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LETTER TO THE EDITOR

ISO-Chal 52: a weakly-accreting young stellar object with a dipper light curve.*

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ABSTRACT

We report on the discovery of periodic dips in the multiband lightcurve of ISO-Chal 52, a young stellar object in the Chamaeleon I dark cloud. This is one among the peculiar objects that display very low or negligible accretion both in their UV continuum and spectral lines, although they present a remarkable infrared excess emission characteristic of optically-thick circumstellar disks. We have analyzed a VLT/X-Shooter spectrum with the tool ROTFIT to determine the stellar parameters. The latter, along with photometry from our campaign with the REM telescope and from the literature, have allowed us to model the spectral energy distribution and to estimate the size and temperature of the inner and outer disk. From the rotational period of the star/disk system of 3.45 days we estimate a disk inclination of 36° . The depth of the dips in different bands has been used to gain information about the occulting material. A single extinction law is not able to fit the observed behavior, while a two-component model of a disk warp composed of a dense region with a gray extinction and an upper layer with an ISM-type extinction provides a better fit of the data.

Key words. stars: pre-main sequence – stars: low-mass – accretion, accretion disks – protoplanetary disks

1. Introduction

A key issue in the study of planet formation is to explain how optically thick accretion disks surrounding the youngest solar-mass stars and giving rise to a remarkable infrared (IR) excess (class II IR sources) evolve into optically thin debris disks, passing through a phase where only a mild (or null) IR excess is visible (class III sources). Normally, thick disks are observed around the classical T Tauri (CTT) stars, which display strong emission lines produced by mass accretion, while the weak-line T Tauri (WTT) stars, with negligible signatures of accretion, very often show up as class III sources (Hartmann et al. 2005; Lada et al. 2006, and references therein). Generally, these disks persist for a few million years, during which part of the material is accreted onto the star, part is lost via outflows and photo-evaporation (e.g., Ercolano & Pascucci 2017; Nisini et al. 2018), and part condenses into centimeter-sized and larger bodies or planetesimals (e.g., Testi et al. 2014). A possible intermediate stage of T Tauri disk evolution is observationally identified with the so-called *transitional disks* (TDs), which are characterized by inner holes and gaps in their

dust distribution (e.g., Espaillat et al. 2014, and references therein).

A further category of objects is recently emerging with apparently very little or no evidence for accretion in optical ($\lambda > 3400 \text{ \AA}$) spectra, yet with the near-infrared (NIR) emission characteristic of optically thick dust in the inner (few AU) regions of the disk, so that their spectral energy distribution looks like that of class II sources (e.g., Wahhaj et al. 2010; Alcalá et al. 2019; Thanathibodee et al. 2019). The existence of such objects might be explained by slightly different timescales for the decline of disk and accretion processes in young stars (e.g., Fedele et al. 2010). Another possibility is that the accretion is highly variable and occurs mainly during bursts (e.g., Cody et al. 2017). Weak accretion, in general, is not easily detectable in the region of the Balmer jump, hence other diagnostics like modeling of the $H\alpha$ line profile (e.g., Espaillat et al. 2008; Thanathibodee et al. 2019) and/or measures of excess emission in near-UV/far-UV spectra (e.g., Alcalá et al. 2019, and references therein) are necessary in these cases of very low accretion rates.

Recently, in our X-Shooter surveys of the Chamaeleon I (Cha I, Manara et al. 2016, 2017) and Lupus star forming regions (Alcalá et al. 2014, 2017), we have detected sev-

* Based on observations collected at the ESO REM telescope (La Silla, Chile) and at the ESO VLT (ID 084.C-1095).

eral such weak accretors. Thereby we have confirmed that some of these disk-bearing objects lack the X-Shooter continuum UV excess, typical of accreting objects. The recent ALMA surveys of Pascucci et al. (2016) in Cha I and of Ansdell et al. (2016b) in Lupus show that the disks around some of these objects are bright in the sub-mm and still host a large amount of dust.

YSOs display luminosity variations on different timescales due to geometric and intrinsic effects. The variability of non-accreting objects is mainly related to stellar magnetic activity (cold photospheric spots, flares, etc.). For accreting objects, a variety of processes, including hot spots, variable circumstellar extinction and/or burst of accretion, can be at the origin of such variations. A way to efficiently characterize such processes is by photometric multi-band imaging techniques. Long-term ground-based observations (e.g., Herbst et al. 2002; Frasca et al. 2009) and, more recently, high-precision, high-cadence space photometry (e.g., Venuti et al. 2017; Stauffer et al. 2017), enabled the exploration of different scenarios for the physical process involved in the variability of YSOs. An intriguing observed behavior is the presence of recurrent luminosity dips that are likely due to the periodic occultation of the central star by the magnetically-warped inner disk edge (e.g., Bouvier et al. 1999, 2003; McGinnis et al. 2015).

In this letter we report on the discovery of quasi-periodic luminosity dips in the weakly accreting object ISO-ChaI 52 (= BYB 18 = 2MASS J11044258-7741571) that were observed simultaneously from the optical to the NIR with the Rapid Eye Mount (REM) telescope at the La Silla observatory. We have also analyzed an intermediate-resolution spectrum of this source ($R \simeq 18\,000$) taken on 18 Dec 2010 with X-Shooter at the ESO VLT (program ID 084.C-1095) with the aim of deriving stellar parameters in support of this study.

2. Observations

The photometric observations were performed with the 60-cm robotic REM telescope located at the ESO-La Silla Observatory (Chile), on 80 nights from 3 April to 1 October 2019. By means of a dichroic, REM feeds simultaneously two cameras at the two Nasmyth focal stations, one for the NIR (REMIR) and one for the optical (ROSS2). The cameras have nearly the same field of view of about $10' \times 10'$ and use wide-band filters (J , H , and K' for REMIR and Sloan/SDSS g' , r' , i' , and z' for ROSS2). Due to a technical failure in the REMIR camera, we have NIR data only for the first 20 nights of the campaign. In total, we collected 260, 277, 280, 285, 103, 102, and 108 usable images in g' , r' , i' , z' , J , H , and K' bands, respectively. Exposure times were 180 sec for ROSS2, which acquires simultaneously images in the four Sloan bands, while five ditherings of 7 sec each were adopted for each filter of REMIR. Details on the reduction of the photometric data are reported in Appendix A. The reduction of the X-Shooter spectrum is performed and described in Manara et al. (2016).

3. Results

3.1. Stellar parameters and accretion diagnostics

We analyzed the X-Shooter spectrum with the code ROTFIT (Frasca et al. 2017), which allows us to derive

Table 1. Stellar parameters of ISO-ChaI 52 derived in this work.

T_{eff} (K)	$\log g$ (dex)	RV (km s ⁻¹)	$v \sin i$ (km s ⁻¹)	R_* ($R_{\odot}$)	L_* ($L_{\odot}$)	M_* ($M_{\odot}$)
3195±70	4.20±0.35	18.1±3.6	13±6	1.14±0.04	0.123±0.011	0.20±0.05

the atmospheric parameters (T_{eff} , $\log g$), the radial velocity (RV), the projected rotational velocity ($v \sin i$), and the veiling (r). Details are given in Appendix B. The results of the ROTFIT analysis are summarized in Table 1. The veiling in the red spectral regions analyzed by us is $r < 0.2$. Our T_{eff} is fully consistent, within the errors, with the value of 3270 K that was derived by Manara et al. (2016) from the M4 spectral type (SpT) and the SpT- T_{eff} calibration relation of Luhman et al. (2003). Moreover, this spectral type corresponds to $T_{\text{eff}}=3200$ K and $T_{\text{eff}}=3190$ K according to the SpT- T_{eff} relations of Pecaut & Mamajek (2013) and Herczeg & Hillenbrand (2014), respectively, which are in perfect agreement with our T_{eff} determination. Both the H α line width at 10% of the peak, $W_{10\%} = 180 \pm 18$ km s⁻¹, and the H α flux of 1.2×10^6 erg cm⁻² s⁻¹, which we measure on the X-Shooter spectrum, indicate ISO-ChaI 52 as a non-accreting object (see, e.g., Fig. 11 in Frasca et al. 2015) as already noticed by Manara et al. (2016, 2017). They observed a small excess in the Balmer continuum (see Fig. C.1 in Manara et al. 2016), which translates into an upper limit of -10.34 (rescaled to the *Gaia* DR2 distance, Manara et al. 2019) for $\log \dot{M}_{\text{acc}}$ ($M_{\odot} \text{yr}^{-1}$), and considered it as a doubtful accretor, at a level compatible with the typical chromospheric emission line activity. We note that the profiles of the Balmer lines and Ca II K line are all rather narrow and symmetric (see Fig. C.1) with no sign of redshifted absorption components or reversals that are frequently observed in the line profiles of accretors (e.g., Thanathibodee et al. 2019; de Albuquerque et al. 2020). The only notable feature is a wing emission, which is stronger in the red side of the H β and H γ profiles and extends up to $\simeq 200$ km s⁻¹. This is reminiscent of mass flows or turbulence in the upper atmosphere, which are mostly observed during flare events (see, e.g., Doyle et al. 1988). Moreover, as shown in Fig. C.1, the He I lines $\lambda\lambda 5876, 6876$ Å are not clearly detected and the Ca II IRT lines display only a filling in of their cores (Fig. C.2) that resembles a purely chromospheric emission. These diagnostics support ISO-ChaI 52 as a non-accreting or, based on the evidence for UV continuum excess described above, a weakly-accreting object.

3.2. Spectral energy distribution

We use the average values of $g'r'i'z'JHK'$ magnitudes outside the dips (Fig. 2) to construct the optical/NIR spectral energy distribution (SED). We extended the SED to the blue side and to mid-infrared (MIR) and far-infrared (FIR) wavelengths by adding flux values from the literature. These data are quoted in Table B.1.

We adopted the BT-Settl spectrum (Allard et al. 2012) with $T_{\text{eff}}=3200$ K, $[\text{Fe}/\text{H}]=0.0$, and $\log g=4.0$, i.e. the one with the parameters closest to those found with ROTFIT, to fit the optical-NIR portion (from B to J band) of the

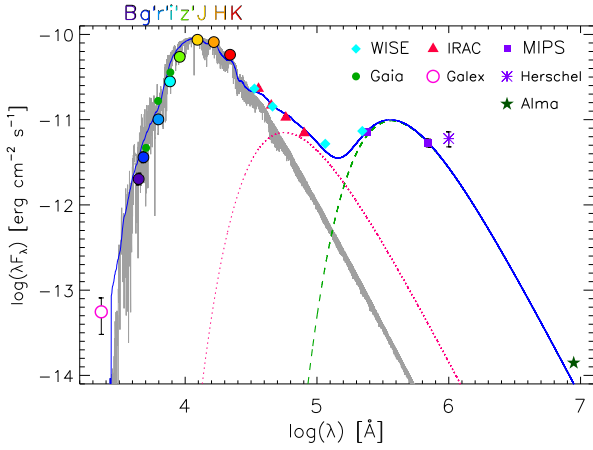


Fig. 1. Spectral energy distribution of ISO-ChaI 52. Mid- and far-infrared fluxes are shown with different symbols, as indicated in the legend. The BT-Settl spectrum (Allard et al. 2012) that provides the best fit to the star photosphere is shown by a gray line. The two black bodies with $T = 650\text{ K}$ and $T = 100\text{ K}$ that fit the mid- and far-infrared disk emission are shown by the red-dotted and green-dashed lines, respectively. The continuous blue line displays the sum of the smoothed photospheric template and the two black bodies.

SED (Fig. 1). Details on the fitting procedure can be found in Appendix B and some derived parameters are reported in Table 1.

The GALEX/NUV flux is clearly in excess with respect to the photosphere. NUV flux excess was observed in older M-type stars and was ascribed to stellar magnetic activity (e.g., Stelzer et al. 2013), but it could be also indicative of a mild accretion onto the central star. If this were the case, ISO-ChaI 52 would be somewhat similar to MY Lup, for which accretion is clearly displayed only by UV line and continuum emission revealed by HST (Alcalá et al. 2019).

The SED also displays a significant IR excess at wavelengths longer than about $3\mu\text{m}$ that is produced by the circumstellar disk. The IR excess can be fitted reasonably well with thermal emission from two sources with two different temperatures. The MIR emission, which is related to the warmer part of the disk, is fitted with a black body of 650 K with an emitting area 53 times larger than the stellar surface (red dotted line in Fig. 1), while the FIR emission is reproduced by a source with $T = 100\text{ K}$ and an area 1.3×10^5 times larger than the stellar surface (green dashed line). The excess IR luminosity estimated as the sum of these two black-body components, amounts to $L_{\text{disk}} \simeq 0.026 L_{\odot}$, i.e. about 22% of the stellar luminosity. This is in close agreement with what has been found for accreting objects in Lupus (Merín et al. 2008) and Cha II (Alcalá et al. 2008), which all display a fractional disk luminosity $L_{\text{disk}}/L_{\star} > 8\%$ that is the limit between passive reprocessing disks and accretion disks proposed by Kenyon & Hartmann (1987). The Hertzsprung-Russell (HR) diagram is shown in Fig. B.2 along with the pre-main sequence evolutionary tracks and isochrones by Baraffe et al. (2015). The position of ISO-ChaI 52 is between the isochrones at 1 and 3 Myr and close to the evolutionary track for a $0.2 M_{\odot}$ star.

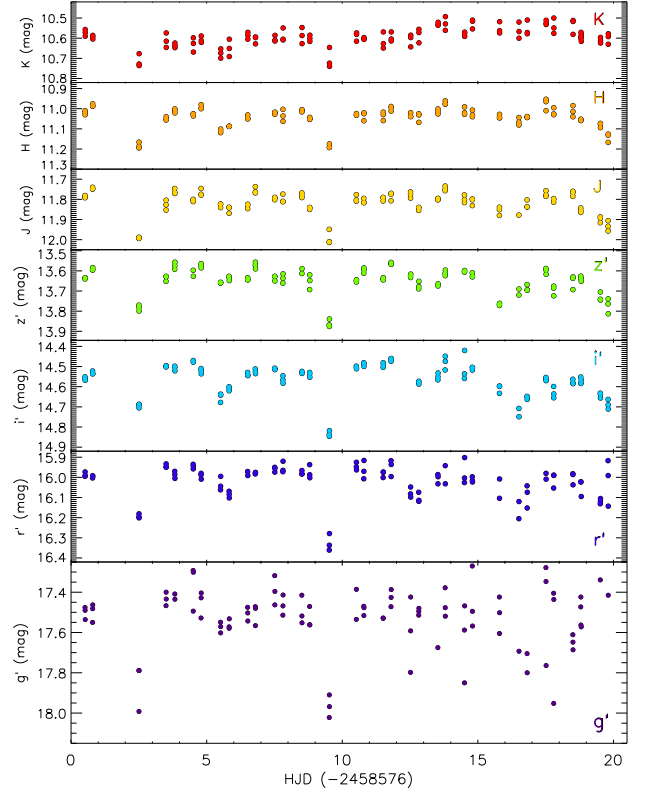


Fig. 2. REM multi-band optical/NIR lightcurves of ISO-ChaI 52 for the first 20 days of the campaign. The scales of the vertical axes have been chosen so as to keep the magnitude ranges constant for a better display of the variation amplitudes.

3.3. REM lightcurves

ISO-ChaI 52 displays quasi-periodic dimmings throughout the photometric monitoring, which seem to occur about every seven days. The light curves observed in the $g'r'i'z'JHK'$ bands are highly correlated and the depth of the dips increases systematically for bluer bands (Fig. 2). These features have been observed in several YSOs both from ground-based and space observations (e.g., McGinnis et al. 2015; Rodríguez et al. 2017; Stauffer et al. 2017, and references therein) and have been ascribed to accretion-driven warps in highly inclined inner disks, which may be misaligned with respect to the outer disks (e.g., Bouvier et al. 2003; Ansdell et al. 2016a; Alencar et al. 2018).

We searched for the period of these variations by applying a periodogram analysis (Scargle 1982) to the $r'i'z'JH$ light curves, which are those in which the dips are best observed and the photometric errors are low enough. To overcome the problem of non-regular data sampling, which introduces aliases in the periodogram, we limited the analysis to the first part of the data (50 days for ROSS2 and 20 days for REMIR) and applied the CLEAN iterative deconvolution algorithm (Roberts et al. 1987). We found for all the bands a peak at about 0.29 d^{-1} , corresponding to a period of about 3.45 d, with a false-alarm probability < 0.01 , i.e. with a confidence level $> 99\%$. The period uncertainty, evaluated following the prescriptions of Horne & Baliunas

(1986), is in the range 0.01–0.05 d. A string-length analysis (Dworetzky 1983) produced similar results, with a first deep minimum in string length detected at around 3.5 d, and subsequent minima of similar or shallower depth at multiples of that period value. The peak of the periodogram is broader in the NIR band, as expected from the shorter time baseline of REMIR observations (see Fig. C.3). We cannot exclude that this period is half of the disk rotation period, because the main dips are about 7 days apart, while smaller dips between them are barely visible. The latter ones could be produced by another feature on the opposite side of the disk. However, a period of $\simeq 3.5$ d would hamper the observation of consecutive dips from Earth, due to the day-night cycle. This would explain why the deeper dimmings are seen at about 7 days from each other in the first part of the data, while they are not clearly visible in the second part (Aug-Oct 2019, see Fig. A.2). The light curve in the i' band folded in phase with the periods of 3.45 and 6.9 days is shown in Fig. C.4.

4. Discussion and conclusions

Assuming that the occulting material is located near the corotation radius, we can take $P_{\text{rot}} = 3.45$ d as the star's rotation period and derive the inclination of the rotation axis as $\sin i = v \sin i \frac{P_{\text{rot}}}{2\pi R_*}$, where we adopted the value of stellar radius given in Table 1, and the more precise value of $v \sin i = 9.9 \pm 0.6 \text{ km s}^{-1}$ reported by Nguyen et al. (2012), which was derived from a high-resolution spectrum. We find $\sin i = 0.59 \pm 0.04$ or $i = 36^\circ \pm 4^\circ$ in which we have considered the errors on $v \sin i$, R_* , and P_{rot} (0.05 d). This value rules out a nearly edge-on inner disk if our best period, and not twice its value (which we say in Sec. 3.3 we cannot exclude), is the real period. A rotation period of 6.9 d in the above equation would give $\sin i = 1.18 \pm 0.15$, which is almost exactly $i = 90^\circ$. However, an edge-on disk would cause a strong obscuration of the central star, making the object subluminal. This possibility is ruled out by the position of ISO-ChaI 52 in the HR diagram (Fig. B.2) close to the isochrone at 2 Myr. Therefore, we consider 3.45 d as the more reliable rotation period.

The size of the disk can be estimated on the basis of the IR excess. As seen in Sect. 3.2, the FIR part of the SED can be reproduced by a region 1.3×10^5 times larger than the stellar surface emitting as a black-body with $T_{\text{eff}} = 100$ K. Under this approximation, such an isothermal disk would have a radius of $R_{\text{disk}} \approx 3$ AU, which would correspond to about 15 mas at the distance of ISO-ChaI 52. This is 30 times smaller than the resolution of the ALMA images collected by Pascucci et al. (2016) in which the disk is not resolved (see, e.g., their Figs. 3 and 4).

The depth of the dips in bands of different wavelengths provides useful information on the material occulting the central star. In particular, in the optical bands, the luminosity dimming is largely due to the dust, which affects in a different way bluer and redder bands depending on the average size of the grains. We display the extinction A_λ , taken as the depth of the best observed dip (JD=2458585), as a function of wavelength in Fig. 3. Bearing in mind the radiative transfer equation for a purely absorbing medium, $I = I_0 e^{-\tau_\lambda}$, the extinction A_λ is proportional to the effective optical depth τ_λ of the layer ($A_\lambda = 1.086 \tau_\lambda$). We have compared the extinction observed during the dip with the extinction law by Cardelli et al. (1989). It is evident from

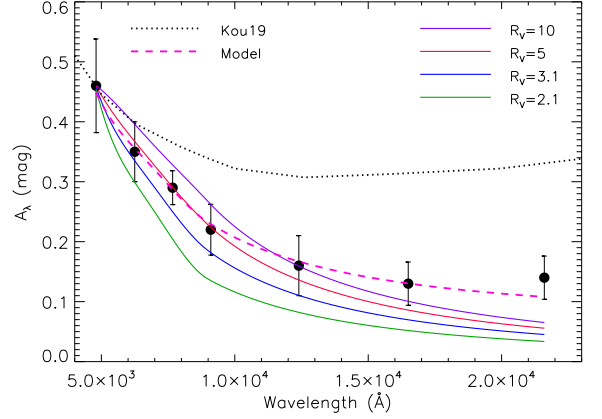


Fig. 3. Extinction taken as the amplitude of the second dip (observed at JD=2458585) as a function of the filter wavelength (dots). The full lines represent extinction laws at different values of $R_V = A_V/E(B - V)$ according to Cardelli et al. (1989). The black dotted line is the model by Koutoulaki et al. (2019). The dashed line represents our two-component model (Eq. 1). All curves are normalized to the observed extinction in the g' band ($\lambda_c = 4800 \text{ \AA}$).

Fig. 3 that the observed extinction is flatter than the one typical of the interstellar medium (ISM), which has a total-to-selective extinction ratio $R_V = A_V/E(B - V) = 3.1$. It is best reproduced in the $g'r'i'z'$ bands by a Cardelli law with $A_V = 0.41$ mag and $R_V = 5.0$, which implies an average grain size larger than in the ISM. However, no value of R_V is able to reproduce the extinction up to the NIR.

A slope in the extinction law flatter than that of the ISM and sometimes a nearly gray extinction have been reported for AA Tau, the prototype of dippers, by Bouvier et al. (1999, 2003) and for the dips of LkCa 15 by Alencar et al. (2018). Koutoulaki et al. (2019) were able to reproduce the extinction curve observed during a dimming event of RW Aur observed with X-Shooter, which is much flatter than the ISM, with a power-law distribution of grain size from a minimum value $a_{\text{min}} = 0.1 \mu\text{m}$ to a maximum $a_{\text{max}} = 150 \mu\text{m}$, including the scattering in the effective optical depth. Neglecting the scattering would require even larger grains. An ISM-type extinction ($R_V = 3.1$) would be instead produced by much smaller grains ($a_{\text{max}} = 0.1 \mu\text{m}$). In Fig. 3, we have overplotted the extinction model by Koutoulaki et al. (2019), scaled to the observed extinction in the g' band. We note that this curve is flatter than our data, suggesting different conditions for the region causing the dips in ISO-ChaI 52, as also expected from the much deeper dips of RW Aur (2–3 mag in the V band) compared to 0.3-mag dips in r' for ISO-ChaI 52. Another dipper that displays an extinction effect similar to ISO-ChaI 52 is V354 Mon (Fonseca et al. 2014). Schneider et al. (2018) used X-Shooter spectra of this star taken outside and within the dips to derive the properties of the obscuring dust in the disk warp. They showed that it is not possible to reproduce the entire dimmed spectrum by applying a single reddening law to the unclipped spectrum. The blue-visible part of the dimmed spectrum could be roughly reproduced by applying a Cardelli law with $A_V \simeq 1.2$ mag and $R_V = 6.0$ to the unclipped spectrum, but this left an IR flux excess,

similar to what we find for ISO-ChaI 52. A more evolved model, which includes an upper disk layer with an ISM-type extinction and an opaque disk region producing a gray extinction, was able to reproduce their data. We have applied to our data a simple two-component model similar to that of Schneider et al. (2018). In this case, the extinction of the stellar light can be expressed as

$$I = I_0(\alpha e^{-\tau_G} + \beta e^{-\tau_\lambda}), \quad (1)$$

where α and β are the fractions of the stellar disk occulted by the opaque-gray and thin region of the circumstellar disk, respectively, and τ_G the optical depth of the gray layer. The best-fitting model (dashed line in Fig. 3) has $\alpha = 0.38$, $\beta = 0.62$, $\tau_G = 0.22$ ($A_G \simeq 0.24$ mag), and $\tau_{\lambda=5500} \simeq 0.51$ ($A_V = 0.55$ mag) with $R_V = 3.1$. Although this is a simple schematic model, it tells us that the feature occulting the central star must contain both small-sized grains ($a < 0.5 \mu\text{m}$), producing an ISM-type extinction, and larger ones giving rise to a much flatter or gray extinction.

The fact that we observe dips in a weakly-accreting object rises the question on the connection between the strength of accretion and the presence of a warp in YSO disks. Stauffer et al. (2015) have analyzed a specific category of dipper stars, which share with ISO-ChaI 52 the late SpT and the short-duration ($\simeq 1$ d) and shallow (0.1–0.4 mag) dips. They explore alternative scenarios to warped inner disks to explain the origin of the dips, including the occultation of the star by spiral-arm overdensities in the inner disk raised by an embedded planet or by dust entrained into an accretion funnel. The latter model can also explain dips for low-inclination objects. Dippers with nearly face-on outer disks have already been found (e.g., Ansdell et al. 2016a; Scaringi et al. 2016). In particular, Ansdell et al. (2020) found the disk inclination distribution to be consistent with isotropic. The possible explanations they propose include dust clouds driven by disk winds (which can determine dips in systems with inclinations as low as $\sim 30^\circ$), or misalignments between inner disk and outer disk (which might be caused by a substellar or planetary companion). Transits of cometary-mass objects have been also proposed for a few non-accreting dippers (e.g. Scaringi et al. 2016; Ansdell et al. 2019). A dipping behavior may also be observed in objects seen at mid-inclinations when the dipole magnetic field of the star exhibits a small tilt angle with respect to its rotation axis, which leads to the formation of accretion streams that extend high above the disk midplane (Bodman et al. 2017). In the following we consider the conditions under which accretion-driven structures, such as disk warps or funnels, can produce dips in ISO-ChaI 52.

Bessolaz et al. (2008) have studied the conditions for a steady accretion flow from a circumstellar disk in the presence of a dipolar stellar magnetic field. There are few measures of photospheric magnetic fields for low-mass ($M_* \leq 0.5 M_\odot$) YSOs with typical values of $B_* \approx 1$ kG (e.g., Hill et al. 2019; Lavail et al. 2019), but fields on the order of 100 G and lower have been also observed (e.g. Donati et al. 2010; Morin et al. 2011). From Eq. 6 of Bessolaz et al. (2008) and adopting the stellar parameters in Table 1 and the upper limit $\dot{M}_{\text{acc}} \leq 4.6 \times 10^{-11} M_\odot \text{yr}^{-1}$ (Manara et al. 2019), we derive for ISO-ChaI 52 a disk truncation radius $R_T \simeq 25 R_\odot$ for a magnetic field $B_* = 1$ kG. For comparison, the Keplerian corotation radius for $P_{\text{rot}} = 3.45$ d, i.e. the period we observe for the dips, is $R_C \simeq 5.6 R_\odot$, which would be $8.9 R_\odot$ adopting $P_{\text{rot}} = 6.9$ d. The case of

$R_T > R_C$ corresponds to a propeller regime of star-disk interaction (Ustyugova et al. 2006), where stable funnel-flow accretion is inhibited. Therefore, to have a steady accretion regime, the photospheric magnetic field should be lower: $B_* \leq 70$ G for $P_{\text{rot}} = 3.45$ d and $B_* \leq 150$ G if the period is twice. Another possible explanation is that the magnetic field is strong, but we are underestimating the mass accretion rate. A truncation radius $R_T \leq 8.9 R_\odot$ with a field $B_* = 1$ kG would require a mass accretion rate $\dot{M}_{\text{acc}} \geq 3 \times 10^{-9} M_\odot \text{yr}^{-1}$, i.e. about two orders of magnitude larger than the observed upper limit; a larger mass accretion rate is needed for $R_T \leq 5.6 R_\odot$. Indeed, there are examples of YSOs with apparently low or no accretion from optical tracers, but with significant accretion as drawn from near-UV and far-UV observations (e.g., Alcalá et al. 2019). We note, however, that these objects are hotter than ISO-ChaI 52, hampering the detection of Balmer continuum excess emission due to a low contrast with respect to the photospheric emission. Should $\dot{M}_{\text{acc}}$ in ISO-ChaI 52 be so high, a much stronger UV excess and Balmer continuum than observed would have been detected.

To conclude, we think that a low-accretion rate coupled with a relatively weak surface magnetic field can give rise to disk warps or accretion structures able to produce dips in this low-mass YSO. This work shows the effectiveness of long-term simultaneous multiband photometry ranging from the optical to the NIR domain for the study of the circumstellar environment in YSOs.

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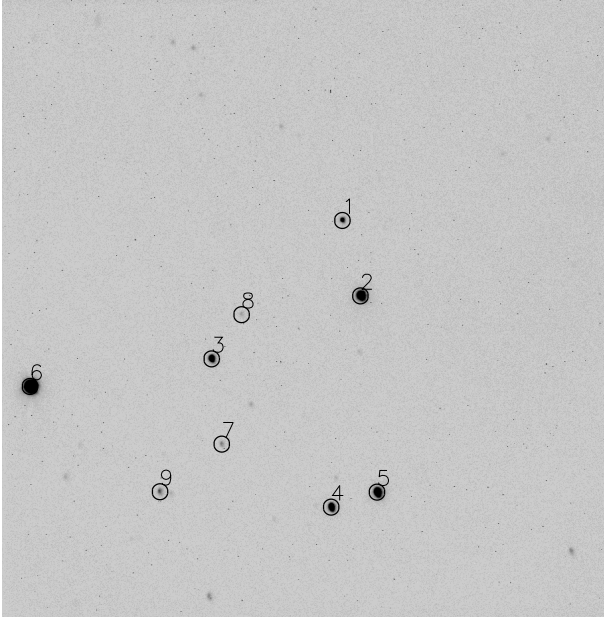


Fig. A.1. Field of ISO-ChaI 52 as observed by ROSS2 camera with the i' filter. The identification code (ID in Table A.1) is written next to the stars for which we extracted magnitudes. ID = 1 for ISO-ChaI 52.

Appendix A: Photometric data reduction

For the ROSS2 camera, we have generated master flats using the twilight flat-fields taken during the observing run, which are available in the REM archive. The latter were used to correct for pixel-to-pixel sensitivity variations, as well as for the vignetting and illumination of the field of view. Each scientific image, after subtraction of the dark-frame, was divided by the proper master-flat, depending on the filter.

The field of view, as observed in the i' filter, along with the identification code for our target and the comparison stars reported in Table A.1, is displayed in Fig. A.1.

The pre-reduction of the REMIR images is automatically done by the AQUA pipeline (Testa et al. 2004) and the co-added and sky-subtracted frames, resulting from five individual ditherings, are made available to the observer.

Aperture photometry for all the stars listed in Table A.1 was performed with DAOPHOT by using the IDL¹ routine APER. The photometric errors based on the photon statistics in the NIR bands are typically in the range 0.008–0.018 mag for ISO-ChaI 52 ($H \simeq 11^m0$) with average values of 0.013, 0.008, and 0.015 mag in J , H , and K' , respectively. They range instead from 0.003 to 0.006 mag for a brighter star like ISO-ChaI 35 ($H \simeq 9^m5$). In the optical bands the average photometric errors for ISO-ChaI 52 are of 0.013, 0.008, 0.020, and 0.045 mag for z' , i' , r' , and g' , respectively.

As a result of the field rotation, the center of the field can vary in different pointings of the telescope by as much as a few arcmin, so that only three stars (#1, #2, and #3) are included in all the useful images. We have therefore chosen #2=2MASS J11042217-7741319, which is the brightest

¹ IDL (Interactive Data Language) is a registered trademark of Harris Corporation.

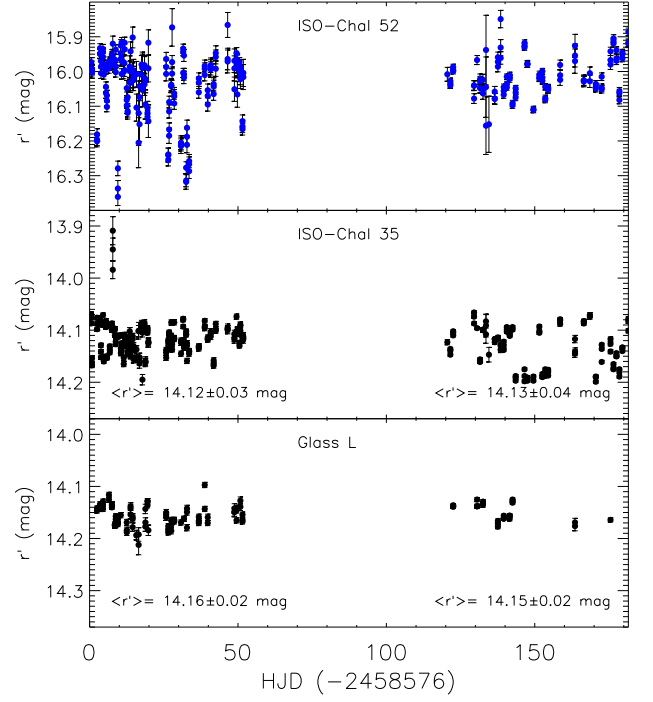


Fig. A.2. ROSS2 r' lightcurve of ISO-ChaI 52 (top panel), ISO-ChaI 35 (middle panel) and Glass L (bottom panel). We have used 2MASS J11042217-7741319 (#2) for the first two stars and Glass M (#5) for the latter, as comparison stars, adopting the magnitudes listed in Table A.1. The average magnitudes of the last two stars and their rms scatter are provided in the respective panels for the two data segments.

among these three stars and the closest to ISO-ChaI 52, as comparison object for the purpose of differential photometry. This object is not reported in the literature as a ChaI member, even if its parallax $\pi = 5.2009$ mas is nearly the same as that of ISO-ChaI 52. For each band, we have added the magnitude of #2 listed in Table A.1 to get the magnitude of ISO-ChaI 52 and the other stars in the field. The latter stars, even if with less data points, allowed us to check for any eventual variability of #2 and to evaluate the *final* data uncertainty as the rms of their magnitude differences. To this aim, we have calculated the magnitude of #4=Glass L using as comparison #5=Glass M, which are close to each other and both background stars unrelated to the ChaI cloud. The rms scatter of the photometry of Glass L all along the observing season is 0^m021 , 0^m022 , 0^m024 , 0^m047 , for the z' , i' , r' , and g' band, respectively.

As an example we show in Fig. A.2 the light curves of ISO-ChaI 52, ISO-ChaI 35 (# 3) and Glass L (#4). The points of ISO-ChaI 35 at JD $\simeq 2458583.8$, which are about 0.15 mag brighter than those taken 7 hours before and 17 hours later, are likely taken during a flare, because a similar event is also observed in the g' band as an enhancement of about 0.35 mag, while it is not visible in the redder bands. If we consider the magnitude difference between #2 and #4, the rms increases slightly to just 0.03 mag and no clear periodicity appears in the data, which reassures us on the use of # 2 as a comparison for ISO-ChaI 52.

Table A.1. Literature data for some stars in the field of ISO-ChaI 52.

ID ^a	Name	2MASS	g' (mag)	r' (mag)	i' (mag)	z (mag)	J (mag)	H (mag)	K' (mag)	π^b (mas)
1	ISO-ChaI 52	J11044258-7741571	...	...	...	13.549	11.814	11.002	10.642	5.18 ± 0.07
2		J11042217-7741319	13.377	12.187	11.702	11.245	9.738	9.167	8.913	5.20 ± 0.02
3	ISO-ChaI 35	J11035902-7743349	15.792	14.224	13.296	12.349	10.323	9.479	9.050	3.16 ± 0.04
4	Glass L	J11032288-7741301	16.587	14.297	13.140	12.155	9.926	8.691	8.264	0.40 ± 0.04
5	Glass M	J11032892-7740518	16.180	13.437	12.058	10.934	8.416	7.044	6.523	0.75 ± 0.05
6	Tyc 9414-768-1	J11034449-7746111	11.385	10.650	10.626	10.503	9.547	9.039	8.934	9.97 ± 0.02
7		J11033587-7743146	...	...	...	14.984	12.668	11.415	11.024	0.49 ± 0.08
8		J11041245-7743144	...	...	...	15.625	13.240	12.037	11.525	2.31 ± 0.10
9		J11032037-7744028	...	...	...	14.572	12.589	11.728	11.351	1.88 ± 0.05

Notes ^a Identification code as in Fig. A.1. $g'r'i'$ magnitudes from APASS (Henden et al. 2015).

z magnitudes from SkyMapper (Wolf et al. 2017). JHK' magnitudes from 2MASS (Cutri et al. 2003; Skrutskie et al. 2006).

^b Parallax from *Gaia* DR2 (Gaia collaboration 2018).

Appendix B: ROTFIT and SED analysis

The code ROTFIT finds the best photospheric template spectrum (here: BT-Settl, Allard et al. 2012) that reproduces the target spectrum by minimizing the χ^2 of the difference between the observed and synthetic spectra in specific spectral segments. The spectral intervals selected for the analysis with ROTFIT are normalized to the local continuum and contain features that are sensitive to the effective temperature and/or surface gravity, such as the Na I doublet at $\lambda \approx 819$ nm and the K I doublet at $\lambda \approx 766$ –770 nm (see Fig. B.1).

The SED was built by complementing the $g'r'i'z'JHK'$ magnitudes observed with REM outside the dips with literature values. Notably, we have added the Johnson B magnitude reported in the TIC (TESS input catalog, Stassun et al. 2019) and the GALEX-DR5 NUV flux of $4.28 \mu\text{Jy}$ at $2316 \text{ \AA}$ (Bianchi et al. 2011) at shorter wavelengths. The mid-infrared (MIR) and far-infrared (FIR) fluxes were retrieved from the WISE data release (Wright et al. 2010), from *Spitzer* IRAC and MIPS data (Dunham et al. 2015), and Herschel/PACS $100 \mu\text{m}$ (Ribas et al. 2017). We have also included the sub-mm flux at $\lambda=887 \mu\text{m}$, $F_\nu = 4.15 \pm 0.16 \text{ mJy}$, reported by Pascucci et al. (2016), which is a disk-integrated value, as the disk is not resolved in the ALMA image (see, e.g., Figs. 3 and 4 in Pascucci et al. 2016). All these values are reported in Table B.1.

The distance $d = 193 \pm 3 \text{ pc}$ was calculated from the *Gaia* DR2 parallax of ISO ChaI 52 ($\pi = 5.18 \pm 0.07 \text{ mas}$) as $d = 1000/\pi$. This large value of parallax allows us to neglect small corrections like those proposed by Lindegren et al. (2018) that would decrease the distance by only 1 pc, which is less than the distance error. In the fitting procedure, applied to the fluxes from B to J band, we fixed the distance and the effective temperature and let the stellar radius, R_* , and the extinction, A_V , vary until a minimum χ^2 was reached. The key parameter affecting the results is the effective temperature, therefore we run the code also fixing T_{eff} to the extreme values given by the T_{eff} error of 70 K. We found $A_V = 0.43 \pm 0.32 \text{ mag}$, which is lower than the value of 1.2 mag reported by Manara et al. (2016), who analyzed the full calibrated X-Shooter spectrum and used real-star spectra of slightly higher T_{eff} as templates. We note that A_V is very sensitive to the intrinsic shape of the flux distribution,

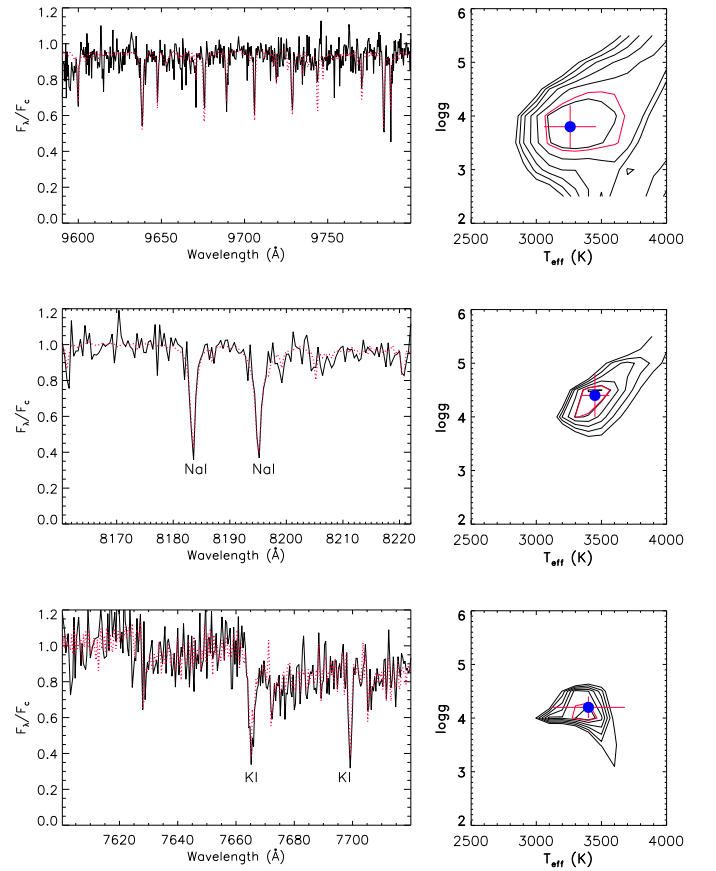


Fig. B.1. Left panels: continuum-normalized VIS X-Shooter spectrum of ISO-ChaI 52 in three regions (black full lines) with the best fitting synthetic spectrum overplotted (red dotted lines). Right panels: χ^2 contour maps in the T_{eff} - $\log g$ plane. In each panel, the 1σ confidence level is denoted by the red contour. The best values and errorbars on T_{eff} and $\log g$ are also indicated.

i.e. to T_{eff} , while the stellar radius, $R_* = 1.14 \pm 0.04 R_\odot$, is strongly dependent on the distance and on the observed flux near the maximum of the SED, where the extinction plays a minor role.

Table B.1. Data for the SED of ISO-ChaI 52.

Band	λ_c (μm)	Magnitude (mag)	Flux ($\text{erg cm}^{-2} \text{s}^{-1} \text{\AA}^{-1}$)	Reference
<i>NUV</i>	0.231	22.32 ± 0.48	$(2.41 \pm 1.10) \text{E-17}$	B2011
<i>B</i>	0.444	18.57 ± 0.16	$(2.68 \pm 0.40) \text{E-16}$	S2019
<i>g'</i>	0.485	17.50 ± 0.15	$(4.68 \pm 0.63) \text{E-16}$	Present work
<i>r'</i>	0.621	15.95 ± 0.10	$(1.16 \pm 0.11) \text{E-15}$	Present work
<i>i'</i>	0.767	14.52 ± 0.07	$(2.88 \pm 0.19) \text{E-15}$	Present work
<i>z'</i>	0.910	13.55 ± 0.07	$(5.02 \pm 0.32) \text{E-15}$	Present work
<i>BP</i>	0.505	17.146 ± 0.014	$(5.77 \pm 0.08) \text{E-16}$	<i>Gaia</i> DR2
<i>G</i>	0.623	15.218 ± 0.003	$(2.09 \pm 0.01) \text{E-15}$	<i>Gaia</i> DR2
<i>RP</i>	0.772	13.892 ± 0.008	$(3.65 \pm 0.03) \text{E-15}$	<i>Gaia</i> DR2
<i>J</i>	1.24	11.75 ± 0.07	$(6.24 \pm 0.42) \text{E-15}$	Present work
<i>H</i>	1.65	10.98 ± 0.06	$(4.59 \pm 0.25) \text{E-15}$	Present work
<i>K'</i>	2.19	10.58 ± 0.06	$(2.51 \pm 0.14) \text{E-15}$	Present work
<i>WISE</i> 1	3.35	10.186 ± 0.023	$(6.89 \pm 0.15) \text{E-16}$	W2010
<i>WISE</i> 2	4.60	9.728 ± 0.020	$(3.10 \pm 0.06) \text{E-16}$	W2010
<i>WISE</i> 3	11.56	7.901 ± 0.019	$(4.50 \pm 0.08) \text{E-17}$	W2010
<i>WISE</i> 4	22.09	5.461 ± 0.030	$(3.33 \pm 0.09) \text{E-17}$	W2010
<i>IRAC</i> 1	3.6	...	$(6.42 \pm 0.33) \text{E-16}$	D2015
<i>IRAC</i> 2	4.5	...	$(3.27 \pm 0.15) \text{E-16}$	D2015
<i>IRAC</i> 3	5.8	...	$(1.83 \pm 0.09) \text{E-16}$	D2015
<i>IRAC</i> 4	8.0	...	$(8.71 \pm 0.41) \text{E-17}$	D2015
<i>MIPS</i> 24	24	...	$(2.95 \pm 0.11) \text{E-17}$	D2015
<i>MIPS</i> 70	70	...	$(7.64 \pm 0.82) \text{E-18}$	D2015
<i>Herschel</i>	100	...	$(6.00 \pm 1.20) \text{E-18}$	R2017
<i>ALMA</i>	887	...	$(1.58 \pm 0.06) \text{E-21}$	P2016

Notes B2011 = Bianchi et al. (2011); S2019 = Stassun et al. (2019); Gaia DR2 = Gaia collaboration (2018); W2010 = Wright et al. (2010); D2015 = Dunham et al. (2015); R2017 = Ribas et al. (2017); P2016 = Pascucci et al. (2016).

The stellar luminosity, calculated as $L_* = 4\pi R_*^2 \sigma T_{\text{eff}}^4$, is $L_* = 0.123 \pm 0.011 L_\odot$, where the error takes into account the R_* and T_{eff} errors and their covariance. This value of L_* is higher than the luminosity of $0.09 L_\odot$ reported by Manara et al. (2016), who adopted a distance $d = 160$ pc. However, the latter becomes $0.13 L_\odot$ with the *Gaia* distance $d = 193$ pc, in good agreement with our determination.

Appendix C: Additional plots

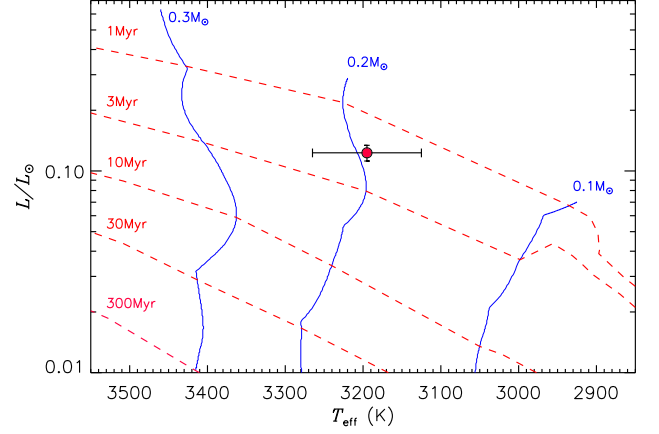


Fig. B.2. ISO-ChaI 52 in the Hertzsprung-Russell diagram. See Table 1 for the values of T_{eff} and L_* derived with the ROTFIT code. Isochrones and evolutionary tracks by Baraffe et al. (2015) are overplotted as dashed and solid lines, with the labels representing their age and mass, respectively.

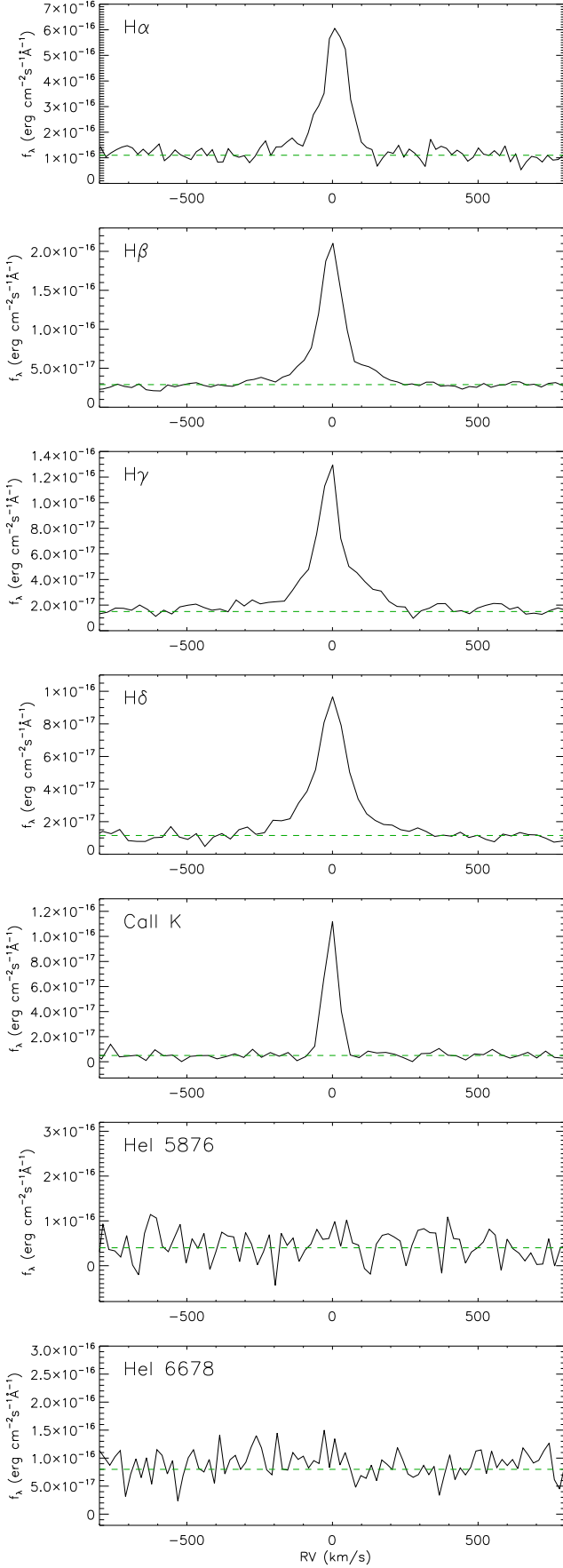


Fig. C.1. Profiles of Balmer lines, Ca II K, and He I lines. In each plot, the horizontal dashed green line denotes the continuum level.

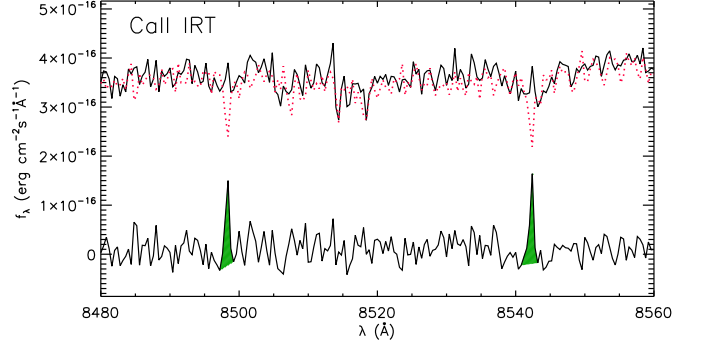


Fig. C.2. X-Shooter spectrum of ISO-ChaI 52 in the region of Ca II IRT (solid black line) along with the inactive template (dotted red line). The difference between observed and template spectrum is shown in the bottom of the box, along with the residual emission in the line cores (hatched green areas).

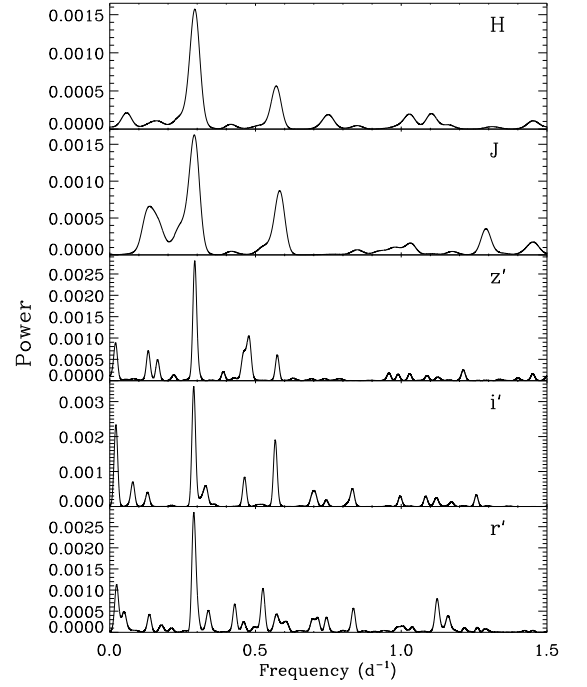


Fig. C.3. Cleaned periodograms for the photometric data of ISO-ChaI 52 in the $HJz'i'r'$ bands (from top to bottom). The highest peak in each band corresponds to a period of about 3.45 days.

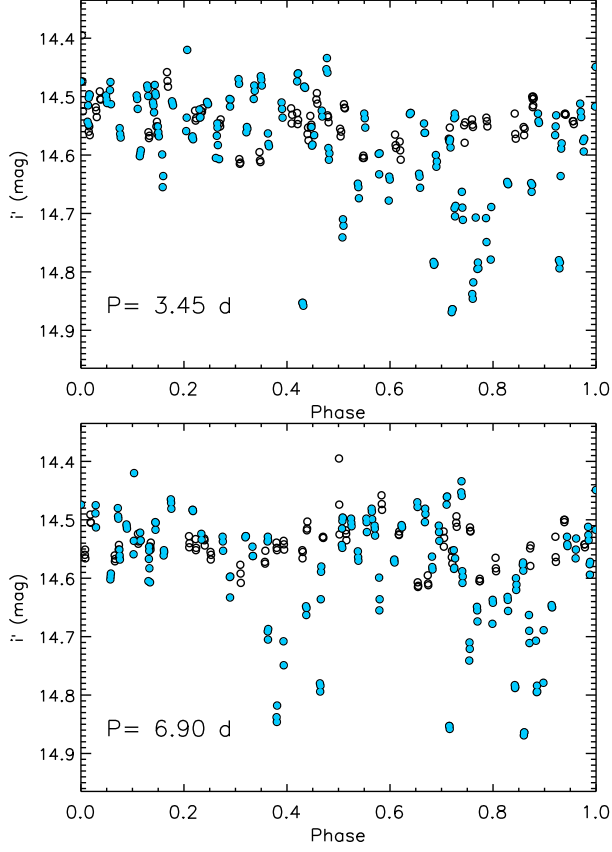


Fig. C.4. Light curve in the i' band folded in phase with the period of 3.45 days (*upper panel*) and 6.9 days (*lower panel*). Filled dots refer to the data acquired before the end of May 2019 ($\text{JD} < 2458630$); open dots are related to the second part of the data (Aug-Oct 2019).