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Authors	Lusso, E., Risaliti, G., NARDINI, EMANUELE, Bargiacchi, G., Benetti, M., Bisogni, Susanna, Capozziello, S., Civano, F., Eggleston, L., Elvis, M., Fabbiano, G., GILLI, Roberto, Marconi, Alessandro, Paolillo, Maurizio, Piedipalumbo, E., SALVESTRINI, Francesco, Signorini, M., Vignali, C.
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Quasars as standard candles

III. Validation of a new sample for cosmological studies

E. Lusso^{1,2*}, G. Risaliti^{1,2}, E. Nardini^{1,2}, G. Bargiacchi¹, M. Benetti^{3,4,5}, S. Bisogni⁶, S. Capozziello^{3,4,5,7,8}, F. Civano⁹, L. Eggleston¹⁰, M. Elvis⁹, G. Fabbiano⁹, R. Gilli¹¹, A. Marconi^{1,2}, M. Paolillo^{3,4,13}, E. Piedipalumbo^{3,4}, F. Salvestrini^{12,11}, M. Signorini¹ and C. Vignali^{12,11}

¹Dipartimento di Fisica e Astronomia, Università di Firenze, via G. Sansone 1, 50019 Sesto Fiorentino, Firenze, Italy

²INAF – Osservatorio Astrofisico di Arcetri, 50125 Florence, Italy

³Dipartimento di Fisica, Università degli studi di Napoli Federico II, via Cinthia, 80126 Napoli, Italy

⁴INFN – Sezione di Napoli, via Cinthia 9, 80126 Napoli, Italy

⁵Scuola Superiore Meridionale, Università di Napoli Federico II, Largo San Marcellino 10, 80138 Napoli, Italy

⁶INAF – Istituto di Astrofisica Spaziale e Fisica Cosmica Milano, via Corti 12, 20133 Milano, Italy

⁷Gran Sasso Science Institute, Via F. Crispi 7, I-67100, L'Aquila, Italy

⁸Laboratory for Theoretical Cosmology, Tomsk State University of Control Systems and Radioelectronics (TUSUR), 634050 Tomsk, Russia

⁹Center for Astrophysics | Harvard & Smithsonian, 60 Garden Street, Cambridge, MA 02138, USA

¹⁰Yale University, Department of Computer Science, 51 Prospect St, New Haven, CT 06511

¹¹INAF – Osservatorio di Astrofisica e Scienza dello Spazio di Bologna, via Gobetti 93/3 - 40129 Bologna - Italy

¹²Dipartimento di Astronomia, Università degli Studi di Bologna, via Gobetti 93/2, 40129 Bologna, Italy

¹³INAF – Osservatorio Astronomico di Capodimonte, via Moiariello 16, 80131 Napoli, Italy

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ABSTRACT

We present a new catalogue of $\sim 2,400$ optically selected quasars with spectroscopic redshifts and X-ray observations from either *Chandra* or *XMM-Newton*. The sample can be used to investigate the non-linear relation between the UV and X-ray luminosity of quasars, and to build a Hubble diagram up to redshift $z \sim 7.5$. We selected sources that are neither reddened by dust in the optical/UV nor obscured by gas in the X-rays, and whose X-ray fluxes are free from flux-limit related biases. After checking for any possible systematics, we confirm, in agreement with our previous works, that (i) the X-ray to UV relation provides distance estimates matching those from supernovae up to $z \sim 1.5$, and (ii) its slope shows no redshift evolution up to $z \sim 5$. We provide a full description of the methodology for testing cosmological models, further supporting a trend whereby the Hubble diagram of quasars is well reproduced by the standard flat Λ CDM model up to $z \sim 1.5$ – 2 , but strong deviations emerge at higher redshifts. Since we have minimized all non-negligible systematic effects, and proven the stability of the $L_X - L_{UV}$ relation at high redshifts, we conclude that an evolution of the expansion rate of the Universe should be considered as a possible explanation for the observed deviation, rather than some systematic (redshift-dependent) effect associated with high-redshift quasars.

Key words. quasars: general – quasars: supermassive black holes – galaxies: active

1. Introduction

Quasars are the most luminous and persistent energy sources in our Universe. As they can be observed up to redshift $z \approx 7.5$ (Bañados et al. 2018), when the age of the Universe was less than ≈ 700 million years, quasars bear an extraordinary potential as cosmological probes. Several techniques making use of empirical correlations between quasar properties have been proposed to determine cosmological parameters such as the dark matter (Ω_M) and dark energy (Ω_Λ) content of the Universe. Examples include the relation between the continuum luminosity and the emission-line equivalent width (Baldwin 1977), or with the radius of the quasar broad-line region established via reverberation mapping (Watson et al. 2011). Another luminosity distance estimator combines the correlation between the quasar X-ray variability amplitude and its black hole (BH) mass (La Franca et al.

2014). None the less, these correlations are affected by too large a dispersion (up to ≈ 0.6 dex), and are typically applicable over a limited redshift range. Other promising methods employ geometric distances through, again, the broad-line region radius via reverberation mapping (Elvis & Karovska 2002), the luminosity of super-Eddington accreting quasars (Wang et al. 2013), the eigenvector formalism in the quasar main sequence (Marziani & Sulentic 2014), or the combination of spectroastrometry and reverberation mapping (Wang et al. 2020). All these techniques still need some level of refinement, and/or much higher sample statistics, to be competitive as cosmological tools.

Since 2015, our group has been developing a new technique that hinges upon the observed non-linear relation between the ultraviolet (at $2500 \text{ \AA}$, L_{UV}) and the X-ray (at 2 keV , L_X) emission in quasars (e.g. Tananbaum et al. 1979, $L_X \propto L_{UV}^\gamma$, with $\gamma \approx 0.6$) to provide an independent measurement of their distances, thus turning quasars into *standardizable* candles and ex-

* e-mail: elisabeta.lusso@unifi.it

tending the distance modulus–redshift relation (or the so-called *Hubble-Lemaître diagram*) of supernovae Ia to a redshift range still poorly explored ($z > 2$; Risaliti & Lusso 2015).

The applicability of our technique is based upon the understanding that most of the observed dispersion in the $L_X - L_{UV}$ relation is not intrinsic to the relation itself but due to observational issues, such as X-ray absorption by gas, UV extinction by dust, calibration uncertainties in the X-rays (Lusso 2019), variability, and selection biases associated with the flux limits of the different samples. In fact, with an optimal selection of *clean* sources (i.e. where we can measure the *intrinsic* UV and X-ray quasar emission), the dispersion drops from 0.4 to ≈ 0.2 dex (Lusso & Risaliti 2016, 2017).

A key consequence of this technique is that the $L_X - L_{UV}$ relation must be the manifestation of a universal mechanism at work in the quasar engines, although the details on the physical process originating this relation are still unknown (e.g. Haardt & Maraschi 1991, 1993; Haardt et al. 1994; Nicastro 2000; Merloni 2003; Lusso & Risaliti 2017; Arcodia et al. 2019).

The main results of our work are that (1) the distance modulus–redshift relation of quasars at $z < 1.4$ is in agreement with that of supernovae Ia and with the concordance Λ CDM model (Risaliti & Lusso 2015, 2019; Lusso et al. 2019), yet (2) a deviation from Λ CDM emerges at higher redshift, with a statistical significance of about 4σ . If we interpret the latter result by considering an evolution of the dark energy equation of state in the form $w(z) = w_0 + w_a \times z/(1+z)$, the data suggest that the dark energy density is increasing with time (Risaliti & Lusso 2019; Lusso et al. 2019).

As our approach may still have shortcomings, we need to demonstrate that the observed deviation from Λ CDM at redshift > 2 is neither driven by systematics in the quasar sample selection nor by the procedure adopted to fit the quasar Hubble-Lemaître diagram. To build a quasar sample that can be utilised for cosmological purposes, both X-ray and UV data are required to cover the rest-frame 2 keV and 2500 Å. The choice of these monochromatic luminosities is rather arbitrary, and mostly based on historical reasons. It is possible that the $L_X - L_{UV}$ relation is tighter with a different choice of the indicators of UV and X-ray emission (see e.g. Young et al. 2010). A careful analysis of this issue may also provide new insights on the physical process responsible for this relation. In the present analysis we will not investigate the possible alternatives, but we will focus on demonstrating that the commonly used relation is calibrated in a robust way and can be safely used as a tool to determine quasar distances. In this third paper of our series, we thus mainly concentrate on the quasar sample selection, whilst we defer a detailed analysis of the cosmographic fitting technique we adopted in a forthcoming publication.

At the time of writing, the most extended spectroscopic coverage in the UV is given by the *Sloan Digital Sky Survey* (Pâris et al. 2018), which supplies more than $\sim 500,000$ quasars with spectroscopic redshift up to $z \sim 7$. This sample needs to be cross-matched with the current X-ray catalogues, namely the *Chandra X-ray Catalogue* (CXC2.0, Evans et al. 2010) and the 4XMM Data Release 9 (Webb et al. 2020), which contain all the X-ray sources detected by the *Chandra* and *XMM-Newton* observatories that are publicly available in the respective archives. These data need to be complemented by dedicated pointed observations of active galactic nuclei¹ (AGN) at both low ($z < 0.1$) and high ($z > 3$) redshifts to extend the coverage and increase the sample statistics in the distance modulus–redshift relation.

The main aims of this manuscript are to discuss in detail all the criteria required to select a homogeneous sample of quasars for cosmological purposes from the above archives, and to present the key steps in fitting the Hubble-Lemaître diagram that can be adopted to reproduce our results. Specifically, we will examine the effects on the sample selection and on the UV and X-ray flux measurements of (1) dust extinction and host-galaxy contamination, (2) gas absorption in the X-rays, and (3) Eddington bias. We will identify the quasars that can be used for testing cosmological models, and investigate in depth all the possible systematics in the quasar Hubble-Lemaître diagram as a function of the contaminants mentioned above.

Since a detailed spectroscopic UV and X-ray analysis can be carried out only for a relatively small number of sources, our latest quasar sample presented here still relies also on broadband photometry in both UV and X-rays, as most of the parameters currently employed in our works, e.g. monochromatic UV and X-ray fluxes, UV colours and X-ray slopes, are derived from the photometric spectral energy distribution (SED) of our sources. In the future, we plan to gradually move towards a full spectroscopic analysis, as spectroscopy can deliver cleaner measurements of the relevant parameters.

The paper is constructed as follows. In Section 2 we present the different data sets employed to build the main quasar sample and the procedures adopted to measure the UV fluxes and slopes from the photometry. Section 3 discusses how the photometric quasar SEDs are constructed. Section 4 is devoted to the presentation of our technique to compute the monochromatic X-ray emission and the photon index from the catalogued broadband flux values. In Section 5 we discuss the several quality filters employed to select a homogeneous sample of quasars for cosmological purposes, whilst in Section 6 we verify that the $L_X - L_{UV}$ relation for the final “best” quasar sample does not evolve with redshift. Section 7 presents the main steps adopted to fit the Hubble diagram using a model independent technique (i.e. cosmography), whilst in Section 8 we fit the Hubble diagram with the most commonly used Λ CDM extension to test our fitting technique and to verify how different choices regarding the fitting method and the quasar sample selection affect the final results. In Section 9 we carry out an in-depth investigation on possible remaining systematics in the residuals of the quasar Hubble diagram, as a function of the parameters involved in the selection of the sample. Finally, we summarise our work and main results in Section 10.

Although we mainly use fluxes, whenever luminosity values are reported we have assumed a standard flat Λ CDM cosmology with $\Omega_M = 0.3$ and $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$.

2. The data set

The broad-line quasar sample we considered for our analysis has been assembled by combining seven different samples from both the literature and the public archives. The former group includes the samples at $z \approx 3.0\text{--}3.3$ by Nardini et al. (2019), $4 < z < 7$ by Salvestrini et al. (2019), $z > 6$ by Vito et al. (2019), and the XMM–XXL North quasar sample published by Menzel et al. (2016). We then complemented this collection by including quasars from a cross-match of optical (i.e. the *Sloan Digital Sky Survey*) and X-ray public catalogues (i.e. *XMM-Newton* and *Chandra*), which we refer to as SDSS–4XMM and SDSS–*Chandra* samples hereafter. Finally, we also added a local subset of AGN with UV (i.e. *International Ultraviolet Explorer*) data and X-ray archival information. The same order in which these samples have been introduced above is adopted as an order

¹ In the following we will refer to AGN and quasars indistinctly.

Table 1. Summary of sample statistics.

Sample	Initial (1)	Main (2)	Selected (3)	Ref
<i>XMM-Newton</i> $z \approx 3$	29	29	14	1
<i>XMM-Newton</i> $z \approx 4$	1	1	1	2
High- z	64	64	35	3
XXL	840	840	106	4
SDSS-4XMM	13,800	9,252	1,644	5
SDSS- <i>Chandra</i>	7,036	2,392	608	6
Local AGN	15	15	13	7
Total	21,785	12,593	2,421	

Notes. (1) These number counts refer to the sample statistics before correcting for overlaps amongst the subsamples. (2) Sample statistics after accounting for overlaps and the quality pre-selection described in § 5. The order of priority decreases from the top to the bottom row. (3) Sample statistics in the final cleaned quasar sample after the various filtering steps: see § 5 for details. References for the various samples: 1: Nardini et al. (2019); 2: see § 2.5; 3: Salvestrini et al. (2019) and Vito et al. (2019); 4: Menzel et al. (2016); 5: see § 2.1; 6: Bisogni et al., to be submitted; 7: § 2.7.

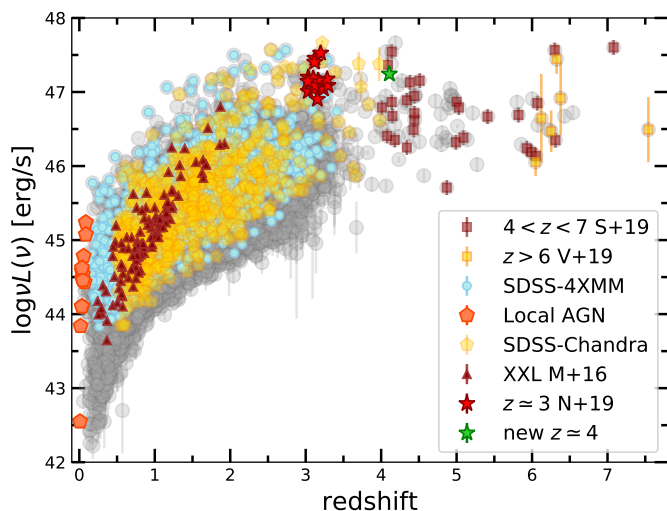


Fig. 1. Distribution of luminosities at rest-frame 2500 Å as a function of redshift for the main (grey points) and the selected (cleaned) samples (see § 5.4). Brown and yellow squares: high- z sample (Salvestrini et al. 2019; Vito et al. 2019, see § 2.6), cyan points: SDSS-4XMM (§ 2.1), brown triangles: XMM-XXL (Menzel et al. 2016, § 2.3), orange pentagons: the local AGN sample (§ 2.7), red stars: $z \approx 3$ quasar sample (Nardini et al. 2019), green star: new $z \approx 4$ quasar (§ 2.5), gold pentagons: SDSS-*Chandra* (§ 2.2).

of priority to take into account all the possible overlaps. X-ray fluxes coming from pointed observations and medium/deep surveys (i.e. XMM-XXL) have been considered first, as they are generally more reliable. The main parent sample is composed by $\sim 19,000$ objects spanning the redshift range $0.009 < z < 7.52$. In Figure 1 we present the distribution of rest-frame 2500 Å luminosities as a function of redshift for all the different quasar subsamples. A summary of the sample statistics is provided in Table 1, whilst below we describe in detail of how each subsample has been constructed.

2.1. The SDSS-4XMM sample

The bulk of the data belongs to the *Sloan Digital Sky Survey* quasar catalogue, Data Release 14 (Pâris et al. 2018; SDSS-DR14 hereafter). The catalogue contains 526,356 optically selected quasars detected over 9376 deg² with robust identification and spectroscopic redshift. Firstly, we removed from the catalogue all quasars flagged as broad absorption lines (BALs, where sources with the BI_CIV=0 flag are non-BALs) and kept all the objects with a measurement in all the SDSS magnitudes. This preliminary selection leads to 503,744 quasars.

We note that the BAL classification in the SDSS-DR14 quasar catalogue is based on a fully automated detection procedure on C IV absorption troughs for sources at $z > 1.57$. Hence, a number of BAL quasars might still be included in this preliminary sample. BAL quasars are often found in galaxies with red optical/UV colours and hard X-ray spectra (e.g. Gallagher et al. 2006), the latter suggesting that their relative X-ray weakness could be primarily due to gas absorption. The selection criteria discussed in Section 5 efficiently remove red/X-ray absorbed quasars, possibly excluding most unclassified BALs from the final sample.

The photometric rest-frame spectral energy distribution (SED), whose derivation is discussed in Section 3, is then used to define the parameters required to exclude radio-loud, dust-absorbed or host-galaxy contaminated sources. Also the rest-frame monochromatic luminosities at 2500 Å are obtained from the photometric SEDs. For comparison purposes, as in our previous works on this topic, we discard bright radio quasars through the same radio loudness parameter, R , as that used in Shen et al. (2011), which is defined as the ratio of the rest-frame fluxes at 6 cm and 2500 Å (i.e. $R = L_{\nu,6\text{cm}}/L_{\nu,2500\text{Å}}$). A quasar is then classified as radio-loud if $R > 10$. We computed R for the 17,561 objects with a FIRST detection, 16,315 of which are indeed radio-loud and have been therefore excluded from our sample.

To further remove powerful radio-loud quasars we considered the catalogue published by Mingo et al. (2016), which is currently the largest available Mid-Infrared (WISE), X-ray (3XMM) and Radio (FIRST+NVSS) collection (MIXR) of AGN and star-forming galaxies: 2,753 sources, 918 of which are considered radio-loud based on multiwavelength diagnostics (we refer to their paper for details). We excluded 349 quasars in our sample flagged as radio-loud in the MIXR catalogue within a matching radius of 3 arcsec. This yields 487,080 SDSS radio-quiet quasars with a L_{UV} measurement.

This SDSS quasar sample is then cross-matched with the latest *XMM-Newton* source catalogue 4XMM-DR9 (Webb et al. 2020). 4XMM-DR9 is the fourth generation catalogue of serendipitous X-ray sources, which contains 810,795 detections (550,124 unique X-ray sources) made publicly available by 2018 December 18². The net sky area covered (taking into account overlaps between observations) is ~ 1152 deg², for a net exposure time of ≥ 1 ksec.

To select reliable X-ray detections, we have applied the following quality cuts in the 4XMM-DR9 catalogue: SUM_FLAG<3 (low level of spurious detections), OBJ_CLASS≤3 (quality classification of the whole observation³) and EP_TIME>0 (EPIC exposure time available). These filters lead to 692,815 X-ray detections. We have adopted a maximum separation of 3 arcsec to provide optical classification

² http://xmmssc.irap.omp.eu/Catalogue/4XMM-DR9/4XMM-DR9_Catalogue_User_Guide.html

³ For more details the reader should refer to the 4XMM catalogue user guide.

and spectroscopic redshift for all the cross-matched objects. This yields 22,196 *XMM-Newton* observations: 13,858 unique sources (3,976 of which have ≥ 2 observations) covering the redshift range $0.056 < z < 4.306$.

Following the results presented by Lusso & Risaliti (2016, LR16 hereafter; see their Section 4), we decided to average all X-ray observations for sources with multiple detections that meet our selection cuts, including that associated with the Eddington bias (see § 5.3 for details). In this way, we reduce the effect of X-ray variability on the dispersion (~ 0.12 dex, see § 4 in LR16) by using only unbiased detections.

For each *XMM-Newton* observation, we have computed the EPIC sensitivity (5σ minimum detectable flux) at 2 keV following a similar approach as in LR16. We first estimated the minimum detectable flux in the soft band for both pn and MOS as a function of the on-time⁴ exposure following the relations plotted in Figure 3 by Watson et al. (2001). The total MOS on-time exposure is the one with the largest exposure value between the two individual cameras, MOS1 and MOS2. We then corrected this sensitivity for the pn and MOS, using the same vignetting correction for both cameras at the energy of 1.5 keV, as a function of their respective off-axis angles, where the smaller value between the two individual cameras is again assumed for the MOS. The sensitivity at 2 keV ($F_{\min}$) is then estimated for both pn and MOS, assuming a photon index of 1.7, following the same approach as in LR16. Finally, we have prioritized the pn sensitivity flux values over the MOS when available.

2.2. The SDSS–Chandra sample

To further increase the statistics, we also cross-matched the SDSS–DR14 quasar catalogue with the second release of the *Chandra Source Catalog* (CSC2.0). The CSC2.0⁵ (Evans et al. 2010) contains $\sim 315,000$ X-ray sources observed in 10,382 *Chandra* ACIS and HRC-I imaging observations publicly released prior to 2015. A cross-match of these two catalogues, with a matching radius of 3 arcsec, leads to 7,036 unique objects. The detailed analysis of this sample will be presented in a forthcoming publication (Bisogni et al., to be submitted). Briefly, from this sample we excluded radio-loud and BAL quasars following the same approach adopted with the SDSS–4XMM sample. SEDs were also compiled for all the quasars (see § 3), which were then used to estimate both L_{UV} and optical colours and thus select objects with low levels of dust reddening and host-galaxy contamination.

CSC2.0 provides photometric information and data products for each source, already reduced and ready to use for spectroscopic analysis⁶. We selected all the AGN with at least one measure of the flux in the soft band and with an off-axis angle < 10 arcmin (3,569 quasar observations, 2,392 single quasars). We performed a full spectral analysis with the *XSPEC* v.12.10.1b X-ray spectral fitting package (Arnaud 1996). For each observation, we assumed a model consisting of a power law with Galactic absorption, as provided by the CSC 2.0 catalogue at the source location. The spectral analysis provides us with the rest-frame flux at 2 keV and its uncertainty. Overall, the *Chandra* data have a reasonable signal-to-noise ratio ($S/N > 5$ in the soft band) that ensures uncertainties on F_X on the order of 0.15 dex or better.

⁴ The total good (after flares removal) exposure time (in seconds) of the CCD over which the source is detected.

⁵ <http://cxc.harvard.edu/csc2/>

⁶ All the X-ray info can be downloaded from the CSCview application <http://cda.harvard.edu/cscview/>

Flux limits are estimated for any given *Chandra* observation by computing the percentage of net counts to deduce the significance of the source detection, and a factor that takes into account the level of background, P_{bkg} . The 0.5–2 keV and 2–7 keV fluxes are then multiplied by P_{bkg} to obtain an approximate value of the background flux in these energy bands. The flux limit in each energy band is then estimated from the background flux by assuming a minimum signal-to-noise ratio of 3. The flux limit at 2 keV is finally inferred by interpolation (or extrapolation) of the band flux limits, depending on the redshift of the source.

2.3. The XMM–XXL sample

We also considered the AGN sample published by Menzel et al. (2016) from the equatorial subregion of the *XMM-Newton* XXL survey (XMM–XXL, PI: Pierre), i.e. XMM–XXL North (in the following we will refer to this sample as XXL for simplicity), which overlaps with the SDSS–DR8 imaging survey. XMM–XXL North is a medium-depth (10 ks per pointing) X-ray survey distributed around the area of the original 11 deg² XMM–LSS survey. The total catalogue contains 2,570 X-ray AGN with optical counterparts, spectroscopic redshifts and emission lines information. From the main sample, we considered only the AGN classified as (point-like) optically unobscured (flagged as BLAGN1; 1,353 sources). To have consistent measurements of optical/UV luminosities and redshifts amongst the different samples, we cross-matched the XXL BLAGN1 with the SDSS–DR14 quasar catalogue (with 3 arcsec matching radius) finding 1,067 objects. We have then included only the AGN with available SDSS photometry and classified as non-BAL, leading to 915 AGN. For this sample, we compiled the photometric SEDs following the same approach as in Section 3, and computed luminosities at various rest-frame wavelengths (e.g. 2500 Å, 1450 Å, 6 cm), optical/UV colours (Γ_1 , Γ_2 , see § 5.1 for details) and radio loudness. The latter parameter further excludes 75 AGN, leading to a final sample of 840 sources.

2.4. The $z \sim 3$ quasar sample

We included a sample of 29 bright ($L_{\text{bol}} > 10^{47}$ erg s^{−1}) quasars at $z \simeq 3$ with X-ray observations obtained from an extensive campaign performed with *XMM-Newton* (cycle 16, proposal ID: 080395, PI: Risaliti). This campaign targeted 30 quasars⁷ in the $z = 3.0$ – 3.3 redshift range for a total exposure of 1.13 Ms. This sample, selected in the optical from the SDSS Data Release 7 to be representative of the most luminous, intrinsically blue quasar population, boasts by construction a remarkable degree of homogeneity in terms of optical/UV properties. The X-ray data have been extensively analysed: the interested reader should refer to Nardini et al. (2019) for details.

2.5. New $z \sim 4$ quasar

We also included one new optically-selected SDSS quasar at $z = 4.109$, J074711.14+273903.3, whose X-ray observation was obtained as part of a proposed large programme with *XMM-Newton* (cycle 18, proposal ID: 084497, PI: Lusso). This is the only target actually observed from its parent sample, which consisted of 19 quasars in the $z \simeq 4$ redshift range for a total exposure of 1.34 Ms. The sample had been selected in the optical

⁷ One quasar in this sample turned out to be radio-loud, although not flagged as such in the SDSS–DR7 catalogue, so we exclude it from the present analysis.

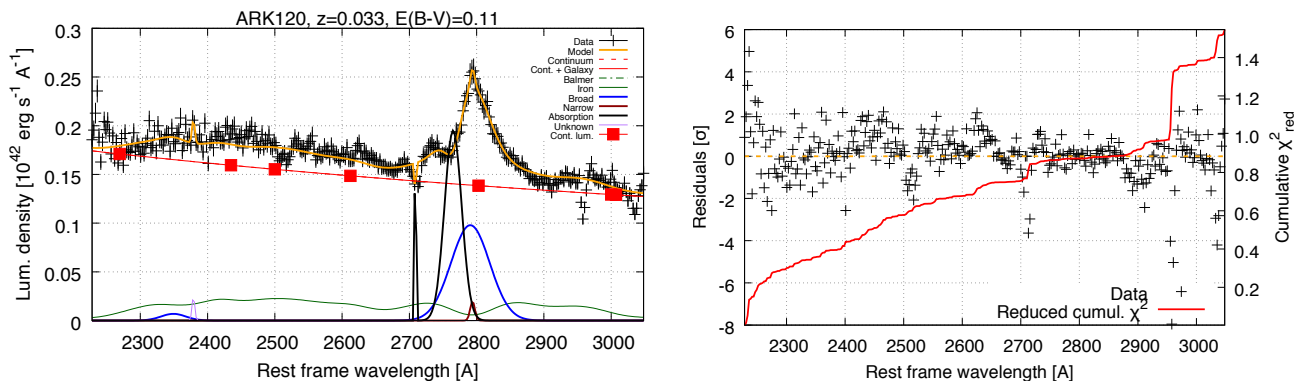


Fig. 2. Example of UV spectral fitting. Best-fit model (left panel, orange curve) and residuals (right panel, in units of 1σ uncertainties on the data and cumulative reduced χ^2) are presented for Ark 120. Redshift and Galactic extinction at the source location are shown on top of the plot. The main model components are plotted with different colours: the solid red line is the sum of continuum and host galaxy components; the black solid line is the absorption component; the dot-dashed green line is the Balmer component; the solid green line is the UV iron template; the sum of all broad and narrow emission-line components is shown with blue and brown lines, respectively. The red square symbols are the continuum luminosities estimated by QSFit.

from the SDSS Data Release 14 with the same criteria described in Section 2.4. The X-ray spectrum of this quasar has been analysed following the same procedure as presented by Nardini et al. (2019), and we decided to include this source in the current sample as a proof of feasibility for future campaigns.

2.6. The high redshift sample

To improve the coverage at high redshifts, we considered two additional samples of $z > 4$ quasars with pointed X-ray observations published by Salvestrini et al. (2019) and Vito et al. (2019).

The Salvestrini et al. (2019) quasar sample consists of 53 objects in the redshift range $4.01 < z < 7.08$, which benefit from a moderate-quality coverage in the UV and X-ray energy bands. Of the 53 quasars, 47 objects were observed with *Chandra* and 9 with *XMM-Newton*. The galaxies ULAS J1120+0641, SDSS J114816.7+525150.4 and SDSS 1030+0524 have been observed by both satellites. The authors performed a full X-ray spectral analysis of the archival data, we thus refer to their paper for details. The majority of the quasars in this sample (33 out of 53) have L_{UV} measurements from the SDSS-DR7 quasar catalogue. For the remaining quasars, the L_{UV} values are computed by extrapolating the UV spectra to longer wavelengths with a fixed continuum slope (see their Section 4 and Appendix B for further details).

Vito et al. (2019) published a sample of 25 quasars at $z > 6$ with either archival data (15 objects) or new *Chandra* observations (10 sources), which were selected to have virial black-hole mass estimates from Mg II line spectroscopy. All the X-ray data were reprocessed by the authors (see their Section 3.1), whilst the L_{UV} values were computed from the 1450 Å magnitude assuming a power-law spectrum ($F_\nu \propto \nu^{-\alpha}$) with $\alpha = -0.3$ (see their Section 4.1). We excluded from their sample 3 BAL candidates, 1 weak line quasar, all sources with an upper limit in α_{ox} (i.e. X-ray undetected) and all radio-loud sources, for a total of 9 quasars. For the remaining 16 sources, we found five overlaps with the Salvestrini et al. (2019) sample, so the final number of quasars included from Vito et al. (2019) is 11 sources. This sample also contains the highest-redshift quasar observed so far, i.e. ULAS J134208.10+092838.61 at $z = 7.54$ (Bañados et al. 2018).

2.7. The local quasar sample

To anchor the normalization of the quasar Hubble diagram with Type Ia supernovae, we need to extend the coverage at very low redshifts ($0.009 < z < 0.1$). We searched for all the local AGN with ultraviolet data from the *International Ultraviolet Explorer* (IUE) in the Mikulski Archive for Space Telescopes (MAST). We chose to use the reduced spectra from the long-wavelength prime (LWP) camera of IUE, which spans the wavelength interval 1845–2980 Å, thus always covering the rest-frame 2500 Å at the redshifts of interest. We then considered all AGN with X-ray data available in the *XMM-Newton* archive or the with X-ray flux values in the literature, finding 17 objects, 11 of which with ≥ 2 UV spectra (although the majority consists of consecutive observations). In this sample, NGC 1566 and NGC 7603 are well known highly variable/changing look sources, so we excluded them from the starting sample. Multiple UV spectra for the remaining AGN have been stacked, verifying that the inclusion of non consecutive observations does not change the final composite for each AGN.

We then carried out a detailed spectral fitting of all the UV spectra using the publicly available QSFit package (Calderone et al. 2017). We modelled each spectrum as follows: the Mg II emission line is reproduced by a combination of a broad (with a full-width at half-maximum, FWHM, larger than 2000 km/s) and a narrow (FWHM < 2000 km/s) profile, whilst the continuum includes the contributions from the host galaxy, the iron complex, the Balmer continuum and the AGN continuum. Spectra are also corrected for Galactic extinction using the parametrization by Cardelli et al. (1989) and O'Donnell (1994), with a total selective extinction $A(V)/E(B-V) = 3.1$ (Calderone et al. 2017). The rest-frame 2500 Å luminosity is finally measured from the AGN continuum component only. An example of a UV spectral fit on one of the objects in the local AGN sample is shown in Figure 2.

The X-ray information (soft and hard fluxes, photon index) has been taken from the literature. Most of the sources in the local sample have been drawn from the CAIXA catalogue, which consists of radio-quiet, X-ray unobscured ($N_H < 2 \times 10^{22} \text{ cm}^{-2}$) AGN observed by *XMM-Newton* in targeted observations (Bianchi et al. 2009). For two AGN, ESO 141-G055 and IRAS 09149-6206, the X-ray fluxes and Γ_X values are given by de Marco et al. (2009) and Ricci et al. (2017), respectively. The

Table 2. Properties of the local AGN sample.

Name	z	$\log F_{\text{UV}}$	$\log F_{\text{X}}$	Ref.
Ark 120	0.0327	-24.88 ± 0.01	-28.37 ± 0.002	1
Mrk 841	0.0364	-25.46 ± 0.01	-29.15 ± 0.006	1
NGC 4593	0.0090	-25.79 ± 0.01	-28.38 ± 0.002	1
HE 1029–1401	0.0858	-25.10 ± 0.01	-28.77 ± 0.005	1
ESO 141–G055	0.0371	-25.10 ± 0.01	-28.43 ± 0.005	2
IRAS 09149–6206	0.0573	-25.19 ± 0.01	-28.86 ± 0.011	3
HE 1143–181	0.0329	-25.37 ± 0.01	-28.57 ± 0.005	1
NGC 7469	0.0163	-25.02 ± 0.01	-28.55 ± 0.003	1
Mrk 205	0.0708	-25.73 ± 0.01	-29.33 ± 0.016	1
Mrk 926	0.0469	-25.25 ± 0.02	-28.54 ± 0.007	1
Fairall 9	0.0470	-25.34 ± 0.01	-29.01 ± 0.016	1
Mrk 1383	0.0866	-25.28 ± 0.01	-29.02 ± 0.017	1
Mrk 509	0.0344	-24.89 ± 0.01	-28.43 ± 0.003	1
Mrk 478	0.0791	-25.52 ± 0.01	-29.51 ± 0.023	1
Mrk 352	0.0149	-26.84 ± 0.01	-29.03 ± 0.011	1

Notes. UV and X-ray fluxes are in units of $\log(\text{erg s}^{-1}\text{cm}^{-2}\text{Hz}^{-1})$. References for the X-ray data: 1: [Bianchi et al. \(2009\)](#); 2: [de Marco et al. \(2009\)](#); 3: [Ricci et al. \(2017\)](#).

rest-frame 2 keV monochromatic flux is then estimated following the procedure described in Section 4.

A summary of the properties of the local AGN sample is provided in Table 2.

3. The quasar SED compilation

To compile the quasar SEDs, we used all the multicolour information as reported in the SDSS–DR14 catalogue. The catalogue includes multiwavelength data from radio to UV: the FIRST survey in the radio ([Becker et al. 1995](#)), the Wide-Field Infrared Survey (WISE, [Wright et al. 2010](#)) in the mid-infrared, the Two Micron All Sky Survey (2MASS, [Cutri et al. 2003](#); [Skrutskie et al. 2006](#)) and the UKIRT Infrared Deep Sky Survey (UKIDSS; [Lawrence et al. 2007](#)) in the near-infrared, and the Galaxy Evolution Explorer (GALEX, [Martin et al. 2005](#)) survey in the UV. Galactic reddening has been properly taken into account by utilising the selective attenuation of the stellar continuum $k(\lambda)$ from [Fitzpatrick \(1999\)](#), whilst Galactic extinction is estimated from [Schlegel et al. \(1998\)](#) for each object in the SDSS catalogue. For each source, we computed the observed flux and the corresponding frequency in all the available bands. The data used in the SED computation were blueshifted to the rest-frame, and no K-correction was applied. All the rest-frame luminosities were then determined from a first-order polynomial between two adjacent points. At wavelengths bluer than about 1400 Å, we expect significant absorption by the intergalactic medium (IGM) in the continuum ($\sim 10\%$ between the Ly α and C IV emission lines, see [Lusso et al. 2015](#) for details). Hence, when computing the relevant parameters, we excluded from the SED all the rest-frame data at $\lambda < 1500\text{Å}$. The rest-frame monochromatic luminosities are finally obtained by interpolation whenever the reference frequency is covered by the photometric SED. Otherwise, the value is extrapolated by considering the slope between the luminosity values at the closest frequencies. Thanks to this broad photometric coverage, we can compute the rest-frame luminosity at 2500 Å (L_{UV}) via interpolation for the majority of the SDSS quasars. Indeed, we were not able to estimate L_{UV} due to a sparse photometric SED coverage (i.e. when the SED is composed by a single rest-frame data point) for only 130 quasars.

Uncertainties on monochromatic luminosities ($L_{\nu} \propto \nu^{-\gamma}$) from the interpolation (extrapolation) between two values L_1 and L_2 are computed as:

$$\delta L = \sqrt{\left(\frac{\partial L}{\partial L_1}\right)^2 (\delta L_1)^2 + \left(\frac{\partial L}{\partial L_2}\right)^2 (\delta L_2)^2}. \quad (1)$$

Examples of photometric SEDs for two quasars at different redshifts in the SDSS–4XMM sample are shown in Figure 3. The red circles in the figure mark all the available photometry from the SDSS–DR14 catalogue, whilst the ones used to construct the SEDs are highlighted with black circles. The magenta squares represent the luminosities at 2500 Å and 1450 Å. The cyan and blue solid lines are the composite SDSS quasar spectrum from [Vanden Berk et al. \(2001\)](#) and the average SDSS quasar SED from [Krawczyk et al. \(2013\)](#), respectively. Both composites are shown for reference, for the AGN continuum plus line emission and continuum only, and are normalised to 2500 Å. The AGN SED in the left panel shows a case of both host-galaxy contamination and dust absorption. This source, indeed, does not fulfil our selection criteria (described in detail in § 5), as opposed to the object in the right panel, which represents an object with low levels of contaminants.

3.1. On the use of photometric rest-frame 2500 Å fluxes

In this work, we focus on rest-frame 2500 Å monochromatic fluxes as derived from photometry for two main reasons. The first one is based on the physics of the $L_{\text{X}} - L_{\text{UV}}$ relation, whilst the other on the fact that broadband photometry allows us to build much bigger samples over a larger redshift and luminosity range than spectroscopy alone.

Concerning the former, the 2500 Å monochromatic flux has been adopted since the first studies on the topic, yet its choice was mainly based upon observational considerations. Indeed, the rest-frame 2500 Å at $z \approx 1$ corresponds to the observed V band, which was available for a significant number of sources. Additionally, this rest-frame UV wavelength is less affected by host-galaxy contamination (dominant at low luminosity for $\lambda > 4000\text{Å}$) and intergalactic absorption ($\lambda < 1450\text{Å}$) than other regions

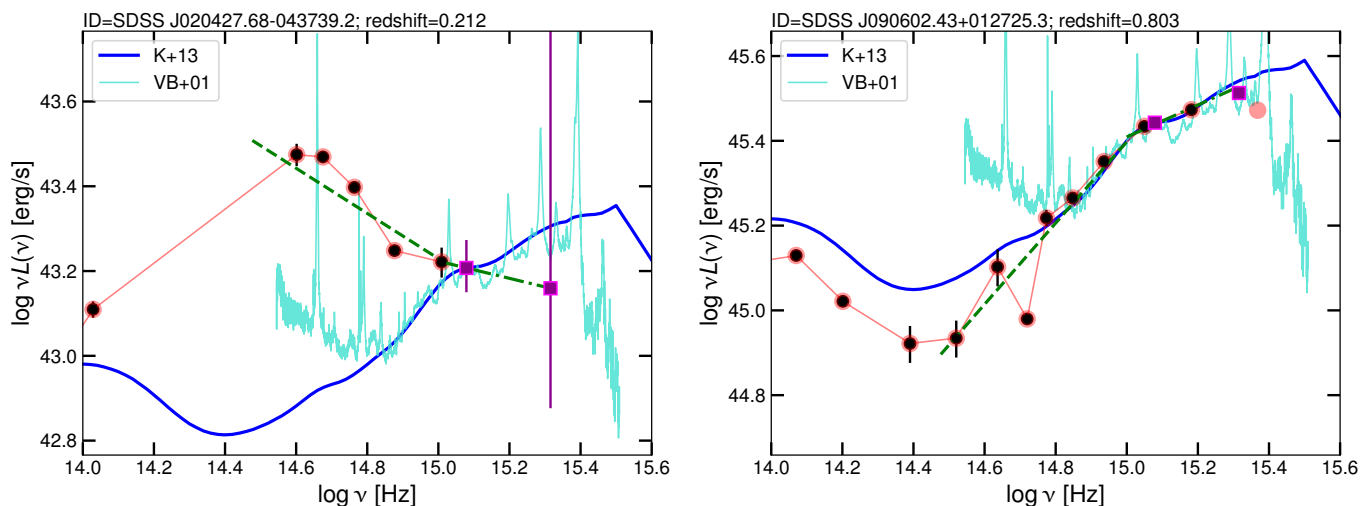


Fig. 3. Examples of AGN SEDs. The red circles mark all the available photometry from the SDSS catalogue, whilst the ones used to construct the SEDs are highlighted with black circles. The magenta squares represent the luminosities at 2500 Å and 1450 Å. The cyan line is the composite SDSS quasar spectrum from [Vanden Berk et al. \(2001\)](#), whilst the blue solid line is the average SDSS quasar SED from [Krawczyk et al. \(2013\)](#). Both composites are normalised to 2500 Å. The green dashed and dot-dashed lines represent the two near-infrared/optical slopes Γ_1 and Γ_2 in the 0.3–1 μm and 1450–3000 Å range (rest frame), respectively. The SED in the left panel is representative of an AGN that does not fulfil our selection criteria (see § 5), as it is affected by both host-galaxy contamination and dust absorption in the UV. The SED in the right panel, instead, is characteristic of an object with low levels of contaminants, which therefore belongs to the clean AGN sample.

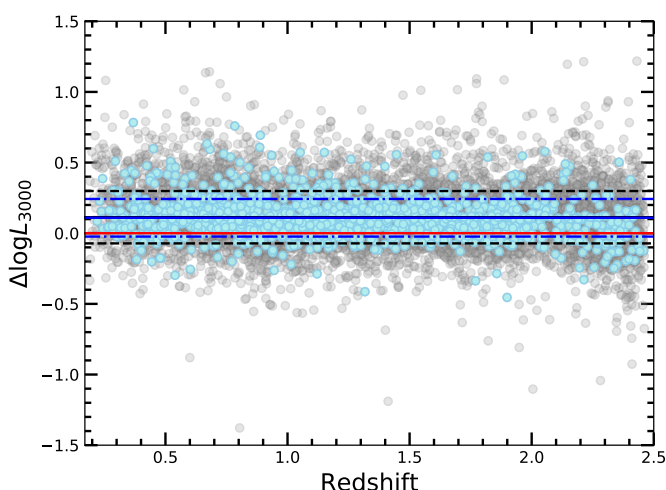


Fig. 4. Comparison between the photometric SDSS luminosity values at 3000 Å and the spectroscopic ones (i.e. $\Delta \log L_{3000} = \log L_{3000, \text{photo}} - \log L_{3000, \text{spectro}}$) as a function of redshift for all the quasars in our sample with available 3000 Å monochromatic luminosities from [Rakshit et al. \(2020\)](#). Grey and cyan points represent the initial ($\sim 12,000$) and selected ($\sim 1,500$) quasar samples from SDSS-4XMM, respectively. The black and blue lines mark the mean (solid) and 1σ dispersion (dashed and dot-dashed) of $\Delta \log L_{3000}$ for the initial and selected samples, respectively (black and blue solid lines overlap at $\langle \Delta \log L_{3000} \rangle = 0.11$). The red solid line is the $\Delta \log L_{3000} = 0$.

of the AGN SED, thus representing the ideal proxy of the intrinsic disc emission. This notwithstanding, the photometric 2500 Å flux might be contaminated by strong Fe II line emission (see Figure 2), which can introduce systematic uncertainties on the photometrically derived F_{UV} values (e.g. [Netzer 2019](#)).

[Rakshit et al. \(2020\)](#) recently published spectroscopic measurements for more than 500,000 quasars selected from the SDSS-DR14 quasar catalog. They performed a homogeneous

analysis of the SDSS spectra to estimate the continuum and line properties (e.g. H α , H β , Mg II, C IV, and Ly α) of these sources. This catalogue also provides a measurement of the 3000 Å luminosity, which is the closest wavelength to the one of our interest. We have therefore estimated the 3000 Å monochromatic luminosities from the photometric SEDs for all the SDSS-4XMM quasars in the initial sample of 13,800 quasars, similarly to what we have done at 2500 Å. Figure 4 shows the comparison between the photometric and the spectroscopic SDSS luminosity values (i.e. $\Delta \log L_{3000} = \log L_{3000, \text{photo}} - \log L_{3000, \text{spectro}}$) as a function of redshift for the objects within the SDSS-4XMM sample with a good quality 3000 Å monochromatic luminosity value (QUALITY_L3000=0) available from spectroscopy, i.e. $\approx 12,000$ quasars (where $z \approx 2.5$ represents the higher redshift for which the rest-frame 3000 Å is covered by the BOSS spectrograph). The $\Delta \log L_{3000}$ distribution shows a mean at 0.1 dex, a dispersion around the mean of 0.18 dex, and no trend with redshift. Although a systematic offset in the $\Delta \log L_{3000}$ measurements is expected, as our L_{3000} could be contaminated by the Fe II emission, this is reassuringly small (only a flux factor of ~ 1.3). We also note that, any redshift independent offset in the UV fluxes would not be an issue, since it is balanced out in the cross-calibration between the Hubble diagram of supernovae and quasars.

Figure 4 also shows the $\Delta \log L_{3000}$ distribution as a function of redshift for the 1,473 quasars in the clean SDSS-4XMM sample (1,644 total sources, see Section 5 for details) with a measurement of L_{3000} . The average $\Delta \log L_{3000}$ is perfectly consistent with the initial sample, the dispersion around the mean is 0.13 dex, and again there is no trend with redshift, implying that our selection criteria have the only effect of singling out the most reliable luminosity measurements as proxy of the nuclear emission.

We finally note that, regardless of the details of the physical mechanism driving the $L_X - L_{\text{UV}}$ relation, the characteristic UV flux wavelength should be the one most closely related to the global emission of accretion disc, and might not be precisely

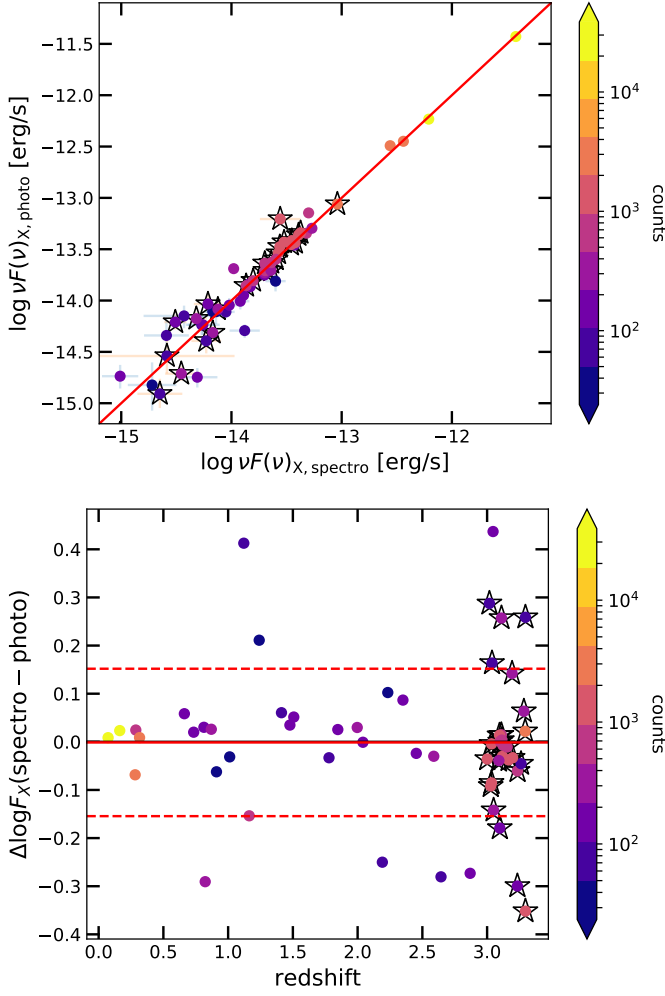


Fig. 5. Top panel: comparison between the spectroscopic and photometric 2 keV monochromatic fluxes for 30 quasars randomly extracted from the SDSS-4XMM sample and the quasars of the $z \approx 3$ sample (Nardini et al. 2019, marked with open star symbols). Bottom panel: difference between the spectroscopic and photometric 2 keV monochromatic fluxes as a function of redshift. The mean $\Delta \log F_X = \log F_{X,\text{spectro}} - \log F_{X,\text{photo}}$ value and its 1σ dispersion are shown with the red and dashed lines, respectively. Points are colour-coded by the number of net counts. The $\Delta \log F_X$ distribution is scattered around $\Delta \log F_X \sim 0$ with no clear trend with redshift.

the rest-frame 2500 Å (see Section 6 in Risaliti & Lusso 2019). It is even possible that nuclear emission should be combined with other AGN spectral properties (e.g. emission-line FWHM, continuum slope; see Lusso & Risaliti 2017). We are currently exploring these possibilities, and results will be published in a forthcoming work.

Since we are still far from grasping the nature of the $L_X - L_{UV}$ relation, our photometric fluxes are seen as a more conservative representation of the broadband disc emission, capturing the “true” dependence between disc and X-ray corona in AGN in a statistical sense.

4. The rest-frame 2 keV monochromatic flux

Given the large source statistics in the SDSS-4XMM and XXL samples, a detailed X-ray spectral analysis of all the quasars is impractical. Therefore, to compute the rest-frame 2 keV monochromatic flux (F_X), we follow the same approach as de-

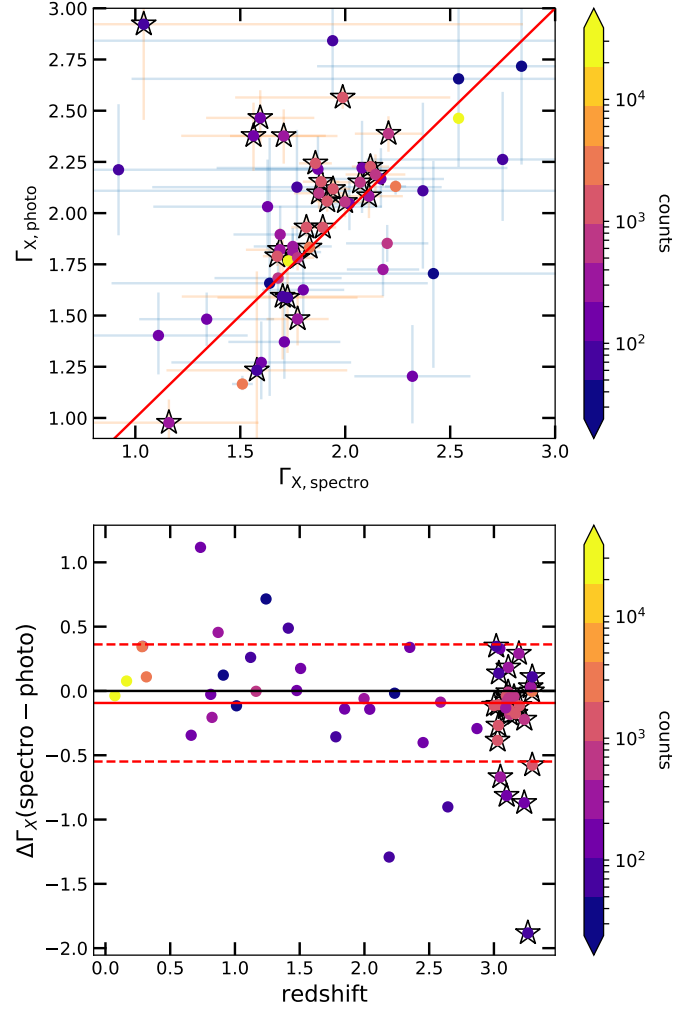


Fig. 6. Top panel: comparison between the spectroscopic and photometric photon indices for 30 quasars randomly extracted from the SDSS-4XMM sample and the quasars of the $z \approx 3$ sample (Nardini et al. 2019, marked with open star symbols). Bottom panel: difference between the spectroscopic and photometric photon index values as a function of redshift. The mean $\Delta \Gamma_X = \Gamma_{X,\text{spectro}} - \Gamma_{X,\text{photo}}$ value and its 1σ dispersion are shown with the red and dashed lines, respectively. Points are colour-coded by the number of net counts. Although the Γ_X distribution along the one-to-one relation is rather scattered, the $\Delta \Gamma_X$ does not seem to show a clear trend with redshift.

scribed in Risaliti & Lusso (2019). For the SDSS-4XMM sample, we derived the rest-frame 2 keV fluxes and the relative (photometric) photon indices, Γ_X (along with their 1σ uncertainties), from the tabulated 0.5–2 keV (soft, F_S) and 2–12 keV (hard, F_H) fluxes reported in the 4XMM-DR9 serendipitous source catalogue. These band-integrated fluxes are blueshifted to the rest-frame by considering a pivot energy value of 1 keV (E_S) and 3.45 keV (E_H), respectively, and by assuming the same photon index used to derive the fluxes in the 4XMM catalogue (i.e. $\Gamma_X = 1.42$). For the soft band, the monochromatic flux at E_S is then:

$$F_E(E_S) = F_S \frac{(2 - \Gamma_X) E_S^{1-\Gamma_X}}{(2 \text{ keV})^{2-\Gamma_X} - (0.5 \text{ keV})^{2-\Gamma_X}}, \quad (2)$$

in units of $\text{erg s}^{-1} \text{ cm}^{-2} \text{ keV}^{-1}$. An equivalent expression holds for the hard band, with the obvious modifications. Flux values are corrected for Galactic absorption.

The photometric photon index is then estimated from the slope of the power law connecting the two soft and hard monochromatic fluxes at the rest-frame energies corresponding to the observed pivot points. The rest-frame photometric 2 keV flux (and its uncertainty) is interpolated (or extrapolated) based on such a power law.

To justify the employed soft and hard pivot energy values and to ensure that our photometric F_X values are accurate, we performed on the one hand several simulations, and on the other hand full X-ray spectral fitting of a number of random objects at different redshifts.

Regarding the former approach, we simulated a high-quality power-law spectrum, assuming both a typical average background and calibrations for *XMM-Newton*, with the same photon index assumed by the 4XMM-DR9 catalogue. We fitted the data in the soft band with a power law parametrized as $F(E) = F(E_0)(E/E_0)^{-\Gamma_X}$, with $F(E_0)$ and Γ_X as free parameters, with E_0 ranging from 0.5 to 1.5 keV in steps of 0.05 keV. In each case, we derived the $F(E_0) - \Gamma_X$ confidence contours. In general, the covariance between these two parameters is non-zero (i.e. the contours are elongated and tilted), so we explored which value of energy E_S returns covariance zero between $F(E_S)$ and Γ_X (i.e. the contours become 2D Gaussians). This *pivot energy* represents the energy value dividing the soft band in two regions having the same statistical weight. As such, this value is not located at exactly the centre of the energy band because of the dependence of the effective area on energy.

As a result of the $F(E_S) - \Gamma_X$ zero-covariance, our photometric F_X values are independent of the specific Γ_X assumed in the 4XMM-DR9 catalogue. This also implies that, even if our photometric Γ_X deviates from the *true intrinsic* value, the resulting F_X will be accurate in any case. Finally, the relative error on the monochromatic flux, $\Delta F(E)/F(E)$, at the pivot energy is the same as the one of the band flux, whereby the absolute value of $\Delta F(E)$ at the pivot energy is the smallest possible.

The same procedure is also applied to the XXL sample using their catalogued soft and hard band fluxes.

In parallel, we performed a full spectral analysis on a number of random objects. The top panels in Figures 5 and 6 present the comparison between the inferred spectroscopic and photometric F_X and Γ_X values, respectively, for 30 random quasars in the SDSS-4XMM sample. We also considered for this comparison the 27 sources in the $z \simeq 3$ quasar sample (Nardini et al. 2019) that have an entry in the 4XMM-DR9 catalogue⁸. The points are colour-coded by their number of net counts. Whilst the values of Γ_X display a large scatter (up to ~ 0.46 dex), our photometric F_X values are in remarkable agreement with the spectroscopic ones (with a scatter of just ~ 0.15 dex). The most obvious outlier in the bottom panel of Figure 6 is J1425+54, a marginally detected $z \simeq 3$ quasar with 22 ± 13 net counts in the pn (see Table 1 in Nardini et al. 2019; for the same camera, 4XMM-DR9 gives a consistent number of counts). The observed soft flux reported in the 4XMM-DR9 catalogue for this object is $(2.0 \pm 0.5) \times 10^{-15}$ erg s⁻¹ cm⁻², whilst it is virtually undetected in the observed hard band, with a S/N of ~ 0.5 and $F_H = (9 \pm 20) \times 10^{-16}$ erg s⁻¹ cm⁻². None the less, even with a nominally large discrepancy between the spectroscopic and photometric Γ_X values⁹, the F_X estimates are well within a factor of 1.2.

⁸ Out of the 30 objects in the $z \simeq 3$ sample, one had no public data on 2018 December 18, and two (J0945+23 and J1159+31; see Section 4.3 in Nardini et al. 2019) are not detected in 4XMM-DR9.

⁹ Interestingly, J1425+54 would not have met the selection criterion on Γ_X in either case (see § 5.2).

Overall, we have a consistency within a flux factor of 1.6 for about 80% of the sample (only 12/57 quasars lie outside $|\Delta \log F_X| = 0.2$) and, as expected, the higher the number of counts, the better the agreement, with the most deviant points having less than 100 counts¹⁰.

We have also investigated whether the difference between spectroscopic and photometric fluxes ($\Delta F_X = F_{X,\text{spectro}} - F_{X,\text{photo}}$) and photon indices ($\Delta \Gamma_X = \Gamma_{X,\text{spectro}} - \Gamma_{X,\text{photo}}$) displays any trend with redshift. The bottom panels of Figures 5 and 6 show such distributions and, despite the limited statistics, both ΔF_X and $\Delta \Gamma_X$ are scattered around zero with no clear trend.

The rather poor comparison between $\Gamma_{X,\text{spectro}}$ and $\Gamma_{X,\text{photo}}$ shown in Figure 6 might cast some doubts on the reliability of the photon indices derived from the broadband (soft and hard) fluxes. However, we believe that our technique of computing photometric Γ_X values can be safely employed for large sample of quasars and that it provides robust results, for the following reasons: (1) the spectroscopic and photometric X-ray fluxes are in very good agreement, meaning that our distance measures are not strongly affected by the use of photometric Γ_X values; and (2) we performed a series of checks by varying the photometric Γ_X range employed to select the final sample, finding that our main results are not significantly modified (see § 5.2 for further details). Summarizing, our $\Gamma_{X,\text{photo}}$ may not be correct on an object-by-object basis, but they are reliable in a *statistical sense* for large enough samples.

4.1. X-ray non-detected quasars

Quasar samples that include X-ray non-detections are likely to be unbiased, but the analysis of both the $L_{\text{UV}} - L_X$ and the distance modulus–redshift relations is far from straightforward, since it strongly depends on the weights assumed in the fitting algorithm. In the case of flux-limited surveys, objects with an expected emission (based on the observed $L_{\text{UV}} - L_X$ relation) close to the flux limit will be observed only in case of positive fluctuations, and this effect is likely redshift-dependent (see § 5.3). Considering only detections might thus introduce some bias in the $L_{\text{UV}} - L_X$ relation, and this should be more relevant to the X-rays, since the relative observed flux interval is narrower than in the UV.

Lusso & Risaliti (2016) investigated the effect of the inclusion of X-ray non-detections in the study of the $L_X - L_{\text{UV}}$ relation for an optically selected sample of quasars, whose selection was very similar to the one employed in the present analysis. Their main conclusion was that there were no statistically significant variations on slope, intercept and dispersion (within their uncertainties) between X-ray detected and censored quasar samples across the different selection steps, with the slope being rather constant around 0.6. Further analysis was performed in Risaliti & Lusso (2019) (see Section 3 of their Supplementary Material), where the fraction of X-ray non-detected quasars was on the order of 2% in their final cleaned sample. Such a fraction of censored data has negligible statistical weight in the fitting procedure, so their inclusion does not change the results of the statistical analysis.

Here, we have adopted a similar strategy as the one presented by Risaliti & Lusso (2019) to obtain a sample where biases are minimised even without the inclusion of non-detections, which is discussed at length in Section 5.3. Additionally, we explored whether any possible remaining bias in our X-ray de-

¹⁰ Note that this is the typical threshold for spectral analysis to return reasonably accurate results (see e.g. Nardini et al. 2019).

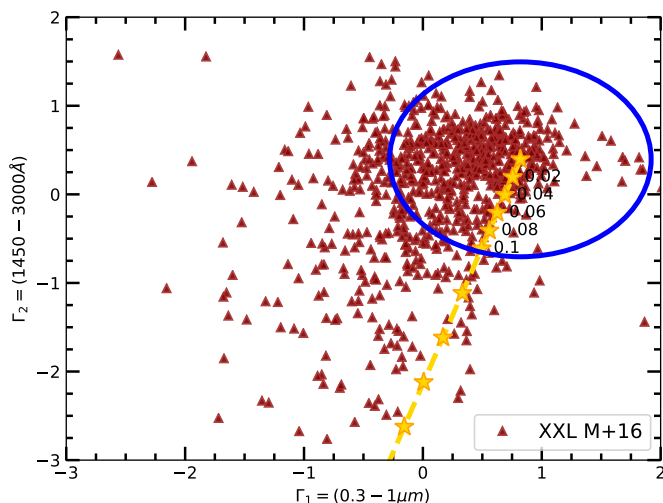


Fig. 7. Example of the $\Gamma_1 - \Gamma_2$ distribution for the XXL quasar sample, where Γ_1 and Γ_2 are the slopes of a power law in the $\log(\nu) - \log(\nu L_\nu)$ plane in the $0.3-1 \mu\text{m}$ and $1450-3000 \text{ \AA}$ intervals, respectively (see § 5.1). The stars represent the $\Gamma_1 - \Gamma_2$ values of the quasar SED by Richards et al. (2006) with increasing dust reddening (following the extinction law of Prevot et al. 1984), with $E(B - V)$ in the range $0-0.3$. We selected all the quasars inside the blue circle (i.e., with minimum host-galaxy and dust reddening contamination).

tected quasar sample is present in the residuals of the quasar Hubble diagram (see § 9.1). All these checks motivated us to analyse the Hubble diagram where non-detections are neglected.

5. Selection of a clean quasar sample

Our aim is to select a subsample with accurate estimates of L_{UV} and L_{X} , covering a redshift range as wide as possible, by removing systematic effects and low-quality measurements. For the latter, we applied a couple of preliminary filters that ensure good measurement quality. These filters mainly involve the X-ray data, since these are affected by larger uncertainties. Specifically, we considered only soft and hard flux measurements with a relative error smaller than 1 (i.e. a minimum S/N of 1 on both band fluxes): $\Delta F_{\text{S}}/F_{\text{S}} < 1$ and $\Delta F_{\text{H}}/F_{\text{H}} < 1$. A similar filter is currently not applied to UV fluxes since the S/N at these wavelengths is typically much higher than 1. Overall, these two filters exclude about 30% of the X-ray detections in the initial sample.

The main possible sources of contamination/systematic error are: dust reddening and host-galaxy contamination in the optical/UV, gas absorption in the X-rays, and *Eddington bias* associated with the flux limit of the X-ray observations. Here we briefly discuss each of these effects, and describe the filters we applied to obtain the final ‘best’ sample for a cosmological analysis.

5.1. Dust reddening and host-galaxy contamination

To retain the quasars with minimum levels of dust reddening and host galaxy contamination, we follow a similar approach to the one presented in our previous works (Risaliti & Lusso 2015; Lusso & Risaliti 2016; Risaliti & Lusso 2019). We used the rest-frame photometric SEDs discussed in Section 3 to compute, for each object, the slope Γ_1 of a $\log(\nu) - \log(\nu L_\nu)$ power law in the rest frame $0.3-1 \mu\text{m}$ range, and the analogous slope Γ_2 in the $1450-3000 \text{ \AA}$ range (see also Hao et al. 2013). Figure 3 shows two examples, where the green dashed and dot-dashed lines rep-

resent the $0.3-1 \mu\text{m}$ and $1450-3000 \text{ \AA}$ (rest frame) slopes Γ_1 and Γ_2 , respectively. The wavelength intervals for these slopes are chosen based on the fact that the SED of an *intrinsically* blue quasar is very different from the one of an inactive galaxy or a dust-reddened source. The intrinsic SED of a quasar presents a dip around $1 \mu\text{m}$, where the galaxy has the peak of the emission from the passive stellar population (e.g. Elvis et al. 1994; Richards et al. 2006; Elvis et al. 2012; Krawczyk et al. 2013). Dust reddening is wavelength dependent and the UV portion of the quasar SED will be attenuated differentially. These two concurrent factors impact on the quasar SED shape, allowing us to define a set of slopes that single out the majority of quasars with minimum levels of both host-galaxy emission and dust reddening (see Figure 1 in Hao et al. 2013).

The $\Gamma_1 - \Gamma_2$ distribution for the XXL subset of quasars is shown, as an example, in Figure 7. We assumed a standard SMC extinction law $k(\lambda)$ after Prevot et al. (1984), with $R_V = 3.1$ (as appropriate for unobscured AGN; Hopkins et al. 2004; Salvato et al. 2009), to estimate the $\Gamma_1 - \Gamma_2$ correlation as a function of extinction, parametrised by the colour excess $E(B - V)$. We obtained the red dashed line shown in Figure 7, where the starting point corresponds to the SED of Richards et al. (2006, i.e. $\Gamma_1 = 0.82$, $\Gamma_2 = 0.40$) with zero extinction. The distribution of $\Gamma_1 - \Gamma_2$ towards low values along the red dashed line is indicative of possible dust reddening, whilst sources towards more negative Γ_1 values are objects with possible host-galaxy contamination. The $\Gamma_1 - \Gamma_2$ plane is also very useful to identify unusual SEDs or SEDs characterised by bad photometry, which are then excluded from the sample.

We selected all the sources in the (Γ_1, Γ_2) plane within a circle centred at the reference values for a standard quasar SED (see Risaliti & Lusso 2015; Lusso & Risaliti 2016; Risaliti & Lusso 2019 for further details), with a radius corresponding to a reddening $E(B - V) \simeq 0.1$.

We note that our quasar selection based on photometry could still be affected by some contamination from the light of the host, especially in low redshift ($z \lesssim 0.7$) AGN, whose flux values at $2500 \text{ \AA}$ are located at the edge of the SDSS photometric coverage. Hence, any uncertainties in the estimate of the quasar UV continuum from the optical can make the $2500 \text{ \AA}$ monochromatic fluxes less reliable and possibly overestimated. Low-redshift AGN are on average less luminous in the optical/UV, with $L_{\text{bol}} \lesssim 10^{44} \text{ erg s}^{-1}$, thus the contrast between nuclear continuum and host-galaxy emission is smaller with respect to higher luminosity objects. Moreover, the data quality of low-redshift/low-luminosity AGN is, on average, lower. Host-galaxy contamination can be minimised through a source-by-source spectral fitting, as we did for the local AGN sample, but this procedure is rather time consuming for samples of several hundred thousands of objects. We will further discuss possible issues for cosmology related to our selection in Section 8.

5.2. X-ray absorption

Since X-ray fluxes may contain some level of absorption, which is naturally heavier in the soft band, we included only X-ray detections with a photon index Γ_{X} that falls within a range representative of unobscured quasars. For the majority of the sample, we adopted the following selection criterion, which also takes into account the uncertainties on Γ_{X} , i.e. $\Gamma_{\text{X}} - \delta\Gamma_{\text{X}} \geq \Gamma_{\text{X,min}}$ and $\Gamma_{\text{X}} \leq \Gamma_{\text{X,max}}$. The values of $\Gamma_{\text{X,min}}$ and $\Gamma_{\text{X,max}}$ are chosen based on two considerations: the average Γ_{X} within that interval should roughly correspond to $\Gamma_{\text{X}} \sim 2$ with a dispersion of $0.2-0.3$ (con-

sistent with e.g. [Young et al. 2009a](#)), and the $L_X - L_{UV}$ relation should not present any systematic deviation from the assumed *true* slope of 0.6 (within uncertainties).

We thus proceeded as follows. We evaluated the $F_X - F_{UV}$ relation in narrow redshift bins (so the effect of cosmology is negligible) for different choices of $\Gamma_{X,min}$ and $\Gamma_{X,max}$. For the SDSS-4XMM sample, we started by assuming a reasonable $\Gamma_{X,max}$ of 2.8 and a varying $\Gamma_{X,min}$ in the interval 1.4–1.9 with steps of 0.1, and converged to a $\Gamma_{X,min} = 1.7$. We checked that a smaller value of $\Gamma_{X,min}$ (i.e. 1.6) would not change the results of our analysis, but we prefer to be conservative, even at the expenses of sample statistics. For the SDSS-4XMM sample, we selected only X-ray observations with a photon index satisfying the condition $\Gamma_X - \delta\Gamma_X > 1.7$, and excluded the (few) objects with $\Gamma_X > 2.8$. The latter filter on Γ_X is needed to avoid strong outliers ($\sim 5\%$) which may be due to observational issues such as incorrect background subtraction in one of the two bands. This Γ_X interval roughly corresponds to an average $\Gamma_X \sim 2.1 - 2.2$ and a dispersion of 0.3. The same selection is applied to all the other subsamples at $z < 4$.

For the higher redshift ($z > 4$) sample, such a stringent criterion on Γ_X would exclude the majority of the objects, given their higher uncertainties. We thus decided to simply select all the objects with $\Gamma_X \geq 1.7$.

Given the observed Γ_X range (up to 2.8), some soft-excess (e.g. [Sobolewska & Done 2007](#); [Gliozzi & Williams 2020](#)) contribution for low- z quasars might be still present. We have thus repeated the analysis further imposing an upper limit to the Γ_X range of 2.5, but, besides losing statistics, our results are not affected.

5.3. Eddington bias

Owing to X-ray variability, AGN with an average X-ray intensity close to the flux limit of the observation will be observed only in case of a positive fluctuation. This introduces a systematic, redshift-dependent bias towards high fluxes, known as *Eddington bias*, which has the effect to flatten the $L_X - L_{UV}$ relation.

To reduce this bias, we excluded all X-ray detections below a threshold defined as κ times the intrinsic dispersion of the $L_X - L_{UV}$ relation (LR16; [Risaliti & Lusso 2019](#)), specifically:

$$\log F_{2\text{keV},\text{exp}} - \log F_{\min} < \kappa\delta, \quad (3)$$

where $F_{2\text{keV},\text{exp}}$ is the monochromatic flux at 2 keV expected from the observed rest-frame quasar flux at 2500 Å with the assumption of a *true* γ of 0.6, and it is calculated as follows:

$$\log F_{2\text{keV},\text{exp}} = (\gamma-1) \log(4\pi) + (2\gamma-2) \log d_L + \gamma \log F_{UV} + \beta, \quad (4)$$

where d_L is the luminosity distance calculated for each redshift with a fixed cosmology, and the parameter β represents the pivot point of the non-linear relation in luminosities, $\beta = 26.5 - 30.5\gamma \approx 8.2^{11}$. $F_{\min}$ in equation (3) is obtained as detailed in Sections 2.1 and 2.2, whilst the product $\kappa\delta$ is a value estimated for all the subsamples that we constructed from archives (SDSS-4XMM, SDSS-*Chandra*) or surveys (XXL).

Specifically, we first computed the flux limit of each X-ray observation, for both the SDSS-4XMM and SDSS-*Chandra* samples (see Sections 2.1 and 2.2). We minimised the Eddington bias by including only X-ray detections for which the minimum detectable flux $F_{\min}$ in that given observation is lower than

the expected X-ray flux $F_{2\text{keV},\text{exp}}$ by a factor that is proportional to the intrinsic dispersion in the $L_X - L_{UV}$ relation (we refer to Appendix A in LR16 and [Risaliti & Lusso 2019](#)). On average, the minimum detectable monochromatic fluxes at 2 keV are approximately $4.6 \times 10^{-32} \text{ erg s}^{-1} \text{ cm}^{-2} \text{ Hz}^{-1}$ and $3 \times 10^{-32} \text{ erg s}^{-1} \text{ cm}^{-2} \text{ Hz}^{-1}$ for the SDSS-4XMM and SDSS-*Chandra* samples, respectively. However, we caution that these values should not be considered as the “survey limiting fluxes”, since both the 4XMM and CSC2.0 catalogues are not proper flux-limited samples, but rather a collection of all X-ray observations performed over a certain period. It is thus not trivial to estimate the expected minimum flux for these catalogues. The XXL sample is, instead, a “standard” flux-limited sample, so we applied a soft-band flux threshold to the data ($F_S > 10^{15} \text{ erg s}^{-1} \text{ cm}^{-2}$), which corresponds to a flux limit at 2 keV of $5 \times 10^{-32} \text{ erg s}^{-1} \text{ cm}^{-2} \text{ Hz}^{-1}$. We considered $\kappa\delta = 0.9$ for SDSS-4XMM and XXL, whilst we used $\kappa\delta = 0.5$ for the SDSS-*Chandra* sample. All the other subsamples rely on pointed observations, so we did not apply any flux threshold to the data.

In principle, the effects of this bias could be further reduced if also non-detections were considered. Yet, this would not only complicate the statistical analysis, but also make the estimate of the intrinsic dispersion of the observed relations (e.g. $L_{UV} - L_X$, Hubble diagram) much more uncertain. Moreover, we have shown that there is no significant variation in both the slope and the intercept of the $L_{UV} - L_X$ correlation (within their uncertainties) among censored and detected samples once the Eddington bias is taken into account (see Appendix A in LR16). We therefore decided to include only detections in this work. This choice implies that we have to be very conservative in the correction for the Eddington bias, at the expense of sample statistics.

It is worth noting that our procedure to minimise the Eddington bias is slightly circular: we need the $L_X - L_{UV}$ relation (i.e. we assumed $\gamma = 0.6$) in order to estimate the ‘expected’ X-ray flux. Yet, our simulations show that we are able to retrieve the assumed cosmology (using different input values for Ω_M and Ω_Λ), when the selection criteria are applied to mock quasar samples.

5.4. The final cleaned sample

Summarizing, we applied a series of selection criteria to filter all the data that are likely contaminated by dust reddening, host-galaxy contamination, and X-ray absorption, or affected by the Eddington bias. We first selected all quasars within a circle centered at $(\Gamma_1, \Gamma_2) = (0.82, 0.4)$, i.e. $E(B-V) = 0$, and with a radius such as:

$$\sqrt{(\Gamma_1 - 0.82)^2 + (\Gamma_2 - 0.40)^2} \leq 1.1, \quad (5)$$

which corresponds to an $E(B-V) \lesssim 0.1$. The equation above filters out all quasar SEDs that show reddening in the UV, significant host-galaxy contamination in the near-infrared, as well as bad photometry (see § 5.1). We then applied an additional cut to keep only the X-ray observations where photon indices are indicative of low levels of X-ray absorption, and to exclude the X-ray data characterized by peculiar photon indices, especially at low/moderate redshifts (see § 5.2). Specifically, we required that:

$$\begin{cases} \Gamma_X + \delta\Gamma_X \geq 1.7 \text{ and } \Gamma_X \leq 2.8, & \text{if } z < 4 \\ \Gamma_X \geq 1.7, & \text{if } z \geq 4. \end{cases} \quad (6)$$

To correct for the Eddington bias, we further selected all observations that satisfy equation (3) where the product $\kappa\delta$ is 0.9

¹¹ The value of the luminosity normalizations are chosen based on the average values for the entire sample.

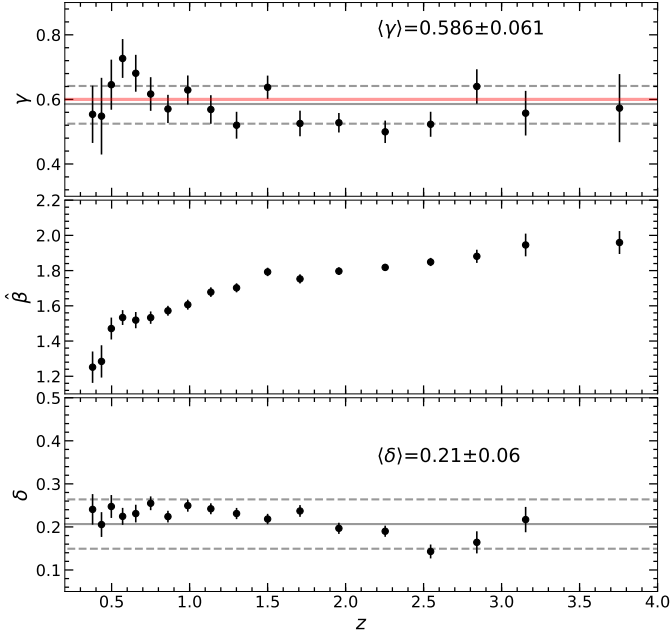


Fig. 8. Redshift evolution of the slope γ , the intercept $\hat{\beta}$ and the dispersion δ of the $F_X - F_{UV}$ relation. To perform the regression fit, X-ray and UV fluxes have been normalised to 10^{28} and 10^{30} erg s $^{-1}$ cm $^{-2}$ Hz $^{-1}$, respectively. The data points in each panel are averages in narrow redshift bins ($\Delta z \approx 0.06 - 0.45$). Error bars represent the 1σ uncertainty on the mean in each bin. The grey solid and dashed lines are the means and 1σ uncertainties, respectively, on the slope (γ) and on the dispersion (δ). The red line marks $\gamma = 0.6$.

for the SDSS–4XMM and XXL subsamples, and 0.5 for SDSS–*Chandra*. Pointed observations are available for the local, $z \approx 3$, and high-redshift samples (see § 5.3). For any quasar, all the multiple X-ray observations that survive the filters above are finally averaged to minimize the effects of X-ray variability (e.g. LR16, see also Lusso 2019).

The final cleaned sample is composed by 2,421 quasars spanning a redshift interval $0.009 \leq z \leq 7.52$, with a mean (median) redshift of 1.442 (1.295). Table 1 summarizes the statistics of each subsample, whilst a more detailed summary of the various subsamples after a given selection is provided in Table A.1. The main UV and X-ray properties of the final sample are presented in Table 3.

6. Analysis of the $F_X - F_{UV}$ relation with redshift

Before building the Hubble diagram, we need to check whether the $F_X - F_{UV}$ relation for the clean quasar sample shows any trend with redshift. We thus divided the sample in narrow redshift bins, with a variable step $\Delta z = 0.06 - 0.45$ within the redshift range 0.45–4 to have enough statistics. The redshift step is chosen to ensure that the dispersion in distances over each interval is smaller than the one of the relation in luminosities. In this way, we can consider fluxes as proxies of luminosities.

The best-fit parameters (slope, intercept and dispersion) of the $F_X - F_{UV}$ relation and their uncertainties are shown in Figure 8, whilst all the fits of the $F_X - F_{UV}$ relation in the chosen redshift bins are presented in Figure B.1. They are obtained through the Python package *emcee* (Foreman-Mackey et al. 2013), which is a pure-Python implementation of Goodman & Weare’s affine invariant Markov chain Monte Carlo (MCMC) ensemble sampler. To perform the regression fit, X-ray and UV fluxes were

normalised to 10^{28} and 10^{30} erg s $^{-1}$ cm $^{-2}$ Hz $^{-1}$, respectively. On average, the $F_X - F_{UV}$ slope does not show any clear trend with redshift within the analysed interval. Conversely, the trend of the intercept $\hat{\beta}$ of the normalized $F_X - F_{UV}$ relation observed in the middle panel of Figure 8 just reflects the overall shape of the quasar Hubble diagram (see Section 7). We note that, the trend of $\hat{\beta}$ with redshift, is not exactly the same as the one in Figure 9 since such a parameter does not have a simple direct proportional dependence on the distance modulus (equation 7) because of the different dependence between slope and normalization in each redshift bin.

The sample statistics is so sparse at redshift higher than 4 that we cannot provide a meaningful fit of the relation. None the less, we have checked that the data points at $z > 4$ do not show any trend with redshift in the residuals of the Hubble diagram (see Section 9). In fact, these data points are extremely useful to set the shape of the Hubble diagram, thus providing better constraints on the measurements of the expansion rate of the Universe. We thus confirm that the slope of the X-ray to UV relation shows no redshift evolution up to $z \sim 4$, in agreement with our previous works (e.g. Risaliti & Lusso 2015; Lusso & Risaliti 2016, 2017; Risaliti & Lusso 2019; Lusso et al. 2019; see also Salvatrin et al. 2019 for a high redshift analysis).

7. The quasar Hubble diagram

To fit the Hubble diagram we first need to derive the distance modulus for each object. We start by computing the luminosity distance (e.g. see Risaliti & Lusso 2015, 2019) as:

$$\log d_L = \frac{[\log F_X - \beta - \gamma(\log F_{UV} + 27.5)]}{2(\gamma - 1)} + \frac{1}{2} \log(4\pi) + 28.5, \quad (7)$$

where F_X and F_{UV} are the flux densities (in erg s $^{-1}$ cm $^{-2}$ Hz $^{-1}$). F_{UV} is normalised to the (logarithmic) value of 27.5 in the equation above, whilst d_L is in units of cm and is normalised to 28.5 (in logarithm). The slope of the $F_X - F_{UV}$ relation, γ , is a free parameter, and so is the intercept β ¹². The distance modulus, DM, is thus:

$$\text{DM} = 5 \log d_L - 5 \log(10 \text{ pc}), \quad (8)$$

and the uncertainty on DM, $d\text{DM}$, is:

$$d\text{DM} = \frac{5}{2(\gamma - 1)} \left[(d \log F_X)^2 + (\gamma d \log F_{UV})^2 + (d\beta)^2 + \left(\frac{d\gamma [\beta + \log F_{UV} + 27.5 - \log F_X]}{\gamma - 1} \right)^2 \right]^{1/2}, \quad (9)$$

where $d \log F_X$ and $d \log F_{UV}$ are the logarithmic uncertainties on F_X and F_{UV} , respectively. Equation 9 assumes that all the parameters are independent, and takes into account also the uncertainties on β and γ . The fitted likelihood function, LF , is then defined as:

$$\ln LF = -\frac{1}{2} \sum_i^N \left(\frac{(y_i - \psi_i)^2}{s_i^2} - \ln s_i^2 \right) \quad (10)$$

¹² The intercept β of the $L_X - L_{UV}$ relation is related to the one of the $F_X - F_{UV}$, $\hat{\beta}$ (see § 6), as $\hat{\beta}(z) = 2(\gamma - 1) \log d_L(z) + (\gamma - 1) \log 4\pi + \beta$.

Table 3. Properties of the final quasar sample

Name	ra	dec	z	$\log F_{\text{UV}}$	$\log F_{\text{X}}$	Group	Γ_{X}	DM
030341.04–002321.9	45.92103	–0.38942	3.235	-27.00 ± 0.04	-31.38 ± 0.04	1	1.87 ± 0.12	46.42 ± 0.31
030449.85–000813.4	46.20775	–0.13708	3.296	-26.98 ± 0.01	-31.24 ± 0.03	1	1.99 ± 0.29	45.62 ± 0.20
090508.88+305757.3	136.28702	30.96593	3.034	-26.97 ± 0.01	-31.14 ± 0.02	1	2.12 ± 0.09	44.90 ± 0.17

Notes. This table is presented in its entirety in the electronic edition; a portion is shown here for guidance.

Fluxes are in units of $\log(\text{erg s}^{-1}\text{cm}^{-2})$. The Group column flags the different subsamples: 1 = *XMM–Newton* $z \approx 3$ sample, 2 = new *XMM–Newton* $z \approx 4$ quasar, 3 = High- z sample, 4 = XXL, 5 = SDSS – 4XMM, 6 = SDSS – *Chandra*, 7 = local AGN. The column DM reports the distance moduli (with uncertainties) to reproduce the top panel of Figure 9.

where N is the number of sources, $s_i^2 = dy_i^2 + \gamma^2 dx_i^2 + \exp(2 \ln \delta)$ takes into account the uncertainties on both the x_i ($\log F_{\text{UV}}$) and y_i ($\log F_{\text{X}}$) parameters of the fitted relation, whilst δ represents its intrinsic dispersion. The variable ψ is the modelled X-ray monochromatic flux, defined as:

$$\psi = \log F_{\text{X, mod}} = \beta + \gamma(\log F_{\text{UV}} + 27.5) + 2(\gamma - 1)(\log d_{\text{L, mod}} - 28.5), \quad (11)$$

and is dependent upon the data, the redshift and the cosmological model assumed for the distances (e.g. Λ CDM, w CDM or a polynomial function). We fitted the data with a luminosity distance described by a fifth-grade polynomial of $\log(1+z)$, where the cosmographic function is:

$$d(z)_{\text{L, mod}} = k \ln(10) \frac{c}{H_0} [\log(1+z) + a_2 \log^2(1+z) + a_3 \log^3(1+z) + a_4 \log^4(1+z) + a_5 \log^5(1+z)] + O[\log^6(1+z)], \quad (12)$$

where k, a_2, a_3, a_4 and a_5 are free parameters.

For any analysis that involves a detailed test of cosmological models, we should cross-calibrate quasar distances making use of the *distance ladder* through Type Ia supernovae. In fact, the DM values of quasars are not *absolute*, thus a cross-calibration parameter (k) is needed. The parameter k should be fit separately for SNe Ia and quasars (i.e. k is a rigid shift of the quasar Hubble diagram to match the one of supernovae).

Whilst in our previous works we kept γ fixed, in this analysis we have marginalized over the slope γ of the $L_{\text{X}} - L_{\text{UV}}$ relation. The latter approach is preferred to check whether any degeneracy of the slope with the other parameters is present, and whether the statistical significance of the deviation from the Λ CDM model can be affected by the assumption of a γ value that slightly deviates from the *true* one. The marginalization on γ is a more conservative procedure, hence it might reduce the significance of the deviation with respect to the same MCMC analysis with γ fixed. Therefore, if a statistical deviation persists even allowing for a variable γ , its significance should be considered as an indicative lower limit with respect to the case where γ is fixed.

We finally note that the Hubble constant H_0 in equation (12) is degenerate with the k parameter, so it can assume any arbitrary value and was fixed to $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$ (see also Lusso et al. 2019).

Figure 9 shows the Hubble diagram for the clean quasar sample, combined with the most updated compilation of Type Ia supernovae from the *Pantheon* survey (Scolnic et al. 2018). The best MCMC cosmographic fit is shown with the red line, whilst black points are the averages (along with their uncertainties) of the distance modulus in narrow (logarithmic) redshift intervals, plotted for clarity purposes only. The residuals are displayed in the middle panel with the same symbols, and do not

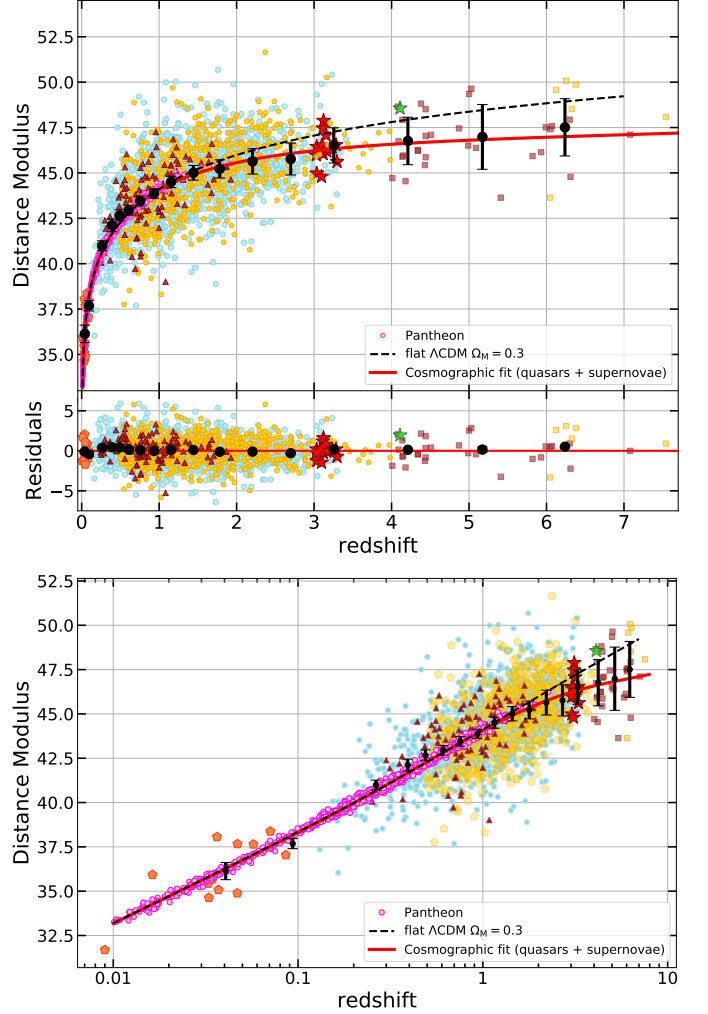


Fig. 9. Top panel: distance modulus–redshift relation (Hubble diagram) for the clean quasar sample and Type Ia supernovae (*Pantheon*, magenta points). Symbol keys are the same as in Figure 1. The red line represents a fifth order cosmographic fit of the data, whilst the black points are averages (along with their uncertainties) of the distance moduli in narrow (logarithmic) redshift intervals. The dashed black line shows a flat Λ CDM model fit with $\Omega_{\text{M}} = 0.3$. The middle panel shows the residuals with respect to the cosmographic fit and the black points are the averages of the residuals over the same redshift intervals. Bottom panel: distance modulus–redshift relation plotted with a logarithmic horizontal axis scale to better visualise the agreement between Type Ia SNe and quasars in the low-redshift range.

reveal any apparent trend with redshift. The MCMC fit assumes uniform priors on the parameters. More details on our cosmo-

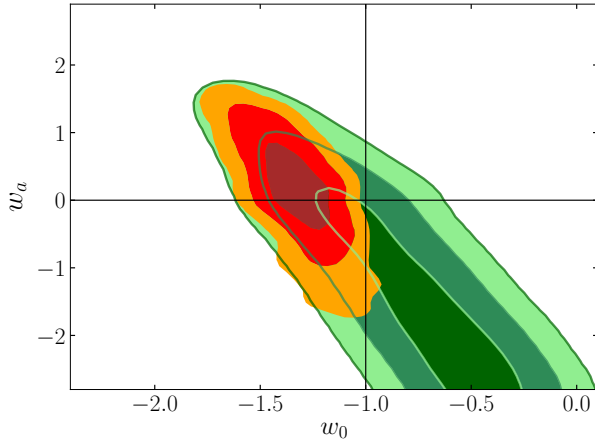


Fig. 10. Results from a fit of a w_0w_a CDM model to the combined Hubble diagram of supernovae and the “best” quasar sample, i.e. removing objects at $z < 0.7$ with a photometric determination of the UV flux (see the text for details). The green contours refer to the CMB results from *Planck* (Planck Collaboration et al. 2018). The orange (3σ), red (2σ), and brown (1σ) contours are obtained by adding the constraints from the Hubble diagram of supernovae and quasars.

graphic technique will be provided in a companion publication (Bargiacchi et al., in preparation).

We confirm that, while the Hubble diagram of quasars is well reproduced by a standard flat Λ CDM model (with $\Omega_M = 0.3$) up to $z \sim 1.5$, as shown in the top panel of Figure 9, a statistically significant deviation emerges at higher redshifts, in agreement with our previous works (e.g. Risaliti & Lusso 2015, 2019; Lusso et al. 2019) and other works on the same topic (e.g. Di Valentino et al. 2020).

The detailed discussion of the cosmological implications of this deviation and its statistical significance is not the main aim of this analysis. Here we want to focus on the study of possible systematic effects that could drive this deviation instead.

8. Cosmological fits of the Hubble diagram

In this Section we want to test our quasar sample and our method by fitting a “physical” cosmological model. Our aim is not to fully explore the consequences of our new Hubble diagram for the determination of cosmological parameters, or for the tests of different cosmological models, which will be presented in subsequent papers. Here we only intend to verify how different choices regarding the fitting method and the quasar subsample affect the final results. We choose to perform these tests with a flat w_0w_a CDM model, which is the simplest and most commonly used extension of the standard Λ CDM model, where the parameter w of the equation of state of the dark energy is assumed to vary with redshift according to the parametrization $w(z) = w_0 + w_a \times (1 - a)$, where $a = 1/(1 + z)$. Based on the analysis presented in the previous sections, three points deserve further consideration regarding cosmological fits:

- We can use the full quasar sample or add a filter of $z > 0.7$ for the sources with photometric determination of the UV flux. As discussed in Section 5.1, the possible uncertainties in the extrapolation from the optical of the quasar UV continuum at low redshift, where the host-galaxy contamination

can be important, make the 2500 Å monochromatic luminosities less reliable at $z < 0.7$. In particular, this effect is likely to be more severe at lower fluxes/luminosities, where the data quality is also lower. If the continuum slope at UV wavelengths becomes steeper than in the optical, the actual 2500 Å flux would be underestimated and this could explain the higher average values of the slope of the $L_X - L_{UV}$ relation at $z < 0.7$ (see Figure 8). This point deserves further investigation, which is deferred to a subsequent paper. To define the optical sample for cosmological applications in a conservative way, we thus prefer to cut the quasar sample at $z > 0.7$, with the exception of the local sources discussed in Section 2.7, whose 2500 Å flux is determined from the UV spectra without extrapolations. The results of the fit of this “best” quasar sample with the flat w_0w_a CDM model are shown in Figure 10. We can see that considering our data significantly reduces the $w_0 - w_a$ parametric space with respect to the CMB analysis only (Planck Collaboration et al. 2018)¹³, still being compatible with the latter data at 1σ . At the same time, the Λ CDM model (recovered for values $w_0 = -1$ and $w_a = 0$) is in tension with our data at more than 3σ , in agreement with RL19.

Finally, we note that the role of quasars at $z < 1-1.3$ is mainly to set the absolute calibration with supernovae in cosmological fits, with only a small contribution to the determination of the values and uncertainties of the cosmological parameters, given the much higher statistical weight of the supernovae. Removing quasars at $z < 0.7$ should not affect the final results significantly, and the number of quasars at redshifts overlapping with supernovae remains high enough for a precise calibration. In order to test these expectations, we repeated the fit with the whole sample, obtaining the results shown in the first panel of Figure C.1, where the contours are nearly indistinguishable from those of Figure 10.

- To fit the Hubble diagram, we can either adopt a fixed value of the slope of the relation (with its uncertainty), $\gamma = 0.59 \pm 0.06$, as determined in Section 6, or marginalize on γ as a free parameter, as discussed in Section 7. In general, the latter choice is more conservative and should be preferred. This is what we did for our reference fit, and also for the cosmographic fit used to analyze the residuals. Yet, it is worth noting that, in case of a mismatch in the shape of the Hubble diagrams of quasars and supernovae in the common redshift interval, leaving γ as a free parameter allows us to partly alleviate this problem by slightly “bending” the $L_X - L_{UV}$ relation in order to obtain a better agreement. In our case, the fit of the “best” sample gives $\gamma = 0.600 \pm 0.015$, consistent within 1.5σ with the value found from the fits of the relation in narrow redshift bins. In order to check the possible effects of this choice, we repeated the fit with the w_0w_a CDM model, again obtaining a contour plot totally consistent with the reference case (second panel in Figure C.1).
- The residuals in the Hubble diagram show a moderate, but statistically significant, redshift dependence on the X-ray slope, Γ_X . It is important to understand whether this introduces a bias in the fits of the Hubble diagram. We checked this possibility by splitting the sample in two parts, with $\Gamma_X < 2.2$ and $\Gamma_X > 2.2$, respectively, and repeating the fit with the w_0w_a CDM model. The results are shown in the last two panels of Figure C.1. These contours are slightly larger

¹³ Baseline Λ CDM chains with baseline likelihoods: https://wiki.cosmos.esa.int/planck-legacy-archive/index.php/Cosmological_Parameters

than in the previous cases, as expected given the lower statistics, but no systematic trend is observed. We conclude that, while the possible dependence on Γ_X deserves further analysis in order to understand its physical and/or observational origin and to reduce the dispersion of the $L_X - L_{UV}$ relation (Signorini et al., in preparation), no systematic effect related to Γ_X is introduced in the Hubble diagram of quasars.

9. Study of systematics in the Hubble diagram

Since the main aim of our analysis is to check whether any systematic is present in the residuals of the quasar Hubble diagram, at this stage we avoid the inclusion of Type Ia supernovae. As noted above, when only quasars are involved the DM values should not be considered as proper absolute distances. In this Section we thus present an in-depth investigation of possible systematics in the residuals of the quasar Hubble diagram, unaccounted for in the selection of the sample. In particular, we explored whether our procedure (1) to correct for the Eddington bias (Section 5.3), (2) to neglect quasars with possible gas absorptions (Section 5.2), and (3) to select blue quasars based on their SED shape, where dust absorption and host-galaxy contamination are minimised (Section 5.1), introduces spurious trends in the Hubble diagram residuals as a function of redshift and for different intervals of the relevant parameters. For each variable, we divided the sample between the sources that fall below and above the average value of the variable itself, and examined each subset separately as any hidden dependence should lead to a systematic difference between the two.

9.1. Residuals as a function of the Eddington bias

To verify whether our adopted technique to correct for the Eddington bias, based on the assumption that the *true* slope of the $L_X - L_{UV}$ is $\gamma = 0.6$, leaves some hidden trends in the residuals of the Hubble diagram as a function of redshift, we defined an Eddington bias parameter, Θ , as the difference between the expected X-ray monochromatic flux at 2 keV and the sum of the flux sensitivity value at 2 keV ($F_{\min}$) and the product $\kappa\delta$ (see equation 3):

$$\Theta = \log F_{2\text{keV}, \text{exp}} - (\log F_{\min} + \kappa\delta). \quad (13)$$

Given the fact that the source has survived the Eddington bias filter (see equation 3, where $\kappa\delta = 0.5$ for SDSS-*Chandra* and 0.9 for XXL and SDSS-4XMM, respectively), Θ is always positive: the higher its value, the lower the bias due to the flux limit of the specific X-ray observation.

The other samples benefit from pointed observations, for which the bias due to the X-ray flux sensitivity is fully negligible.

We then defined two subsets, above (1069 objects) and below (1289) the mean of this distribution, $\langle\Theta\rangle = 1.17 \pm 0.01$ (with a 1σ dispersion of 0.33), and plotted the residuals of the Hubble diagram as a function of redshift in Figure 11. The three quasar samples span a rather large redshift interval, showing no clear trend of the residuals with distance and a similar dispersion around zero (≈ 1), with an average residual value of about -0.07 (0.04) for $\Theta > 1.2$ ($\Theta < 1.2$). These results imply that our selection of X-ray observations discussed in Section 5.3 does not introduce any systematic trends, at least up to $z \approx 4$.

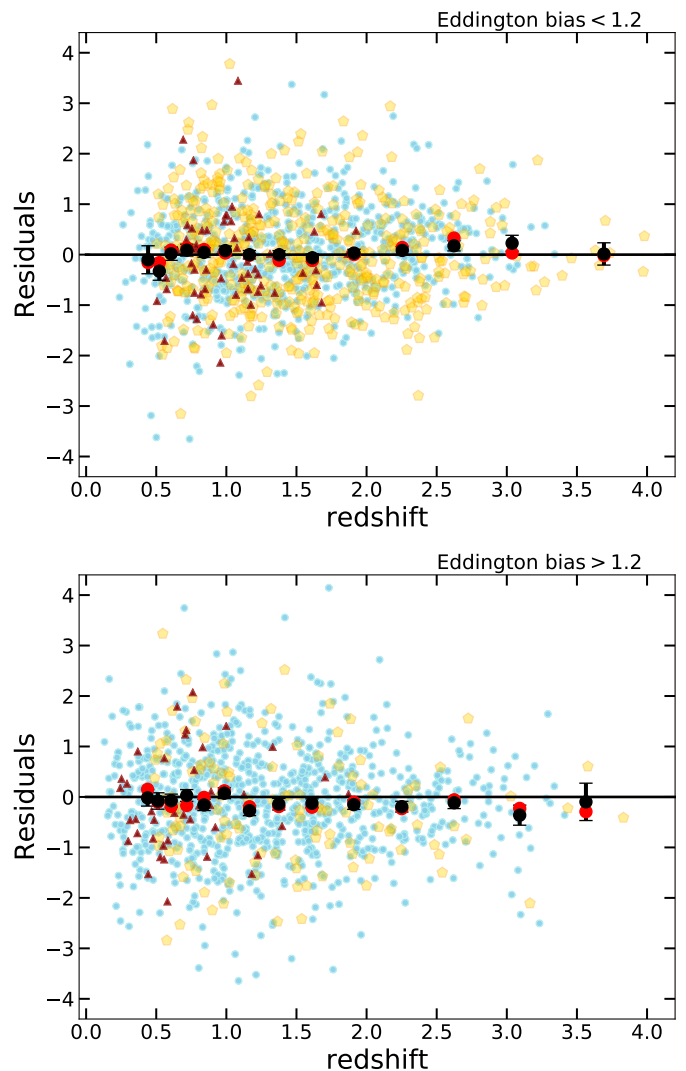


Fig. 11. Distribution of the Hubble diagram residuals (middle panel of Figure 9) as a function of redshift for the quasars with Θ lower (higher) than the average (i.e. $\langle\Theta\rangle = 1.2$). The black and red points represent the mean and median of the residuals in narrow redshift intervals, respectively. Symbol keys as in Figure 1.

9.2. Residuals as a function of the photon index

Gas absorption in the X-ray band and strong outliers with extremely steep Γ_X are minimized by neglecting all quasars outside a given interval of photon index values ($1.7 \leq \Gamma_X \leq 2.8$, see Section 5.2). In Figure 12 we present the distribution of the Γ_X values for the clean quasar sample. The mean (median) value for the sample is $\langle\Gamma_X\rangle \approx 2.202 \pm 0.005$ ($\langle\Gamma_X\rangle \approx 2.173^{+0.006}_{-0.005}$), with a dispersion of about 0.23 (0.30). Statistical errors are quoted.

The average Γ_X is biased towards slightly steeper values with respect to the more typical $\Gamma_X \sim 1.9$ – 2 (e.g. $\Gamma_X = 1.99 \pm 0.01$ with a dispersion of 0.3, Scott et al. 2011; see also Young et al. 2009b; Mateos et al. 2010), which is likely due to our conservative cut at $\Gamma_X = 1.7$ and to the presence of a tail of sources with power laws softer than $\Gamma_X = 2.6$. The residuals of the Hubble diagram do not show a significant trend as a function of redshift when the sample is split in two subsets with Γ_X higher and lower than 2.2 (Figure 13). None the less, the marginalized distribution of residuals shown in Figure 14 presents an offset with respect to zero, with an average value for the two subsamples of

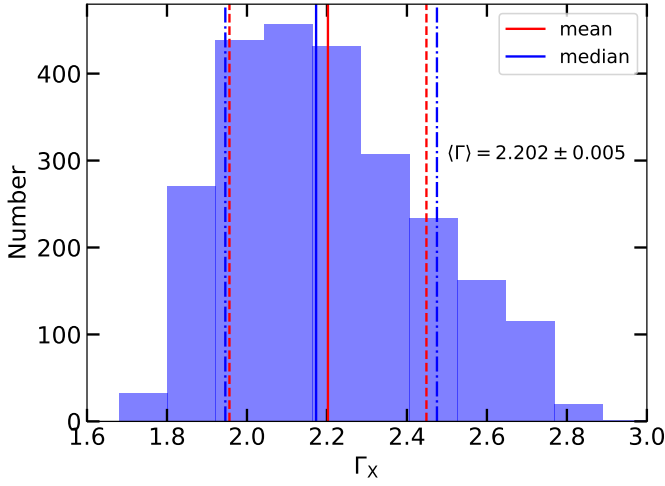


Fig. 12. Distribution of the photon index Γ_X for the clean final sample. The mean and the statistical error on the mean are also quoted. The red (blue) solid line represents the mean (median) of Γ_X , with its 1σ dispersion around the average marked with the dashed (dot-dashed, 16% and 84%) lines.

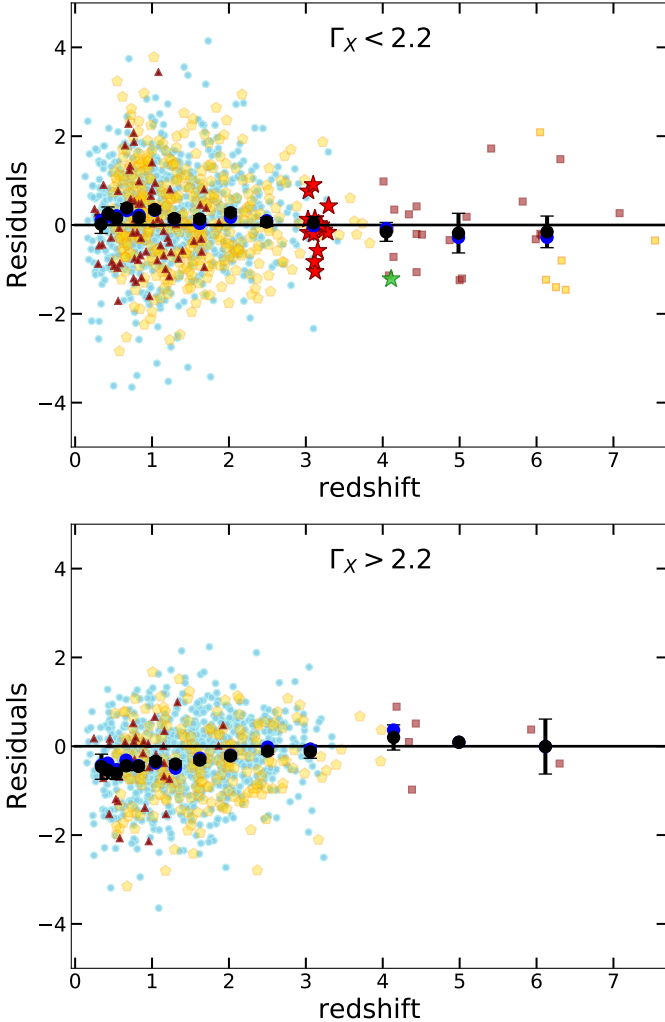


Fig. 13. Distribution of the Hubble diagram residuals (middle panel of Figure 9) as a function of redshift for the quasars with Γ_X lower (higher) than the average (i.e. $\langle\Gamma_X\rangle = 2.2$, see Figure 12). The black and red points represent the mean and median of the residuals in narrow redshift intervals, respectively. Symbol keys as in Figure 1.

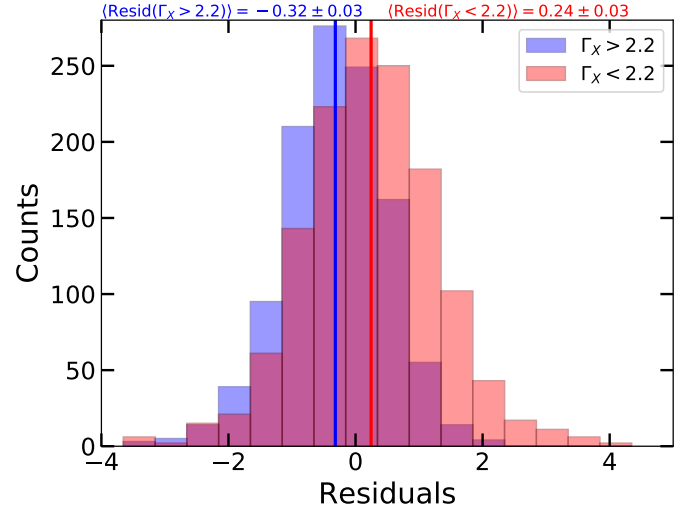


Fig. 14. Distribution of the Hubble diagram residuals (middle panel of Figure 9) for the quasars with Γ_X higher (blue) and lower (red) than the average (i.e. $\langle\Gamma_X\rangle = 2.2$, see Figure 12). The blue and red solid lines represent the mean of the residuals in each Γ_X distribution, respectively (averages are reported on top).

-0.32 ± 0.03 and 0.24 ± 0.03 for Γ_X higher and lower than 2.2, respectively. The dispersion around the average values is 0.82 (1.05) for $\Gamma_X > 2.2$ ($\Gamma_X < 2.2$).

As we pointed out in Section 5.2, our photometric Γ_X values may not be always accurate for individual objects, but they are reliable in a *statistical sense* over a large sample of quasars. Consequently, the main drawback of not using the spectroscopic values is likely to increase the dispersion, rather than to introduce a strong systematic with redshift that may affect the cosmological analysis. Moreover, even considering the presence of a small redshift trend in Figure 13 for $z < 2$, this is counterbalanced by the use of Type Ia supernovae in the same redshift interval. In fact, the statistical significance of the deviation of the quasar Hubble diagram from the standard Λ CDM reported in our previous works starts to be important for $z > 2$, where all the residuals do not present any obvious systematic with redshift.

9.3. Residuals as a function of SED colours

As a further check, we have explored whether the SED colours produce some systematic trends with redshift in the residuals of the Hubble diagram. We have selected two subsamples in different regions of the (Γ_1, Γ_2) plane (see Figure 7). The first one considers all the quasars within the circle having a radius of 0.45 (i.e. $\sqrt{(\Gamma_1 - 0.82)^2 + (\Gamma_2 - 0.40)^2} \leq 0.45$; 1271 quasars), which corresponds to a colour excess $E(B - V)$ of 0.04 (using an SMC-like reddening law). The second is the contiguous annulus with outer radius of 1.1 (1087 quasars) corresponding to an $E(B - V) \simeq 0.1$. As the majority of the quasars in the final sample are drawn from the SDSS-DR14 catalogue, the subsample with a colour excess lower than 0.04 should represent the bluest objects, whilst the sample in the outer annulus with $0.04 \leq E(B - V) \leq 0.1$ could show some pattern in the residuals in case of a colour-related systematic with redshift.

Figure 15 presents the distribution of the residuals as a function of redshift for the two subsamples defined as above for the clean SDSS-4XMM, XXL and SDSS-*Chandra* samples, since we can construct homogeneous SEDs (and thus retrieve Γ_1 and

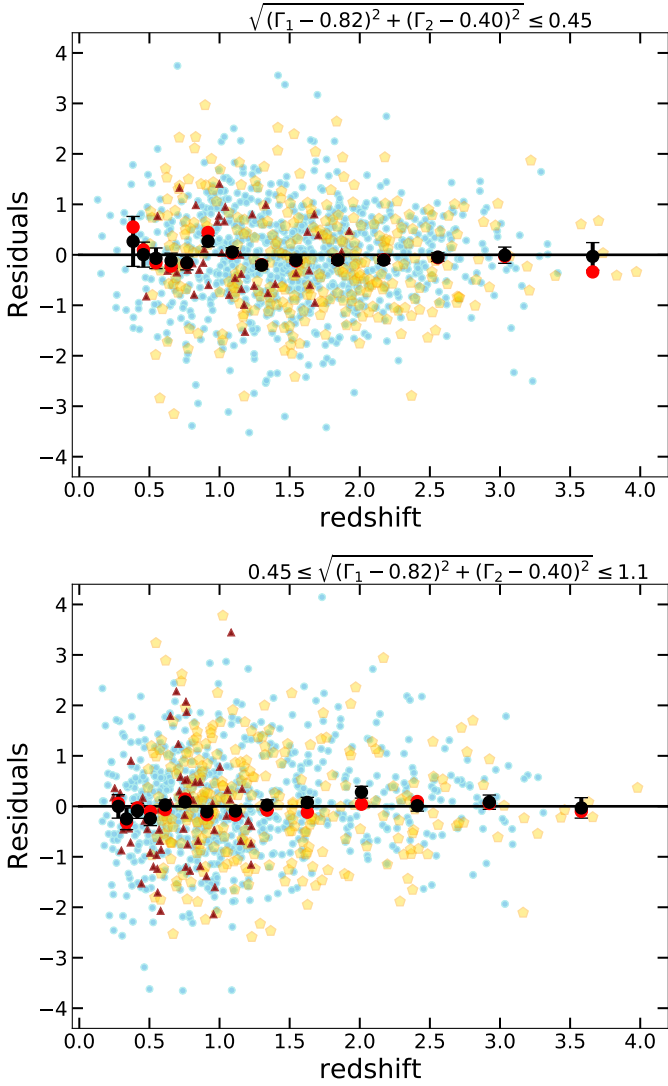


Fig. 15. Distribution of the Hubble diagram residuals (middle panel of Figure 9) as a function of redshift for the quasars with $\sqrt{(\Gamma_1 - 0.82)^2 + (\Gamma_2 - 0.40)^2} \leq 0.45$, i.e. $E(B - V) \leq 0.04$ (top panel, the bluest quasars in the clean sample) and $0.45 \leq \sqrt{(\Gamma_1 - 0.82)^2 + (\Gamma_2 - 0.40)^2} \leq 1.1$, i.e. $0.04 \leq E(B - V) \leq 0.1$ (bottom panel) in the clean SDSS-4XMM, XXL and SDSS-*Chandra* samples. The black and red points represent the mean and median of the residuals in narrow redshift intervals, respectively. Symbol keys as in Figure 1.

Γ_2 slopes) for them all using the SDSS photometry from the DR14 quasar catalogue. It is clear that there is no obvious trend of the residuals with redshift in either subsample, implying that our colour selection does not introduce any redshift-dependent bias in the Hubble diagram. Moreover, in both cases there is no difference in the dispersion of the residuals around zero (≈ 1), with average values of about -0.04 for $E(B - V) \leq 0.04$ and 0.03 for $0.04 \leq E(B - V) \leq 0.1$. This further proves that even the inclusion of sources with possible (yet modest) contamination from dust and/or host galaxy does not lead to any systematics with redshift up to $z \approx 4$.

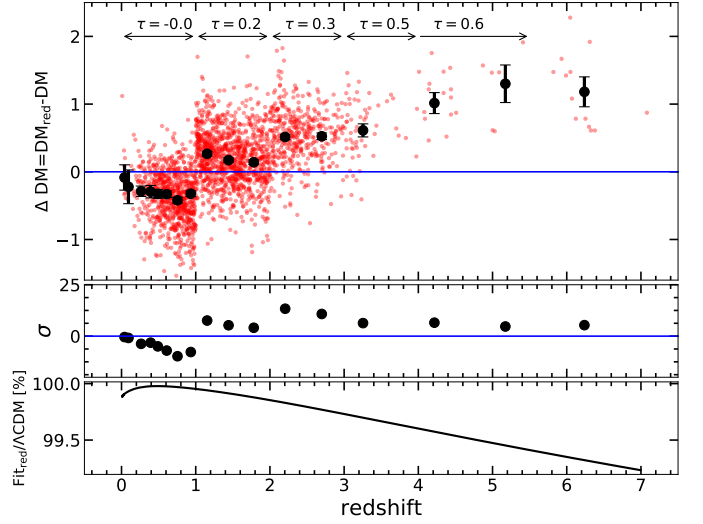


Fig. 16. Analysis of the possible contribution of dust at $2500 \text{ \AA}$. Top: distribution of the difference of the distance moduli (ΔDM , red points; $\Delta DM = 0$ represents the Λ CDM) as a function of redshift. The parameter ΔDM is defined as $DM_{\text{red}} - DM$, where DM is computed from the quasar data as described in § 7. The DM_{red} values are estimated by correcting the observed F_{UV} for an additional intrinsic (and redshift-dependent) extinction. τ is the optical depth at the rest-frame $2500 \text{ \AA}$ that would be needed to ascribe the discrepancy to an underestimated dust extinction in the UV flux. The black points represent the average (and 1σ uncertainties) of ΔDM in the same narrow redshift intervals used in Figures 8 and 9. Middle: deviation in σ of the average points with respect to Λ CDM. Bottom: ratio between the cosmographic fit obtained from the data, where F_{UV} is corrected for additional extinction, and the Λ CDM. The cosmographic fit perfectly matches Λ CDM for the chosen values of dust reddening (see § 9.4).

9.4. Is dust/gas absorption driving the deviation in the quasar Hubble diagram?

Among the possible residual (and redshift-dependent) observational systematics in the Hubble diagram, we must also consider the presence of an additional contribution of dust reddening in the UV band. As we move to higher redshifts, the rest-frame optical/UV spectra shift to higher (shorter) frequencies (wavelengths), where the dust absorption cross-section is higher. This might lead to an underestimate of F_{UV} , which would imply an intrinsically larger value of the luminosity distance (and thus DM) than the one we measured.

We thus evaluated the amount of dust extinction required to make the cosmographic fit shown in Figure 9 coincide with a standard flat Λ CDM model with $\Omega_M = 0.3$ (i.e. the black dashed line in Figure 9). We defined five redshift intervals with increasing values of $E(B - V)$, specifically, $E(B - V) = 0, 0.1, 0.15, 0.2$, and 0.25 at $z = 0-1$, $z = 1-2$, $z = 2-3$, $z = 3-4$, and $z > 4$. We then assumed the standard extinction law by Prevot et al. (1984) with $R_V = 3.1$ and corrected F_{UV} by the amount dictated by the chosen $E(B - V)$ in each redshift interval, and we finally fitted the “reddening corrected” distance modulus (DM_{red})–redshift relation.

Figure 16 shows the resulting distribution of the difference of the distance moduli, ΔDM (red points) between the DM_{red} values computed above and the observed DM in Figure 9 (see § 7), as a function of redshift. The cosmographic fit obtained from the $DM_{\text{red}} - z$ relation perfectly matches the Λ CDM curve (bottom panel). We also report the values of optical depth, τ , at the rest-frame $2500 \text{ \AA}$ that would be needed to entirely ascribe the dis-

crepancy to underestimated dust extinction on the 2500 Å quasar fluxes.

We note that the negative ΔDM values at $z < 1$ (where no dust extinction correction is applied) are simply caused by the different overall cosmographic fit obtained for the $\text{DM}_{\text{red}} - z$ relation with respect to the observed one. Therefore, the ΔDM values in this redshift range should not be taken at face value.

At 2500 Å, it is $k(\lambda) \simeq 7$, which corresponds to an increase in flux by a factor of $\simeq 2.6 - 5$ for $E(B - V) = 0.15 - 0.25$ at $z \geq 2$. The bolometric luminosities of the $z > 4$ quasars in the clean sample are in the range $10^{46.5-48.1} \text{ erg s}^{-1}$, assuming a bolometric correction of 2.75 at 2500 Å (Krawczyk et al. 2013). If the extra correction were effectively required, this would imply an *intrinsic* bolometric luminosity for these sources of the order of $10^{47.3-49} \text{ erg s}^{-1}$. For example, the highest redshift quasar in the sample, ULAS J134208.10+092838.61, would have a bolometric luminosity by at least a factor of 4 higher than reported in the literature (i.e. $1.6 \times 10^{47} \text{ erg s}^{-1}$, Bañados et al. 2018). This would also imply a larger black-hole mass, thus a heavier seed, which needs to be interpreted within current models of black-hole growth (e.g. Latif et al. 2013; Dijkstra et al. 2014; Pacucci et al. 2015).

For the spectroscopic *XMM-Newton* $z \simeq 3$ sample, the F_{UV} values would increase by a factor of at least 3.5, thus shifting L_{bol} into the range $10^{48-48.8} \text{ erg s}^{-1}$. We note that the average stack of their SDSS spectra does not suggest any significant levels of dust absorption when compared to other composites obtained from bright, blue quasar spectra (see Figure 2 in Nardini et al. 2019). As a result, the presence of any additional dust component does not seem to be justified.

For completeness, we have also employed a different reddening law (i.e. Fitzpatrick 1999), and modified both the $E(B - V)$ values and the redshift intervals (always verifying that the cosmographic fit of the modified Hubble diagram is consistent with the standard ΛCDM), finding equivalent results.

In conclusion, the intrinsic UV fluxes required to offset the observed tension between the cosmographic fit of the Hubble diagram and the ΛCDM through redshift-dependent dust absorption should be, on average, brighter by 0.4–0.8 mag. This would imply a steeper extrapolated slope than observed at longer wavelengths, which would be hard to reconcile with a standard α -disc model (Shakura & Sunyaev 1976).

We note that concurrent, and redshift-dependent, absorption is found at X-ray energies. The observed 0.5–2 keV energy band moves towards higher rest-frame energies as the redshift increases, where the effect of any gas absorption should become progressively negligible. Spectra with flat photon indices ($\Gamma_X < 1.7$) are removed from the sample, so we expect a minimal degree of residual gas absorption in the X-rays. In the high-redshift ($z > 6$) quasars for which a spectral analysis was carried out (see e.g. Vito et al. 2019), the level of intrinsic gas absorption is always suggested to be rather low (of the order of $\times 10^{23} \text{ cm}^{-2}$ or less). Assuming the presence of a local column density with $N_{\text{H}} = 10^{23} \text{ cm}^{-2}$ at $z = 6$, the exact correction to the inferred rest-frame 2 keV flux density depends on the details (flux, observed slope, data quality) of the single observation, but it is expected to be within a factor of 2, i.e. comparable to the typical measurement uncertainty. Such high N_{H} values are not required in any of the X-ray spectral fits for the high-redshift quasars that belong to the clean sample. By shifting the data to higher DM values, this correction would have the effect of increasing the departure from the ΛCDM (i.e. opposite to the one for the dust absorption in the UV).

9.5. On the possible dependence of the $L_X - L_{\text{UV}}$ relation on black-hole mass and accretion rate

One can argue that the observed $L_X - L_{\text{UV}}$ relation might be a secondary manifestation of some other, more fundamental relations, involving, for instance, a possible dependence on black-hole mass and accretion rate (usually parametrised by the Eddington ratio, λ_{edd} , defined as the ratio of the bolometric and the Eddington luminosity). Yet, when estimated from single-epoch spectroscopy, both parameters are derived quantities, i.e. a combination of continuum luminosity and emission-line FWHM. Therefore, the presence or absence of a correlation between the residuals of the Hubble diagram and M_{BH} and/or the Eddington ratio could be misleading and potentially hide systematics.

Since the $L_X - L_{\text{UV}}$ relation already has a dependence on the nuclear luminosity (which we have discussed in detail in Section 3.1), we can explore possible additional correlations with the FWHM of a given emission line. This issue was already investigated in Lusso & Risaliti (2017), who found that not only is the dependence of the $L_X - L_{\text{UV}}$ relation on the FWHM (of Mg II) statistically significant, but also that such a dependence has the effect of further reducing the dispersion of the $L_X - L_{\text{UV}}$ correlation. However, building a statistically significant quasar sample that both covers a wide redshift range (as the one presented here) and relies on a measurement of FWHM from a single line (to ensure homogeneity in the estimate of the line parameters, see § 2.1 in Lusso & Risaliti 2017) is impractical, as it would require a series of dedicated spectroscopic near-infrared campaigns to probe the same line (e.g. Mg II) at high redshifts ($z > 2.5$). The statistical relevance of the supplementary dependence of the $L_X - L_{\text{UV}}$ relation on the FWHM, as well as on other parameters inferred from spectroscopy (e.g. continuum slope, iron content, line strength) will be further investigated in a forthcoming publication, entirely dedicated to the physics of the relation.

While here we prefer to consider the simple $L_X - L_{\text{UV}}$ relation for cosmological purposes, as the present analysis aims at verifying that our selection criteria do not introduce systematics in the quasar Hubble diagram, we can still check whether the sources in the clean sample with a measurement of the FWHM from Mg II (which provides the widest possible redshift coverage in BOSS) show any systematic trend and/or a reduced scatter in the residuals. We thus cross-matched the clean sample with the catalogue of spectral quasar properties of Rakshit et al. (2020), and selected the quasars with a broad (FWHM $> 2000 \text{ km s}^{-1}$) component of the Mg II emission line, finding 1,858 quasars (77% of the clean sample).

The distribution of the residuals as function redshift for this sample is shown in Figure 17. The amplitude of the scatter of the residuals is similar to the other cases discussed in Sections 9.1, 9.2, and 9.3. Moreover, there is no trend with redshift up to $z \simeq 2.5$. These results further prove that our selection criteria are already effective in selecting an optimal sample for cosmology, without introducing strong systematics.

10. Discussion and conclusions

Our group has presented a new technique that makes use of the observed non-linear relation between the 2500 Å and the 2 keV emission in quasars to provide an independent measurement of their distances, thus turning quasars into *standardizable* candles. Our method extends the distance modulus–redshift relation (or the so-called *Hubble-Lemaître diagram*) of supernovae Ia to a redshift range still poorly explored ($z > 2$; e.g. Risaliti & Lusso

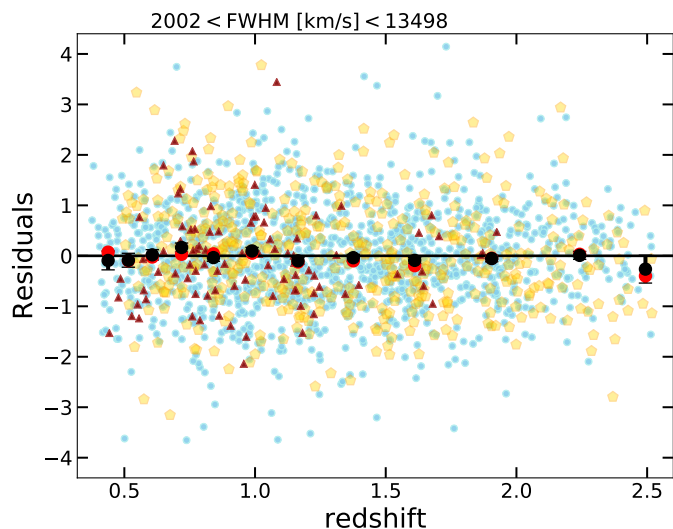


Fig. 17. Distribution of the Hubble diagram residuals (middle panel of Figure 9) as a function of redshift for the quasars with a broad ($\text{FWHM} > 2000 \text{ km s}^{-1}$) component of the Mg II emission line in the clean sample. The black and red points represent the mean and median of the residuals in narrow redshift intervals, respectively. Symbol keys as in Figure 1.

2015, 2019; Lusso et al. 2019), and it relies upon the evidence that most of the observed dispersion in the $L_X - L_{\text{UV}}$ relation is not intrinsic to the relation itself but due to observational issues. When an optimal selection of *clean* sources (i.e. for which we can measure the *intrinsic* UV and X-ray quasar emission) is possible, the dispersion in the $L_X - L_{\text{UV}}$ relation drops to ≈ 0.2 dex (Lusso & Risaliti 2016, 2017).

We have previously demonstrated that the distance modulus–redshift relation of quasars at $z < 1.4$ is in agreement with that of supernovae Ia and with the concordance Λ CDM model (Risaliti & Lusso 2015, 2019; Lusso et al. 2019), yet a deviation from the Λ CDM emerges at higher redshift, with a statistical significance of about 4σ . If we interpret the latter result by considering an evolution of the dark energy equation of state in the form $w(z) = w_0 + w_a \times z/(1+z)$, the data suggest that the dark energy density is increasing with time (Risaliti & Lusso 2019; Lusso et al. 2019).

However, our technique may still have some limitations, and we need to verify that the observed deviation from the Λ CDM at redshift > 2 is not driven by either systematics in the quasar sample selection or the cosmological procedure adopted to fit the distance modulus–redshift relation.

The aim of this manuscript is thus to discuss, on the one hand, all the criteria required to select a homogeneous sample of quasars for cosmological purposes and, on the other hand, the specific procedures adopted to compute the UV and X-ray fluxes and spectral slopes from the available photometry. We identified the quasars that can be used for a cosmological analysis, examined the key steps in fitting the distance modulus–redshift relation, and considered the possible systematics in the quasar Hubble-Lemaître diagram. In particular, we investigated in depth the residuals of the quasar Hubble diagram, in order to unveil any systematics unaccounted for in the selection of the sample. We explored whether our procedure (1) to correct for the Eddington bias, (2) to neglect quasars with possible gas absorption, and (3) to select blue quasars based on their SED shape, where dust absorption and host-galaxy contamination are minimised, introduces spurious trends in the Hubble diagram residuals as

a function of redshift and for different intervals of the relevant parameters.

Our main results are the following:

- We verified that the $L_X - L_{\text{UV}}$ relation (i.e. slope and dispersion) for the final “best” quasar sample does not evolve with redshift.
- We confirmed that, while the quasar Hubble diagram is well reproduced by a standard flat Λ CDM model (with $\Omega_M = 0.3$) up to $z \sim 1.5$, a statistically significant deviation emerges at higher redshifts, in agreement with our previous works (e.g. Risaliti & Lusso 2015, 2019; Lusso et al. 2019).
- We found that none of the adopted filters introduce strong systematics in the Hubble diagram residuals, and specifically where the quasars become the only contributors and the deviation from the standard Λ CDM is more significant, i.e. at $z > 1.5$.

Even if our analysis shows that both the quasar selection criteria and the cosmological fitting technique are robust, we can already envisage several further improvements, especially in the quasar sample selection (e.g. by using spectra instead of photometry in both X-rays and UV to better measure the X-ray and UV fluxes and slopes). None the less, given the remarkable absence of systematics in the residuals, the main effect of these refinements will be on the reduction of the dispersion, thus allowing a better estimate of the cosmological parameters.

With currently operating facilities, dedicated observations of well-selected high- z quasars (similarly to what our group has done at $z \approx 3$) will greatly improve the test of the cosmological model and the study of the dispersion of the $L_X - L_{\text{UV}}$ relation, especially at $z \approx 4$. The *extended Roentgen Survey with an Imaging Telescope Array* (eROSITA, Predehl 2012; Merloni et al. 2012), flagship instrument of the ongoing Russian *Spektrum-Roentgen-Gamma* (SRG) mission, will represent an extremely powerful and versatile X-ray observatory in the next decade. The sky of eROSITA will be dominated by the AGN population, with ~ 3 million AGN with a median redshift of $z \sim 1$ expected by the end of the nominal 4-year all-sky survey at the sensitivity of $F_{0.5-2 \text{ keV}} \approx 10^{-14} \text{ erg s}^{-1} \text{ cm}^{-2}$, for which extensive multi-wavelength follow-up is already planned. Concerning the constraints on the cosmological parameters (such as Ω_M , Ω_Λ , and w) through the Hubble diagram of quasars, we predict that the 4-year eROSITA all-sky survey alone, complemented by redshift and broadband photometric information, will supply the largest quasar sample at $z < 2$ (average redshift $z \approx 1$), but a relatively small population should survive the Eddington bias cut at higher redshifts (see e.g. Medvedev et al. 2020), thus being available for cosmology. Indeed, eROSITA will sample the brighter end of the X-ray luminosity function (Lusso 2020, but see also section 6.2 in Comparat et al. 2020). None the less, the large number of eROSITA quasars at $z \approx 1$ will be pivotal for both a better cross-calibration of the quasar Hubble diagram with supernovae and a more robust determination of Ω_Λ , which is sensitive to the shape of the low redshift part of the distance modulus–redshift relation.

In the mid/long term, surveys from *Euclid* and LSST in the optical/UV, and *Athena* in the X-rays, will also provide samples of millions of quasars. With these samples it will be possible to obtain constraints on the observed deviations from the standard cosmological model, which will rival and complement those available from the other cosmological probes.

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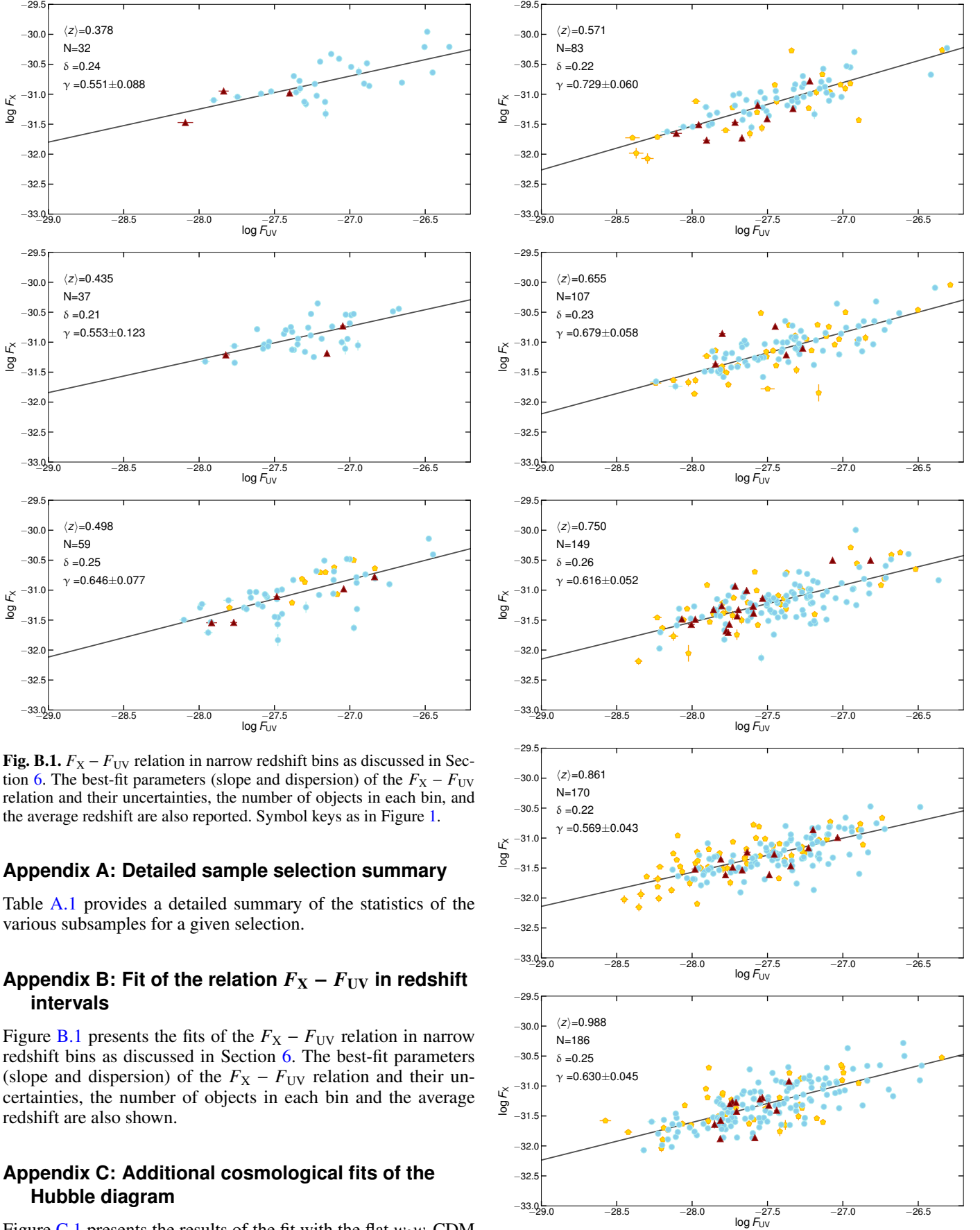


Fig. B.1. Continued.

Fig. B.1. $F_X - F_{UV}$ relation in narrow redshift bins as discussed in Section 6. The best-fit parameters (slope and dispersion) of the $F_X - F_{UV}$ relation and their uncertainties, the number of objects in each bin, and the average redshift are also reported. Symbol keys as in Figure 1.

Appendix A: Detailed sample selection summary

Table A.1 provides a detailed summary of the statistics of the various subsamples for a given selection.

Appendix B: Fit of the relation $F_X - F_{UV}$ in redshift intervals

Figure B.1 presents the fits of the $F_X - F_{UV}$ relation in narrow redshift bins as discussed in Section 6. The best-fit parameters (slope and dispersion) of the $F_X - F_{UV}$ relation and their uncertainties, the number of objects in each bin and the average redshift are also shown.

Appendix C: Additional cosmological fits of the Hubble diagram

Figure C.1 presents the results of the fit with the flat $w_0 w_a$ CDM model of the quasar sample in four different cases: (1) the full sample of $\sim 2,400$ quasars where the slope, γ , of the $F_X - F_{UV}$ relation is a variable parameter, (2) the “best” $z > 0.7$ quasar

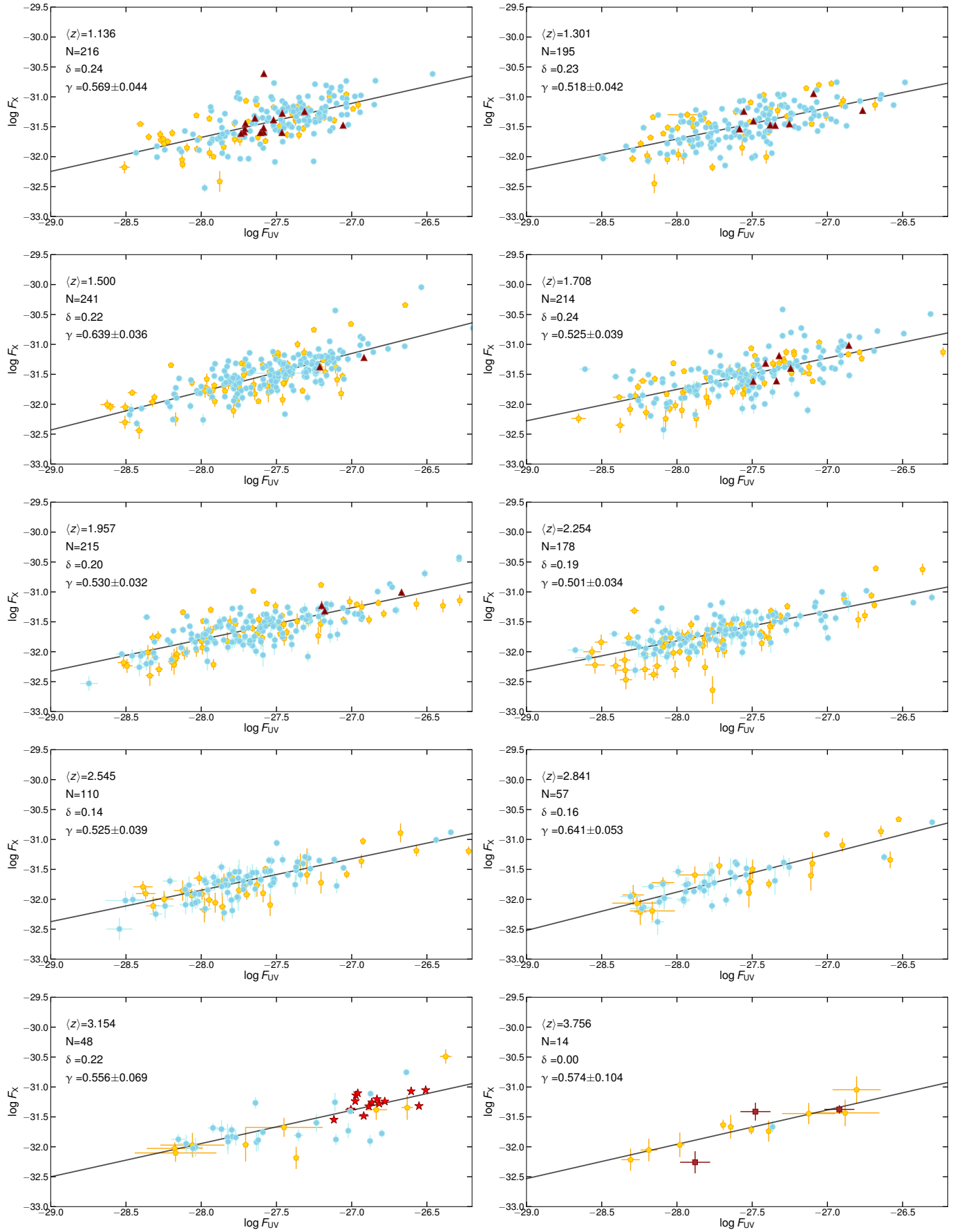


Fig. B.1. Continued.

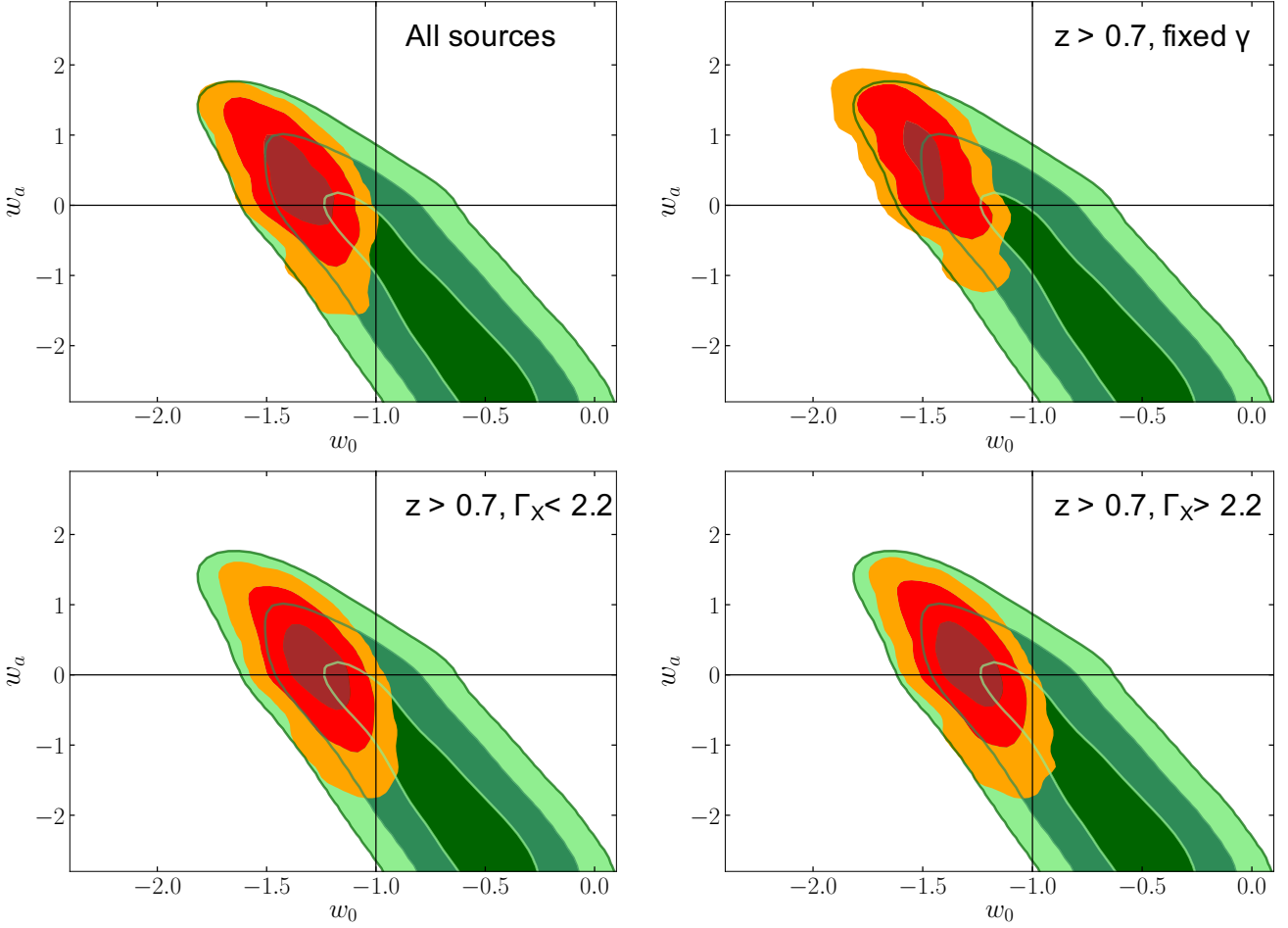


Fig. C.1. Results from a fit of a w_0w_a CDM model to our Hubble diagram of supernovae and quasars in four different cases. Top left: whole quasar sample, with the slope of the $L_X - L_{UV}$ relation as a free parameter; top right: quasars at $z > 0.7$ with a fixed $L_X - L_{UV}$ slope, $\gamma = 0.59$, i.e. the average value from the fit of the relation in narrow redshift intervals, as shown in Section 6; bottom panels: results for two subsamples with $\Gamma_X < 2.2$ and $\Gamma_X > 2.2$. Contours are at the 1σ , 2σ , and 3σ confidence levels. The green contours refer to the CMB results from *Planck* (Planck Collaboration et al. 2018). The orange, red, and brown contours are obtained by adding the constraints from the Hubble diagram of supernovae and quasars.

sample with γ fixed to the average value of the entire sample estimated in narrow redshift bins (i.e. $\gamma = 0.59$, see § 6), (3) the “best” $z > 0.7$ quasar sample with Γ_X lower than the average photon index as discussed in § 9.2 (see Figure 12) and finally (4) the “best” $z > 0.7$ quasar sample with Γ_X higher than the average photon index. In both fits with the (3) and (4) samples γ is left free. This analysis confirms that the Λ CDM model is in tension with our data at more than 3σ , in agreement with the results in RL19, irrespectively of the sample selection.

Table A.1. Sample statistics of the various subsamples in each give selection.

Selection	Samples							
	<i>XMM-Newton</i> $z \simeq 3$	<i>XMM-Newton</i> $z \simeq 4$	High- z	XXL	SDSS-4XMM	SDSS- <i>Chandra</i>	Local AGN	Total
Initial sample (1)	29	1	64 ^a	840	13,800	7,036	15	21,785
Main sample (2)	29	1	64	840	9,252	2,392	15	12,593
Dust & Host correction (§ 5.1)	29	1	64	535	7,319	1,136	15	9,099
X-ray absorption (§ 5.2)	14 [†]	1	35 ^b	259	2,576	994	13	3,921
Eddington bias (§ 5.3)	14	1	35	106	1,644	608	13	2,421

Notes. (1) These number counts refer to the sample statistics before correcting for overlaps amongst the subsamples. (2) Sample statistics after accounting for overlaps and some preliminary selection (e.g. RL, BAL, off-axis angle in the X-rays). The order of priority decreases from the leftmost to the rightmost column. ^a See § 2.6. ^b 35 quasars = 29 sources from [Salvestrini et al. \(2019\)](#) and 6 from [Vito et al. \(2019\)](#). [†] Seven X-ray weak quasars are amongst the 15 objects excluded (see [Nardini et al. 2019](#) for details).