  $k = 0.103$ , respectively).

## 6.2 Dust mass–size relation

It is now interesting to explore how the total mass of the dust disc depends on its size. In Fig. 10 we consider the dependence of the dust mass on the effective radius for the Sérsic model (left-hand panel) and the scale length (middle), and scale height (right-hand panel) for the 3D disc model. These can be considered as scaling relations since we can expect that a more massive dust disc should have a larger radius given that the mean surface density does not change dramatically from galaxy to galaxy (indeed, the average dust surface density profile in fig. 5 from Smith et al. 2016 does not show a significant spread within the optical radius for 45 large spiral galaxies). Therefore, the correlation between the effective radius and the dust mass naturally occurs for both the edge-on galaxies ( $\rho = 0.86$ ) and the whole DustPedia sample ( $\rho = 0.80$ ). Similarly, we can see a correlation between the disc scale length and the dust mass, albeit with a larger scatter ( $\rho = 0.52$ ). The noisier correlation compared to the effective radius correlation can be explained by the more complex structure of the broken-exponential model as compared to the simple single Sérsic law with only one scaling parameter  $r_e$ . The stellar discs follow a similar relation ( $\rho = 0.75$ ).

Interestingly, there is no significant correlation ( $\rho = 0.31$ ) between the scale height and the total mass of the dust disc (see Fig. 10, right-

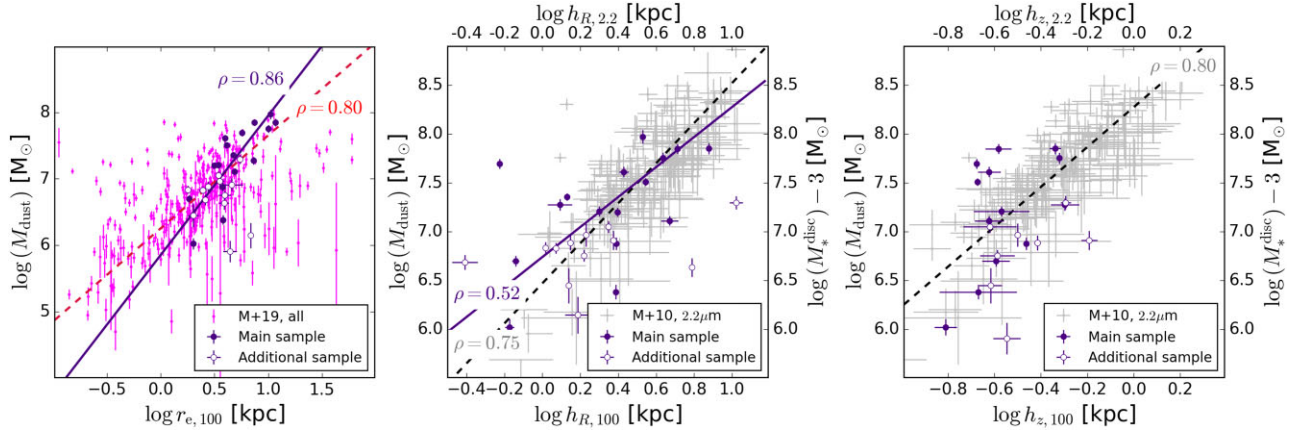


**Figure 9.** Relation between the disc scale length and scale height at both 3.4 and 100  $\mu\text{m}$ . The thick solid line depicts the linear regression line for the 100  $\mu\text{m}$  data, whereas the red dashed line corresponds to the linear regression fit for 165 edge-on galaxies in the  $K_s$  band from Mosenkov et al. (2010).

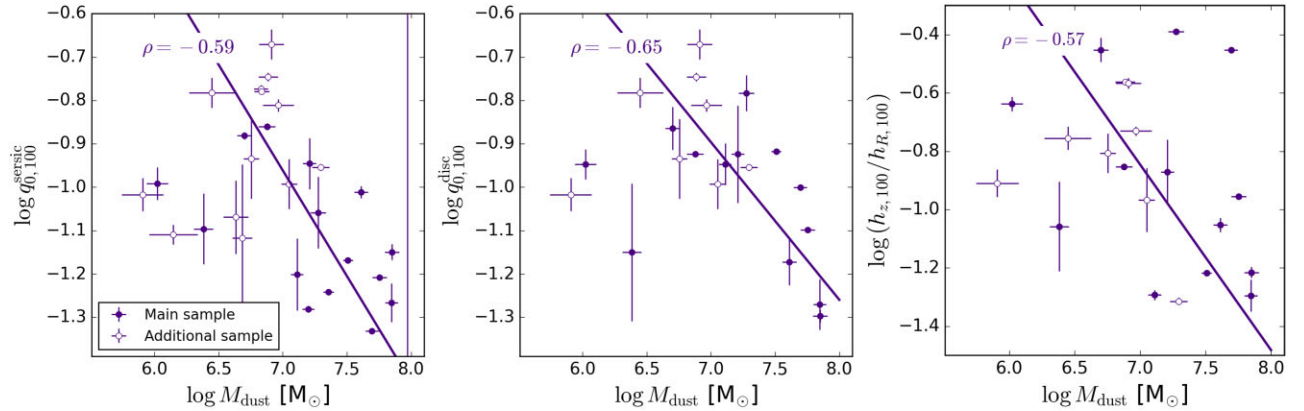
hand plot): the most important structural parameter related to the total dust mass is the radial extent of the dust disc, not its vertical size. This is in contrast to what is observed for the stellar discs for which the scale height increases with the mass (as shown by grey symbols based on the results taken from Mosenkov et al. 2010), with a similar scatter as for the stellar mass–scale length relation. This is one of the main differences between the scaling relations for the dust and stellar discs in our study.

## 6.3 Flattening of the dust disc

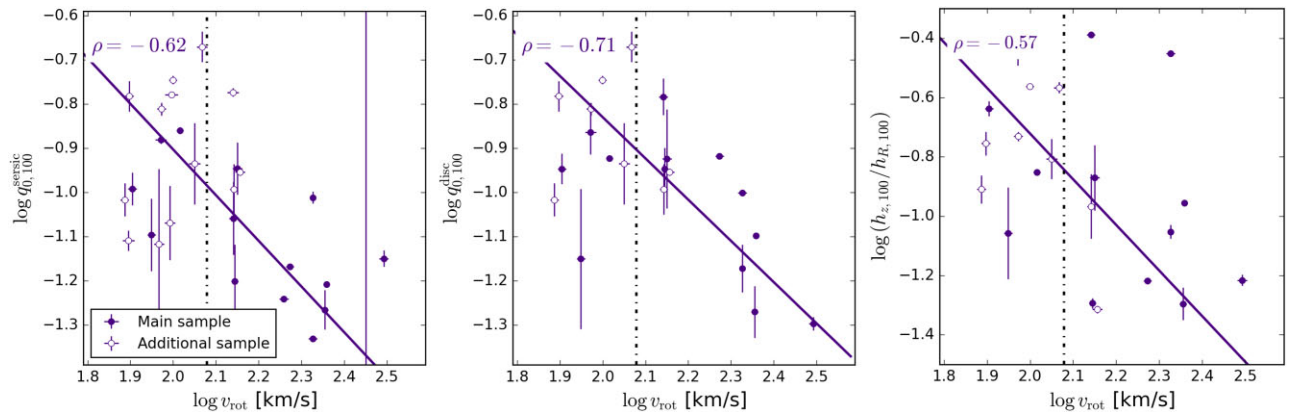
Finally, let us consider how the flattening of the dust disc relates to its mass. In Fig. 11 we show the intrinsic flattening for the Sérsic (left-hand panel) and 3D disc (middle panel) models as well as the ratio of the disc scale height to scale length (right-hand panel). In all these scatter plots we can see a certain trend which indicates that more massive dust discs tend to be flatter ( $\rho = -0.59$  for the Sérsic flattening,  $-0.65$  for the disc flattening,  $-0.57$  for the scale height-to-scale length disc ratio). Furthermore, the relative thickness of the disc and the maximum rotation velocity of the galaxy also negatively correlate with each other (see Fig. 12) –  $\rho = -0.62$ ,  $-0.71$ , and  $-0.57$ , for the Sérsic flattening, 3D disc flattening, and the scale height-to-scale length disc ratio, respectively. This suggests that the flatness of the general dust disc is also related to the total mass of the galaxy ( $v_{\text{rot}}^2$  is proportional to the mass): flatter dust discs tend to be found in more massive spiral galaxies compared to relatively thick diffuse dust discs in less massive galaxies. Finally, we find that the ratio between the scale height of the dust disc and that of the stellar disc also negatively correlates with the dust mass ( $\rho = -0.65$ ) and maximum rotation velocity ( $\rho = -0.62$ , see Fig. 13): less massive dust discs have a larger scale height with respect to the host stellar discs, and for slow-rotating spiral galaxies the scale height of the dust disc can be similar or even larger than that of the stellar disc (this has been shown in Fig. 7). In other words, in low-mass galaxies the dust is relatively more distributed throughout the height



**Figure 10.** Dependences between the total dust mass from Nersesian et al. (2019) on the Sérsic effective radius (left-hand panel), disc scale length (middle), and disc scale height (right-hand panel) – all for 100  $\mu\text{m}$ . In the left-hand panel we also show galaxies from the whole DustPedia sample in the same PACS 100 waveband. In the middle and right-hand panels, we also display the scaling relations for the stellar discs from Mosenkov et al. (2010) in the 2MASS  $K_s$  (2.2  $\mu\text{m}$ ) band. The stellar masses are shifted downwards by 3 dex in both plots. The dashed lines are linear regression lines for the large samples of galaxies, whereas the solid lines represent linear regression lines for our samples.



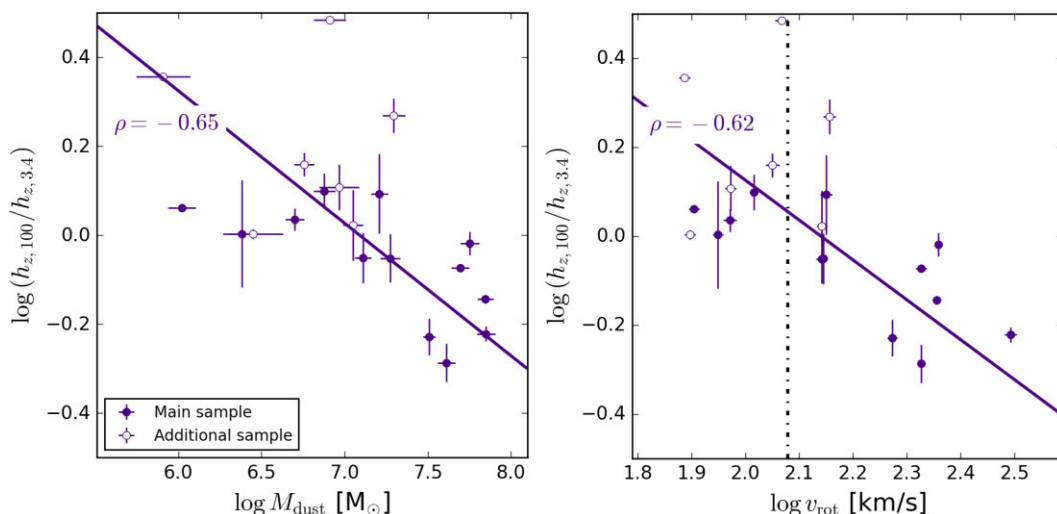
**Figure 11.** Dependences between the total dust mass from Nersesian et al. (2019) on the intrinsic flattening for the Sérsic (left-hand panel) and 3D disc (middle) models and the disc scale height-to-length ratio (right-hand panel). The solid lines depict linear regression lines.



**Figure 12.** Dependences between the rotation velocity from HyperLeda and the intrinsic flattening for the Sérsic (left-hand panel) and disc (middle) models and the disc scale height-to-length ratio (right-hand panel). The solid lines depict linear regression lines. The dash-dotted lines show 120  $\text{km s}^{-1}$ .

of the hosting stellar disc, whereas in high-mass galaxies, the dusty ISM is concentrated in the plane of the thin disc and appears as a prominent dust lane if the galaxy is oriented edge-on with respect to the observer.

This observational fact was first revealed by Dalcanton et al. (2004). They concluded that in slowly rotating galaxies with  $v_{\text{rot}} < 120 \text{ km s}^{-1}$  (i.e. in low-mass spiral galaxies), the dust distribution has a much larger scale height than in more massive, more rapidly



**Figure 13.** Correlations between the ratio of the scale heights at 100 and 3.4  $\mu\text{m}$  and the total dust mass (upper panel) and the rotation velocity (bottom panel). The solid lines depict linear regression lines. The dash-dotted line shows 120  $\text{km s}^{-1}$ .

rotating spiral galaxies (see also Holwerda et al. 2019). In the optical, such low-mass galaxies, being observed edge-on, have ill-defined or no mid-plane dust lanes at all. They attributed this phenomenon to the vertical stability of the gas and stellar disc: in galaxies with  $v_{\text{rot}} > 120 \text{ km s}^{-1}$ , the high stellar mass surface density creates instabilities in the cold ISM, which vertically collapses into a thin stellar disc with the formation of a thin dust disc. For galaxies with a lower circular velocity, the dust settles in a clumpy structure, which is well-seen in NGC 4244 and NGC 4437 (the brightest clumps were fitted on top of the Sérsic and disc components) from our main sample. For the other two galaxies with low rotation velocities (ESO 373-008 and NGC 5023), the resolution is not sufficient to see the details, and we only trace the global, smoothed structure of the dust disc.

Holwerda et al. (2019) examined the frequency of the occurrence of dust lanes in edge-on galaxies and found that bulge morphology correlates with the presence of dust lanes: dust lanes are present more often in galaxies with large, round bulges, whereas the presence of boxy bulges anticorrelates with the identification of dust lanes in such galaxies. They also concluded that dust lanes are mostly observed in galaxies where no signs of interaction with a neighbour are found. Our sample is too small to make any conclusions on this matter in general, but several examples contradict the mentioned findings of Holwerda et al. (2019). For instance, NGC 3628 is a well-known galaxy with a boxy/peanut-shaped (B/PS) bulge and with apparent signs of a merger and a well-defined dust disc (see e.g. Mosenkov et al. 2020). IC 2531, NGC 0891, NGC 4013, NGC 5529, and NGC 5746 also demonstrate the presence of a B/PS bulge, but, at the same time, exhibit highly contrasted dust lanes.

## 7 CONCLUSIONS

In this paper, we have extended the study by M19 with a main focus on the vertical structure of the dust discs using FIR/submm *Herschel* observations. We considered 16 large (main sample) and 13 less spatially resolved (additional sample) edge-on spiral galaxies in each of the five 100–500  $\mu\text{m}$  *Herschel* bands where possible. We carefully examined the background in the images to ensure that it is flat and does not affect the results of our modelling. We used both Sérsic and 3D (broken) exponential disc profiles to fit the galaxy images. For galaxies with bright sources within the plane,

they were simultaneously fitted along with the target galaxy to minimize the influence of the light from such sources on the galaxy fitting parameters. This analysis enabled us to directly study not only the radial distribution of dust, but also the distribution that extends out perpendicular to the galaxy mid-plane. The cumulative galaxy profiles with the superimposed models are presented in Fig. A1 and Table A1 (both published in their entirety as online material). We also used *WISE* 3.4  $\mu\text{m}$  imaging of the galaxies to retrieve the structural parameters of their old stellar distribution. We can summarize our results as follows.

- (i) The obtained Sérsic and 3D (broken) exponential disc models describe the observed emission profiles equally well, albeit the advantage of using a 3D disc over a Sérsic model is in retrieving the scale height, an independent parameter of the 3D disc geometry.
- (ii) We find that the effective radius of the dust emission increases with wavelength, whereas the apparent flattening of the Sérsic model remains almost constant (Fig. 6, two upper panels). We also point out an increase of both the disc scale length and scale height with wavelength (Fig. 6, two bottom panels). The first effect is related to the dust heating by the diffuse ISRF (coupled with the fact that dust discs are more widespread than their stellar counterparts), the second one – with a dramatic drop in dust temperature within the disc scale height and/or the presence of unresolved clumps or other unresolved global structural components (ring, superthin disc, spiral arms) in the FIR images.
- (iii) For 6 (and possibly additional four) of the 16 large galaxies in our sample, we clearly detect a second vertically extended component which can be interpreted as a thick dust disc or a halo (see Fig. A1). These galaxies are NGC 0891, NGC 3628, NGC 4013, ESO 209-009, NGC 4437, NGC 4631, and NGC 4244, ESO 373-008, NGC 5529, NGC 5746 with a possible extraplanar emission. In the additional sample, we identify six galaxies (IC 2233, NGC 3592, UGC 07321, UGC 07387, UGC 07522, and UGC 09242) demonstrating an excess of the FIR emission.
- (iv) Similar to the stellar discs, the dust discs grow proportionally in the vertical and radial directions (Fig. 9).
- (v) We found moderate positive correlations between the disc scale lengths in the NIR and FIR and between the scale heights in the same domains (Fig. 7), confirming that the dust mass distribution correlates

with the total stellar distribution (Alton et al. 1998a; Muñoz-Mateos et al. 2009; Hughes et al. 2014; Casasola et al. 2017), probably through the population of AGB stars.

(vi) Based on the dust emission in the FIR, we found an anticorrelation between the intrinsic flattening (relative thickness) of the dust disc and the dust mass (Fig. 11). Also, we report that the dust disc flatness and the ratio between the dust disc scale height and the stellar disc scale height anticorrelate with the maximum galaxy rotation velocity (Figs 12 and 13, respectively): spiral galaxies with higher dust (and total) masses possess flatter dust discs, whereas low-mass spiral galaxies host puffed-up diffuse dust discs distributed throughout the stellar disc. This is a confirmation of the findings by Dalcanton et al. (2004) and Holwerda et al. (2019), but is seen for the first time in dust emission.

In our next articles in this series dedicated to the study of the dust distribution in edge-on galaxies, we will carry out a detailed analysis of the structure for the most resolved galaxies in our sample which, as we showed in this paper, demonstrate an extraplanar dust emission above a single exponential disc. Also, it is particularly important to study how contemporary dust RT models, which are built upon the observed dust attenuation in edge-on galaxies in the UV, optical, and NIR, compare with the results of this study. For example, the well-known dust-energy balance problem, which states that RT models underestimate the observed FIR emission by a factor of two to four (see e.g. Popescu et al. 2000b; De Geyter et al. 2015; Mosenkov et al. 2018), has not been entirely solved.

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*Herschel* is an ESA space observatory with science instruments provided by European-led Principal Investigator consortia and with important participation from NASA. The *Herschel* spacecraft was designed, built, tested, and launched under a contract to ESA managed by the *Herschel/Planck* Project team by an industrial consortium under the overall responsibility of the prime contractor Thales Alenia Space (Cannes), and including Astrium (Friedrichshafen) responsible for the payload module and for system testing at spacecraft level, Thales Alenia Space (Turin) responsible for the service module, and Astrium (Toulouse) responsible for the telescope, with an excess of a hundred subcontractors.

## DATA AVAILABILITY

The data underlying this article are available in the article and in its online supplementary material.

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## SUPPORTING INFORMATION

Supplementary data are available at [MNRAS](https://academic.oup.com/mnras/article/515/4/5698/6652508) online.

**Figure A1.** Cumulative radial ( $r$ ) and vertical ( $v$ ) profiles of galaxies from the main and additional samples for the PACS 100 waveband.

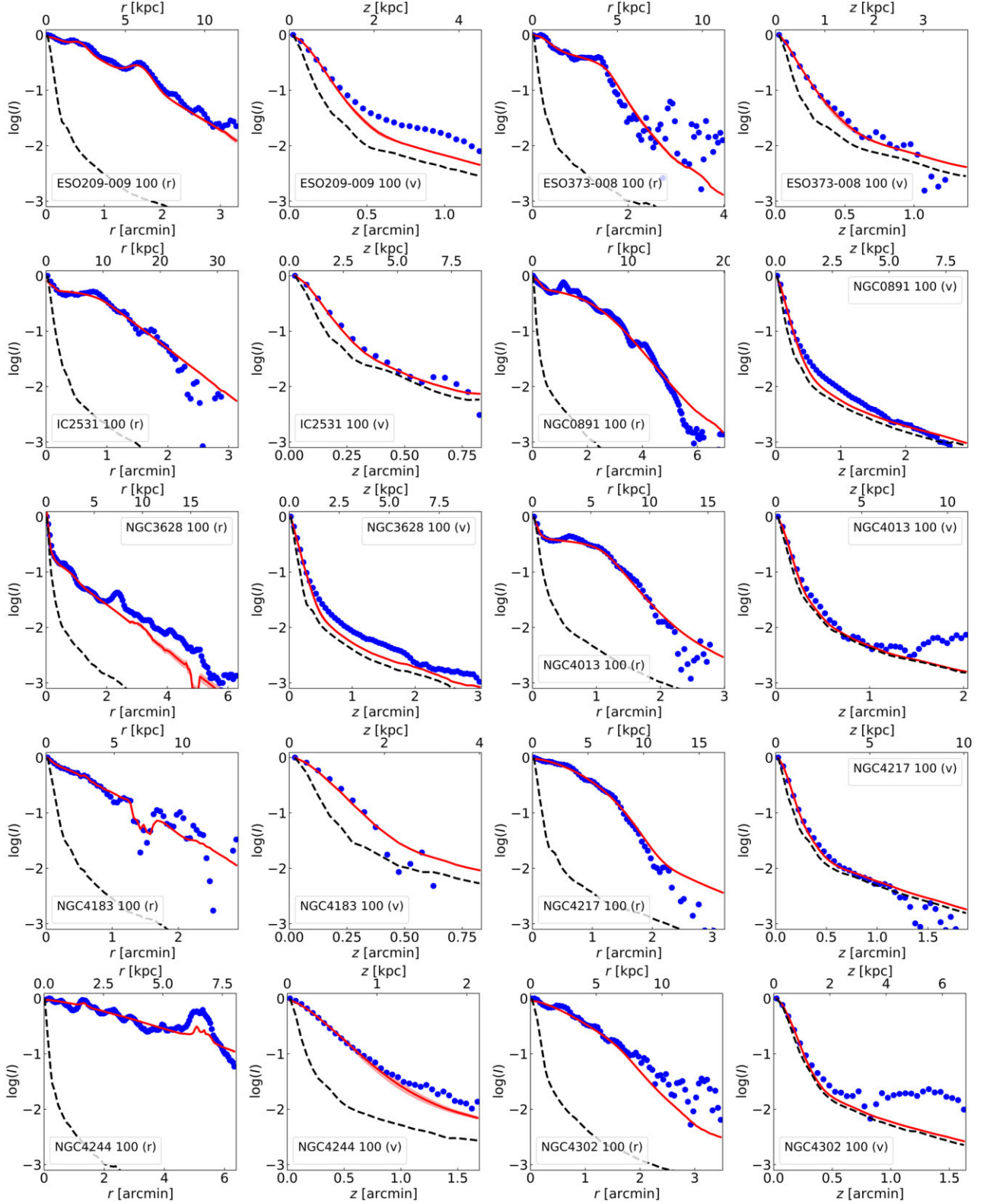
**Table A1.** The results of our fitting for galaxies from the main and additional samples.

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## APPENDIX A: RESULTS OF THE FITTING

**Table A1.** The results of our fitting for galaxies from the main and additional samples. The results for the Sérsic modeling at  $3.4 \mu\text{m}$  are taken from [M19](#). \* means that this parameter can be unreliable due to the poor vertical resolution.  $\infty$  denotes that the inner disc profile is a plateau with a very large inner disc scale length  $h_{\text{in}}$ . The complete table is available online.

| Galaxy     | Band | $n$             | $r_e$<br>(kpc)  | $q_0^{\text{serisc}}$ | $h_{\text{in}}$<br>(kpc) | $h_{\text{out}}$<br>(kpc) | $R_b$<br>(kpc) | $h_z$<br>(kpc)  | $q_0^{\text{disc}}$ |
|------------|------|-----------------|-----------------|-----------------------|--------------------------|---------------------------|----------------|-----------------|---------------------|
| ESO209-009 | 3.4  | 1.38            | 3.66            | $0.13 \pm 0.02$       | $1.89 \pm 0.07$          | —                         | —              | $0.22 \pm 0.02$ | —                   |
|            | 100  | 1.00            | $3.45 \pm 0.01$ | $0.11 \pm 0.02$       | $2.00 \pm 0.05$          | —                         | —              | $0.27 \pm 0.07$ | $0.12 \pm 0.03$     |
|            | 160  | 1.00            | $3.63 \pm 0.02$ | $0.12 \pm 0.01$       | $2.21 \pm 0.03$          | —                         | —              | $0.31 \pm 0.08$ | $0.12 \pm 0.03$     |
|            | 250  | $1.00 \pm 0.08$ | $4.12 \pm 0.02$ | $0.13 \pm 0.01$       | $2.42 \pm 0.02$          | —                         | —              | $0.34 \pm 0.05$ | $0.12 \pm 0.02$     |
|            | 350  | 1.00            | $4.67 \pm 0.03$ | $0.13 \pm 0.01$       | —                        | —                         | —              | —               | —                   |
| ...        | ...  | ...             | ...             | ...                   | ...                      | ...                       | ...            | ...             | ...                 |



**Figure A1.** Cumulative radial ( $r$ ) and vertical ( $v$ ) profiles of galaxies from the main and additional samples for the PACS 100 waveband. In each plot, the blue dots represent the galaxy profile, the red solid line – its 3D disc model, the reddish spread shows a  $1\sigma$  spread of the model profiles for varied inclination angles (if the spread is not seen, its width is less than the width of the red solid line), and the black dashed line – the corresponding LSF profile. All profiles are normalized by the maximum of the galaxy profile. Profiles for the other *Herschel* wavebands and for the Sérsic modelling are available online.

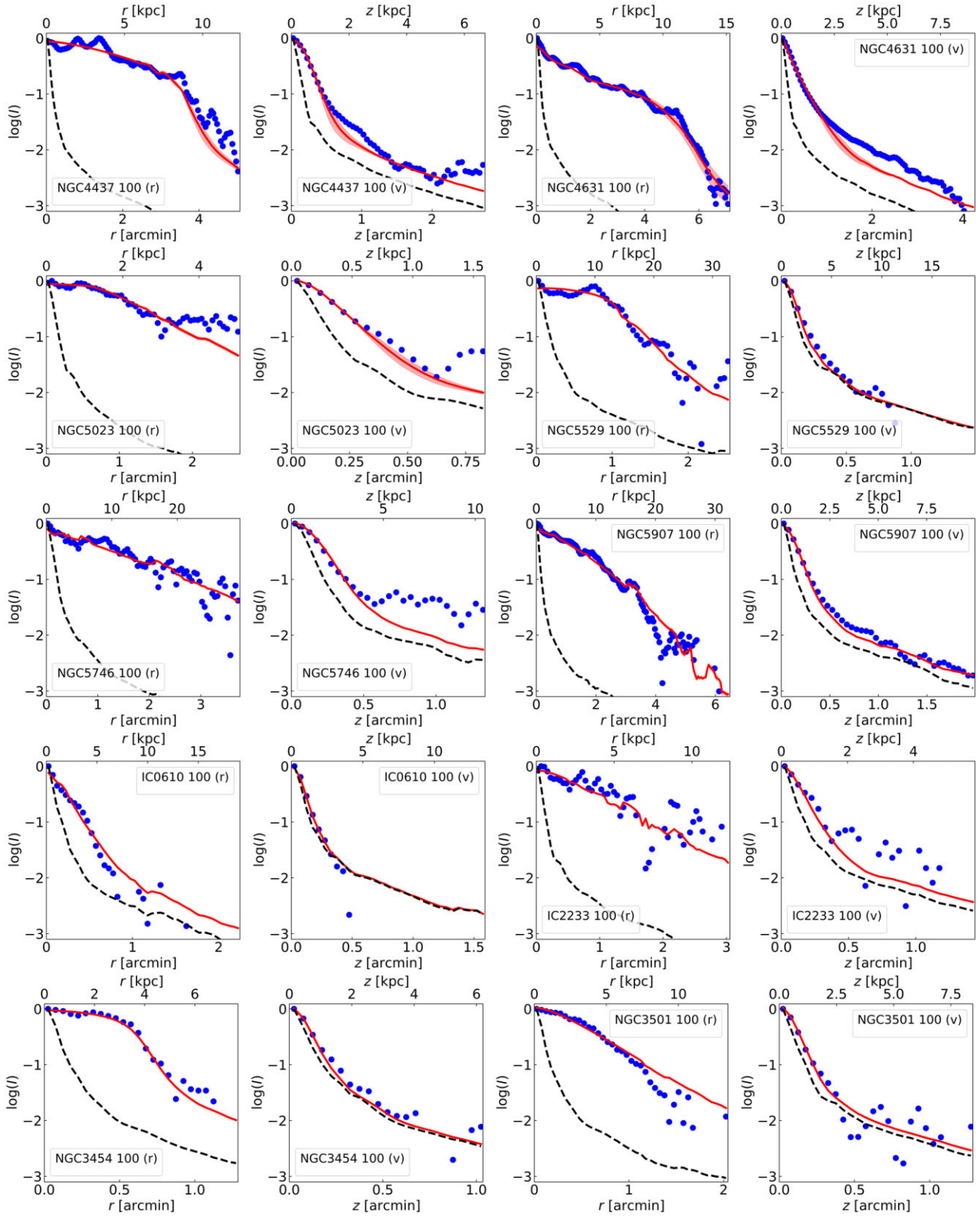


Figure A1 – continued

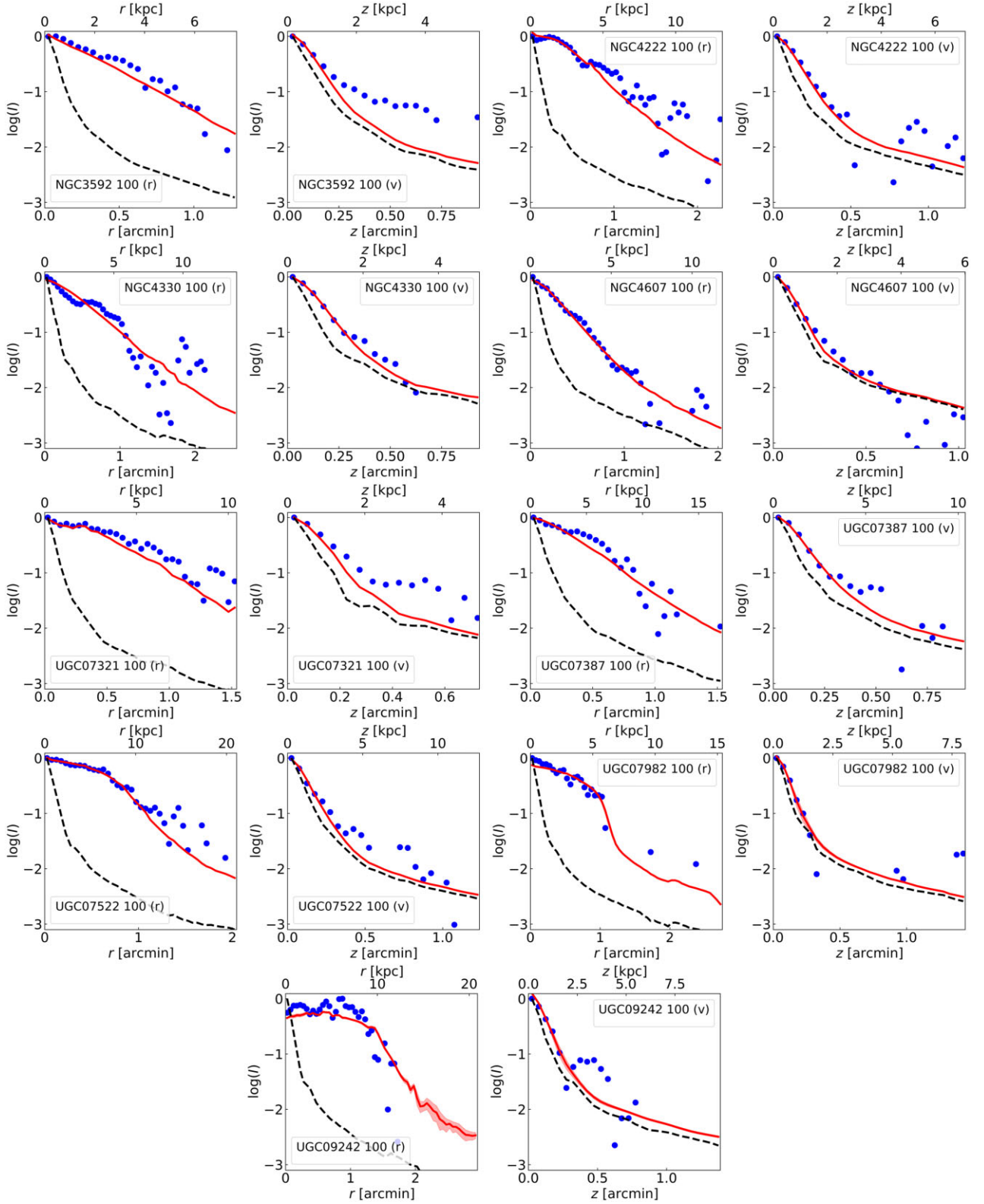


Figure A1 – continued

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