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The dependence of gamma-ray burst X-ray column densities on the model for Galactic hydrogen

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ABSTRACT

We study the X-ray absorption of a complete sample of 99 bright *Swift* gamma-ray bursts (GRBs). In recent years, a strong correlation has been found between the intrinsic X-ray absorbing column density ($N_{\text{H}}(z)$) and the redshift. This absorption excess in high- z GRBs is now thought to be due to the overlooked contribution of the absorption along the intergalactic medium (IGM), by means of both intervening objects and the diffuse warm-hot intergalactic medium along the line of sight. In this work we neglect the absorption along the IGM, because our purpose is to study the eventual effect of a radical change in the Galactic absorption model on the $N_{\text{H}}(z)$ distribution. Therefore, we derive the intrinsic absorbing column densities using two different Galactic absorption models: the Leiden Argentine Bonn H_I survey and the more recent model that includes molecular hydrogen. We find that if, on the one hand, the new Galactic model considerably affects the single column density values, on the other hand, there is no drastic change in the distribution as a whole. It becomes clear that the contribution of Galactic column densities alone, no matter how improved, is not sufficient to change the observed general trend and it has to be considered as a second order correction. The cosmological increase of $N_{\text{H}}(z)$ as a function of redshift persists and, to explain the observed distribution, it is necessary to include the contribution of both the diffuse intergalactic medium and the intervening systems along the line of sight of the GRBs.

Key words. gamma-ray burst: general – X-rays: general

1. Introduction

Long gamma-ray bursts (LGRBs) are associated with type Ib/c supernova explosions (Woosley & Bloom 2006) and observed in galactic regions with high star formation rate (Fruchter et al. 2006): this enables us to infer that LGRBs have massive stars as progenitors. Moreover, their circumburst medium is thought to be denser than the typical star formation regions, as we can see from the high absorbing column density values measured in X-rays (Galama & Wijers 2001; Stratta et al. 2004; Gendre et al. 2006). Both metals in the circumburst medium and within the whole host galaxy may contribute to the intrinsic column density $N_{\text{H}}(z)$. In addition, X-ray absorption occurs within our Galaxy (parametrized by $N_{\text{H}}(\text{Gal})$) and along the line of sight owing to the intergalactic medium (IGM). This occurs by means of a warm-hot diffuse medium that pervades it or, mainly, by means of compact discrete intervening systems (galactic halos, galaxy groups) casually placed along the line of sight of the GRB (Behar et al. 2011; Eitan & Behar 2013; Starling et al. 2013; Campana et al. 2015, and references therein).

Campana et al. (2012, hereafter C12) worked out a redshift distribution of the intrinsic column density of a complete sample of 58 *Swift* LGRBs, assuming that the absorption along the IGM is negligible. An increasing trend of $N_{\text{H}}(z)$ emerged with the redshift of the GRB, indicating a lack of non-absorbed GRBs at high redshift or, alternatively, an absorption excess in high- z GRBs. The sample in C12, named BAT6 (Salvaterra et al. 2012), has been selected above a specific threshold of the gamma peak flux

and has a high completeness in redshift, which is fundamental not only to exclude the hypothesis that the $N_{\text{H}}(z) - z$ correlation is due to observative biases, but also to provide the real intrinsic column density values calculated at the GRB redshift, taking account of the scaling cosmological factor $\sim(1+z)^{2.4}$ (see Campana et al. 2014). Recently, it has been shown that a model, which includes the (previously overlooked) contributions of both the intervening objects and the diffuse warm-hot intergalactic medium along the line of sight, could explain qualitatively and quantitatively the observed $N_{\text{H}}(z)$ distribution (Campana et al. 2015; see also Behar et al. 2011; Eitan & Behar 2013; Starling et al. 2013).

We now derive the intrinsic column densities while still neglecting the contribution along the IGM (then allocating all the absorption in excess to the host galaxy), because our purpose is to study the eventual effect on the $N_{\text{H}}(z)$ distribution of a radical change in the Galactic absorption model. We keep the Galactic component fixed, using the Leiden Argentine Bonn (LAB) H_I survey (Kalberla et al. 2005, which is largely adopted by previous literature, including C12) to verify the $N_{\text{H}}(z) - z$ correlation and, in addition, we consider the new Galactic absorption model introduced by Willingale et al. (2013, hereafter W13). This new model completes the H_I distribution of the LAB survey, obtained in radio, adding the contribution of the molecular hydrogen that was extracted from the IR dust absorption in our Galaxy. Therefore, this model reports generally greater values than the LAB survey and, consequently, the new intrinsic column density values are expected to be smaller. Moreover, we derive our

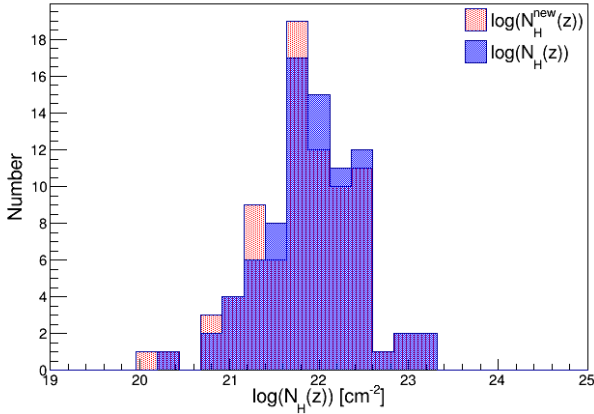


Fig. 1. Intrinsic X-ray absorbing column density distributions of the BAT6ext sample. The dotted (red) histogram is related to the $N_H^{\text{new}}(z)$ distribution, while the thicker dotted (blue) histogram shows the $N_H(z)$ distribution.

results using an extended complete sample of 99 *Swift* LGRBs, which is called BAT6ext and updated to GRB 140703A with the bursts detected by *Swift* that have satisfied the same selection criteria since the construction of the BAT6 (Pescalli et al. 2016). With this extension, we enlarge our sample from 58 to 99 LGRBs, but we slightly lose redshift completeness from 95% to 82%.

2. X-ray absorbing column density distribution

The intrinsic column densities were computed from GRBs X-ray afterglow spectral fits using the UK *Swift* Science Data Centre spectra repository, which provides *Swift*/XRT data suitable for scientific analysis (Evans et al. 2009). All the results, listed in Table A.1, are obtained by selecting a specific time interval when there are no strong spectral variations. This is obtained by strictly selecting time intervals when the hardness ratio (between counts in the 1.5–10 keV and in the 0.3–1.5 keV bands) is constant. For an afterglow spectrum modelled as a power law, this condition ensures that no strong spectral changes are occurring. This is particularly important because any curvature estimate of the X-ray spectrum would reverberate into the column density, which in turn would produce biased values (Butler & Kocevski 2007).

Each spectrum was then fitted by a power law that was absorbed by both a Galactic component, held fixed, and an intrinsic component at the redshift of the GRB, which were modelled with TBABS and ZTBABS, respectively, within XSPEC (version 12.9.0). We note that the results in C12 were obtained by using different models, namely PHABS and ZPHABS (the Leicester *Swift* automatic spectral analysis tool adopted the TBabs models only in 2014¹). Therefore, the intrinsic column density values of the BAT6 GRBs, which are used by both works, are slightly different. Abundances from Wilms et al. (2000) and cross-sections from Verner et al. (1996) were used. Table A.1 shows the intrinsic column density values obtained using the Galactic absorption models provided by both LAB and W13 surveys, the latter being labelled as new. In addition, for the 18 GRBs of the BAT6ext sample that have no spectroscopic redshift, we fixed z to be zero and, consequently, the intrinsic column densities have to be considered as lower limits.

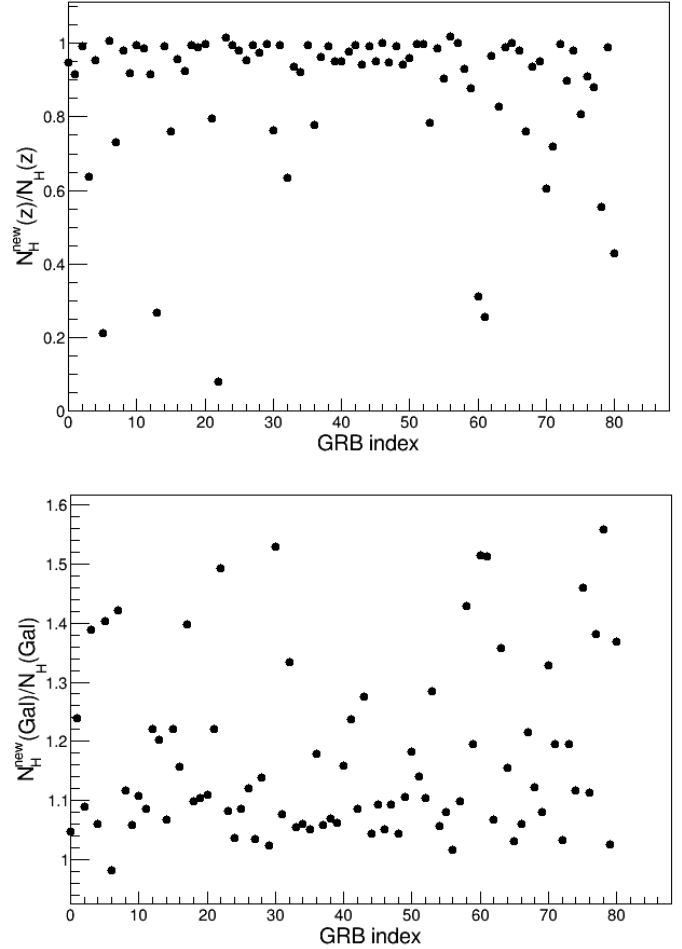


Fig. 2. *Top panel:* ratio between new and old intrinsic column densities, as a function of an index representing the 81 GRBs of Table A.1 with a measure of redshift. *Bottom panel:* the comparable distribution of the ratio between the new (W13) and the old (LAB) Galactic column densities is shown.

The two $N_H(z)$ distributions are shown in Fig. 1 and can be described by a Gaussian, with a mean value of $\log(N_H/\text{cm}^{-2}) = 21.93 \pm 0.54$ and $\log(N_H/\text{cm}^{-2}) = 21.84 \pm 0.61$, for the LAB survey and the W13 Galactic model, respectively. The latter mean value is only slightly smaller but fully consistent with the former. This probably indicates that the change in the Galactic column densities plays a minor role in shaping the intrinsic column density distribution. Both values are consistent with the distribution of column densities reported in C12, which is $\log(N_H/\text{cm}^{-2}) = 21.7 \pm 0.5$. This indicates that the impact of different *Swift*/XRT response matrices² and absorption model (PHABS was adopted) again do not play a major role in the $N_H(z)$ distribution.

In Fig. 2, the ratio between the new and the old intrinsic column densities is reported for all GRBs of Table A.1. While it is clear that every new single value shows a decrease (with a maximum of a factor ~ 5), most of them are only marginally affected by the different Galactic model that was used. This is because, as we expected, the ratio between the new and the old Galactic column densities is, in almost all cases, just slightly greater than 1 (see bottom panel of Fig. 2).

¹ See the change log here www.swift.ac.uk/xrt_spectra/docs.php

² See the release note at http://www.swift.ac.uk/analysis/xrt/files/SWIFT-XRT-CALDB-09_v20.pdf, prepared by Beardmore and collaborators, and references therein.

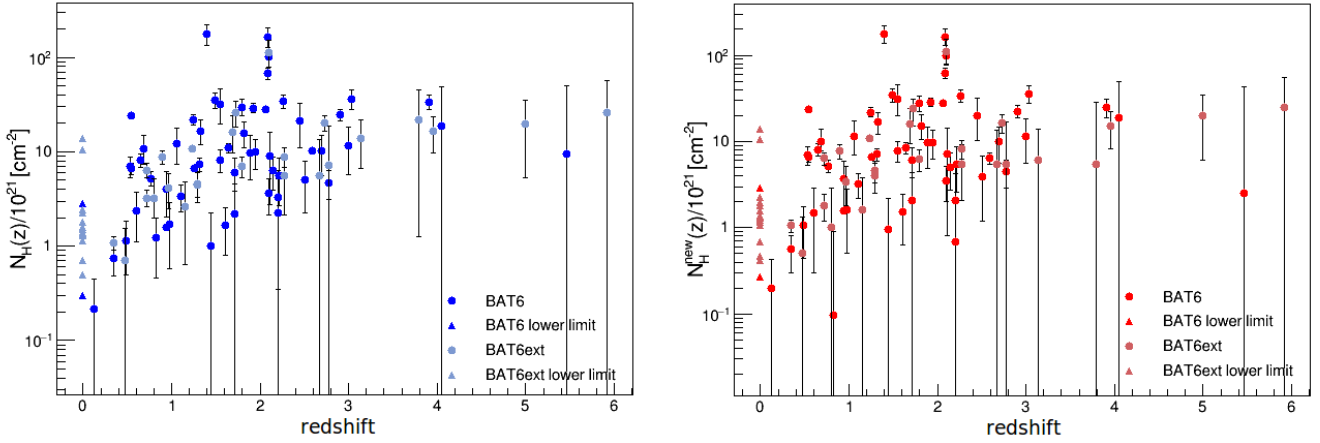


Fig. 3. Intrinsic X-ray column density distribution as a function of redshift for both Galactic models. The *panel on the left* shows the $N_H(z)$ distribution obtained with the LAB survey. We note that the 41 GRBs of the extended sample, shown in a lighter colour with respect to the (blue) dots of the BAT6 sample, do not affect the increasing trend that came out in C12. The *panel on the right* reports the $N_H^{\text{new}}(z)$ distribution, computed with the new W13 Galactic absorption model. The triangles, representing $N_H(z)$ for the 18 GRBs with no redshift measure, should be considered as lower limits, which increase with redshift as $\sim(1+z)^{2.4}$. The error bars were calculated within XSPEC at 90% confidence level ($\Delta\chi^2 = 2.71$), as were the upper limits.

The redshift distributions of the X-ray intrinsic column density are shown in Fig. 3. The increasing trend of $N_H(z)$ with redshift is still evident, not only in the distribution obtained with the LAB Galactic model, but also in the distribution obtained with the new W13 Galactic model. This states that, even if the W13 model changes single values, the whole $N_H(z)$ distribution follows the same increasing trend found in previous works. Furthermore, we note that the supplementary 41 GRBs of the extended BAT6 sample are distributed evenly among the other column densities, with no influence on the observed pattern. Hence, it seems that neither a radical improvement of the Galactic model nor an extension of the sample affect the observed general trend in the $N_H(z)$ distribution. This suggests that those should be considered as second order corrections.

2.1. Monte Carlo simulations and K-S test

It seems that the uniform trend of absorbing column densities at low redshift does not persist at high redshift. To see if there is a real difference, whereby the increasing trend could not be attributed to statistical fluctuations, we cut the sample at $z = 1$ (where the distribution bends) and carried out a Kolmogorov-Smirnov test on the two $N_H(z)$ subsamples. Monte Carlo simulations have been used to provide 10 000 $N_H(z)$ values for each GRB, extracted from a Gaussian distribution that peaked at the intrinsic column density values listed in Table A.1 (and with a width of 1σ). Furthermore, for the 18 GRBs of the BAT6ext sample, which do not have a redshift measure, we also simulated z , extracting the values from the redshift probability density function that was obtained from the other 81 GRBs of the sample. We then scaled the corresponding $N_H(z = 0)$, previously simulated, with the factor $(1+z)^{2.4}$. For each iteration, we computed the K-S test and obtained a distribution of (log) probabilities, with mean values $\log P = -5.02 \pm 0.95$ and $\log P = -4.68 \pm 0.99$, for the intrinsic column densities obtained with the LAB survey and with the W13 Galactic model, respectively. These values are extremely low, and, as we can see in Fig. 4, even the tails of the two distributions never reach a 1% probability.

In addition, a much more brutal technique has been used to include those 18 GRBs in the simulations. We fixed z to be

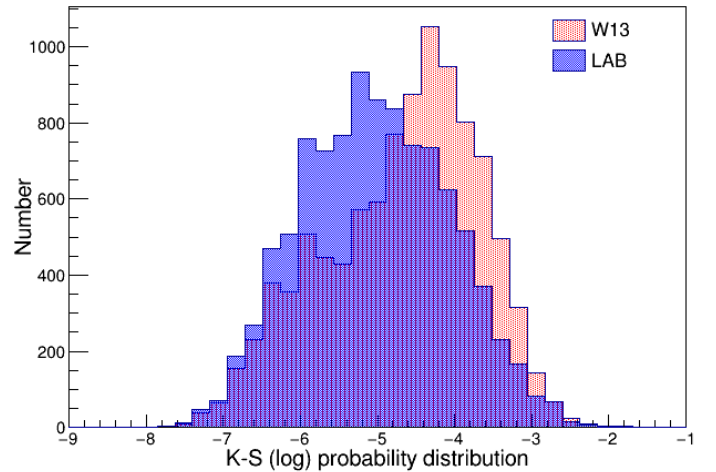


Fig. 4. Kolmogorov-Smirnov (log) probability distributions. The dotted (red) histogram relates to the $N_H^{\text{new}}(z)$ values (obtained with W13 Galactic model), while the thicker dotted (blue) histogram relates to the probability distribution of $N_H(z)$ values (obtained with the LAB survey). In both cases, it is evident that the probability of having the two subsamples descending from the same parent distribution is extremely low.

1 and we simulated the $N_H(z)$ values, first attributing all the $z = 1$ GRBs to the subsample of GRBs with redshift under $z = 1$, and then to the other subsample. Then, the K-S test was carried out. This obtained, in the former case, mean values $\log P = -3.02 \pm 0.52$ and $\log P = -2.60 \pm 0.44$ (for the values obtained with the LAB survey and the W13 model, respectively). In the latter case, mean values $\log P = -3.44 \pm 0.69$ and $\log P = -3.42 \pm 0.66$ were obtained (for LAB survey and W13 model, respectively). This technique was seen as useful proof, demonstrating that even in the most extreme cases, i.e., allocating all the GRBs around the bending redshift, the difference between the two sub-distributions does not seem to be due to adverse statistics.

In all these instances, the probability of having the two subsamples descend from the same parent distribution is extremely low. This indicates that there is a real difference between the

two sub-distributions of column densities $N_{\text{H}}(z)$ in excess of the Galactic value. Therefore the increasing trend of the intrinsic column density as a function of redshift can not be due to statistical fluctuations. With this test, we can not attribute all the absorption in excess to the host galaxy, which leads us to suppose that every single value of these $N_{\text{H}}(z)$ distributions actually includes a contribution that clearly increases with the redshift. One possibility to account for this is to invoke the overlooked component of the absorption along the IGM, which obviously increases with the distance of the GRB (Behar et al. 2011; Campana et al. 2015). Alternatively, one would have to postulate that the GRB environment gets denser with redshift.

In addition, we performed a K-S test to show that GRBs with an upper limit in the measure of $N_{\text{H}}(z)$ are uniformly distributed in redshift. This is important for assuring that the increasing trend of $N_{\text{H}}(z)$ with the redshift is not due to an incapability of measuring low N_{H} at high z . In our sample, a measure of $N_{\text{H}}(z)$ is provided essentially for almost all GRBs, but the eventual presence of most of the upper limit measures at high redshift could give rise to some objections. Hence, we separated into two subsamples the redshift of GRBs with a measure of $N_{\text{H}}(z)$, and the redshift of GRBs with an upper limit (denoted with a $\downarrow$ in Table A.1). We then computed a K-S test for each iteration of the Monte Carlo simulation. We obtained mean probabilities of $P = 30 \pm 8\%$ and $P = 25 \pm 9\%$ (for measures obtained with the LAB Galactic model and W13, respectively), which indicates that the two sets of redshift originate from the same parent distribution.

2.2. Correlations of $N_{\text{H}}(z)$

We already highlighted the fact that the only selection effect of our sample is the peak flux limit and that it is also highly complete in redshift, which means we can compare the rest frame physical properties of GRBs in an unbiased way. In this work, we analyzed the intrinsic X-ray absorbing column density redshift distribution. Moreover, the rest-frame peak energy $E_{\text{p}}^{\text{rest}}$ and the bolometric equivalent isotropic luminosity L_{iso} have been computed by Pescalli et al. (2016) for each GRB of the BAT6ext with a redshift measure.

We now want to study the eventual correlation between $N_{\text{H}}(z)$ and these quantities³. However, given that the sample was selected above a peak flux threshold, $E_{\text{p}}^{\text{rest}}$ and, primarily, L_{iso} have a tight correlation with redshift. This is due to the fact that, for a flux-limited sample, the only sources detected out to great distances are necessarily powerful (see Blundell et al. 1999). However, we do not expect any correlation between $N_{\text{H}}(z)$ and luminosities (or energies), but, since all these quantities correlate with z , we must take the precaution of doing a correlation test that excludes the dependence on redshift. Indeed, the spurious correlation between $N_{\text{H}}(z)$ and $E_{\text{p}}^{\text{rest}}$ (L_{iso}) can be avoided by using a partial correlation analysis. First, we computed the Spearman rank correlation coefficient r (Spearman 1904) for each case, namely r_{12} , r_{13} , and r_{23} , where 1, 2 represent $N_{\text{H}}(z)$ and $E_{\text{p}}^{\text{rest}}$ (L_{iso}), respectively, and 3 is z . Then, we obtained the correlation coefficient between 1 and 2, excluding the redshift bias:

$$r_{12,3} = \frac{r_{12} - r_{13}r_{23}}{\sqrt{1 - r_{13}^2} \sqrt{1 - r_{23}^2}} \quad (1)$$

³ Note that in this work we considered the redshift of GRB 121123A to be zero, and that no energy or luminosity measure has been obtained for GRB 070306 in Pescalli et al. (2016). Hence, the following analysis involves 80 GRBs.

(see Kendall & Stuart 1979, Chap. 27; Padovani 1992).

From the Spearman rank coefficient one can derive the associated probability value to calculate the significance of the correlation. The coefficient between $N_{\text{H}}(z)$ ($N_{\text{H}}^{\text{new}}(z)$) and $E_{\text{p}}^{\text{rest}}$ is $r_{12,3} = 0.056$ (0.093) and the related probability is $P = 62\%$ (41%), while the correlation coefficient between $N_{\text{H}}(z)$ ($N_{\text{H}}^{\text{new}}(z)$) and L_{iso} is $r_{12,3} = 0.052$ (0.075), with probability $P = 65\%$ (51%). Hence, the results of the partial correlation analysis state that $N_{\text{H}}(z)$ does not correlate with $E_{\text{p}}^{\text{rest}}$ and L_{iso} . Being $N_{\text{H}}(z)$ connected to the distribution of matter along the line of sight and $E_{\text{p}}^{\text{rest}}$ and L_{iso} to the GRB prompt emission, this result is not surprising.

Furthermore, our data can be used to test the anticorrelation between the X-ray absorbing column density in excess of the Galactic value (evaluated in the observer rest frame) and redshift, as pointed out by Grupe et al. (2007). In our work, $N_{\text{H}}(z = 0)$ was obtained for the 81 GRBs of the sample with a redshift measure, dividing the intrinsic column density values listed in Table A.1 by the scaling factor $\sim(1+z)^{2.4}$. Afterwards, we computed the Spearman rank correlation coefficient between $N_{\text{H}}(z = 0)$ and $1+z$, obtaining $r = -0.28$ and $r = -0.26$ for the values obtained with the LAB survey and the W13 Galactic model, respectively. The related probabilities are $P = 1\%$ and $P = 1.8\%$, which are too ambiguous to draw any significant conclusion from about the presence of a correlation. Hence, with our results we can not confirm that such a correlation exists.

3. Conclusions

In this work we focused on the eventual effect of a radical change in the Galactic absorption model on the $N_{\text{H}}(z)$ distribution. We derived the intrinsic absorbing column densities, neglecting the absorption along the IGM and using two different Galactic absorption models: the Leiden Argentine Bonn H_I survey and the most recent model provided by W13. The two $N_{\text{H}}(z)$ distributions have a mean value of $\log(N_{\text{H}}/\text{cm}^{-2}) = 21.93 \pm 0.54$ and $\log(N_{\text{H}}/\text{cm}^{-2}) = 21.84 \pm 0.61$, respectively. As expected, the new single intrinsic column density values are smaller, even by a factor ~ 5 . Nonetheless, if on the one hand the new Galactic model considerably affects the single column density values, on the other hand there is no radical change in the distribution as a whole. Therefore, the contribution of Galactic column densities alone, no matter how improved, is not sufficient to change the observed general trend. To explain the cosmological increase of $N_{\text{H}}(z)$ as a function of redshift, it is necessary to include the contribution of both the diffuse intergalactic medium and the intervening systems along the line of sight of the GRBs.

This study clearly leaves unanswered the question about the high column density in GRB spectra. Several studies (Campana et al. 2012, 2015; Schady et al. 2011) show that the intrinsic column density of GRBs is high. A small contribution of absorption can come from the intervening matter (Campana et al. 2015), but this is clearly not enough to account for the observed distribution. Several mechanisms have been envisaged that involve photoionization of the circumburst medium (by the GRB itself or pre-existing) or dense H₂ regions (Campana et al. 2007, 2010, 2012, 2015; Watson et al. 2007, 2013; Schady et al. 2011; Starling et al. 2013; Krongold & Prochaska 2013). Our study demonstrates that the absorbing contribution of our Galaxy plays a minor role in that.

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Appendix A

Table A.1. Intrinsic column densities of the BAT6ext sample.

GRB	z	$N_{\text{H}}(z)$ (10^{21} cm^{-2})	Γ	$N_{\text{H}}(Gal)$ (10^{20} cm^{-2})	Time-interval (s)	Obs. mode (exposure time in ks)	$N_{\text{H}}^{\text{new}}(z)$ (10^{21} cm^{-2})	$N_{\text{H}}^{\text{new}}(Gal)$ (10^{20} cm^{-2})	Γ^{new}
050318	1.44	$\downarrow 1.00^{+1.24}_{-1.00}$	1.94 ± 0.09	1.87	$3000-7 \times 10^4$	PC(23.4)	$\downarrow 0.95^{+1.24}_{-0.95}$	1.96	1.94 ± 0.09
050401	2.90	$24.47^{+3.57}_{-3.41}$	1.86 ± 0.05	4.40	200–2000	WT(1.4)	$22.4^{+3.6}_{-3.4}$	5.45	1.86 ± 0.05
050416A	0.654	$8.07^{+1.40}_{-1.28}$	1.97 ± 0.10	2.45	$400-3 \times 10^5$	PC(92.3)	$8.0^{+1.4}_{-1.3}$	2.67	1.97 ± 0.10
050525A	0.606	$2.36^{+1.38}_{-1.25}$	1.97 ± 0.16	9.08	$6000-3 \times 10^5$	PC(18.7)	$1.5^{+1.4}_{-1.2}$	12.6	1.96 ± 0.16
050802	1.71	$\downarrow 2.20^{+2.32}_{-2.20}$	1.80 ± 0.11	1.85	400–2000	PC(1.5)	$\downarrow 2.1^{+2.3}_{-2.1}$	1.96	1.80 ± 0.11
050922C	2.198	$3.25^{+3.60}_{-2.90}$	2.12 ± 0.10	5.40	$400-3 \times 10^5$	PC(35.1)	$\downarrow 0.69^{+3.60}_{-2.90}$	7.58	2.14 ± 0.10
060206	4.048	$\downarrow 18.89^{+18.89}_{-18.89}$	2.00 ± 0.31	0.93	$700-10^5$	PC(4.5)	$\downarrow 19.0^{+30.0}_{-19.0}$	0.917	1.99 ± 0.30
060210	3.91	$33.77^{+6.20}_{-5.99}$	1.99 ± 0.05	6.08	$3000-10^6$	PC(76.8)	$24.7^{+6.3}_{-6.0}$	8.64	2.00 ± 0.05
060306	1.55	$31.60^{+15.00}_{-11.66}$	1.77 ± 0.26	3.44	$300-925$	PC(0.6)	$31.0^{+15.0}_{-11.0}$	3.84	1.77 ± 0.25
060614	0.125	$\downarrow 0.22^{+0.23}_{-0.22}$	1.77 ± 0.08	1.88	$4400-10^6$	PC(204.3)	$\downarrow 0.20^{+0.23}_{-0.20}$	1.99	1.77 ± 0.07
060814	1.923	$28.93^{+3.41}_{-3.21}$	1.96 ± 0.07	2.33	$250-500$	WT(0.2)	$28.8^{+3.4}_{-3.2}$	2.58	1.96 ± 0.07
060904A	–	$>2.84^{+0.79}_{-0.71}$	3.42 ± 0.37	1.33	$185-225$	WT(0.04)	$>2.85^{+0.79}_{-0.71}$	1.39	3.42 ± 0.37
060908	1.884	$9.73^{+5.28}_{-4.71}$	2.07 ± 0.19	2.35	$200-10^5$	PC(11.6)	$9.6^{+5.3}_{-4.7}$	2.55	2.07 ± 0.19
060912A	0.94	$4.04^{+2.00}_{-2.00}$	1.92 ± 0.18	3.85	$200-2000$	PC(1.7)	$3.7^{+2.2}_{-2.2}$	4.70	1.93 ± 0.18
060927	5.467	$\downarrow 9.39^{+40.86}_{-9.39}$	1.85 ± 0.20	4.61	$100-10^4$	PC(3.5)	$\downarrow 2.5^{+40.8}_{-2.5}$	5.54	1.86 ± 0.20
061007	1.261	$6.71^{+0.35}_{-0.35}$	1.89 ± 0.02	1.77	$90-2000$	WT(1.9)	$6.65^{+0.35}_{-0.35}$	1.89	1.89 ± 0.02
061021	0.346	$0.74^{+0.26}_{-0.26}$	1.89 ± 0.06	4.53	$3000-3 \times 10^5$	PC(83.1)	$0.56^{+0.26}_{-0.26}$	5.53	1.89 ± 0.09
061121	1.314	$7.43^{+1.24}_{-1.18}$	1.80 ± 0.06	4.00	$600-3 \times 10^5$	PC(43.1)	$7.1^{+1.2}_{-1.2}$	4.63	1.80 ± 0.06
061222A	2.088	$67.70^{+9.29}_{-8.68}$	1.99 ± 0.09	9.02	$3 \times 10^4-2 \times 10^5$	PC(49.5)	$65.5^{+9.2}_{-8.6}$	12.60	2.00 ± 0.09
070306	1.496	$34.99^{+6.62}_{-5.98}$	1.77 ± 0.11	2.85	$10^4-4 \times 10^4$	PC(6.0)	$34.8^{+6.6}_{-6.0}$	3.13	1.77 ± 0.11
070328	2.063	$27.93^{+2.00}_{-1.94}$	2.05 ± 0.04	2.61	$350-1000$	WT(0.7)	$27.6^{+2.0}_{-1.9}$	2.88	2.05 ± 0.04
070521	2.087	$164.51^{+35.48}_{-35.48}$	1.70 ± 0.19	2.92	$3000-10^4$	PC(1.9)	$164.0^{+35.0}_{-35.0}$	3.24	1.70 ± 0.19
071020	2.146	$6.27^{+2.42}_{-2.30}$	1.78 ± 0.07	5.12	$70-300$	WT(0.2)	$5.0^{+2.4}_{-2.3}$	6.25	1.79 ± 0.07
071112C	0.82	$1.22^{+0.79}_{-0.76}$	1.73 ± 0.08	7.44	$90-280$	WT(0.2)	$\downarrow 0.1^{+0.8}_{-0.1}$	11.1	1.75 ± 0.08
071117	1.33	$16.54^{+5.39}_{-4.60}$	1.95 ± 0.18	2.33	$2900-6.2 \times 10^4$	PC(18.2)	$16.8^{+4.8}_{-4.6}$	2.52	1.97 ± 0.20
080319B	0.937	$1.59^{+0.14}_{-0.14}$	1.73 ± 0.02	1.11	$800-2000$	WT(0.9)	$1.58^{+0.14}_{-0.14}$	1.15	1.74 ± 0.02
080319C	1.95	$9.90^{+3.71}_{-3.42}$	1.56 ± 0.10	2.20	$200-3 \times 10^5$	PC(24.3)	$9.7^{+3.7}_{-3.4}$	2.39	1.56 ± 0.10
080413B	1.10	$3.34^{+1.07}_{-1.02}$	1.89 ± 0.08	3.06	$6 \times 10^3-10^6$	PC(98.5)	$3.18^{+1.06}_{-1.00}$	3.43	1.89 ± 0.08
080430	0.767	$5.20^{+0.91}_{-0.86}$	1.96 ± 0.09	0.96	$5000-3 \times 10^5$	PC(44.6)	$5.18^{+0.91}_{-0.86}$	0.99	1.96 ± 0.09
080602	1.82	$15.69^{+3.15}_{-4.61}$	1.91 ± 0.14	3.51	$200-800$	PC(0.6)	$15.3^{+3.1}_{-4.6}$	4.00	1.91 ± 0.14
080603B	2.69	$10.12^{+4.81}_{-4.50}$	1.79 ± 0.10	1.24	$100-250$	WT(0.2)	$10.1^{+4.8}_{-4.5}$	1.27	1.79 ± 0.10

Notes. A solid horizontal line separates BAT6 GRBs from the extended sample. GRBs with an updated redshift estimate with respect to C12 are marked in italics. The $N_{\text{H}}(z)$ values for the GRBs with no redshift are evaluated at $z = 0$. The results with an upper limit of column density are labelled by $\downarrow$. ^(a) This GRB has an inferred photometric redshift of $z = 2.7$. Since this is not a spectroscopic measure in data analysis we considered z to be zero, but for completeness we reported also the intrinsic column densities computed at $z = 2.7$. ^(b) In this case the column density fitted was an upper limit (a potential $\downarrow$), but being relative to a GRB with no redshift, this should be also a lower limit. The symbol $\approx$ indicates this ambiguous situation.

Table A.1. continued.

GRB	z	$N_{\text{H}}(z)$ (10^{21} cm^{-2})	Γ	$N_{\text{H}}(\text{Gal})$ (10^{20} cm^{-2})	Time-interval (s)	Obs. mode (exposure time in ks)	$N_{\text{H}}^{\text{new}}(z)$ (10^{21} cm^{-2})	$N_{\text{H}}^{\text{new}}(\text{Gal})$ (10^{20} cm^{-2})	Γ^{new}
080605	1.64	$11.01^{+1.33}_{-1.27}$	1.66 ± 0.04	6.67	100–800	WT(0.6)	$8.4^{+1.3}_{-1.3}$	10.2	1.67 ± 0.04
080607	3.036	$36.07^{+8.79}_{-8.14}$	2.04 ± 0.12	1.69	$4000-6 \times 10^4$	PC(19.7)	$35.9^{+8.8}_{-8.2}$	1.82	2.04 ± 0.12
080613B	–	$\geq 0.3^b$	1.21 ± 0.10	3.17	105–190	WT(0.1)	$\geq 0.27^b$	3.55	1.20 ± 0.11
080721	2.591	$10.11^{+0.92}_{-0.91}$	1.72 ± 0.02	6.95	100–2000	WT(1.3)	$6.40^{+0.93}_{-0.91}$	9.27	1.73 ± 0.02
080804	2.20	$\downarrow 2.24^{+2.89}_{-2.24}$	1.73 ± 0.09	1.63	$200-3 \times 10^5$	PC(41.0)	$\downarrow 2.1^{+2.9}_{-2.1}$	1.72	1.73 ± 0.09
080916A	0.689	$10.73^{+4.16}_{-3.33}$	2.09 ± 0.26	1.84	$2 \times 10^4-10^6$	PC(110.2)	$9.9^{+4.0}_{-2.2}$	1.95	2.04 ± 0.18
081007	0.53	$6.94^{+1.81}_{-1.59}$	1.94 ± 0.16	1.37	$5000-4 \times 10^5$	PC(46.9)	$6.9^{+1.8}_{-1.6}$	1.44	1.94 ± 0.16
081121	2.512	$5.01^{+2.86}_{-2.74}$	1.85 ± 0.07	4.03	$3000-2 \times 10^6$	PC(128.0)	$3.9^{+2.8}_{-2.7}$	4.75	1.85 ± 0.07
081203A	2.10	$3.63^{+1.54}_{-1.48}$	1.69 ± 0.05	1.71	200–600	WT(0.4)	$3.5^{+1.6}_{-1.5}$	1.81	1.69 ± 0.05
081221	2.26	$34.16^{+5.66}_{-5.26}$	1.96 ± 0.09	2.03	300–500	WT(0.2)	$33.9^{+5.6}_{-5.3}$	2.17	1.96 ± 0.09
081222	2.77	$4.74^{+1.62}_{-1.59}$	1.93 ± 0.04	2.23	60–1000	WT(0.8)	$4.5^{+1.6}_{-1.5}$	2.37	1.93 ± 0.04
090102	1.547	$8.20^{+2.24}_{-2.14}$	1.66 ± 0.08	4.05	$700-7 \times 10^4$	PC(19.3)	$7.8^{+2.3}_{-2.1}$	4.69	1.67 ± 0.08
090201	2.1	$101.21^{+21.08}_{-21.93}$	1.81 ± 0.14	4.94	3000–6000	WT, (0.2)	$99.0^{+21.0}_{-19.0}$	6.11	1.81 ± 0.15
090424	0.544	$6.58^{+0.90}_{-0.84}$	1.95 ± 0.08	1.86	$2000-3 \times 10^6$	PC(213.1)	$6.55^{+0.89}_{-0.83}$	2.02	1.95 ± 0.08
090709A	1.8	$29.45^{+6.20}_{-5.61}$	1.84 ± 0.11	6.62	$4000-1.5 \times 10^4$	PC(4.4)	$27.7^{+6.2}_{-5.6}$	8.44	1.84 ± 0.11
090715B	3.00	$11.71^{+6.48}_{-6.02}$	1.92 ± 0.12	1.34	$4000-10^5$	PC(28.8)	$11.6^{+6.5}_{-6.0}$	1.40	1.92 ± 0.12
090812	2.452	$21.12^{+11.56}_{-9.97}$	2.08 ± 0.24	2.26	$10^4-7 \times 10^4$	PC(8.6)	$20.1^{+11.6}_{-9.9}$	2.47	2.07 ± 0.24
090926B	1.24	$21.90^{+2.77}_{-2.60}$	1.88 ± 0.08	1.92	130–300	WT(0.2)	$21.9^{+2.8}_{-2.6}$	2.02	1.88 ± 0.08
091018	0.971	$1.69^{+1.21}_{-1.11}$	1.76 ± 0.11	2.81	150–1000	PC(0.9)	$1.6^{+1.2}_{-1.1}$	3.07	1.76 ± 0.11
091020	1.71	$6.05^{+2.22}_{-2.12}$	1.78 ± 0.10	1.39	200–400	WT(0.2)	$6.0^{+2.2}_{-2.1}$	1.45	1.78 ± 0.10
091127	0.49	$1.15^{+0.69}_{-0.65}$	1.77 ± 0.12	2.81	$6000-2 \times 10^4$	PC(1.6)	$1.08^{+0.69}_{-0.64}$	3.11	1.74 ± 0.11
091208B	1.06	$12.10^{+5.90}_{-4.69}$	2.05 ± 0.25	4.86	200–600	PC(0.4)	$11.6^{+5.8}_{-4.6}$	5.75	2.06 ± 0.25
100615A	1.398	$176.15^{+46.54}_{-41.09}$	2.25 ± 0.34	3.28	200–2000	PC(1.4)	$176.0^{+47.0}_{-40.0}$	3.74	2.25 ± 0.34
100621A	0.542	$23.88^{+1.75}_{-1.17}$	2.67 ± 0.09	2.89	130–200	WT(0.1)	$23.8^{+1.7}_{-1.7}$	3.19	2.68 ± 0.09
100728B	2.106	$9.05^{+7.18}_{-6.32}$	1.98 ± 0.19	6.17	$4000-3 \times 10^4$	PC(8.5)	$7.1^{+7.1}_{-6.3}$	7.93	1.99 ± 0.19
110205A	2.22	$5.57^{+3.05}_{-2.88}$	2.02 ± 0.11	1.61	$5000-5 \times 10^4$	PC(18.8)	$5.5^{+3.1}_{-2.9}$	1.70	2.02 ± 0.11
110503A	1.613	$1.67^{+0.90}_{-0.87}$	1.71 ± 0.05	2.63	200–700	WT(0.5)	$1.51^{+0.92}_{-0.87}$	2.84	1.71 ± 0.05
110709A	–	$> 10.49^{+1.31}_{-1.32}$	1.92 ± 0.12	1.54	600–9000	PC(3.9)	$> 10.7^{+1.3}_{-1.3}$	1.61	1.93 ± 0.13
110709B	< 4	$> 1.50^{+0.20}_{-0.19}$	2.04 ± 0.06	5.58	$10^4-2 \times 10^5$	PC(45.6)	$> 1.37^{+0.20}_{-0.19}$	6.91	2.04 ± 0.06
110915A	–	$> 2.43^{+1.06}_{-0.92}$	1.90 ± 0.25	1.61	5000–8000	PC(1.5)	$> 1.95^{+1.08}_{-0.93}$	6.67	1.90 ± 0.25
111008A	4.99	$19.63^{+15.40}_{-14.38}$	1.81 ± 0.10	0.98	$5000-3 \times 10^4$	PC(4.8)	$20.0^{+15.0}_{-14.0}$	0.992	1.81 ± 0.10
111228A	0.714	$6.39^{+1.12}_{-1.03}$	2.05 ± 0.09	2.96	$4000-2 \times 10^6$	PC(144.8)	$6.4^{+1.0}_{-1.1}$	3.25	2.05 ± 0.08
120102A	–	$> 2.40^{+0.36}_{-0.34}$	1.91 ± 0.08	10.30	$4000-7 \times 10^4$	PC(18.1)	$> 1.98^{+0.36}_{-0.34}$	14.7	1.91 ± 0.08
120116A	–	$> 1.32^{+0.75}_{-0.66}$	1.94 ± 0.23	4.64	300–6000	PC(1.5)	$> 1.24^{+0.76}_{-0.66}$	5.53	1.94 ± 0.23
120119A	1.728	$25.83^{+8.72}_{-7.50}$	1.65 ± 0.13	7.91	$7000-2 \times 10^5$	PC(18.9)	$24.0^{+7.3}_{-8.9}$	11.30	1.67 ± 0.13
120326A	1.798	$7.07^{+1.79}_{-1.71}$	1.79 ± 0.06	5.24	$4000-10^5$	PC(21.4)	$6.2^{+1.8}_{-1.7}$	6.26	1.80 ± 0.06
120703A	–	$> 1.57^{+0.44}_{-0.41}$	1.82 ± 0.14	1.21	$3000-5 \times 10^4$	PC(9.8)	$> 1.58^{+0.45}_{-0.41}$	1.27	1.78 ± 0.13

Table A.1. continued.

GRB	z	$N_{\mathrm{H}}(z)$ (10^{21} cm^{-2})	Γ	$N_{\mathrm{H}}(\text{Gal})$ (10^{20} cm^{-2})	Time-interval (s)	Obs. mode (exposure time in ks)	$N_{\mathrm{H}}^{\text{new}}(z)$ (10^{21} cm^{-2})	$N_{\mathrm{H}}^{\text{new}}(\text{Gal})$ (10^{20} cm^{-2})	Γ^{new}
120729A	0.8	$3.21^{+1.94}_{-1.29}$	1.68 ± 0.13	14.2	300–7000	PC(2.8)	$\downarrow 1.0^{+1.9}_{-1.0}$	21.5	1.70 ± 0.13
120802A	3.796	$21.56^{+20.29}_{-20.29}$	1.83 ± 0.17	9.65	$100-1.5 \times 10^4$	PC(5.1)	$\downarrow 5.5^{+23.2}_{-5.5}$	14.6	1.87 ± 0.17
120811C	2.671	$\downarrow 5.60^{+8.02}_{-5.60}$	1.58 ± 0.15	2.09	200–1100	PC(0.9)	$\downarrow 5.4^{+5.0}_{-5.4}$	2.23	1.58 ± 0.15
120907A	0.97	$4.11^{+1.77}_{-1.62}$	1.81 ± 0.13	5.35	280–8000	PC(3.2)	$3.4^{+1.8}_{-1.6}$	7.26	1.82 ± 0.13
121123A ^a	–	$>0.49^{+0.30}_{-0.28}$	1.85 ± 0.11	4.01	4500 – 42000	PC(14.9)	$>0.42^{+0.30}_{-0.28}$	4.75	1.85 ± 0.11
121125A	–	$8.64^{+5.11}_{-4.77}$	1.82 ± 0.10	1.27	$4000-3.5 \times 10^4$	PC(13.2)	$7.37^{+5.12}_{-4.78}$	1.32	1.82 ± 0.10
121209A	2.1	$>1.15^{+0.33}_{-0.31}$	2.04 ± 0.12	3.81	5000–7000	PC(1.4)	$>1.15^{+0.33}_{-0.31}$	4.40	2.04 ± 0.12
130420A	1.297	$112.15^{+40.69}_{-33.11}$	1.79 ± 0.25	1.26	$1000-2 \times 10^5$	PC(35.5)	$111.0^{+41.0}_{-33.0}$	1.30	1.79 ± 0.25
130427A	0.34	$4.60^{+1.35}_{-1.28}$	2.16 ± 0.10	1.80	$2 \times 10^4-2.5 \times 10^6$	PC(165.0)	$4.6^{+1.3}_{-1.3}$	1.91	2.16 ± 0.10
130427B	2.78	$1.08^{+0.18}_{-0.18}$	1.68 ± 0.04	4.40	$200-2 \times 10^4$	PC(5.3)	$1.06^{+0.18}_{-0.18}$	5.35	1.68 ± 0.04
130502A	–	$\downarrow 7.11^{+11.48}_{-7.11}$	1.71 ± 0.19	2.49	$100-10^4$	PC(1.7)	$\downarrow 5.4^{+11.4}_{-5.4}$	2.73	1.71 ± 0.19
130505A	2.27	$>1.80^{+0.98}_{-0.85}$	2.05 ± 0.30	3.59	$3000-10^5$	PC(16.3)	$>1.79^{+0.99}_{-0.84}$	4.03	2.05 ± 0.30
130527A	–	$8.86^{+2.08}_{-2.01}$	1.77 ± 0.05	3.73	$450-2000$	PC(1.2)	$8.3^{+2.1}_{-2.0}$	4.34	1.77 ± 0.05
130606A	5.913	$>1.37^{+0.74}_{-0.65}$	1.46 ± 0.17	1.98	$5000-6 \times 10^4$	PC(16.5)	$>1.32^{+0.75}_{-0.66}$	2.14	1.46 ± 0.17
130609B	–	$\downarrow 26.28^{+30.77}_{-26.28}$	1.87 ± 0.13	1.45	5000–8000	PC(2.4)	$\downarrow 25.0^{+31.0}_{-25.0}$	1.51	1.87 ± 0.13
130701A	1.155	$>1.30^{+0.37}_{-0.34}$	1.87 ± 0.12	6.93	$1850-3.5 \times 10^4$	PC(6.6)	$>1.30^{+0.37}_{-0.35}$	9.20	1.87 ± 0.12
130803A	–	$2.64^{+2.20}_{-2.00}$	1.77 ± 0.14	4.15	$4000-6 \times 10^4$	PC(11.3)	$\downarrow 1.6^{+2.2}_{-1.6}$	5.12	1.77 ± 0.14
130831A	0.479	$>13.83^{+2.88}_{-2.88}$	2.67 ± 0.31	4.80	$11000-1.2 \times 10^5$	PC(11.5)	$>14.0^{+3.0}_{-2.7}$	5.74	2.67 ± 0.31
130907A	1.238	$\downarrow 0.71^{+0.82}_{-0.71}$	1.74 ± 0.14	0.997	$7000-1.5 \times 10^5$	PC(37.3)	$\downarrow 0.51^{+0.82}_{-0.51}$	1.03	1.74 ± 0.14
131030A	1.293	$10.86^{+0.90}_{-0.87}$	1.82 ± 0.04	4.70	4000–90000	PC(9.6)	$10.84^{+0.92}_{-0.86}$	5.62	1.82 ± 0.04
131105A	1.686	$4.45^{+1.62}_{-1.52}$	1.92 ± 0.10	2.91	350–1100	PC(0.8)	$4.0^{+1.6}_{-1.5}$	3.25	1.92 ± 0.10
140102A	–	$16.30^{+8.03}_{-6.66}$	1.83 ± 0.21	2.71	950–2300	WT(1.3)	$16.0^{+8.0}_{-6.6}$	3.04	1.83 ± 0.21
140206A	2.73	$>0.70^{+0.22}_{-0.21}$	1.77 ± 0.07	4.79	10^4-10^5	PC(22.7)	$>0.68^{+0.22}_{-0.21}$	6.99	1.77 ± 0.07
140215A	–	$20.28^{+4.07}_{-3.89}$	1.81 ± 0.06	9.71	300–700	PC(0.4)	$16.4^{+4.1}_{-3.9}$	13.1	1.83 ± 0.06
140419A	3.956	$>1.41^{+1.01}_{-0.88}$	1.65 ± 0.23	3.53	1500–20000	PC(7.8)	$>1.08^{+1.02}_{-0.88}$	3.93	1.65 ± 0.23
140506A	0.889	$16.57^{+6.75}_{-6.75}$	1.83 ± 0.07	7.67	$5000-10^5$	PC(18.9)	$15.1^{+7.1}_{-6.8}$	10.6	1.84 ± 0.07
140512A	0.725	$8.74^{+1.45}_{-1.45}$	1.77 ± 0.08	9.43	$4500-7 \times 10^4$	PC(15.0)	$7.7^{+1.6}_{-1.4}$	14.7	1.78 ± 0.08
140619A	–	$3.24^{+0.64}_{-0.62}$	1.81 ± 0.06	1.58	$1700-25000$	PC(6.2)	$1.80^{+0.65}_{-0.61}$	1.65	1.83 ± 0.06
140628A	–	$>2.24^{+0.54}_{-0.50}$	2.09 ± 0.15	2.92	$750-2.2 \times 10^4$	PC(6.9)	$>2.25^{+0.54}_{-0.50}$	3.22	2.09 ± 0.15
140629A	2.275	$>0.50^{+0.54}_{-0.47}$	1.86 ± 0.20	0.909	$4000-1.5 \times 10^4$	PC(5.1)	$>0.47^{+0.55}_{-0.47}$	0.932	1.86 ± 0.20
140703A	3.14	$5.56^{+3.67}_{-3.41}$	1.92 ± 0.12	9.36	$3800-7 \times 10^4$	PC(11.9)	$5.5^{+3.7}_{-3.4}$	12.8	1.91 ± 0.12
		$14.00^{+7.80}_{-7.27}$	1.80 ± 0.09				$\downarrow 6.0^{+7.8}_{-6.0}$		1.82 ± 0.09