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A dusty compact object bridging galaxies and quasars at cosmic dawn

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Understanding how super-massive black holes form and grow in the early Universe has become a major challenge^[12] since the discovery of luminous quasars only 700 million years after the Big Bang^[34]. Simulations indicate an evolutionary sequence of dust-reddened quasars emerging from heavily dust-obscured starbursts that then transition to unobscured luminous quasars by expelling gas and dust^[5]. Although the last phase has been identified out to a redshift of 7.6,^[6] a transitioning quasar has not been found at similar redshifts owing to their faintness at optical and near-infrared wavelengths. Here we report observations of an ultraviolet compact object, GNz7q, associated with a dust-enshrouded starburst at a redshift of $z = 7.1899 \pm 0.0005$. The host galaxy is more luminous in dust emission than any other known object at this epoch, forming 1,600 solar masses of stars per year within a central radius of 480 parsec. A red point source in the far-ultraviolet is identified in deep, high-resolution imaging and slitless spectroscopy. GNz7q is extremely faint in X-rays, which indicates the emergence of a uniquely ultraviolet compact star-forming region or a Compton-thick super-Eddington black-hole accretion disk at the dusty starburst core. In the latter case, the observed properties are consistent with predictions from cosmological simulations^[7] and suggest that GNz7q is an antecedent to unobscured luminous quasars at later epochs.

In recent uniform reprocessing of all archival *Hubble Space Telescope* (*HST*) imaging and slitless spectroscopy (see Methods), GNz7q was identified in the GOODS-North extra-galactic field as a luminous galaxy candidate at a redshift $z > 6.5$ with the F160W band AB magnitude of 23.09 ± 0.05 . Although the source was detected and highlighted as a potential high-redshift galaxy by previous authors, it was never spectroscopically confirmed^[8]. However, the full suite of *HST* data reveals an unambiguous continuum break at $\sim 1.0 \mu\text{m}$ which is best explained by a Lyman 1216-Å break at $z = 7.23 \pm 0.05$ (Fig. 1).

GNz7q is distinct in the rest-frame UV when compared to any other object currently known at similar redshifts ($z > 6$). Its luminosity falls between typical quasars and galaxies and it is quite red in colour (Extended Data Fig. 1), with a rest-frame 1450 Å luminosity, M_{1450} , of -23.2 mag and continuum slope, $F_\lambda \propto \lambda^{\alpha_\lambda}$, of $\alpha_\lambda = 0.1 \pm 0.3$ (see Methods). This is the reddest continuum slope found among objects at similar redshifts, ($\alpha_\lambda \lesssim -1.5$,^[9,10]) but is comparable to lower-redshift red quasars identified in the Sloan Digital Sky Survey (SDSS)^[11] (Fig. 1). GNz7q is also spatially unresolved in all *HST* bands (Extended Data Fig. 2) and is bright at rest-frame $3 \mu\text{m}$, detected with *Spitzer*/MIPS. These characteristics suggest that GNz7q is a distant, red quasar.

However, GNz7q is not detected in the extremely deep 2 Ms X-ray map of the Chandra Deep Field North^[12]. Even accounting for obscuration, we obtain an upper limit (99% confidence level) on the X-ray luminosity of $L_X < 3.9 \times 10^{42} \text{ erg s}^{-1}$ (see Methods). This is several orders of magnitude lower than what would be predicted by assuming the correlation between L_X and optical luminosity observed for other quasars (Extended Data Fig. 3). GNz7q is therefore strikingly faint in 2–10 keV X-rays, apparently in tension with its being a quasar. The absence of strong, broad UV emission lines (Fig. 1) in addition to this unique X-ray faintness raises the possibility that GNz7q could instead be an extreme UV-compact star-forming object.

1- and 3-millimeter (mm) observations were carried out with the NOthern Extended Millimeter Array (NOEMA) between June 2020 and February 2021 (see Methods). The [C II] 158 μm line was robustly detected at 17σ peak intensity at a redshift of $z = 7.1899 \pm 0.0005$, consistent with the Lyman-break redshift. The underlying 1 mm and 3 mm continua are also detected at 16σ and 3.9σ respectively. The sky positions of the 1 and 3 mm continua and emission line are consistent with the *HST* source. The [C II] line is spatially resolved with an effective radius of $r_e = 1.4 \text{ kpc}$. The 1 mm continuum is unresolved with an upper limit of $r_e \leq 0.5 \text{ kpc}$, suggesting that the majority of the far-infrared (IR) emission is arising from a compact region of $\leq 0.7 \text{ kpc}^2$. No close neighbors are detected in the mm line or continuum maps. However, a second source (“ND1”) is seen in the 1 mm continuum map that is undetected in the deep *HST* images and can be best explained as a dusty companion $\sim 16 \text{ kpc}$ from GNz7q (see Methods).

The precise redshift determination and the rich multi-wavelength datasets in the GOODS-North field provide a unique opportunity to constrain the host galaxy properties separate from the UV-luminous core. Fits to the optical-to-mm spectral energy distribution (SED) yield an IR lumi-

nosity (rest-frame 8–1,000 μm) of $L_{\text{IR}} = (1.2 \pm 0.6) \times 10^{13} L_{\odot}$ (where $L_{\odot}$ is the luminosity of the Sun) (Fig. 2). This corresponds to a star-formation rate (SFR) of $1,600 \pm 700 M_{\odot} \text{ yr}^{-1}$ after removing the potential contribution of the emission associated with the active galactic nucleus (AGN). Regardless of the interpretation, the host of GNz7q is the most vigorously star-forming galaxy at $z > 7$ found to date^[6]. It has a SFR surface density of $\geq 1,100 M_{\odot} \text{ yr}^{-1} \text{ kpc}^{-2}$, which is at the Eddington limit for star-forming galaxies^[13]. Treating GNz7q as a UV-compact star-forming object instead of an AGN would increase the SFR estimate. The dust in the host has a peak temperature of $T_d \sim 80 \text{ K}$, warmer than typical high-redshift quasar hosts by a factor of ~ 1.5 ^[14]. The $L_{\text{[C II]}}/L_{\text{IR}}$ ratio also shows one of the lowest values so far seen. The high dust temperature and relatively faint [C II] emission may be due to the maximal SFR surface density, as the strong radiation field produced by the intense starburst increases T_d and decreases the abundance of singly-ionized carbon. Dust and gas masses are estimated at $M_{\text{dust}} = 1.6 \times 10^8 M_{\odot}$ and $M_{\text{gas}} = 2.0 \times 10^{10} M_{\odot}$ making GNz7q one of the most dust- and gas-rich systems known at $z > 6$ (see Methods). The presence of the proximate ND1 galaxy is consistent with the high abundance of companion galaxies reported around luminous quasars at $z > 6$ ^[15].

The shape of the NIR–MIR SED of GNz7q—especially the excess emission at rest-frame 3 μm —cannot easily be explained by emission from stars and ionized gas associated with star-formation activity alone (Fig. 2 and Extended Data Fig. 8). Moreover, fitting a profile to the *HST* images provides a stringent upper limit for the effective radius of only 60 pc for the UV emission (see Methods). If this compact emission were attributed to star-formation, the SFR surface density from the UV alone would reach $\gtrsim 5,000 M_{\odot} \text{ yr}^{-1} \text{ kpc}^{-2}$, two orders of magnitude higher than the UV-luminous compact galaxies reported at $z \sim 2$ –3 (Extended Data Fig. 7)^[16]. Taken together, the properties of GNz7q favor the interpretation of the UV source as a red quasar. Its properties are in excellent agreement with the transition phase of the evolutionary paradigm of super-massive black holes (SMBHs): a low-luminosity, dust-obscured quasar emerging in a vigorously star-bursting host. The detection of dust-obscured, super-Eddington accretion objects hosted by starburst galaxies has so far been reported up to $z = 4.6$ ^[17]. GNz7q at $z = 7.2$ is found at a cosmic time that is 500 million years earlier, close to the earliest SMBH known at $z = 7.6$ ^[6]. The AGN in GNz7q is two orders of magnitude fainter than its lower- z analog at $z = 4.6$, but with a host SFR ~ 3 times higher, suggesting that GNz7q is experiencing an early stage of its transition phase.

GNz7q’s extreme X-ray faintness is a strong indicator of the young age of the quasar. Extrapolating the anti-correlation^[18] between X-ray luminosity and AGN Eddington ratio (λ_{Edd}) to the X-ray upper limit of GNz7q, we obtain an Eddington ratio significantly (5.5σ) greater than unity and a black hole mass of only $M_{\text{BH}} \sim 10^7 M_{\odot}$ (Fig. 3a, see also Methods). X-ray-weak quasars are found to be abundant among weak emission line quasars (rest-frame equivalent width of C IV $\text{EW(CIV)} < 16 \text{ \AA}$)^[19,20] with more powerful nuclear winds^[21]. These trends can be explained by a scenario where the inner region of the accretion disk is strongly inflated to a substantial height due to the unusually high accretion, which blocks the nuclear ionizing continuum and the X-rays from reaching the broad line region and external observers^[19]. GNz7q is indeed lacking the C IV line in our spectroscopy ($\text{EW(CIV)} < 10 \text{ \AA}$), consistent with this scenario.

These observational results can be compared with cosmological semi-analytic models for progenitors of high- z quasars [7, 22]. Among simulated merger histories, several progenitors with multi-wavelength properties similar to GNz7q indeed have relatively low mass SMBHs ($M_{\text{BH}} \sim 10^{6.5-7.5} M_{\odot}$), but still reside in the most massive halos of $\sim 10^{11.5-12.5} M_{\odot}$ at $z = 7.2$ (Fig. 4). These simulations show that all of these progenitors will evolve into optically-luminous blue quasars harboring a SMBH with $M_{\text{BH}} > 10^8 M_{\odot}$ at $z = 6.4$. This indicates that GNz7q could be the direct progenitor of an optically-luminous quasar, although models do not rule out the possibility that GNz7q will fail to finalize its transition at later epochs because of possible mergers with other halos hosting more massive BHs. The simulations and recent observations also predict a tight correlation between M_{BH} and the X-ray luminosity normalized by infrared luminosity $L_{\text{IR, SF}}$, confirming that the unique X-ray faintness of GNz7q corresponds to the regime of $M_{\text{BH}} < 10^8 M_{\odot}$ (Fig. 3b).

Given the short-lived nature of a transitioning red quasar and the intrinsically low sky density of the quasar population [23], it is remarkable to find GNz7q within the 170 arcmin² GOODS-North field. Assuming instead a total survey area of the entire *HST* archive of nearly 3 deg², the identification of GNz7q suggests a sky density of 0.33 deg⁻¹ and a lower limit of 3.3×10^{-3} deg⁻¹ based on the Poisson uncertainty at the 99% single-sided confidence level [24]. However, the quasar luminosity function (QLF) and the red quasar fraction at $z \sim 6$ suggest a much lower predicted sky density of 6.8×10^{-4} deg⁻¹, even for less luminous red quasars similar to GNz7q [25, 26]. A recent *HST* study also reports a potentially higher density of less-luminous quasars at $z \sim 8$ than the QLF at $z \sim 6$ [27]. Together, these results may imply that the red and/or less-luminous quasar population is more common at $z > 7$ than our understanding to date up to $z \sim 6$ [28]. We note in passing that classical colour selections for high- z quasars in ground-based surveys would recover the identification of GNz7q (see Methods). This implies that these quasar populations could have been missed in previous surveys due to their faint nature in the mid-IR (MIR) and X-rays and in their rest-frame UV lines, that are here overcome by the uniquely deep and rich multi-wavelength datasets of the GOODS-North field. A systematic high-resolution, deep imaging survey in the optical–MIR bands may discover additional objects similar to GNz7q. Furthermore, follow-up spectroscopy of broad Balmer lines for $z > 7$ objects will become possible with the launch of the *James Webb Space Telescope*. This will have the power to decisively determine whether the quasar classification is correct as well as determine how common such quasars truly are. Even a non-detection of broad lines would imply intriguing conclusions, i.e. the existence of extraordinarily luminous and compact star-forming regions or stark differences between the first quasars and their descendants.

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Author contributions G.B. reduced and analyzed the optical–NIR data of *HST* and *Spitzer* and discovered GNz7q. S.F., G.B., S.T., G.M., D.W., F.V., C.S., J.F., L.C., R.M. M.V., and F.W. discussed and planned the follow-up observing strategy and the data analysis. G.B., G.M., and V.K. conducted the SED analysis and wrote the relevant Methods section. G.B. produced Figs. 1, 2, and Extended Data Figs. 2 and 8. D.W. analyzed the X-ray properties from the *Chandra* data and wrote the relevant Methods section, T.G. reduced and analyzed the SCUBA2 data, and M.K. and I.C. reduced the NOEMA data. R.V., M.G., and R.S. performed the cosmological semi-analytical simulation *GAMETE/QSOdust* and wrote the relevant Methods section. F.R. worked on the 3D modeling for the NOEMA [C II]-line data cube. P.O. investigate the properties of the dust-continuum object identified near GNz7q. All authors discussed the results and commented on the manuscript. S.F. led the team, being Principal Investigator of the follow-up NOEMA programs, analyzed the NOEMA data, wrote the main text and the Methods section, produced Figs. 3 and 4, Extended Data Tables, and Figs. 1, 3–7, 9–11.

Competing interests The authors declare no competing interests.

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Figures

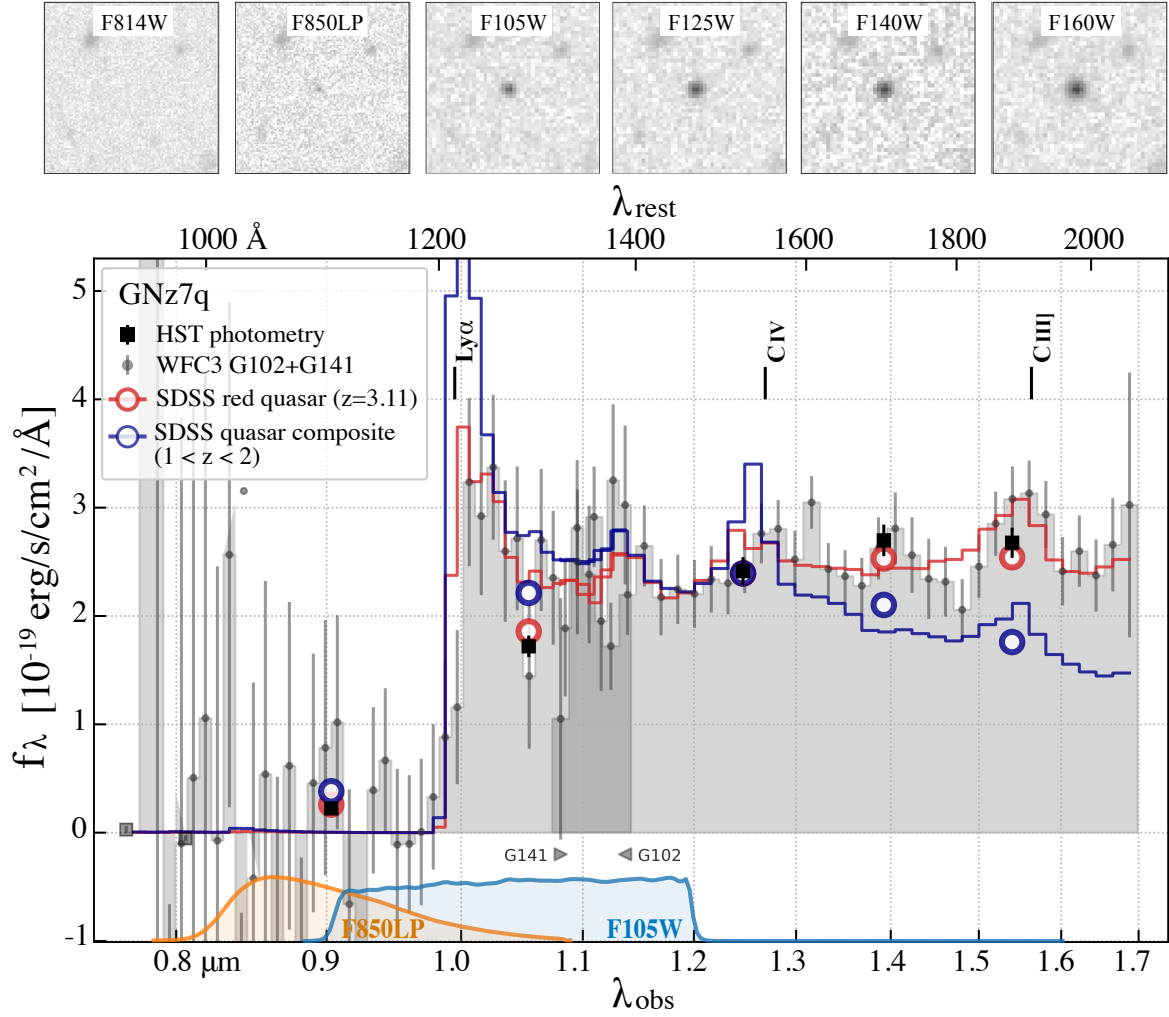


Figure 1: | ***Hubble Space Telescope* near-infrared images and spectrum of GNz7q.** The spectrum and photometry show a strong Lyman break at $\lambda_{\text{obs}} \sim 1.0 \mu\text{m}$. The top panels show the *HST* image cutouts ($5'' \times 5''$). The source is unresolved in all deep *HST* images up to the reddest filter available at $1.6 \mu\text{m}$ (WFC3/IR F160W). In the bottom panel, the black squares and gray dots respectively show the broadband photometry and the slitless spectrum binned by a factor of 4 relative to the nominal pixel scale. The error bars denote 1σ uncertainties. The labeled black bars indicate the expected wavelengths for the main emission lines based on the [C II] $158\text{-}\mu\text{m}$ line redshift of GNz7q at $z = 7.1899$. The blue curve represents a composite spectrum of SDSS optically luminous blue quasars^[10] at $1 < z < 2$, while the red curve shows a red quasar at $z = 3.11$ (SDSS spec-6839-56425-146)^[11] whose FUV spectrum resembles that of GNz7q. Both of the lower- z quasar spectra are shifted to $z = 7.1899$, normalized at $1.2 \mu\text{m}$, and binned to the same spectral resolution as the GNz7q spectrum. The large open circles show the quasar templates integrated through the *HST* filter passbands. The bandpasses of the ACS/F850LP and WFC3/F105W filters shown at the bottom straddle the spectral break, explaining the faint detection in the former and the suppressed flux density relative to the continuum in the latter.

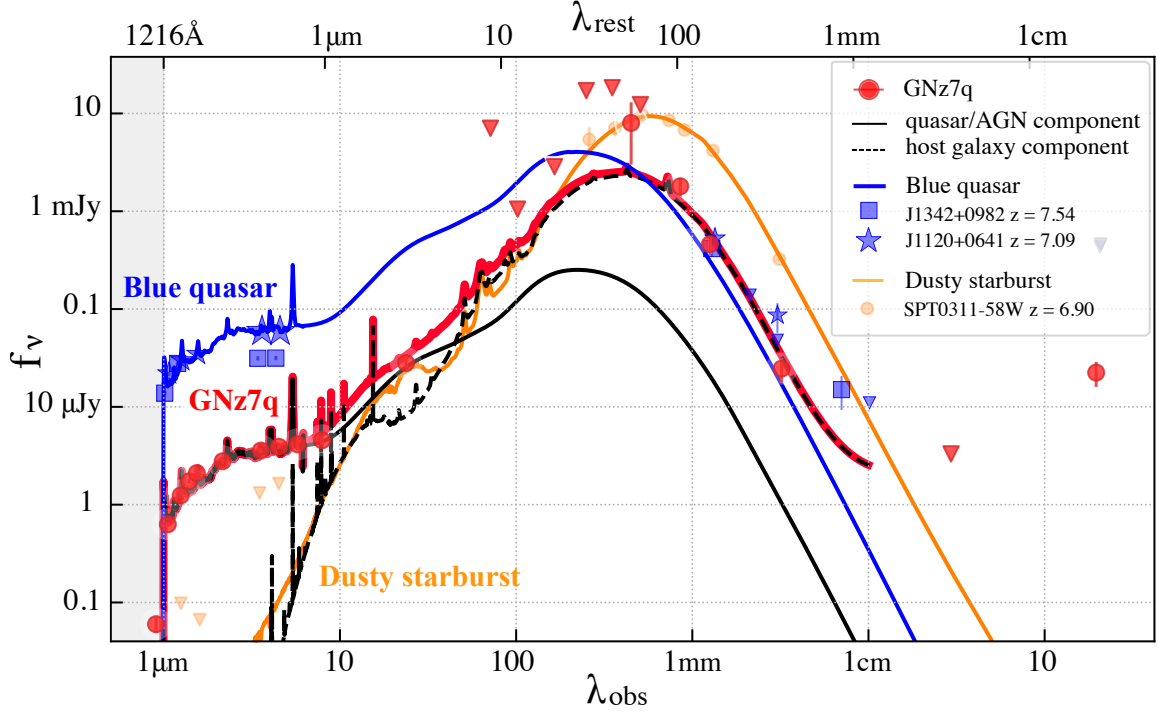


Figure 2: | **The spectral energy distribution of GNz7q from optical to radio wavelengths.** Photometry is shown for data from *HST* (0.8–1.6 μ m), *Spitzer* (3.6–24 μ m), *Herschel* (80–500 μ m), JCMT (450 and 850 μ m), NOEMA (1 and 3 mm) and VLA (3 and 20 cm) in the GOODS-North field (Extended Data Table 1). Triangles indicate 3σ upper limits. The sum of the best-fit quasar/AGN (black solid) and galaxy (black dashed) templates is shown as a red curve. The radio detection at 20 cm is consistent with the enormous implied SFR of the host galaxy (see Methods). For comparison, we also show the SEDs of other source populations at similar redshifts: optically-luminous blue quasars at $z = 7.54$ (J1342+0928^[4], blue squares) and $z = 7.08$ (J1120+0641^[3], blue stars), and a dusty starburst at $z = 6.90$ (SPT0311-58W^[29], orange circles). The blue curve is drawn with the quasar/AGN template normalized to J1120+6410’s rest-frame UV emission. The orange curve is the best-fit SED for SPT0311-58W, taken from the literature^[29]. The SED of GNz7q falls between these two categories of the dusty starburst and the blue quasar, representing a transient phase between them.

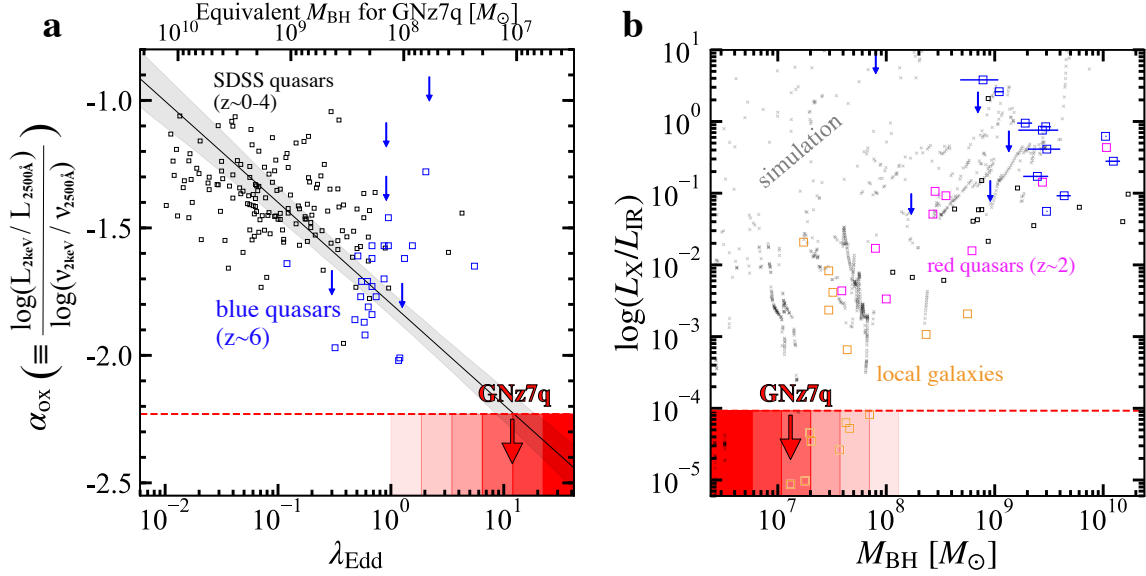


Figure 3: | **The unique X-ray faintness of GNz7q.** The 2 Ms deep *Chandra* data [\[12\]](#) provide a stringent upper limit (red dashed line) for the X-ray luminosity, suggestive of a very high accretion rate ($\lambda_{\text{Edd}} \gtrsim 1$) onto a less massive black hole ($M_{\text{BH}} \lesssim 10^8 M_{\odot}$). **a.** Optical to X-ray spectral index α_{ox} as a function of Eddington ratio (λ_{Edd}). The dust-corrected optical luminosity $L'_{2,500}$ is used in the upper limit estimate of α_{ox} for GNz7q. SDSS quasar measurements (black squares) and the best-fit relation (black line) with its 1 σ confidence level (gray shaded region) are taken from the literature (see Methods). The upper horizontal axis shows the equivalent black hole mass for GNz7q as a function of λ_{Edd} based on its AGN bolometric luminosity from the UV to mm SED fitting (see Methods). The red shaded region shows the λ_{Edd} regime of GNz7q extrapolated from the best-fit relation, where the shading becomes darker with increasing λ_{Edd} . **b.** X-ray luminosity (L_X) normalized by L_{IR} . We show other populations for comparison: blue quasars at $z \sim 6$ (blue squares), red quasars at $z \sim 2$, and dusty starbursts at $z \sim 0$ (orange squares) taken from the literature (see Methods). Gray circles are plotted for simulated galaxies (see Methods) with AGN bolometric luminosity of $L_{\text{bol}} > 10^{42} \text{ erg s}^{-1}$. The colour scale and the horizontal range of each red shaded region of M_{BH} corresponds to those of panel **a**.

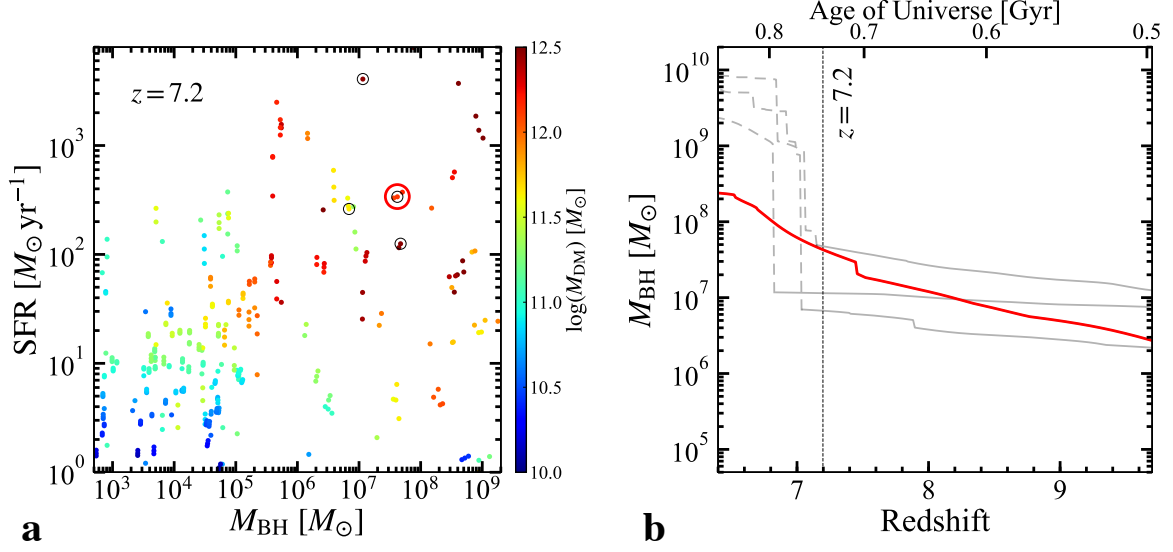


Figure 4: | **SFR and M_{BH} relations for progenitors of luminous quasars in a cosmological simulation.**
a. SFR and BH masses predicted by the semi-analytical model GQd^[22] for selected $z = 7.2$ progenitors of a luminous quasar at $z = 6.4$. Each system is represented by a circle colour-coded by the dark matter halo mass (M_{DM}). The four black circles mark systems that have X-ray, optical, and host galaxy properties similar to GNz7q (See Methods). **b.** BH mass assembly history for the systems marked with the black circles in panel **a**. At $z = 7.2$, one of them has already grown to $4.0 \times 10^7 M_{\odot}$ (red circle in panel **a**) and thereafter it continues to grow by gas accretion and mergers with other BHs to form a SMBH of $M_{\text{BH}} = 2.5 \times 10^8 M_{\odot}$ at $z = 6.4$ (red line). The other systems also have relatively low-mass BHs, down to $\sim 10^{6.5-7.5} M_{\odot}$ at $z = 7.2$ (grey lines), but these systems undergo mergers with galaxies hosting more massive BHs. As a result, these BHs are not the most massive progenitors of the final SMBH, which grows to $\sim 10^{9.3-10} M_{\odot}$ by $z = 6.4$ through gas accretion and mergers with other BH progenitors (grey dashed line). The black vertical line indicates the redshift of GNz7q.

Methods

In this paper, error values represent the 1σ uncertainty, where σ denotes the root-mean-square or standard deviation; upper limits are indicated at the 3σ level; red symbols in figures denote GNz7q, unless otherwise specified.

1. Cosmology We adopt cosmological parameters measured by the Planck mission^[30], i.e. a Λ cold dark matter (Λ CDM) model with total matter, vacuum and baryonic densities in units of the critical density, $\Omega_\Lambda = 0.76$, $\Omega_m = 0.24$, $\Omega_b = 0.04$, and Hubble constant, $H_0 = 100 h \text{ km s}^{-1} \text{ Mpc}^{-1}$, with $h = 0.73$. Based on these parameters, we adopt the angular size distance of 5.32 kpc/arcsec at the source redshift of $z = 7.2$ in this paper.

2. Definition of quasar categories In this paper, we make a distinction between a reddened Type 1 quasar and a Type 2 quasar, where we refer to the former as the red quasar. A Type 1 quasar is defined to have at least one broad emission line ($\text{FWHM} \gtrsim 1,000\text{--}2,000 \text{ km s}^{-1}$) in the spectrum, while Type 2 quasars are defined by those that do not satisfy it, but has too bright narrow UV-optical or IR emission lines, X-ray, or radio continuum for a galaxy^[31]. In the classical AGN unification models, this difference is generally explained by different viewing angles toward the central accretion disk^[32,33], where the observer's line of sight in Type 2 quasars penetrates through optically-thick dusty material due to the nearly edge-on view of the dust torus that blocks the nuclear broad emission lines. In this context, red quasars are generally defined to have at least one broad emission line in the spectrum, distinguishing them from Type 2 quasars, but to be substantially reddened by dust. The extent of the dust reddening has been estimated a number of ways in the literature based on optical^[23,25,34], optical-MIR^[35,36], NIR-radio^[37,38], and MIR^[39,40]. For example, the red quasar is characterized by the optical colour excess of $E(B - V) \gtrsim 0.1$ in the optical colour selection^[23,25,34]. Following these definitions, the SDSS quasar at $z = 3.11$ (Fig. 1, SDSS spec-6839-56425-146) is classified as a red quasar with $E(B - V) = 0.13$ ^[41]. W2246-0526 is also classified as a red quasar that is a super-Eddington, extremely luminous, and dust-obscured quasar at $z = 4.6$ with broad MgII and CIV emission lines and $E(B - V) = 5$.^[17,42,43] Because several sub-millimeter galaxies (SMGs; flux density at $850 \mu\text{m} \gtrsim$ a few mJy) at $z \sim 2$ are reported to have broad emission lines and $E(B - V) \sim 0.4$, we regard these SMGs also as red quasars^[45] in this paper. Local (ultra) luminous infrared galaxies ((U)LIRGs) and high-redshift SMGs that are not confirmed to have broad emission lines are referred to as dusty starbursts, while we refer to the hot dust-obscured quasars confirmed with the broad emission lines^[44] as red quasars in this paper. We refer to Type 1 quasars from the literature that do not exceed the red colour threshold of $E(B - V) > 0.1$ as blue quasars in this paper.

3. Uniform processing of archival *HST* and *Spitzer* data GNz7q was identified in a project nicknamed the Complete Hubble Archive for Galaxy Evolution (CHArGE). CHArGE aims to perform uniform processing and analysis of all archival *HST* and *Spitzer* data taken away from the Galactic midplane. *HST* Advanced Camera for Surveys (ACS) optical and the Wide Field Camera 3 (WFC3) near-infrared and *Spitzer*/Infrared Array Camera (IRAC) observations covering GNz7q were carried out by a variety of large extra-galactic surveys and individual programs. *HST* images

were obtained in the F435W (6 exposures; 7.2 ks integration), F606W (23 exp.; 9.3 ks), F775W (35 exp.; 19 ks), F814W (42 exp.; 24 ks), F850LP (82 exp.; 37 ks), F105W (6 exp.; 3.3 ks), F125W (8 exp.; 4.4 ks), F140W (6 exp.; 1.2 ks), and F160W (8 exp.; 5.4 ks) filters. The IRAC channel 1 ($3.6\ \mu\text{m}$) and channel 2 ($4.5\ \mu\text{m}$) integrations are 345 ks and 330 ks, respectively. We aligned all of the *HST* exposures to sources in the *GAIA* DR2 catalog^[46] and created final mosaics in a common pixel frame with 50 mas and 100 mas pixels for the ACS/WFC and WFC3/IR filters, respectively. We aligned the individual *Spitzer* exposures to the same astrometric frame as the *HST* frame and generated final drizzled IRAC mosaics with $0.''5$ pixels. Further details of the *HST* (*Spitzer*) image processing with the *grizli* (*golfir*) software will be presented in Kokorev (in prep.). In Fig. 1, we present the *HST* images of GNz7q.

Archival *HST* slitless spectroscopy of GNz7q is available with integration times 8.8 ks (12.7 ks) in the G102 (G141) grisms from HST General Observer (GO) program 13420 (11600). Together the two grisms cover $0.8 < \lambda < 1.7\ \mu\text{m}$ without any gaps (Fig. 1). The slitless spectra are reduced and extracted with the *grizli* software^[47].

The *HST* broad-band images and slitless spectroscopy show an unambiguous continuum break at $\sim 1.0\ \mu\text{m}$. By using a SDSS red quasar template at $z = 3.11$ (SDSS spec-6839-56425-146)^[11,41], we obtain the rest-frame UV redshift at $z_{\text{UV}} = 7.23 \pm 0.05$. This is consistent with the [C II] $158\text{-}\mu\text{m}$ line redshift of $z_{[\text{CII}]} = 7.1899 \pm 0.0005$ within the uncertainties (Section 4). We fit a power-law ($F_\lambda \propto \lambda^{\alpha_\lambda}$) model to the rest-frame UV continuum of the G141 slitless spectrum at $1.10\text{--}1.66\ \mu\text{m}$, excluding $1.50\text{--}1.60\ \mu\text{m}$ to avoid potential contribution from a C III] $\lambda 1909$ line and the noisy edge of the spectrum (Fig. 1), and measure a best-fit UV continuum slope $\alpha_\lambda = 0.1 \pm 0.3$. By subtracting the best-fit power-law continuum and optimizing the integration range of the spectrum, the broad Ly α and C III] lines are tentatively detected at 3.7σ and 3.4σ levels, respectively. Assuming a C IV line width of full-width-half-maximum (FWHM) = $3,000\ \text{km s}^{-1}$ typical among $z > 6$ quasars from the literature^[48], we derive a 3σ upper limit of $\text{EW}_{\text{CIV}} = 10\ \text{\AA}$ from the grism spectrum.

In Extended Data Fig. 1, we compare the rest-frame UV properties of GNz7q with other populations at similar redshifts. The absolute UV luminosity at $1450\ \text{\AA}$ (M_{1450}) is estimated from the best-fit power-law model. We find that the UV luminosity of GNz7q falls between typical quasars and galaxies in the literature^[2,6,48,49], being ~ 10 times fainter and brighter than typical quasars and galaxies, respectively, where several faint quasars and luminous galaxies have been reported^[50,57]. The UV continuum slope of GNz7q is redder than that of any other object in either comparison population.

Extended Data Fig. 2 shows the rest-frame UV morphology in all *HST* WFC3/IR bands and the radial profile of GNz7q observed with *HST*/F125W. The instrumental point spread function (PSF) model can fully explain the rest-frame UV morphology / radial profile in all bands. The Sérsic profile fitting with *galfit*^[58] provides almost the same profile as the PSF model in all bands. Based on the F125W band, we measure an effective radius of $r_e = 0.06 \pm 0.07$ pixel (pixel scale = $0.''06$) and adopt a 2σ upper limit of $r_e < 0.02\ \text{pix} \simeq 60\ \text{pc}$. We obtain similar results in

the Sérsic profile fitting to the other bands.

3. NOEMA observations, data reduction, and measurements We observed [C II] with band 3 (1 mm), and CO(7–6), CO(6–5), and [C I](2–1) lines with band 1 (3 mm) of Institute Radio Astronomie Millimétrique (IRAM) Northern Extended Millimeter Array (NOEMA). The observations were carried out between 2020 June 17 and 2021 February 24 in various visits with the AC and D array configurations for the 1 mm and 3 mm observations, respectively, using 9–10 antennas. The data were processed in the standard manner with the pipeline using the latest version of the GILDAS software. We used CASA version 5.6 for the imaging⁵⁹.

For the [C II] observations, the upper side band (USB) of the 1 mm band receiver was tuned to 231.8 GHz in the first execution with the C configuration to cover the [C II] line at the source redshift estimated from the FUV Lyman continuum break (Fig. 1). After confirming the [C II] line detection, we tuned lower side band (LSB) with the A configuration to cover a wide frequency range for the continuum emission. For the 3 mm observations, the 3 mm band receiver was tuned to 97.7 GHz to cover CO(7–6) and [C I](2–1) in the USB and CO(6–5) in the LSB. In both observations, 0851+202 served as band-pass phase calibrator. Additional targets 1300+580 and 1044+719 were used for the phase and amplitude calibrations. We calibrated the absolute flux scale against MWC349 whose flux is regularly monitored at NOEMA. We adopt conservative uncertainties on the absolute scale of 20% and 10% in the 1 mm and 3 mm observations, respectively. The total integration time on-source was 6.8 and 13.5 hours in the 1 mm and 3 mm bands, respectively.

To maximize sensitivity, we used natural weighting for the imaging. The resulting 1 mm and 3 mm maps have synthesized beam FWHM of $0.''64 \times 0.''44$ and $4.''7 \times 4.''1$, with 1σ sensitivities of 21 and $6.3 \mu\text{Jy beam}^{-1}$ for the continuum, and 0.28 and $0.17 \text{ mJy beam}^{-1}$ for the line per 60 km s^{-1} channel, respectively. We produce the 1 mm and 3 mm continuum map from the all line-free channels, except for noisy channels around the central frequency channels of LSB and USB, and the line cubes with several velocity bins in the range of $40\text{--}60 \text{ km s}^{-1}$. The central wavelengths of the 1 mm and 3 mm continuum maps are 1.284 mm and 3.276 mm, respectively.

The continuum is detected both from the 1 mm and 3 mm maps with the peak intensity at 16σ and 3.9σ , respectively. The [C II] line is robustly detected at $232.060 \pm 0.013 \text{ GHz}$ with $\text{FWHM}=280 \pm 40 \text{ km s}^{-1}$ and 17σ peak intensity in the velocity-integrated map. This provides a precise determination of the source redshift of $z = 7.1899 \pm 0.0005$. At this source redshift, the CO(7–6) line is also detected at 5.7σ peak intensity in the velocity integrated map with FWHM of $770 \pm 230 \text{ km s}^{-1}$. While the velocity-integrated CO(7–6) line shows a signal at a marginal significance level, we regard it as a tentative detection due to the difference in its line width from that of the [C II] line beyond the 1σ uncertainties. With optimised apertures, we estimate emission line luminosities of $L_{[\text{CII}]} = (1.1 \pm 0.3) \times 10^9 L_{\odot}$ and $L_{\text{CO}(7-6)} = (1.3 \pm 0.7) \times 10^8 L_{\odot}$. For area-integrated line intensities, we measure $\log L'_{[\text{CII}]} = (5.0 \pm 1.4) \times 10^9 \text{ K km s}^{-1} \text{ pc}^2$ and $\log L'_{\text{CO}} = (6.3 \pm 3.0) \times 10^9 \text{ K km s}^{-1} \text{ pc}^2$, respectively. CO(6–5) and [C I](2–1) lines are not detected with 5σ upper limits of $L_{\text{CO}(6-5)} < 5.0 \times 10^7 L_{\odot}$ and $L_{[\text{CI}]2-1} < 7.9 \times 10^7 L_{\odot}$. In Extended Data Fig. 3 and 4, we summarize the continuum, velocity-integrated maps, and the line spectra. We do

not identify a clear velocity gradient in the [C II] line. We list the continuum flux densities and the line luminosities in Extended Data Table 1 and 2, which includes statistical errors based on the optimized apertures and the absolute flux uncertainty of 10–20%, as discussed above.

We measure the effective radius of the 1.3 mm continuum ($r_{\text{e,FIR}}$) and the [C II] line ($r_{\text{e,[CII]}}$) emission from GNz7q (Extended Data Fig. 5) using the visibility-based fitting CASA task UVMODELFIT. We adopt the 2-dimensional elliptical Gaussian model and leave all parameters free in the fitting routine. We obtain a best-fit major axis radius of $r_{\text{[CII]}} = 0.''26 \pm 0.''04$ (1.40 ± 0.21 kpc). On the other hand, the continuum fit did not converge and indicates that the continuum is not spatially resolved. The image-based fitting CASA task IMFIT provides consistent results for the spatially resolved [C II] line and unresolved continuum. We estimate an upper limit to the continuum size based on an approximate formula for the smallest resolvable scale (θ_{min}) sampled by the interferometric data^[60]:

$$\theta_{\text{min}} = \beta \left(\frac{\lambda_c}{2S/N^2} \right)^{1/4} \theta_{\text{beam}}, \quad (1)$$

where β is a coefficient, taking a range of 0.5–1.0, λ_c is a probability cutoff, and θ_{beam} is a synthesized beam FWHM. Based on the 2σ cutoff in the same manner as previous studies^[61–63] and $\beta = 1.0$, we obtain the upper limit in FWHM of $0.''18$ ($\sim 0.''09$ in effective radius) for the 1.3 mm continuum, corresponding to $r_{\text{FIR}} \leq 0.48$ kpc). In Extended Data Fig. 5, we present the observed and residual maps and the best-fit uv visibility plot for the continuum and [C II] line emission.

A nearby object is detected (9.4σ) in the 1.3 mm continuum map $\sim 3.''1$ northeast of GNz7q, corresponding to a projected separation of ~ 16 kpc if the two sources are at the same redshift. We obtain the 1.3 mm band photometry of $247 \pm 65 \mu\text{Jy}$ for this object within a $1.''0$ radius aperture. We do not identify any line features in either the 1 mm or 3 mm band spectra extracted at the position of the nearby object, nor do we identify any counterpart in the deep multi-wavelength imaging data from *HST* and *Spitzer*. We refer to this NIR-dark nearby object as “ND1” and consider two interpretations for it being either a chance projection of a foreground galaxy or a companion object associated with GNz7q. Such a population of NIR-dark galaxies has been recently reported, perhaps representing the most massive galaxies at $z \sim 3$ –5.^[64] However, given the low surface density of such galaxies, there is a very low probability to identify a ND1-like galaxy by chance within a small radius ($\sim 3.''1$) from another particularly unique object like GNz7q. In fact, similarly NIR-dark galaxies have been recently identified at the same redshift as nearby massive galaxies at $z \sim 7$.^[65] Therefore, we speculate that ND1 is a companion system associated with GNz7q which could eventually merge into a single system. Note that if we assume a typical modified blackbody (MBB) with $T_{\text{d}} = 35$ K and $\beta_{\text{d}} = 1.8$ for the FIR SED of ND1 and a typical $L_{\text{[CII]}}/L_{\text{IR}}$ ratio of 10^{-3} , the expected [C II] line luminosity agrees with the non-detection of the [C II] line from ND1 based on the current 1 mm data. This FIR SED for ND1 is also consistent with the non-detection in the 3 mm continuum due to its faintness and the cosmic microwave background (CMB) effect^[66–68]. ND1 was subtracted before the UVMODELFIT analysis, and thus the presence of ND1 does not affect our size measurements of the primary target.

4. Multi-band photometry The optical and NIR photometry are mainly obtained from the CHArGE data sets. We further obtain multi-band photometry of GNz7q from X-ray to radio from the literature. Given the spatial resolution of each instrument, we adopt the cross-match radii in the range from $1.''0$ to $4.''0$ with the public catalogs in the literature. We identify that GNz7q is robustly detected also at $2.2\ \mu\text{m}$ (Subaru/Multi-Object InfraRed Camera and Spectrograph (MOIRCS)), $24\ \mu\text{m}$ (*Spitzer*/MIPS), $850\ \mu\text{m}$ ((Submillimetre Common-User Bolometer Array 2 (SCUBA2)), and $20\ \text{cm}$ (JVLA) ^[69,72]. For other data at $100\text{--}500\ \mu\text{m}$ (*Herschel*/Photodetector Array Camera and Spectrometer (PACS) and the Spectral and Photometric Imaging Receiver (SPIRE)) and $3\ \text{cm}$ (JVLA) ^[73,75], we do not identify any counterparts. We measure the actual value at the source position in the JVLA $3\ \text{cm}$ map, while we conservatively adopt 3σ upper limits for the *Herschel* maps from the instrumental and confusion sensitivities. We summarize the multi-band photometry and the literature in Extended Data Table 1. Note that we identify three nearby objects in multiple *HST* maps within a radius of $\sim 2''$ (see top panels of Fig. 1). They are all detected in the F814W band and thus regarded as lower- z objects, instead of companion galaxies at $z \sim 7.2$. Their contribution to the *Spitzer* photometry of GNz7q is subtracted or negligible in our analysis. They are resolved in the IRAC images with de-blended $4.5\ \mu\text{m}$ flux densities of 0.08 ± 0.03 , 0.08 ± 0.02 and $0.19 \pm 0.02\ \mu\text{Jy}$, clockwise from the lower left. Besides, they are bluer than GNz7q, where galaxy templates that include re-emitted dust emission in the MIR have $24\ \mu\text{m} / 4.5\ \mu\text{m}$ ratios between 0.5 and 10. Thus, the faint neighboring sources together could contribute a maximum of 10%, likely much less, of the unresolved $24\text{-}\mu\text{m}$ source at the position of GNz7q.

For the SCUBA2 $450\ \mu\text{m}$ photometry, we produced the continuum map by utilizing the all archive data existing in this field. We reduced the data and performed the imaging in the standard manner with the pipeline based on the `starlink` software. We identify a tentative (2.6σ) detection in the $450\ \mu\text{m}$ map including the flux-boosting correction based on the previous studies ^[76]. Because we identify the nearby object of ND1 in the NOEMA $1\ \text{mm}$ map whose offset is smaller than the SCUBA2 beam, we subtract the potential contribution from ND1 in the $450\ \mu\text{m}$ photometry. We estimate the expected $450\text{-}\mu\text{m}$ flux density of ND1 by assuming the typical MBB ($T_d = 35\ \text{K}$ and $\beta_d = 1.8$) at $z = 7.2$. In the same manner, we also subtract the contribution from ND1 in the $850\ \mu\text{m}$ photometry. We then obtain the 450 and $850\ \mu\text{m}$ photometry of 8.0 ± 5.0 and $1.80 \pm 0.39\ \text{mJy}$, respectively, that are also listed in Extended Data Table 1. Note that we confirm the consistency between the results with and without the tentative detection at the $450\ \mu\text{m}$ band in the following analysis. Note that the contribution of the nearby faint *HST* objects to these FIR band photometry should be negligible due to the absence of the detection in the deep NOEMA $1\ \text{mm}$ map. Although ND1 is detected in the deep NOEMA $1\ \text{mm}$ map, a higher resolution $3\ \text{mm}$ map with the briggs weighting (robust = 0.0) whose beam size is smaller than the offset of ND1 from GNz7q shows $-4 \pm 9\ \mu\text{Jy}$ at the position of ND1. Thus, the $3\ \text{mm}$ photometry of GNz7q is not affected by ND1.

In X-rays, there are zero events in the relevant pixel of GNz7q, even with the *Chandra* $2\ \text{Ms}$ integration, some of the deepest X-ray data ever taken ^[72]. The *Chandra* $2\ \text{Ms}$ data have 0.171 mean background counts in the full band. Based on the continuous Poisson distribution with the mean of 0.171 , we compute an upper limit on the net counts of < 1.1 at the 99% confidence level. Assuming

an average photon index $\Gamma = 2.0$ obtained among high-redshift quasars up to $z = 7.5$,^{[4][77][78]} we use the online Portable Interactive Multi-Mission Simulator (see code availability) and estimate upper limits for the X-ray luminosity to be $L_{2\text{keV}} < 5.1 \times 10^{24} \text{ erg s}^{-1} \text{ Hz}^{-1}$ at 2 keV and $L_X < 3.9 \times 10^{42} \text{ erg s}^{-1}$ at 2–10 keV, including a correction for the Galactic absorption in this direction with the hydrogen column density of $N_H = 9.64 \times 10^{19} \text{ cm}^{-2}$.^[79] Note that the upper limit estimate depends on the choice of Γ . The typical range of spectral indices for X-ray AGN at this redshift is $\Gamma = 1.7\text{--}2.3$.^[80] Assuming even an extremely soft spectrum of $\Gamma = 2.3$ increases the luminosity upper limit by $\sim 30\%$, still well below the expected value.

Following the definition of the optical to X-ray spectral index

$$\alpha_{\text{ox}} \equiv \frac{\log(L_{2\text{keV}}/L_{2,500})}{\log(\nu_{2\text{keV}}/\nu_{2,500})}, \quad (2)$$

we then obtain the upper limit of $\alpha_{\text{ox}} < -2.23$. Note that here we use the dust corrected $L'_{2,500}$ value for $L_{2,500}$. In Extended Data Fig. 6, we show the tight correlation between α_{ox} and $L_{2,500}$ previously observed for local and high- z quasars.^{[18][77][81][82][85]} The stringent upper limit of the X-ray luminosity makes GNz7q deviate from this correlation by more than 5σ . In Fig. 3a, we show another tight correlation between α_{ox} and λ_{Edd} known to exist among local and high- z quasars.^{[18][77][81][84]} The black line shows the best-fit relation estimated in Lusso et al. (2010), and the gray shaded region represents the 68th percentile of the relation evaluated by propagation from the 1σ uncertainties of the parameters that define the best-fit relation. This best-fit relation yields $\log(\lambda_{\text{Edd}}) = 1.1 \pm 0.2$ at the upper limit of $\alpha_{\text{ox}} = -2.23$. This indicates that GNz7q has $\log(\lambda_{\text{Edd}}) > 0$ at the 5.5σ level, and if α_{ox} is much smaller than -2.23 , the significance level could be much increased. In Fig. 3b, the values of $L_X/L_{\text{IR,SF}}$ and M_{BH} for other populations are taken from the literature.^{[45][77][82][86][88]} We list the upper limits of the X-ray luminosity and α_{OX} in the Extended Data Table 2.

Note that it is difficult to explain the weak X-ray flux with selective absorption. This is because the redshift being so great means that we are observing very hard X-rays in the rest-frame. To extinguish the X-rays by a factor of 10 in the 2–10 keV observed frame requires a column density that is Compton-thick. This corresponds to a typical dust column of $A_V \gtrsim 1,000$, which is orders of magnitude larger than our estimate of $A_{V,\text{qso}} = 0.3 \pm 0.1$ based on the rest-frame UV, optical, and NIR emission (Section 6). Furthermore, even if we were to assume that all of the gas was highly-ionised and the dust sublimated, so as to have no effect at optical wavelengths, the electron scattering depth would still be substantial, extinguishing the UV light by the same factor. Therefore, the uniquely faint X-ray limits of GNz7q suggests either intrinsically weak X-ray emission, or Compton-thick material that covers only the inner part of the accretion disk and causes different extinction between the X-ray and the UV–NIR emission. The latter is aligned with the scenario that the inner region of the accretion disk is strongly inflated to a substantial height due to the high accretion, a so-called slim disk scenario.^{[19][89]}

5. Interpretations for GNz7q Based on the multi-wavelength observation results, there are two possible interpretations for GNz7q: a Type 1 red quasar or a very compact UV-luminous starburst region. The quasar interpretation is supported by the compact and luminous UV emission and the

bright detection at rest-frame $3\ \mu\text{m}$ (MIPS $24\ \mu\text{m}$). The starburst interpretation is motivated by the stringent upper limit on the X-ray luminosity and the absence of the clear detection of broad FUV emission lines between $\text{Ly}\alpha$ and $\text{C III}] \lambda 1909$.

In Extended Data Fig. 7, we compare the rest-frame UV size and luminosity of GNz7q with those of galaxies at $z > 5.5$ and UV-compact galaxies at $2 \lesssim z \lesssim 3$ ^[16, 50, 90, 91]. For a more direct comparison, we show the observed UV luminosity without any dust correction. If the UV emission of GNz7q is interpreted as arising from hot, recently-formed stars, we find that the implied SFR surface density reaches $\gtrsim 5,000\ M_{\odot}\ \text{yr}^{-1}\ \text{kpc}^{-2}$, exceeding that of even the UV-compact galaxies by two orders of magnitude; any dust correction will make this more extreme still. Such a luminous and compact object can more reasonably be explained by emission from a hot accretion disk in an active nucleus. In this case, the prominent FUV emission and its red continuum slope suggest that GNz7q is a Type 1 red quasar. The full SED analysis (Section 6) indicates an optical colour excess $E(B - V) = 0.11 \pm 0.03$, which would satisfy the colour threshold for red quasars used in the literature^[23, 25, 34].

Extended Data Fig. 8 shows NIR–MIR SED and colour properties of GNz7q, galaxies, and both local and high-redshift quasars. The most striking feature of the GNz7q SED is the red color between the observed *Spitzer* IRAC $8\ \mu\text{m}$ and MIPS $24\ \mu\text{m}$ band-passes, which probe rest-frame ~ 1 and $3\ \mu\text{m}$, respectively, at $z = 7.2$. A galaxy stellar population^[92, 93] fit to the GNz7q photometry at $\lambda_{\text{obs}} < 10\ \mu\text{m}$ is shown in the dark blue curve in Extended Data Fig. 8, along with additional stellar population templates that generally span the galaxy color space^[94] and that include highly obscured starbursts^[95]. Also shown is the average MIR spectrum of local quasars^[96] and a collection of broad-band SEDs of quasars at $z \sim 6$ ^[97]. The flux enhancement in the MIPS $24\ \mu\text{m}$ band cannot be explained by reasonable galaxy stellar population models, which are relatively blue between rest-frame wavelengths $1\text{--}3\ \mu\text{m}$ largely independent of star formation history and dust attenuation (i.e., the “ $1.6\ \mu\text{m}$ ” bump feature). However, the MIR colors of GNz7q are exactly consistent with the typical SEDs of quasars and active nuclei, where the rest-frame $3\ \mu\text{m}$ emission is thought to arise from hot dust associated with the nucleus^[98, 99].

Although the broad UV emission lines are not clearly detected (Fig. 1), a notable group of quasars with exceptionally weak UV broad emission lines have been identified at lower redshifts^[100, 101]. Moreover, the X-ray observations show that these weak emission line quasars also have weak X-ray emission with harder X-ray spectra and high-ionization lines that are more blue-shifted than those of normal quasars^[21, 102, 103]. This suggests that the X-ray absorption is due to Compton-thick material covering the inner part of the accretion disk, consistent with the slim disk scenario (Section 4). A similar implication of the small, Compton-thick absorber has also been obtained from the weak X-ray emission in one of the nearest quasars embedded in the ULIRG-class dusty starburst Mrk 231,^[104, 105] which also shows weak high-ionization lines and a very high nuclear outflow velocity of $\sim 8,000\ \text{km s}^{-1}$ ^[106, 107]. Therefore, the extremely faint X-ray luminosity and the lack of broad UV emission lines do not necessarily contradict the quasar interpretation. In summary, the combined multi-wavelength properties favor the quasar interpretation of GNz7q, although we cannot rule out the possibility that we are witnessing an emergence of the extremely compact and

luminous star-forming region that is without precedent in either the local or distant universe.

6. UV–millimeter SED fitting We fit the available photometry from UV to mm wavelengths with the SED-fitting code STARDUST^[108], which models the emission from stars, active galactic nucleus (AGN), and infrared emission arising from dust heated by star formation. The method does not rely on energy balance assumptions, and rather fits an independent linear combination of templates. To produce a fit we utilise the Draine & Li (2007)^[109] (hereafter DL07) models to fit the dust reprocessed stellar light, the AGN templates of Mullaney et al (2012)^[110] as well as stellar and a quasar UV to NIR templates adopted from EAZY and Shen et al. (2016)^[111], respectively. For the quasar component, we include the SMC dust attenuation law to reproduce the red colours of GNz7q in the UV and NIR.

The SED-fitting yields the basic global UV–IR properties of the source: the observed ($L_{2,500}$) and dust-corrected ($L'_{2,500}$) monochromatic rest-frame optical luminosity, the dust attenuation for the quasar ($A_{V,qso}$) and the host galaxy ($A_{V,host}$), the total infrared luminosity (L_{IR}) along with the relative contribution of the AGN ($L_{IR,AGN}$) and of the star formation ($L_{IR,SF}$) in the IR output of our target, the dust mass (M_{dust}), and the intensity of the mean radiation field ($\langle U \rangle \propto L_{IR}/M_{dust}$) of the source. The L_{IR} , $L_{IR,SF}$ and $L_{IR,AGN}$ are estimated by integrating the corresponding best-fit components over 8–1,000 μm . The bolometric luminosity of the AGN components that dominate the rest-frame FUV–NIR (observed 1–24 μm) is $L_{bol} = 1.7 \pm 0.1 \times 10^{46}$ erg s^{−1}, where the small uncertainty is the random normalization uncertainty of the fixed templates constrained by the high S/N *HST* and 24 μm detections. The systematic uncertainty due to the relatively unconstrained MIR SED shape is somewhat larger. For example, if we rather adopt the “high luminosity” MIR AGN template from Mullaney et al (2012)^[110] as opposed to the “low luminosity” template favored by the full fit, we obtain $L_{bol} = 1.3 \pm 0.1 \times 10^{46}$ erg s^{−1}.

For completeness and to facilitate comparison with the literature, we also fit the $\lambda_{rest} > 50$ μm data points with a MBB fit with a fixed $\beta_d=1.8$ in order to infer a luminosity-weighted dust temperature (T_d) and $L_{IR,SF}$. We obtain the consistent $L_{IR,SF}$ estimates from both fitting approaches. We use the average value of the $L_{IR,SF}$ estimates and evaluate its uncertainty from the difference of the both estimates from the average. During the fitting process of both the multi-component analysis and of the MBB modeling, we have taken into account CMB effects that could be non-negligible at this redshift, especially in the mm regime^[66,68]. The AGN/SF decomposition of the IR emission allows for an independent estimate of the obscured star-formation rate of the source. Using the conversion of Murphy et al. (2011)^[112] based on the Kroupa initial mass function (IMF)^[113], the inferred $L_{IR,SF}$ corresponds to an SFR of $1,600 \pm 700 M_{\odot} \text{ yr}^{-1}$, where the uncertainty of $L_{IR,SF}$ is propagated to the uncertainty. Note that if we assume that GNz7q is rather a compact star-forming region, the obscured and un-obscured SFRs can be derived from L_{IR} and the observed UV luminosity, respectively. Using the same conversion, we obtain an even higher estimate of $1,900 \pm 900 M_{\odot} \text{ yr}^{-1}$, but we adopt the former estimate throughout this paper based on the discussion in Section 5. With the upper limit of the rest-frame FIR emitting region from the NOEMA 1 mm data, we also obtain a lower limit of the SFR density at $\geq 1,100 M_{\odot} \text{ yr}^{-1} \text{ kpc}^{-2}$, which is comparable to its maximal value predicted from the balance between the radiation pressure from the

star-formation and the self gravitation^{[13][114]}. Note that DL07 and MBB models assume the optically thin dust emission, while the compactness of the FIR-emitting region in GNz7q could in fact require an optically-thick treatment. With the optically-thick assumption, $L_{\text{IR, SF}}$ and subsequently the SFR estimate will be increased by $\sim 10\%$ ^[66].

Note that the uniquely deep X-ray data and general L_X and SFR relations^[115] expect the X-ray detection from such an enormous SFR of the host galaxy, while the non-detection of X-ray could be explained by the metallicity dependence of the L_X and SFR relation: L_X at a given SFR decreases in high-metallicity systems^[116]. Based on the recent calibration of Fornasini et al. (2020)^[117], we find that at least the solar value is required for the gas-phase metallicity of the host galaxy to meet the X-ray upper limit (Section 4). Since the SFR estimate for the host galaxy of GNz7q is not much changed regardless of its interpretation, the non-detection of X-ray thus suggests that the host galaxy has experienced a rapid metal enrichment and reached the solar metallicity even at $z = 7.19$. This is consistent with recent ALMA results for a similarly distant quasar at $z = 7.54$ that the ISM gas-phase metallicity of the host galaxy is comparable to the solar value via FIR line diagnostics^[118].

7. M_{gas} and M_{dyn} estimates We obtain molecular gas mass M_{gas} estimates from five empirical calibrations; the metallicity-dependent δ_{GDR} -method^[119] ($M_{\text{gas(dust)}}$), the monochromatic Rayleigh-Jeans (RJ) dust continuum approach^[120] ($M_{\text{gas(RJ)}}$), the CO line luminosity, the [C II] line luminosity^[122] ($M_{\text{gas([CII])}}$), and the [C I](2-1) line luminosity^[125] ($M_{\text{gas([CI])}}$). First, following the method in the previous studies^[119], we convert the inferred M_{dust} to M_{gas} adopting a typical gas-to-dust ratio at solar metallicity, $\delta_{\text{GDR}}=92$. This yields $M_{\text{gas(dust)}}/M_{\odot} = (1.4 \pm 1.0) \times 10^{10}$ which agrees well with the estimate inferred from the RJ approach, $M_{\text{gas(RJ)}}/M_{\odot} = (2.0 \pm 1.2) \times 10^{10}$. For the latter we adopted a RJ luminosity-to-mass ratio $\alpha_{\nu} \equiv L_{\nu 850\mu\text{m}}/M_{\text{gas}} = 6.7 \times 10^{19} \text{ erg s}^{-1} \text{ Hz}^{-1} M_{\odot}^{-1}$ calibrated with star-forming galaxies including local spirals and $z \sim 2$ SMGs following Scoville et al. (2016)^[120]. Moving to the line tracers, we first estimate the area-integrated CO(1-0) intensity of $L'_{\text{CO(1-0)}} = (1.3 \pm 0.6) \times 10^{10} \text{ Jy km s}^{-1} \text{ pc}^2$ from the CO(7-6) line detection, assuming $L'_{\text{CO(1-0)'}} = 1.5 \times L'_{\text{CO(7-6)}}$ estimated in dusty starburst galaxies at $z > 6$ in the literature^{[123][124]}. We then estimated $(M_{\text{gas(CO)}}/M_{\odot}) = (5.0 \pm 3.0) \times 10^{10}$ assuming $\alpha_{\text{CO}} \equiv M_{\text{gas}}/L'_{\text{CO}} = 4.6 M_{\odot} (\text{K km s}^{-1} \text{ pc}^2)^{-1}$ ^[121]. To convert the [C II] line to M_{gas} , we use the $\alpha_{[\text{CII}]} \equiv M_{\text{gas}}/L_{[\text{CII}]} = 22 M_{\odot}/L_{\odot}$ conversion factor, as calibrated on starburst galaxies at $z \sim 2 - 6$ by^[122] and, adopting a 0.2 dex uncertainty, we estimate $M_{\text{gas([CII])}} = 3.2 \pm 1.8 \times 10^{10} M_{\odot}$. Finally, we consider the measured 3σ upper limit of the [C I](2-1) line that for $\alpha_{[\text{CI}](2-1)} \equiv M_{\text{gas}}/L_{[\text{CI}](2-1)} = 34 M_{\odot} (\text{K km s}^{-1} \text{ pc}^2)^{-1}$ and a 0.32-dex uncertainty^{[125][126]}, we obtain an upper limit of $M_{\text{gas([CI])}}/M_{\odot} < 1.6 \times 10^{11}$. These independent M_{gas} estimates are in excellent agreement with each other within the uncertainties. For the purposes of this work, we determine $M_{\text{gas}} = (2.0 \pm 1.2) \times 10^{10} M_{\odot}$ by adopting the median value among $M_{\text{gas(dust)}}$, $M_{\text{gas(CO)}}$ and $M_{\text{gas([CII])}}$. Note that the $M_{\text{gas(CO)}}$ and $M_{\text{gas(dust)}}$ estimates could be decreased by factors of several based on another assumptions of $\delta_{\text{GDR}} = 30$ and $\alpha_{\text{CO}} = 0.8$ in the super-solar metallicity case.

We estimate the dynamical mass of the system M_{dyn} from the [C II] line results. Given the absence of the clear velocity gradient in GNz7q, we interpret GNz7q's host as a dispersion-

dominated system and assume a virialized body with a radius R kpc and one-dimensional velocity dispersion σ km s⁻¹, which yields^[127]

$$M_{\text{dyn}} [M_{\odot}] = 1.56 \times 10^6 \sigma^2 R. \quad (3)$$

For consistency with previous studies^{[128][130]}, we use $r_{\text{e,[CII]}}$ for R after applying a correction factor of 1.5 to recover the contribution of the diffuse emission. We then obtain $M_{\text{dyn}} = (4.5 \pm 0.9) \times 10^{10} M_{\odot}$ which satisfies the requirement that it should be larger than the M_{gas} estimate. Because of the negligible dark matter contribution to M_{dyn} within a compact scale of $1.5 \times r_{\text{e,[CII]}}$, we subtract M_{gas} from M_{dyn} and derive M_{star} of $(2.5 \pm 1.4) \times 10^{10} M_{\odot}$. If we apply a typical dust-to-stellar mass ratio of 0.01 to M_{dust} , we obtain another estimate of M_{star} of $(1.6 \pm 1.1) \times 10^{10} M_{\odot}$ which is consistent with the above M_{star} estimate within the uncertainties.

8. Comparison with other populations Compared to Type 2 quasars that are almost completely obscured in the UV/optical due to the nearly edge-on view of the dust torus, the reddening of red quasars is more moderate than that of Type 2 quasars, where the sight lines to red quasars may graze the dusty material surrounding the accretion disk. In this context, red quasars are thought to represent an early phase of the quasar life cycle: an obscured phase before the energy output from radiation, winds, and/or jets from the AGN and central star formation expels the obscuring material and transitions to an unobscured blue quasar. This is consistent with hydrodynamical simulations of galaxy mergers and quasar feeding^[5], and supportive observational results have been also reported at $z \sim 1 - 3$.^{[23][131][133]} To investigate whether GNz7q at $z = 7.2$ is also consistent with this scenario, we compare its physical properties with those of dusty starbursts, red quasars, and blue quasars in the literature.

In Extended Data Fig. 9, we compare $L_{\text{[CII]}}$ and L_{IR} properties. We find that GNz7q has a relatively low $L_{\text{[CII]}}/L_{\text{IR}}$ ratio among all populations, while we confirm that there are two cases whose ratios are comparable in a similar composite system of the AGN and starburst population identified in the local universe. We also find that the red quasars at $z \sim 3 - 5$, including super-Eddington accretion red quasar W2246-0526 at $z = 4.6$,^[17] have similarly low $L_{\text{[CII]}}/L_{\text{IR}}$.^[43] These results suggest that the interstellar medium (ISM) conditions of GNz7q are unique compared to similar systems in the local universe and could resemble the super-Eddington accretion red quasars. The $L_{\text{[CII]}}/L_{\text{IR}}$ ratio is known to have a strong anti-correlation with ΣL_{IR} , where the decrease of the [C II] emissivity is explained by the strong radiation field in the high ΣL_{IR} regions.^{[134][136]} Based on the upper limit of the effective radius of the 1.3 mm continuum (Section 4), we find that GNz7q has a lower limit of $\log(\Sigma L_{\text{IR}}) > 13 L_{\odot}/\text{kpc}^2$, aligned with the highest-end of the anti-correlation. This indicates that the host of GNz7q is one of the most vigorously star-forming systems, which produces the intense radiation field and induces the high T_{d} of 80 ± 21 K.

In Extended Data Fig. 10, we compare SFR (a), M_{gas} (b), M_{dust} (c), and the depletion time scale τ_{depl} . (d) between GNz7q and other populations at $z > 6$, including the dusty starburst (orange diamond), the blue quasar (blue squares), the red quasar (magenta circle and shaded region), and Lyman-break galaxies (green triangle) from the literature.^{[6][25][29][43][49][86][118][124][137][143]} For dusty starburst and quasar populations, we estimate SFR in the same manner as GNz7q (i.e., conversion

from L_{IR}). The L_{IR} values are systematically calculated by assuming a single MBB with $T_d = 47$ K and $\beta_d = 1.6$ in the same manner as previous studies^{[86][142]} when the source has been observed only with a single submm or mm band, otherwise taken from the literature^{[124][143][144]}. Because we show GNz7q with $L_{\text{IR, SF}}$ after subtracting the AGN contribution of 8%, we also subtract the same fraction in the L_{IR} estimates for the quasar populations. For the dusty starburst at $z = 6.9$ ^{[29][124]}, the AGN contribution is estimated to be 0–50% with nearly equal probability in the literature, where we assume it to be 25% and evaluate the error with the cases of 0% and 50%. For the galaxy population, we use SFR estimates compiled in the literature^[145] applying corrections for the different IMF assumptions. M_{gas} estimates are taken from CO line spectroscopy results in the literature, while we also derive M_{gas} from $L_{\text{[CII]}}$ with $\alpha_{\text{[CII]}} = 22M_{\odot}/L_{\odot}$ ^[122] for the sources that have not been observed with the CO lines. For M_{dust} , we systematically calculate by using the equation:

$$M_{\text{dust}} = \frac{F_{\nu} D_L^2}{(1+z)\kappa_{\nu}(\beta_d)B_{\nu}(\nu, T_d)}, \quad (4)$$

where F_{ν} is the flux density at a frequency of ν , D_L is the luminosity distance, B_{ν} is the Planck function. The dust opacity coefficient is given by $\kappa_{\nu} = 5.1(\nu/\nu_{250\mu\text{m}})^{\beta_d} \text{ cm}^2 \text{ g}^{-1}$, which is used in the DL07 model. We assume $T_d = 47$ K and $\beta_d = 1.6$ when the source has been observed only with a single submm or mm band, otherwise we use the longest wavelength measurements available and adopt the T_d and β_d estimates in the literature^{[124][139][143][144]}. Because a ± 10 K change in the T_d assumption produces a ~ 0.2 – 0.3 dex difference in the SFR and M_{dust} estimates for the sources whose T_d and β_d are assumed to be 47 K and 1.6, respectively, we add a 0.2 dex uncertainty to the 1σ measurement uncertainty in their error bars in Extended Data Fig. 10. We find that the M_{dust} and M_{gas} values of GNz7q fall in the range probed by blue quasars at the same epoch, while the implied SFR is the highest among $z > 7$ objects so far observed.

This indicates that the intense starburst is taking place in the host galaxy of GNz7q with a very short depletion time scale of ~ 10 Myr, which is consistent with the scenario that the red quasar is forming in the dusty starburst. The higher SFR of GNz7q’s host galaxy than the hosts of lower-redshift red quasars may indicate that GNz7q is experiencing an early stage of its transition phase from the dusty starburst to the blue quasar. In fact, the M_{BH} values of the lower- z red quasars are estimated to fall in the super-massive regime of $\log(M_{\text{BH}}/M_{\odot}) = 9.3$ – 9.6 ^{[17][25]}. This may suggest that these lower- z red quasars are found at the end phase of the SMBH evolution and that the super-Eddington accretion in W2246-0526 is caused by an active quasar duty cycle even at the end phase. Although there is a possibility that the SFR values in the blue quasars are also increased if their T_d values are as high as GNz7q, previous studies of blue quasars at $z > 6$ with multi-band FIR photometry in ALMA Bands 6 and 8 generally show the FIR SED with $T_d \sim 40$ – 50 K^{[118][143][144]}. Furthermore, the rest-frame IR regime of the blue line in Fig. 2 implies that the AGN contribution to L_{IR} could be much larger in these luminous quasars than that of GNz7q, which reduces their SFR estimates.

In Extended Data Fig. 11, we present M_{BH} and M_{dyn} properties. For GNz7q, we show the Eddington-limited M_{BH} estimate in the red circle and the potential M_{BH} range indicated by its extremely faint X-ray property in the red shaded regions whose colour scale and vertical range of

each red shade region correspond to those of Fig. 3. For blue quasars at $z \sim 6-7$, we show M_{dyn} measurements based on a systematic kinematic modeling with the ALMA data^{[146][147]}. We include M_{dyn} measurements for $z \gtrsim 6$ quasars based on the assumption of the rotating disc geometry and the axial ratio of [C II] flux map as proxy of the disc inclination angle^{[128][130][148][151]}. We also show 1) the best-fit relation between the stellar mass in the bulge M_{bulge} and M_{BH} obtained in local quiescent galaxies^[152], where we use M_{bulge} as M_{dyn} , and 2) the M_{BH} and M_{star} relation for red quasars at $z \sim 2$ that generally fall below the local relation^[133]. We find a relatively low fraction of $M_{\text{BH}}/M_{\text{dyn}} < 0.2\%$ that falls below the local relation^[152], similar to the general relation of the red quasars at $z \sim 2$. In contrast to ideas of an “over-massive” SMBH relative to the host galaxy reported in previous optically-luminous quasars at $z > 6$ ^{[128][129][146][147]}, the “under-massive” SMBH of GNz7q at $z = 7.2$ offers an intriguing path to the co-evolution between the SMBH and its host in the early universe: the host galaxies grow earlier than the SMBHs, which is aligned with the predictions of the merger-driven SMBH evolution models^[5]. The “under-massive” SMBH is also argued in recent reports of less luminous quasars^{[130][148]}, but generally their central BHs are already massive $M_{\text{BH}} \sim 10^{8-9} M_{\odot}$ with low $\lambda_{\text{Edd}} \sim 0.1-0.2$ ^{[56][86]}, and thus could be placed at the end-phase of the SMBH evolution after the blue-quasar phase^[86]. We note that it is unclear whether the quasar host galaxies at $z \gtrsim 6$ are the bulge-dominated systems similar to the local quiescent galaxies. We thus also show another best-fit relation between the stellar mass of the entire system and M_{BH} among local AGNs^[153], where the M_{BH} range of GNz7q still falls below or on the best-fit relation.

9. The IR/Radio correlation The correlation between the IR luminosity and the radio emission has been empirically known for several decades. Based on the radio detection of GNz7q at 20 cm/1.5 GHz^[72], we evaluate the IR and radio correlation. From previous studies, the IR and radio correlation is typically evaluated with the parameter q_{IR} given by

$$q_{\text{IR}} = \log \left(\frac{1.01 \times 10^{18} L_{\text{IR}}}{4\pi D_L^2 [L_{\odot}]} \right) - \log \left(\frac{10^{-32} F_{1.4\text{GHz}}}{(1+z)^{\alpha_{\text{radio}}-1} [\mu\text{Jy}]} \right), \quad (5)$$

where $F_{1.4\text{GHz}}$ is the observed flux density at 1.4 GHz, and α_{radio} is the radio spectral index which is defined by $F_{\nu} \propto \nu^{\alpha_{\text{radio}}}$. Due to the non-detection in a deep 10-GHz map^[75], we obtain a constraint of $\alpha_{\text{radio}} < -1.0 \pm 0.6$. Given the large uncertainty for the α_{radio} constraint, we adopt a typical value of $\alpha_{\text{radio}} = -0.75$ ^{[75][154]}, convert the observed 1.5 GHz flux density to the 1.4 GHz flux density, and estimate the q_{IR} value of 2.1 ± 0.3 . This is consistent with typical value range of local starburst galaxies and high-redshift SMGs^{[155][156]} and a recent report of the redshift trend among high-redshift star-forming galaxies^[157]. We thus conclude that the majority of the radio emission of GNz7q is caused by the star-formation and classify the GNz7q as a radio-quiet object. This property agrees with the young quasar interpretation of GNz7q, because the radio loudness is a strong function of M_{BH} and the radio-loud quasars generally have massive $M_{\text{BH}} > 10^9 M_{\odot}$ ^[158].

10. Comparison with simulations We compare our observational results with predictions from a data-constrained, semi-analytic model GAMETE/QSOdust (GQd) aimed at studying the formation and evolution of high-redshift quasars and their host galaxies in a cosmological framework^{[74][22][159][160]}. Here we have analysed the hierarchical merger histories of 10 massive dark matter (DM) halos

with $M_{\text{halo}} = 10^{13} M_{\odot}$ at $z = 6.4$, designed to reproduce the observed properties of the optically-luminous quasar SDSSJ1148 at $z = 6.4$ to investigate if we could identify—among its progenitors at $z = 7.2$ —systems with physical properties similar to GNz7q.

We first produce the 10 merger trees for a $10^{13} M_{\odot}$ dark matter halo, decomposing it into its lower mass progenitors backward in time from $z = 6.4$ to $z = 24$ using a Monte Carlo algorithm based on the Extended Press-Schechter formalism. Then GQd follows the evolution of the baryonic component within each progenitor halo along a merger tree, from $z = 24$ down to $z = 6.4$. At each redshift, in each halo, we follow the formation of stars and BHs (light and heavy seed formation channels are simultaneously implemented in the model) according to the environmental properties (i.e. metallicity of the interstellar medium and the level of illuminating external ionizing and H_2 photo-dissociating radiation field). BHs in the centre of galaxies can then grow via gas accretion and mergers with other BHs during major halo-halo mergers (DM halos pair mass ratio $> 1:4$) while in minor ($< 1:4$) halo-halo mergers the least massive of the two BHs is “ejected”, it is considered as a “satellite”, and we do not further follow its evolution. We account for the effect of stellar- and AGN-driven feedback in the form of energy-driven winds and include calculations for the IR luminosity from the host galaxy and the X-ray luminosity from accreting BHs, considering the primary component of the hot corona and the reflection component of the surrounding neutral medium.

For each $10^{13} M_{\odot}$ halo merger tree GQd generates a catalog of progenitor galaxies. Each catalog contains the properties (e.g., the mass of BH, gas, stars, dust, SFR, L_{IR} , L_{X}) of all the progenitor systems (galaxy+BHs). In Figure 4a, we show all the progenitors at redshift slices of $z = 7.1, 7.2$, and 7.3 in the catalog. We mark four progenitors with black circles whose X-ray, optical, and host galaxy properties are close to GNz7q with criteria of $L_{\text{X}} < 10^{44.5} \text{ erg s}^{-1}$, $\alpha_{\text{ox}} < -1.9$, and $\text{SFR} > 100 M_{\odot} \text{ yr}^{-1}$. The assembly histories of the nuclear BHs for these four progenitors are shown in Figure 4b. We find that one of the four progenitors is a direct progenitor which evolves into a luminous quasar harbouring a SMBH with $M_{\text{BH}} > 10^8 M_{\odot}$ at $z = 6.4$. We also confirm that this progenitor resides in one of the most massive dark matter halos at $z \sim 7.2$ with the dark matter halo mass of $\sim 10^{12.2} M_{\odot}$. In the assembly histories, the remaining three progenitors are subsequently ejected from the centre or become satellites of more massive BHs as a consequence of a minor merger experienced by their host galaxies. Therefore, these BHs are not the direct progenitors of the final SMBHs, although their entire systems also form the SMBHs with $M_{\text{BH}} \sim 10^{9.5-10} M_{\odot}$ at $z = 6.4$.

11. Classical colour selection for high- z quasar search A remarkable aspect is that GNz7q is discovered in a relatively small area coverage of the entire *HST* archive ($\sim 3 \text{ deg}^2$), compared with previous wide-area surveys used in the high-redshift quasar search. Mazzucchelli et al. (2017) use Pan-STARRS1^[163] and UKIDSS^[164] data and select the quasar candidates at $z > 6.5$ with the optical colour criteria of^[162]

$$z - y > 1.4 \quad (6)$$

$$y - J \leq 1.0. \quad (7)$$