



Publication Year	2016
Acceptance in OA	2020-05-21T17:15:37Z
Title	Lithium evolution in metal-poor stars: from pre-main sequence to the Spite plateau
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Publisher's version (DOI)	10.1093/mnras/stv1384
Handle	http://hdl.handle.net/20.500.12386/25064
Journal	MONTHLY NOTICES OF THE ROYAL ASTRONOMICAL SOCIETY
Volume	452

Lithium evolution in metal-poor stars: from pre-main sequence to the Spite plateau

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Accepted 2015 June 18. Received 2015 June 18; in original form 2015 March 11

ABSTRACT

Lithium abundance derived in metal-poor main-sequence (MS) stars is about three times lower than the value of primordial Li predicted by the standard big bang nucleosynthesis when the baryon density is taken from the cosmic microwave background or the deuterium measurements. This disagreement is generally referred as the lithium problem. We here reconsider the stellar Li evolution from the pre-main sequence (PMS) to the end of the MS phase by introducing the effects of convective overshooting (OV) and residual mass accretion. We show that ${}^7\text{Li}$ could be significantly depleted by convective OV in the PMS phase and then partially restored in the stellar atmosphere by a tail of matter accretion which follows the Li-depletion phase and that could be regulated by EUV photoevaporation. By considering the conventional nuclear burning and microscopic diffusion along the MS, we can reproduce the Spite plateau for stars with initial mass $m_0 = 0.62\text{--}0.80 M_\odot$, and the Li declining branch for lower mass dwarfs, e.g. $m_0 = 0.57\text{--}0.60 M_\odot$, for a wide range of metallicities ($Z = 0.00001$ to $Z = 0.0005$), starting from an initial Li abundance $A(\text{Li}) = 2.72$. This environmental Li evolution model also offers the possibility to interpret the decrease of Li abundance in extremely metal-poor stars, the Li disparities in spectroscopic binaries and the low Li abundance in planet hosting stars.

Key words: stars abundances – stars: pre-main sequence.

1 INTRODUCTION

${}^7\text{Li}$ is one of the four isotopes synthesized in the primordial nucleosynthesis. Its primordial abundance depends mainly on the baryon-to-photon ratio with only a minor sensitivity to the universal speed-up expansion rate, namely the number of neutrino families (Fields 2011; Fields, Molaro & Sarkar 2014). The universal baryon density can be obtained either from the acoustic oscillations of the cosmic microwave background (CMB) observation or independently, from the primordial deuterium abundance measured in unevolved clouds of distant quasar spectra (Adam 1976). Observations based on the *Wilkinson Microwave Anisotropy Probe* (Komatsu et al. 2011) predict a primordial ${}^7\text{Li}$ abundance $A(\text{Li}) = 2.72^1$ (Coc et al. 2012). From the baryon density measured by the Planck mission (Planck Collaboration XVI 2014) Coc et al. (2014) calculated the primordial value of ${}^7\text{Li}/\text{H}$ to be $4.56 \sim 5.34 \times 10^{-10}$ ($A(\text{Li}) \approx 2.66\text{--}2.73$).

Population II (POP II) main-sequence (MS) stars show a constant ${}^7\text{Li}$ abundance (Spite & Spite 1982), which was interpreted

as an evidence that these stars carry the primordial ${}^7\text{Li}$ abundance because of their low metallicity. In the past three decades, observations of metal-poor MS stars both in the Milky Way halo (Spite & Spite 1982; Sbordone et al. 2010), and in the globular clusters (Lind et al. 2009) confirm that the ${}^7\text{Li}$ abundance remains $A(\text{Li}) \approx 2.26$ (Molaro 2008). This abundance, which defines the so-called Spite plateau, is three times lower than the predicted primordial value. The discrepancy is the long-standing ‘lithium problem’.

There are several lines of study which have been pursued to provide possible solutions to the problem: (i) nuclear physics solutions which alter the reaction flow into and out of mass-7 (Coc et al. 2012); (ii) new particle physics where massive decaying particles could destroy ${}^7\text{Li}$ (Olive et al. 2012; Kajino et al. 2012); (iii) chemical separation by magnetic field in the early structure formation that reduces the abundance ratio of Li/H (Kusakabe & Kawasaki 2015); (iv) ${}^7\text{Li}$ depletion during MS evolution. It has been argued that certain physical processes, which may occur as the stars evolve on the MS, could cause the observed lithium depletion. Among these processes, we recall gravitational settling (e.g. Salaris & Weiss 2001; Richard, Michaud & Richer 2005; Korn et al. 2006) or possible coupling between internal gravity waves and rotation-induced mixing (Charbonnel & Primas 2005).

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¹ $A(\text{Li}) = 12 + \log [n(\text{Li})/n(\text{H})]$ where n is number density of atoms and 12 is the solar hydrogen abundance.

In this paper, we address the lithium problem on the ground of the pre-main-sequence (PMS) stellar physics, under the assumption that the standard Big Bang Nucleosynthesis (BBN) is correct. Hereafter when not otherwise specified for Li we refer to the isotope 7. The surface Li evolution during the PMS has been analysed in stellar evolution models with solar metallicity by several authors (D’Antona & Mazzitelli 1984; Soderblom, Fedele & Jones 1993; D’Antona & Mazzitelli 1994; Swenson et al. 1994; Ventura et al. 1998; Piau & Turck-Chièze 2002; Tognelli, Degl’Innocenti & Prada Moroni 2012), with the general conclusion that mixing efficiency plays a key role on Li depletion. Recent extensive studies in metal-rich young open clusters do show that the observed lithium abundance is modified during the PMS phase (Somers & Pinsonneault 2014). In this respect, two points are worthy of consideration. On one side, it has been shown that overshoot at the base of the solar convective envelope is more efficient than hitherto believed (Christensen-Dalsgaard et al. 2011), and could be even more in the envelope of a PMS star. This would favour a more efficient Li depletion. On the other side, there is evidence that a residual mass accretion (also called late accretion) persists for several 10^7 yr and is observed up to the early MS phase (De Marchi et al. 2011). This may act to partially restore Li in metal-poor PMS stars. Molaro et al. (2012) first proposed that Li abundance both in the Population I (POP I) and POP II stars could be modified due to the combined effects of efficient overshoot and late mass accretion during the PMS evolution. In this study, we use PARSEC (PAдова and TRieste Stellar Evolution Code; Bressan et al. 2012) to examine the working scenario proposed by Molaro et al. (2012) in a more systematic way.

The structure of the paper is as follows. Section 2 describes the theoretical PMS models, which include envelope overshooting (OV), residual accretion and EUV photoevaporation. Section 3 describes the Li evolution in MS. Section 4 presents the results of the models and the comparison with the observations. An ample discussion and the main conclusions are drawn in Section 5.

2 PRE-MAIN-SEQUENCE LI EVOLUTION

PMS is the direct continuation of the protostellar phase. Initially, as the young stellar object (YSO) evolves along its stellar birth line, its luminosity is mainly supported by an accretion process strong enough to maintain active deuterium fusion (Stahler 1983). Once the accretion ceases, the protostar, surrounded only by a residual disc, descends along its Hayashi line almost vertically in the Hertzsprung–Russell diagram (H-R diagram), undergoing a rapid gravitational contraction.

In this phase, evolution of PMS stars critically depends on their mass, and lithium can be burned at different stages through the reaction ${}^7\text{Li}(p,\alpha){}^4\text{He}$. Lithium is very fragile as the nuclear reaction rate R_{nuclear} of ${}^7\text{Li}(p,\alpha){}^4\text{He}$ becomes efficient already at temperatures of a few million Kelvin. The effective Li-burning temperature for PMS stars, i.e. that needed to consume Li in a time-scale of $\sim 10^7$ yr, is $\sim 4 \times 10^6$ K. The R_{nuclear} is adopted from JINA REACLIB data base (Cyburt et al. 2010).

Very low mass stars with initial mass $m_0 < 0.06 M_\odot$ never reach this temperature. More massive PMS stars experience Li burning, which initially affects the entire stellar structure as long as it is fully convective. Later, at the formation of the radiative core, the extent of Li burning can vary, depending on the mass of the star. Thereafter, the Li evolution is critically affected the temperature at the base of the convective envelope, and the efficiency of overshoot at the base of the convective envelope begins to play a significant role.

We use the stellar evolution code PARSEC to calculate the PMS evolution of low-mass stars with initial mass from 0.50 to $0.85 M_\odot$, with metallicity $Z = 0.0001$ ($\sim [\text{M}/\text{H}] = -2.2$ dex), typical of POP II stars. The adopted helium abundance is $Y_p = 0.249$, based on the helium enrichment law (Bressan et al. 2012):

$$Y = Y_p + \frac{\Delta Y}{\Delta Z} Z = 0.2485 + 1.78Z, \quad (1)$$

where $Y_p = 0.2485$ is the adopted primordial value and $\Delta Y/\Delta Z$ is the helium-to-metal enrichment ratio. Following the Standard Big Bang Nucleosynthesis (SBBN) prediction, we set the initial lithium abundance $A(\text{Li}) = 2.72$. A solar-model-calibrated mixing length parameter $\alpha_{\text{MLT}} = 1.74$ is adopted.

2.1 Convective overshooting

OV is the signature of the non-local convective mixing in the star that may occur at the borders of any convectively unstable region (Bressan et al. 2015). As already anticipated, efficient OV at the base of the convective envelope may significantly affect PMS surface Li depletion, because OV extends the mixed region into the hotter stellar interior.

Recently, Christensen-Dalsgaard et al. (2011) suggest that an OV region of size $\Lambda_e \approx 0.3 \sim 0.5 H_p$ (where H_p is the pressure scale-height), at the base of the convective envelope of the Sun, provides a better agreement with the helioseismology data. These values are higher than those of earlier studies (e.g. Miglio & Montalbán 2005), and indicate that the overshoot mixing could be much more efficient than previously assumed. It should be recalled that by applying an efficient envelope overshoot, many other observational data could be better reproduced, such as the blue loop of intermediate and massive stars (Alongi et al. 1991), the low-mass metal-poor stars (Bressan et al. 2012), the carbon stars luminosity functions in the Magellanic Clouds via a more efficient third dredge-up in asymptotic giant branch stars (Herwig 2000; Marigo & Girardi 2007), and the location of the red giant branch (RGB) bump in old star clusters (Alongi et al. 1991).

Soon after the formation of the radiative core, we assume that the OV at the base of the deep convective envelope is quite efficient, with $\Lambda_e \sim 1.5 H_p$. This value is intermediate between the one suggested for the Sun (Christensen-Dalsgaard et al. 2011) and the larger values ($\Lambda_e = 2 - 4 H_p$) suggested by Tang et al. (2014) for intermediate and massive stars. We will show in Section 5 that this value is not critical, since even assuming a maximum overshoot $\Lambda_e \sim 0.7 H_p$, the model can well reproduce the data. The OV distance is computed using the natural logarithm of the pressure, downward from the Schwarzschild border (P_{sch}) at the bottom of the convective envelope, i.e. an underlying stable layer (P_{lay}) is mixed with the envelope if $\ln(P_{\text{lay}}) - \ln(P_{\text{sch}}) \leq \Lambda_e$. Then, as the star moves towards the zero age main sequence (ZAMS), we vary the OV efficiency proportionally to the mass size of the outer unstable convective region (f_{cz}), until the typical value estimated for the Sun $\Lambda_e = 0.3 H_p$ (Christensen-Dalsgaard et al. 2011) is reached:

$$\Lambda_e = 0.3 + (1.5 - 0.3) * f_{\text{cz}}. \quad (2)$$

This assumption produces an efficient photospheric Li depletion during the PMS phase, while preserving the solar constraints for the MS stars. Fig. 1 illustrates the temporal evolution of the size of the convective and overshoot regions for four values of the initial stellar mass, $m_0 = 0.5, 0.6, 0.7, 0.8 M_\odot$ at metallicity $Z = 0.0001$. The mass coordinate, on the Y-axis, goes from 0, at the centre of the star, to 1 at the surface. The stars are initially fully convective

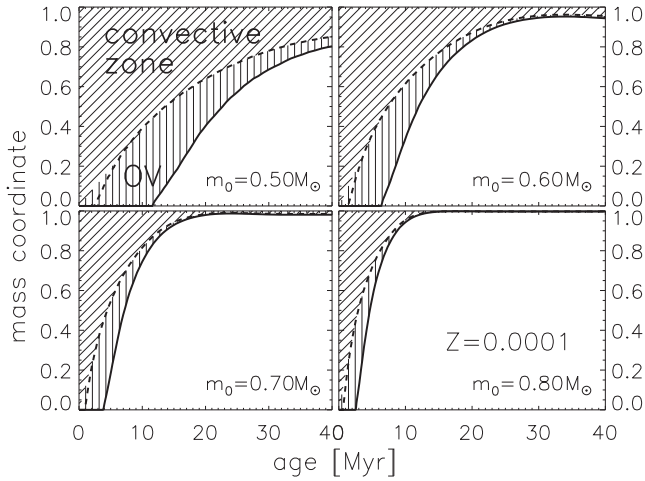


Figure 1. Kippenhahn diagrams for stars with initial metallicity $Z = 0.0001$ and different initial masses. The hatched area corresponds to the mass size of the convective zone, within it the area filled with vertical lines represents the contribution of the envelope overshooting. On the y-axis the mass coordinate of the convective zone starts from 0 (at the centre of the star) which means the star is fully convective, while 1 means the base of the convective zone is at the surface of the star and there is no convection. As the star evolves, the base of the convective zone retreats towards the surface of the star, and the overshoot layer becomes shallower and shallower.

and the temperature at the base of the convective zone (T_{bcz}) is high enough to burn Li so efficiently to begin surface Li depletion. As the central region heats up, the base of the convective envelope moves outwards to the stellar surface, while T_{bcz} decreases. When T_{bcz} drops below the lithium effective burning threshold, the photospheric Li depletion ceases, leading to an almost total or only partial Li depletion, depending on the stellar mass and on the OV efficiency.

Fig. 2 illustrates the effect of efficient OV as the *solo* mechanism applied to the PMS phase for $m_0 = 0.60, 0.70, 0.80 M_\odot$ stars. In standard model without overshoot (left-hand panel), $A(\text{Li})$ almost remains the same as the primordial value, unless the convective zone in the star itself goes deep enough to burn Li, as for $m_0 = 0.60 M_\odot$

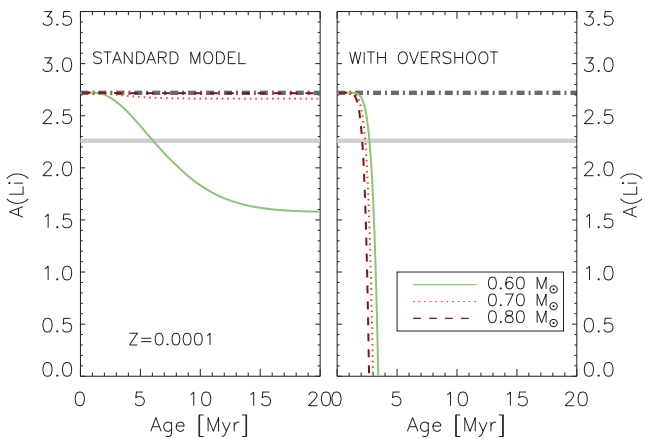


Figure 2. PMS Li evolution in the standard model (left-hand panel) and model with envelope overshoot (right-hand panel) for three stellar models under consideration. The two horizontal lines indicate primordial $A(\text{Li})$ (dark grey dot-dashed line) and Spite plateau $A(\text{Li})$ (light grey solid line), respectively. Overshooting is the *solo* parameter tested, no other mechanism is applied.

star in the left-hand panel. Conversely, assuming efficient envelope OV, Li abundance is found to be fully depleted over a few Myr in the same stellar models (right-hand panel). It follows that if overshoot were the only process at work, a very low Li abundance should be measured at the surface of those stars.

2.2 Late mass accretion

Recent observations of late PMS stars – those that have already abandoned the stellar birth line and are joining the ZAMS – reveal that most of them show H α -line emission, which likely originates in a residual accretion process (De Marchi et al. 2010; Spezzi et al. 2012). This suggests that the disc accretion may last much longer than previously believed: at least tens of Myr and even up to the early MS stars (De Marchi et al. 2011).

This late accretion (also known as residual accretion) is different from the one that maintains the star on the stellar birth line during the protostellar phase. The observed values indicate a median accretion rate of ($\sim 10^{-8} M_\odot \text{ yr}^{-1}$) for YSO T Tauri stars (TTs) with a steady decline as time proceeds (Espaillat et al. 2014). Most residual discs presented in Espaillat et al. (2014) and Gallardo, Del Valle & Ruiz (2012) are associated with accretion rates at the level of $\sim 10^{-10} - 10^{-8} M_\odot \text{ yr}^{-1}$.

Since the accreting material keeps the initial Li abundance, its contribution cannot be neglected, especially if the photospheric value has been already depleted by another process. In this framework, we modify our stellar evolution code PARSEC to account for the effect of such a residual accretion during the PMS phase.

We assume that, after the main accretion phase when the star leaves the stellar birth line by consuming its internal deuterium, a residual accretion keeps going on. Since our PMS models are initially evolved at constant mass from a contracting configuration without nuclear burning, we assume that the residual accretion begins when deuterium burning ends, that is when the photospheric deuterium abundance drops to 1/10 of its initial value, and indicate this time as τ_{acc} . We then apply the accretion to the models, starting at $t = \tau_{\text{acc}}$ and assume that it declines with time following a power law

$$\dot{M} = \dot{M}_0 \left(\frac{t}{\tau_{\text{acc}}} \right)^{-\eta} \quad (M_\odot \text{ yr}^{-1}), \quad (3)$$

where $\dot{M}_0$ is the initial rate and η is the parameter that specifies the rate of decline. In principle, the parameters $\dot{M}_0$, τ_{acc} and η could be treated as adjustable parameters to be constrained with the observed rates, after properly considering a detailed PMS evolution that includes the main accretion phase. However, since in this paper we perform an explorative analysis without a full description of the protostar phase and the initial evolution with large accretion rates, we assume the residual accretion begins when deuterium burning ends as defined before.

Table 1 shows the age at the beginning of the residual accretion (τ_{acc}) for different stellar masses. We note that the duration of the early large accretion, before τ_{acc} , is around $1-2 \times 10^5 \text{ yr}$, hence negligible compared to the PMS lifetime (several 10^7 yr). As to the initial accretion rate $\dot{M}_0$, we assume $\dot{M}_0 = 2 \times 10^{-8} M_\odot \text{ yr}^{-1}$, which is a reasonable value close to the observed median accretion rate for young TTs (Hartmann et al. 1998). The exponent η describes how fast the accretion rate declines with increasing age t . Here, we set $\eta = 0.9$ to recover most of the observed rates from Espaillat et al. (2014), Gallardo et al. (2012) and De Marchi et al. (2011).

Table 1. Relevant parameters for stars with different initial masses. τ_{acc} is the age at the beginning of the late accretion phase; m_* is stellar mass in main sequence; t_{end} is the age at the end of residual accretion; $\Delta n(\text{Li})_{\text{PMS}}$ and $\Delta n(\text{Li})_{\text{MS}}$ denote the decrement in the Li number density during pre-main sequence and main sequence, respectively. Note that the PMS Li depletion is stronger than the MS one.

$m_0 (M_\odot)$	0.57	0.58	0.59	0.60	0.62	0.65	0.70	0.75	0.80
$\tau_{\text{acc}} (10^5 \text{ yr})$	1.876	1.873	1.834	1.787	1.731	1.689	1.595	1.560	1.433
$m_* (M_\odot)$	0.616	0.625	0.633	0.641	0.658	0.685	0.729	0.775	0.821
$t_{\text{end}} (\text{Myr})$	585.55	518.96	460.83	399.76	309.84	211.46	110.98	61.14	35.46
$\Delta n(\text{Li})_{\text{PMS}} (10^{-10})$	3.16	2.81	2.60	2.50	2.33	2.11	1.75	1.17	1.21
$\Delta n(\text{Li})_{\text{MS}} (10^{-10})$	0.70	0.78	0.65	0.49	0.41	0.50	0.86	1.19	1.11

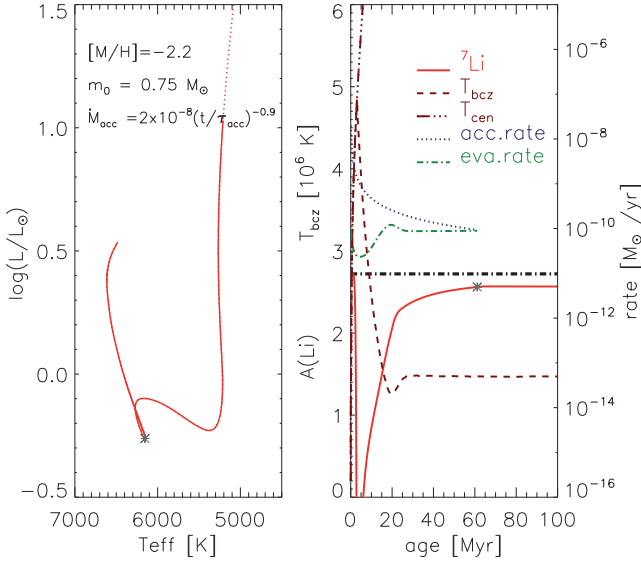


Figure 3. An example of the PMS lithium evolution for a star with $m_0 = 0.75 M_\odot$, $[M/H] = -2.2$. Left-hand panel: evolutionary track at constant mass in H-R diagram. The solid line starts at the end of the dotted-line-indicated deuterium burning phase. At this stage, the accretion is reduced to $2 \times 10^{-8} M_\odot \text{ yr}^{-1}$. The asterisk marks the end of the residual accretion. Right-hand panel: Li evolution starting from an initial abundance $A(\text{Li}) = 2.72 \text{ dex}$ (horizontal black dot-dashed line). The temperatures at the centre (T_{cen} , dark red dot-dashed line) and at the base of the convection zone (T_{bcz} , dark red dashed line) are also shown. The accretion (with rate drawn by the dark blue dotted line) is terminated by the EUV photoevaporation (with rate drawn by the dark green dot-dashed line) when the latter reaches the same value as the former.

We also assume that during this residual accretion phase the original material is accreted following equation (3), irrespective of the geometry of the disc.

The right-hand panel of Fig. 3 illustrates the effect of late accretion on the Li evolution, for a star with initial mass $m_0 = 0.75 M_\odot$. The accretion material falling on to the star contains lithium with the initial abundance and, even if accretion is very small after tens of million years, it restores the surface ^7Li towards the initial value.

2.3 EUV photoevaporation

Late accretion will last until the remaining gas reservoir is consumed or until some feedback mechanism from the star itself is able to clean the nearby disc. In this respect, extremely UV (EUV) radiation (13.6–100 eV, 10–121 nm) photons could be particular important in determining the end of the accretion process because they have energy high enough to heat the disc surface gas. The warm gas can escape from the gravitational potential of the central

star and flow away in a wind (Dullemond et al. 2007). Although no actual evaporative flow observation is available yet, there is a large consensus that this radiation could significantly reduce the disc mass and shorten the lifetime of the disc. The existence of pre-transitional and transitional discs with a gap or a hole in the disc are taken as possible evidence of this process (Dullemond et al. 2007; Espaillat et al. 2014). The magnitude of the disc mass loss caused by EUV evaporation is given by the following relation (Dullemond et al. 2007):

$$\dot{M}_{\text{EUV}} \sim 4 \times 10^{-10} \left(\frac{\Phi_{\text{EUV}}}{10^{41} \text{ s}^{-1}} \right)^{0.5} \left(\frac{M_*}{M_\odot} \right)^{0.5} (M_\odot \text{ yr}^{-1}), \quad (4)$$

where Φ_{EUV} is the EUV photon luminosity (photons s^{-1}) produced by the central star. This effect can be easily included in PARSEC under the assumption that the stars emit as a blackbody at the given effective temperature T_{eff} . In Fig. 4, we show the EUV luminosity evolution for stellar models with initial masses $m_0 = 0.55, 0.65, 0.75$ and $0.85 M_\odot$. The photon luminosity Φ_{EUV} varies with the effective temperature and the radius of the central star, therefore being importantly affected by the stellar evolutionary phase. Most of the EUV luminosities range from $\Phi_{\text{EUV}} \sim 10^{38}$ to $\sim 10^{42}$ photons s^{-1} . The residual accretion does not change these values much because the residual accretion rate is never high enough to significantly increase the mass (and hence the evolution) of the star.

The EUV photoevaporation rates obtained from equation (4) are shown in Fig. 5 (dot-dashed lines) for a grid of low-mass stars,

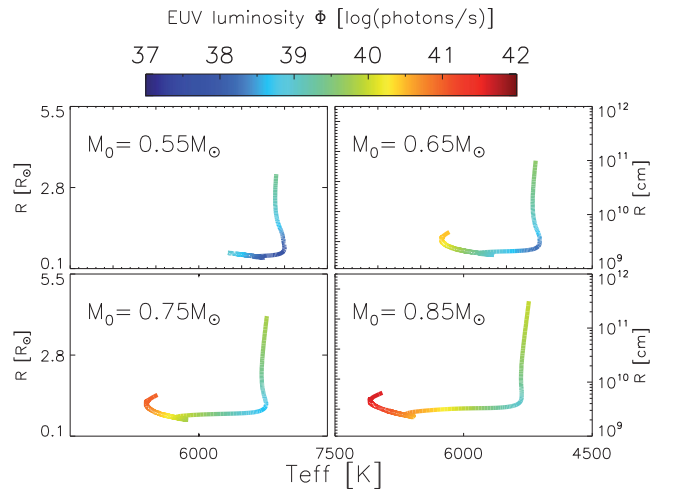


Figure 4. Evolution of the EUV photon luminosity from the central stars (colour-coded according to logarithmic scale shown on the top bar) as a function of the effective temperature and radius, assuming blackbody emission. Note the different units for the radius on the left-hand and right-hand Y-axes.

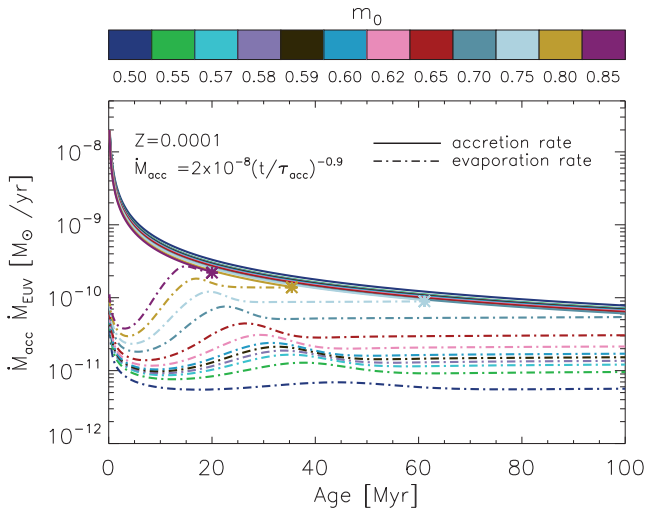


Figure 5. Evolution of the residual accretion rate during the first 100 Myr along the PMS evolution (solid lines). Different colours refer to different initial masses (m_0) as shown in the top colour-bar. The residual accretion begins with the rate $2 \times 10^{-8} \text{ M}_{\odot} \text{ yr}^{-1}$ and drops following equation (3). The EUV photoevaporation (dash-dotted line) rates, calculated with equation (4) are also plotted. From bottom to top are stars with initial mass $0.50\text{--}0.85 \text{ M}_{\odot}$ as labelled in the colour-bar. The asterisks mark the times when the accretion is terminated by the evaporation.

and they are compared with the corresponding accretion rates (solid lines) obtained from equation (3). The EUV photoevaporation rates increase with increasing initial mass. Furthermore, at any given initial mass they depend mainly on the evolution of the EUV luminosity because the total mass of the star is not significantly affected by the accretion rate. At relatively larger masses there is an minimum caused by the initial shrinking of the radius at almost constant effective temperature (Fig. 4). Thereafter, the star evolves at constant luminosity towards the ZAMS and, since the effective temperature increases, the fractional number of EUV photons increases rapidly, so does the evaporation rate. When the evaporation rate is larger than the accretion rate the accretion effectively stops. This point is marked with an asterisk in the figure and the terminal age is shown in Table 1. Beyond this point the star evolves at constant mass. Fig. 5 is meant to provide a schematic diagram of this mechanism for the first 100 Myr.

We would like to emphasize that the quenching of accretion is expected to be much more complex than our model description. For instance, we notice that only a fraction of the total EUV photons from the central star could reach the residual disc, if the latter maintains a small geometrical cross-section. On the other hand, stellar activity in young star (e.g. accretion shock, magnetic field driven chromosphere activities, etc.), which is not included in our model, could also be a source of EUV photons and contribute to evaporation. This effect could balance the geometrical loss of EUV photons though the relative contribution between the two sources is not clear. The end of the late accretion phase could also depend on the mass of the residual disc before the EUV evaporation mechanism becomes effective. For metal-rich star-forming regions, disc masses between 0.01 between $0.2 \text{ M}_{\odot}$ during the burst accretion phase have been estimated by Hartmann et al. (1998). While Bodenheimer (2011) estimate a disc masses from 0.5 down to $0.0001 \text{ M}_{\odot}$ in solar mass YSOs of the Taurus and Ophiuchus star-forming regions. The residual disc masses we are considering for the present work are even more uncertain than those of the main

disc. Here we adopt the working hypothesis that the disc continues to provide material to the central star until EUV evaporation stops it.

2.4 A combined PMS model

A model that combines all the three effects previously discussed is shown in Fig. 3. The left-hand panel shows the PMS evolutionary track of a star with mass $m_0 = 0.75 \text{ M}_{\odot}$ in the H-R diagram. The dotted line shows mass $m_0 = 0.75 \text{ M}_{\odot}$ in the H-R diagram. The dotted line shows the PMS evolution at constant mass up to deuterium exhaustion. The latter point represents the end of the stellar birth line, when the phase of large accretion terminates and the star begins the contraction at almost constant mass downward in the H-R diagram. The solid line begins when the deuterium abundance reduces to $1/10$ of its initial value. At this time, we apply an accretion with a rate starting at $2 \times 10^{-8} \text{ M}_{\odot} \text{ yr}^{-1}$ (dark blue line in the right-hand panel). This value is too low to prevent a rapid contraction and the central temperature T_{cen} (brown dashed line) rises until surface Li depletion begins (red line).

At this stage, the star is fully convective and the surface Li depletion proceeds faster than the restoring effect of the accretion because of the high central temperature. After about 5 Myr, the core becomes radiative and the temperature at the base of the convective region T_{bcz} (brown solid line) begins to decrease followed by a decline of the Li-burning rate. Then the photospheric Li abundance begins to increase again, especially when T_{bcz} falls below $4 \times 10^6 \text{ K}$. This is because on one side Li nuclear burning quenches off and, on the other, the dilution of the infalling Li becomes weaker as the mass contained in the convective envelope becomes smaller. The mass accreted in this phase is relatively little and only slightly affects the total mass of the star with negligible effects on its structure. In practice, the star evolves at constant mass.

The surface Li abundance steadily increases towards its primordial value, and will reach it if accretion does not terminate. This is due to the shrinking of the surface convective zone which becomes so small that even a residual accretion rate less than $10^{-10} \text{ M}_{\odot} \text{ yr}^{-1}$ is enough to engulf the thin surface convective layers. The interplay between the size of the convective regions and the evolution of the accretion rates are critical for reproducing a given pattern in the surface Li abundance as a function of the initial mass on the MS. At an age of about 60 Myr (see Table 1), the residual accretion is terminated by the EUV photoevaporation.

The H-R diagram of the full sample of stars considered in this paper is shown Fig. 6. The area between the two vertical dotted lines illustrates the temperature range of the Spite plateau. The high end of this range is also the temperature threshold of the observable lithium. For stars with higher T_{eff} (the more massive ones), lithium is almost fully ionized in the stellar photosphere. Likewise in Fig. 3, the upper Hayashi lines before the completion of deuterium burning are drawn with dotted lines whilst the solid lines correspond to the stellar evolutionary tracks after deuterium burning. The surface abundance of Li is colour-coded, starting from an initial abundance $A(\text{Li}) = 2.72$. After the end of the accretion (marked with the asterisks), stars evolve with constant masses along MS and Li abundance no longer increases.

3 MAIN-SEQUENCE LI EVOLUTION

During the MS phase, Li is depleted by burning at the base of the convective zone because, even a low nuclear reaction rate at

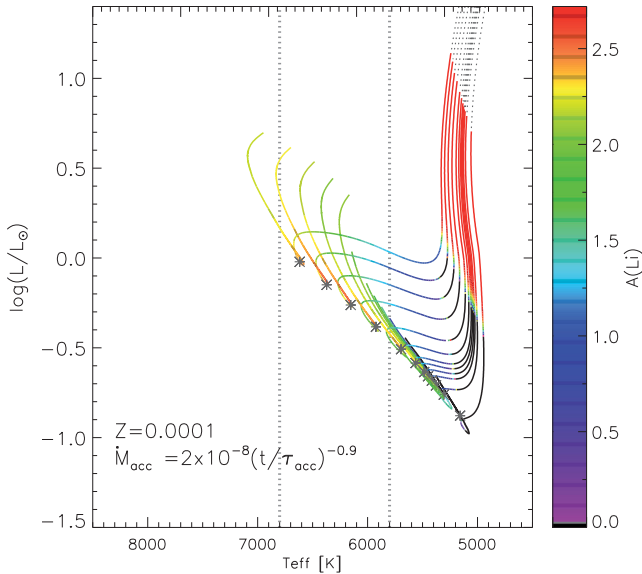


Figure 6. H-R diagram from PMS with overshoot, accretion, and photoevaporation, to the end of the main sequence. The tracks are for stars with initial mass 0.85, 0.80, 0.75, 0.70, 0.65, 0.62, 0.60, 0.59, 0.58, 0.57, 0.55 and 0.50 $M_{\odot}$ from top to bottom. Black dashed lines at the upper-right side are Hayashi line with constant initial masses, which is not the real case for the late accretion. Solid lines are the stellar evolutionary tracks after deuterium burning with an initial lithium value $A(\text{Li}) = 2.72$. Tracks are colour-coded according to their surface lithium abundance, as indicated in the right vertical bar, starting from $A(\text{Li}) = 2.72$ (in red) to $A(\text{Li}) = 0$ (violet). The black colour corresponds to $A(\text{Li}) < 0$. The area between the two vertical dotted lines is the effective temperature range of the Spite plateau.

$T_{\text{bcz}} \sim 2 \times 10^6$ K, could cause significant depletion in a time-scale of several Gyr. However, for masses larger than $m_0 = 0.60 M_{\odot}$, Li burning is insignificant.

Another effect that could modify the photospheric Li abundance in low-mass POP II stars is microscopic diffusion. This long-term stellar process cannot be observed directly in the star, and its efficiency needs to be calibrated. Microscopic diffusion shortens the MS lifetime and leads to a depletion of the surface elements. The solar model requires the inclusion of microscopic diffusion, otherwise even the age of the Sun could not be correctly produced. We calculated evolutionary models up to the end of the MS including pressure diffusion, temperature diffusion and concentration diffusion (Thoul, Bahcall & Loeb 1994). It has already been shown that microscopic diffusion without any correction for radiative levitation and surface convective turbulence cannot reproduce the observed photospheric abundances (Richard et al. 2002). In the case of Li, the surface convective zone of the stars at the plateau high-temperature end (the relatively more massive stars) is so thin that, if microscopic diffusion is at work and not balanced by the radiative levitation, a too strong depletion is produced. For example, according to the model of Salaris & Weiss (2001) we can see that even at $T_{\text{eff}} \sim 6200$ K the predicted Li depletion is too strong compared to the observed plateau.

This does not happen with the standard PARSEC code because gravitational settling is always inhibited in the outermost region of the envelope, $\Delta M = 0.5$ per cent of the stellar mass (Bressan et al. 2013). This choice of ΔM is made according to the suggestion by Chaboyer et al. (2001). They noticed that, while observations indicate that the relatively hot low-mass stars at the turn-off of the globular cluster NGC 6397 show the same surface [Fe/H] abun-

dance of the evolved RGB stars (Gratton et al. 2001), models with gravitational settling that well reproduce the solar data predict that they should show a surface [Fe/H] abundance at least 0.28 dex lower. They thus conclude that gravitational settling is inhibited in the outermost layers of such stars and show that, in order to reconcile the predicted with the observed [Fe/H] abundance, the size of this layer should be $\Delta M \sim 0.5$ –1 per cent $M_{\odot}$ for a star of $M \leq 1 M_{\odot}$. We checked that changing the standard PARSEC parameter in the range from $\Delta M = 0.1$ to $\Delta M = 1$ per cent does not appreciably affect our results.

4 RESULTS

We analyse the importance of overshoot and late accretion during the PMS phase of stellar evolution by computing a set of evolutionary tracks of initial stellar masses 0.85, 0.80, 0.75, 0.70, 0.65, 0.62, 0.60, 0.59, 0.58, 0.57, 0.55 and 0.50 $M_{\odot}$, assuming an efficient overshoot, as described in Section 2.1, and exploring different values for the parameters that describe the form of the residual accretion rate, $\dot{M}_0$ and η . We discuss below the results obtained assuming $\dot{M}_0 = 2 \times 10^{-8} M_{\odot} \text{ yr}^{-1}$ and $\eta = 0.9$ in equation (3), which produces accretion rates compatible with the observed values. The models are calculated from the PMS till the end of the MS.

We will first focus on the evolution of Li abundance during the PMS phase. The surface Li abundance evolution of our selected models during the first 100 Myr of PMS phase is plotted in the left-hand panel of Fig. 7. With our choice of the parameters, the initial Li evolution in these stars is regulated by OV. The convective OV is so efficient that the photospheric Li is significantly depleted after the first few Myr. The depletion caused by OV gets stronger at the lower masses. At later times, the core becomes radiative and the convective envelope starts receding.

At subsequent stages, Li evolution is governed by the competition between late accretion, dilution and nuclear burning at the bottom of the convective envelope. As soon as efficient Li burning ceases, accretion begins to restore the surface Li abundance. In the more massive stars, the convective zone shrinks more rapidly and Li restoring is faster. Conversely, in lower mass stars where Li depletion is more efficient because of the deeper convective zones, the Li restoring time is longer.

Without EUV evaporation, more massive stars rapidly recover the initial surface Li abundance while in lower mass stars this process is slower because nuclear reactions will continue to burn the accreted Li and the convective dilution is larger. The effects of EUV evaporation strongly modify this picture and is regulated by the stellar mass. More massive stars evolve more rapidly towards the ZAMS and their EUV luminosity is able to stop accretion at earlier times, even before the initial abundance is completely restored (see Fig. 7). On the other hand, in lower mass stars where Li depletion is more pronounced, the EUV evaporation rates are lower and accretion, though decreasing, is still able to drive the surface Li abundance towards the initial value, until it is inhibited by evaporation. At even lower masses (e.g. $m_0 = 0.50 M_{\odot}$), the recovery of the initial Li abundance is delayed or even inhibited by the deep convective zone which is still large when the model reaches the ZAMS. In this case, Li restoring is prevented both by an efficient dilution, and possibly by the nuclear burning at the base of such deep convective envelopes.

In summary, our analysis indicates that Li evolution during the PMS phase may be regulated by the competition between the destroying effect of OV at early times, and the restoring effect

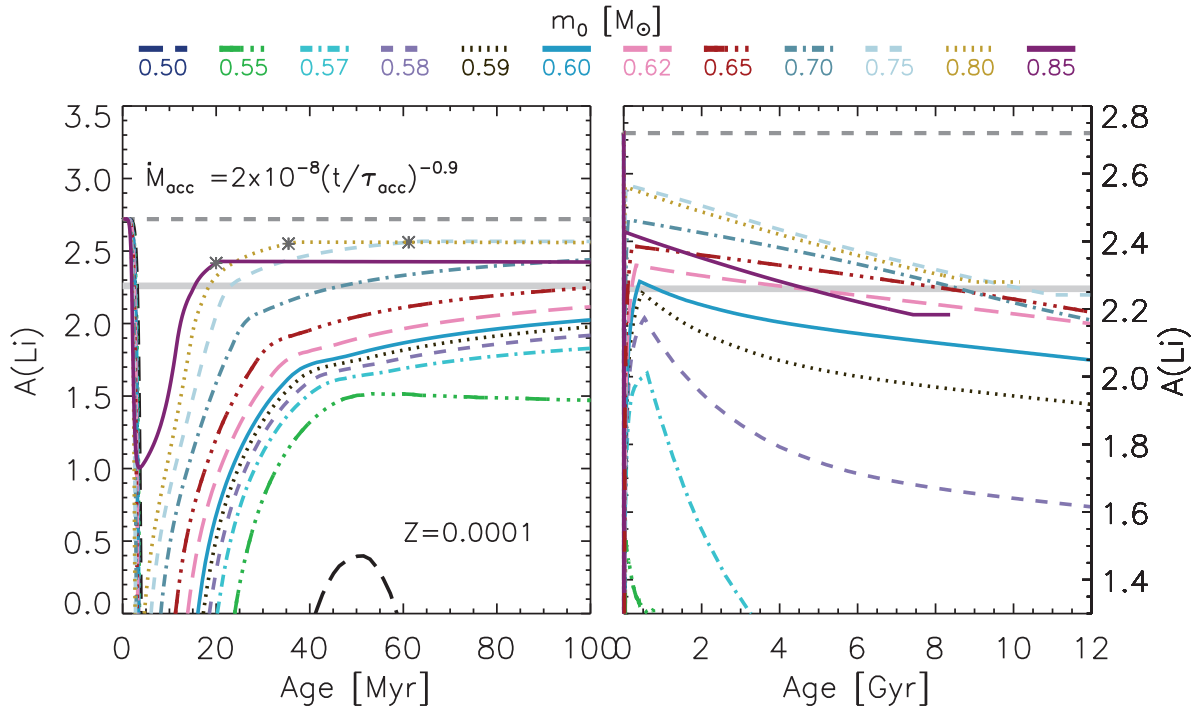


Figure 7. $A(\text{Li})$ as a function of the stellar age for stars with different initial masses. Different colours refer to different initial masses (m_0). The horizontal grey dashed line is the SBBN prediction ($A(\text{Li}) = 2.72$), while the horizontal grey solid one indicates $A(\text{Li})$ for the Spite plateau. Left-hand panel: Li abundance during the PMS phase during the first 100 Myr. The initial mass (m_0) of the models decreases from 0.85 to 0.50 $M_\odot$, from top to bottom in the rapid rising branch on the left of the diagram, as indicated in the legend. The asterisks mark the end of the accretion. Right-hand panel: Li evolution up to an age of 12 Gyr, zooming in around the region around the Spite plateau. Colour codes and the line styles are the same as in the left-hand panel. Note that stars with initial masses greater than 0.80 $M_\odot$, have a main-sequence life shorter than 10 Gyr, and they may have already evolved.

of residual accretion, which is in turn modulated by EUV photoevaporation quenching. This latter effect controls the end result, depending on the mass. At the higher stellar masses Li restoring is faster, but accretion quenching takes place earlier. At lower stellar masses, the time-scale for Li restoring is longer, but the accretion lasts longer because the EUV photoevaporation has a lower efficiency. The combination of these effects in stars with different initial masses tends to level out the Li abundance just after the PMS at a value that is already below the initial one.

Galactic halo POP II stars we observe today are about 10–12 Gyr old (Jofré & Weiss 2011). In order to reproduce the observed Li abundance of these stars, we have evolved our models up to the end of central hydrogen burning (till the turn-off phase). In this phase, we also account for microscopic diffusion which is known to be a long-term effect that can modify the photospheric element abundances. The evolution of the photospheric Li abundance during the MS is depicted in the right-hand panel of Fig. 7. The predicted photospheric Li abundances are compared with those of the Galactic halo POP II stars in Fig. 8. We plot the abundances at the end of the accretion phase (filled triangles) and those at older ages, 10 (filled upper circles) and 12 Gyr (filled lower circles). Stars with initial mass $m_0 \geq 0.85 M_\odot$ are excluded because their effective temperatures on the MS are warmer than the observed ones (see Fig. 6), and also because their MS lifetimes are shorter than the relevant age range. For our standard choice of parameters, stars with initial mass from $m_0 = 0.62$ to $0.80 M_\odot$, nicely populate the Spite plateau ($A(\text{Li}) \approx 2.26$).

Models with lower mass, $m_0 = 0.57$ – $0.60 M_\odot$, fall on the observed declining branch towards lower temperatures. We confirm that the first part of this branch is populated by low-mass MS stars.

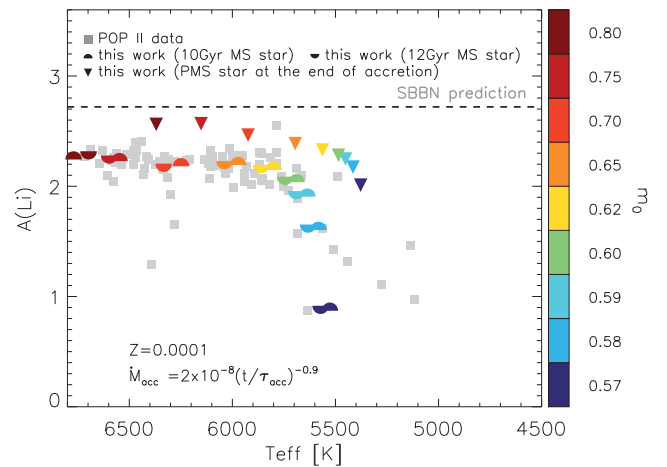


Figure 8. Our results in comparison with the lithium abundance measurements in POP II stars. The grey filled squares are POP II data from Molaro et al. (2012). Our predictions are shown for stars at the end of the late accretion phase (filled triangles), and on the main sequence at 10 (filled upper circle) and 12 Gyr (filled lower circle). Symbols are colour-coded according to the initial stellar mass, from the left to the right are 0.80–0.57 $M_\odot$ as the same as labelled. The black dashed line marks the primordial Li abundance according to the SBBN.

Indeed this is the signature of the strong Li depletion during the PMS phase (filled triangles), followed by further depletion during the MS evolution (filled upper and lower circles).

Both PMS Li depletion and MS diffusion contribute to make the total $A(\text{Li})$ decrease, with the former process playing the main

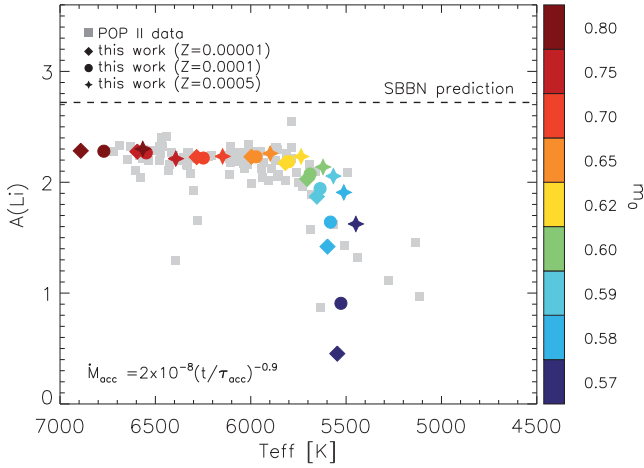


Figure 9. By applying the same parameter of envelope overshooting and accretion, we could reproduce the Spite plateau and the first Li decline branch for a wide range of metallicities (from $Z = 0.00001$ to $Z = 0.0005$). The compared POP II data are the same as in Fig. 8. The model results are all main-sequence stars at age 10 Gyr, with initial mass $0.80\text{--}0.57 M_{\odot}$ from the left to the right for each metallicity as shown in the colour-bar label.

role. Table 1 lists the Li number density decrease during PMS and MS, the former is much more than, or at least the same as, the latter, especially for the lower masses. The first part of Li declining branch is caused mainly by the PMS depletion.

The Spite plateau is populated by metal-poor MS stars with different metallicities. With the same OV and accretion parameters we can indeed reproduce the plateau, including the initial Li declining branch, over a wide range of metallicities, (from $Z = 0.00001$ to $Z = 0.0005$, that is from $[M/H] = -3.2$ to $[M/H] = -1.5$) as shown in Fig. 9.

5 DISCUSSION AND CONCLUSION

In the previous sections, we have shown how the effects of efficient envelope OV, residual accretion during the PMS phase, and microscopic diffusion during the MS may modify the photospheric Li abundance in low-metallicity stars. We have also considered the EUV photoevaporation process which, by terminating the accretion phase, could introduce a sort of self-regulating process.

In order to give a more exhaustive picture of the importance of the various processes, we computed more sets of models under different assumptions concerning the above physical mechanisms. We synthesize the results of these additional models in Fig. 10 where we compare the predicted Li abundances with the observed ones (see Fig. 8).

(i) *Standard models.* We begin with two sets of models where none of the PMS effects have been considered. The first set (orange dashed line) is the standard model without microscopic diffusion while, in the second set of model (green dot-dashed line) microscopic diffusion is introduced as it is in the standard PARSEC model. The model without diffusion is clearly at variance with the observed data. On the other hand, when diffusion is included the Li abundance reaches $A(\text{Li}) = 2.45$ dex, producing a plateau somewhat higher than the observed one, by a factor of 1.6. The pure diffusion model of Salaris & Weiss (2001) is also plotted in this figure (extracted from fig. 1 of Salaris & Weiss (2001) with $[\text{Fe}/\text{H}] = -2.6$ and age 12 Gyr). Contrary to them, our diffusive model does not produce the strong

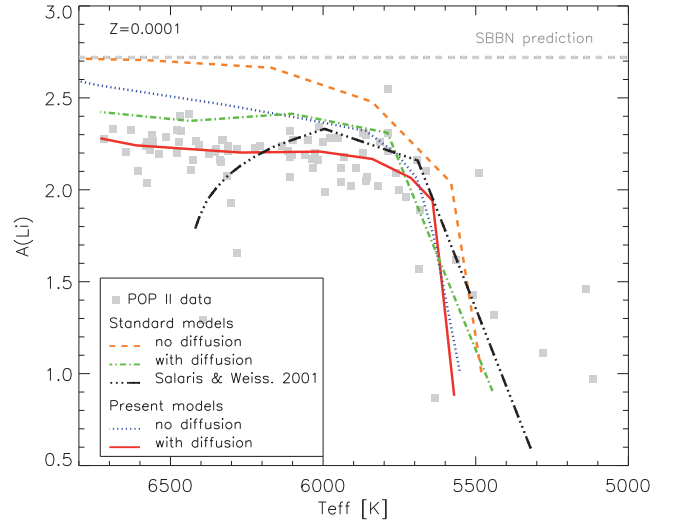


Figure 10. Comparison of the standard models and the present models for metallicity $Z = 0.0001$ together with the POP II data and Salaris & Weiss (2001) pure diffusion model. Different sets of models are classified in the legend. The black dash dot-dot line is extracted from fig. 1 of Salaris & Weiss (2001) with $[\text{Fe}/\text{H}] = -2.6$ and age 12 Gyr.

downturn at high effective temperatures. This is due to the inhibition of diffusion in the outer layers of the star which, as discussed in Section 3, is required to prevent a too strong sedimentation of helium and heavy elements at the stellar surface (Richard et al. 2002; Bressan et al. 2012).

While the plateau could eventually be reproduced by increasing the efficiency of diffusion strongly, we note however that, with the adopted diffusion parameters the solar model is very well reproduced (Bressan et al. 2012).

(ii) *Present models.* The present models, that include PMS effects and diffusion (red solid lines in Fig. 10), provide a good fit to the data. They reproduce both the observed plateau and the declining branch at low temperatures. As a test, we recomputed the same models but switching off microscopic diffusion during the MS phase (blue dotted line). These new models are identical in the PMS evolution and incorporate the Li nuclear burning during MS. The latter is mainly responsible for the behaviour of Li depletion at low effective temperatures. They do not reach the plateau value, terminating somewhat at higher Li abundances, suggesting that some additional depletion is required.

In our model, we do not consider stellar rotation. The additional effects of rotation on Li depletion have been discussed in many works, but mainly in the context of solar-type stars. During the PMS, rotation may affect the internal structure of the stars both because of the presence of the centrifugal force and the effects of rotational mixing at the base of the convective envelope. Eggenberger et al. (2012) conclude that by including the full treatment of rotation in PMS, 0.1 dex more Li depletion is obtained. During the evolution on the MS, rotational mixing tends to oppose to gravitational settling, and this solo effect should decrease the surface elements depletion with respect to non-rotating models. However in the case of Li, rotational mixing may be so efficient that Li is dragged to encounter temperatures which could efficiently destroy it by nuclear burning. This is the case of the slow rotator model of $M = 1.0 M_{\odot}$ of Eggenberger, Maeder & Meynet (2010). García López, Rebolo & Matín (1994) find that because of the rotation-caused structural

difference, stars with faster rotation in open clusters Pleiades and α Persei have higher Li abundance than the slower rotators. This result is confirmed by Pleiades observation of Jones, Fischer & Shetrone (1997). In general, it is found that rotation leads to a large spread in the MS Li abundance of solar-type stars at varying mass (Martin & Claret 1996; Mendes, D’Antona & Mazzitelli 1999), which could be difficult to reconcile with that observed in the Spite plateau of metal-poor stars.

It is important to stress that, as long as the stars are accreting at significant rates, deuterium burning keep their central temperature around 10^6 K. During this phase, the stars evolve along the birth line increasing their masses (e.g. reaching $1 M_{\odot}$ in 0.1 Myr for an accretion rate of $10^{-5} M_{\odot} \text{ yr}^{-1}$ as indicated in Table 1), while preserving the initial Li abundance (Stahler 1983). Once the accretion rates fall below several $10^{-8} M_{\odot} \text{ yr}^{-1}$ the stars abandon the stellar birth line and evolve at almost constant mass. It is from this point that we consider the additional effects of envelope OV, which reduces the Li abundance, and of accretion, which tends to restore the original abundance. We have also shown that EUV evaporation of material falling into the stars can act as a natural self-regulating mechanism.

In the relatively more massive stars, where the original Li abundance could be more quickly restored, EUV evaporation is more efficient and quenches out accretion earlier than in lower mass stars, which thus have more time to accrete the pristine gas. These combined effects tend to level out the Li abundance in stars of different masses. At even lower masses, Li restoring is inhibited by nuclear burning in the deep convective envelopes. After the PMS, Li is slowly depleted by microscopic diffusion and nuclear burning which act during MS phase. For stars with $m_0 \geq 0.62 M_{\odot}$ diffusion drives the main Li depletion whilst, for $m_0 \leq 0.60 M_{\odot}$ stars, Li is significant burned at the base of the convective envelope.

In this scenario, the observed Spite plateau and its falling branch at low temperatures are well reproduced by our models, within a reasonable parametrization of the late accretion mechanism. Very efficient overshoot is only applied during the early PMS to favour strong Li depletion at the beginning of the evolution. We also find that, even for a maximum envelope overshoot $\Lambda_e = 0.7 H_p$, which is the value calibrated from the RGB bump (Alongi et al. 1991), the plateau and its falling branch can be finally well reproduced for the same accretion rate. This indicates that a wide range of envelope overshoot could be applied to the PMS phase. Even for $\Lambda_e = 0.7 H_p$ the PMS Li depletion is very strong; without accretion a very low Li abundance would be observed, e.g. for $m_0 = 0.70 M_{\odot}$ $A(\text{Li})$ drops to 1 dex before Li restoring. It would be important to directly observe whether the very early Li-depletion phase exists or not, because it could confirm that efficient overshoot is at work. Indeed, at higher metallicities there is evidence of a significant drop of the Li abundance during the PMS phase (Somers & Pinsonneault 2014). On the contrary, the lack of nearby metal-poor star-forming regions prevents any observational test at low metallicity. This opportunity will be perhaps offered by the next generation telescopes [European Extremely Large Telescope (E-ELT), Thirty Meter Telescope (TMT) and Giant Magellan Telescope (GMT)] which could observe the details of low-mass metal-poor PMS stars in star-forming regions like those of the Sagittarius dwarf irregular galaxy (metallicity at $Z \sim 0.0004$; Momany et al. 2002).

For stars with initial mass $m_0 = 0.62\text{--}0.80 M_{\odot}$, ages in the range 10–12 Gyr, and a wide range of metallicities ($Z = 0.00001$ to $Z = 0.0005$), our final model is able to predict a present day abundance consistent with the Spite plateau, although starting from an initial value of $A(\text{Li}) = 2.72$ as inferred from the baryon-to-photon

ratio suggested by the CMB and by the deuterium measurements. Thus it offers a likely mechanism to solve the long-standing *lithium problem*. It also reproduces the observed lithium drop at the low temperatures. The proposed solution relies on stellar physics and evolution, assuming the validity of the current SBBN theory. Though this model is quite schematic and rough, it clearly suggests that the Spite plateau could be the net result of mechanisms more complex than those considered so far. In particular, a key role may be attributed to the interaction between the stars and their environments.

If the lithium abundance, we presently observe in POP II stars has been restored in the stellar atmospheres by a residual tail of accretion after a phase of strong depletion, then other puzzling observations could be explained, namely:

(i) *Li at very low metallicities*. Sbordone et al. (2010) found in their observations that for the extreme metal-poor stars, Li abundance drops when $[\text{Fe}/\text{H}] < -3$. Caffau et al. (2011) and Frebel et al. (2005) observed dwarfs with $[\text{Fe}/\text{H}]$ at the lowest levels without any detectable Li. Hansen, Hansen & Christlieb (2014) detected Li in an unevolved star with $[\text{Fe}/\text{H}] = -4.8$ at a level of $A(\text{Li}) = 1.77$ and another recently discovered extremely metal-poor dwarf SDSS J1742+2531 with $[\text{Fe}/\text{H}] = -4.8$ shows $A(\text{Li}) < 1.8$ (Bonifacio et al. 2015). The low Li abundance in these stars could be explained by a failed or weaker late accretion. For the most metal-poor stars, late accretion during the PMS phase might be inhibited and Li could not be restored. As we have already shown in Fig. 2 (right-hand panel), if no accretion takes place after the initial Li is depleted by the convective overshoot, Li in these stars would correspond to a very low, or even undetectable, abundance. Thus, the model presented here provides a possible key to account for this phenomenon, and we will discuss it in detail in a following paper.

(ii) *Spectroscopic binaries and outliers*. Our model might also provide an explanation to the Li discrepancy observed in metal-poor binary stars. The POP II spectroscopic binaries CS 22876-032 (González Hernández et al. 2008) and G 166-45 (Aoki, Ito & Tajitsu 2012) are composed by dwarfs with effective temperatures characteristic of the Spite plateau, but with the primary stars showing slightly higher Li abundances. It could simply be an effect of the competition between the components in the late PMS accretion, with the primary star being the favourite one.

The failure or increase of the accretion process could also provide an explanation for the few POP II ^7Li depleted stars or the few POP II stars that exhibit a ^7Li abundances at the SBBN prediction level, respectively. Among the latter there is BD +23 3912 with $A(\text{Li}) = 2.60$ which stands out from the others in Bonifacio & Molaro (1997).

(iii) *Lithium behaviour in POP I stars*. Li abundance at the Solar system formation time, as obtained from meteorites, is $A(\text{Li})=3.34$ (Anders & Grevesse 1989). The hot F stars of young open clusters never reach this meteoritic value (Ford et al. 2001) despite they should have an initial Li value even higher than the meteoritic one due to the Galactic Li increase with time. This Li drop cannot be explained by MS Li evolution because the stars are very young. A PMS Li modification could be the possible mechanism for the depletion. This will be discussed in detail in a following work dedicated to the metal-rich stars.

(iv) *Low lithium abundance in the planet host stars*. Israelian et al. (2009) and Gonzalez (2014) present the observational evidence that the solar analogues with planet(s) show lower lithium abundance than those without detected planets. To explain this effect Bouvier (2008) and Eggenberger et al. (2010) advance the hypothesis that giant exoplanets prefer stars with slow rotation rates on

the ZAMS due to their longer disc lifetime during the PMS evolution. Then their ZAMS models with rotational mixing, which we already discussed in Section 3, would indicate an enhanced lithium depletion, consistent with observations. Again it is worth noting that a strong PMS depletion at decreasing mass is predicted by solar metallicity models with rotation (Mendes et al. 1999), and it is not yet clear how this could be reconciled with the observations of Israelian et al. (2009) and Gonzalez (2014). In our model, this behaviour could be explained if somehow the planet(s), located within the accreting disc, interfere with the late accretion process. The PMS accretion could end before being terminated by the EUV evaporation, inhibiting efficient lithium restoring. Thus these stars will show lower Li abundance than their analogues without detected planets. This effect could also explain the larger Li abundance scatter observed in metal-rich stars, which are more likely to host planets (Ida & Lin 2004), with respect to what observed in metal-poor stars.

ACKNOWLEDGEMENTS

This research has made use of NASA's Astrophysics Data System. X. Fu. thanks Zhi-yu Zhang for helpful discussions. A. Bressan acknowledges PRIN INAF 2014 'Star formation and evolution in galactic nuclei'. P. Molaro acknowledges the international team #272 lead by C. M. Coppola 'EUROPA-Early Universe: Research On Plasma Astrochemistry' at ISSI (International Space Science Institute) in Bern, and the discussions with Giacomo Beccari, Lorenzo Monaco and Piercarlo Bonifacio. P. Marigo acknowledges support from the University of Padova (*Progetto di Ateneo 2012*, ID: CPDA125588/12).

REFERENCES

- Adams T. F., 1976, *A&A*, 50, 461A
 Alongi M., Bertelli G., Bressan A., Chiosi C., 1991, *A&A*, 224, 95
 Anders E., Grevesse N., 1989, *GeCoA*, 53, 197
 Aoki W., Ito H., Tajitsu A., 2012, *ApJ*, 751, L6
 Bodenheimer P. H., 2011, *Principle of Star Formation*. Springer, Dordrecht, p. 146
 Bonifacio F., Molaro P., 1997, *MNRAS*, 285, 847
 Bonifacio P. et al., 2015, *A&A*, 579, 28
 Bouvier J., 2008, *A&A*, 489, L53
 Bressan A., Marigo P., Girardi L., Salasnich B., Dal Cero C., Rubele S., Nanni A., 2012, *MNRAS*, 427, 127
 Bressan A. et al., 2013, *EPJWC*, 43, 3001
 Bressan A., Girardi L., Marigo P., Rosenfield P., Tang J., 2015, *ASSP*, Vol. 39, Springer International Publishing, Switzerland, p. 25
 Caffau E. et al., 2011, *Nature*, 477, 67
 Chaboyer B., Fenton W., Nelan J., Patnaude D. J., Simon F. E., 2001, *ApJ*, 562, 521
 Charbonnel C., Primas F., 2005, *A&A*, 442, 961
 Christensen-Dalsgaard J., Monteiro M. J. P. F. G., Rempel M., Thompson M. J., 2011, *MNRAS*, 414, 1158
 Coc A., Goriely S., Xu Y., Saimpert M., Vangioni E., 2012, *ApJ*, 744, 158
 Coc A., Uzan J.-P., Vangioni E., 2014, *AIP Conf. Ser. Vol. 1594*, Can mirror matter solve the cosmological lithium problem? *Am. Inst. Phys.*, New York, p. 12
 Cyburt R. H. et al., 2010, *ApJS*, 189, 240
 D'Antona F., Mazzitelli I., 1984, *A&A*, 138, 431
 D'Antona F., Mazzitelli I., 1994, *ApJS*, 90, 467
 De Marchi G., Panagia N., Romaniello M., 2010, *ApJ*, 715, 1
 De Marchi G., Panagia N., Romaniello M., Sabbi E., Sirianni M., Prada Moroni P. G., Degl'Innocenti S., 2011, *ApJ*, 740, 11
 Dullemond C. P., Hollenbach D., Kamp I., DAlessio P., 2007, in Reipurth B., Jewitt D., Keil K., eds, *Protostars and Planets V*. Univ. Arizona Press, Tucson, AZ, p. 555
 Eggenberger P., Maeder A., Meynet G., 2010, *A&A*, 519, L2
 Eggenberger P., Haemmerlé L., Meynet G., Maeder A., 2012, *A&A*, 539, 70
 Espaillat C. et al., 2014, in *Protostars and Planets VI*. Univ. Arizona Press, Tucson, AZ, p. 497
 Fields B., 2011, *Ann. Rev. Nucl. Part. Sci.*, 61, 47
 Fields B. D., Molaro P., Sarkar S., 2014, preprint ([arXiv:1412.1408](https://arxiv.org/abs/1412.1408))
 Ford A., Jeffries R. D., James D. J., Barnes J. R., 2001, *A&A*, 369, 871
 Frebel A. et al., 2005, *Nature*, 434, 871
 Gallardo J., Del Valle L., Ruiz M. T., 2012, *AJ*, 143, 18
 García López R. J., Rebolo R., Matín E. L., 1994, *A&A*, 282, 518
 Gonzáles Hernández J. I. et al., 2008, *A&A*, 480, 233
 Gonzalez G., 2014, *MNRAS*, 441, 1201
 Gratton R. G. et al., 2001, *A&A*, 369, 87
 Hansen T., Hansen C., Christlieb N., 2014, *ApJ*, 787, 162
 Hartmann L., Calvet N., Gullbring E., D'Alessio P., 1998, *ApJ*, 495, 385
 Herwig F., 2000, *A&A*, 360, 952
 Ida S., Lin D. N. C., 2004, *ApJ*, 616, 567
 Israelian G. et al., 2009, *Nature*, 462, 189
 Jofré P., Weiss A., 2011, *A&A*, 533, A59
 Jones B. F., Fischer D., Shetrone M., 1997, *AJ*, 114, 352
 Kajino T., Kusakabe M., Baha Balantekin A., Pehlivan Y., Matthews G. J., Cheoun M.-K., 2012, in Lattanzio J., Karakas A., Lugaro M., Dracoulis G., eds, *Proc. XII International Symposium on Nuclei in the Cosmos (NIC XII)*. Cairns, p. 70
 Komatsu E. et al., 2011, *ApJS*, 192, 18
 Korn A., Grundahl F., Richard O., Barklem P. S., Mashonkina L., Collet R., Piskunov N., Gustafsson B., 2006, *Nature*, 442, 657
 Kusakabe M., Kawasaki M., 2015, *MNRAS*, 446, 1597
 Lind K., Primas F., Charbonnel C., Grundahl F., Asplund M., 2009, *A&A*, 503, 545
 Marigo P., Girardi L., 2007, *A&A*, 469, 239
 Martin E. L., Claret A., 1996, *A&A*, 306, 408
 Mendes L. T. S., D'Antona F., Mazzitelli I., 1999, *A&A*, 341, 174
 Miglio A., Montalbán J., 2005, *A&A*, 441, 615
 Molaro P., 2008, in Knapen J. H., Mahoney T. J., Vazdekis A., eds, *ASP Conf. Ser. Vol. 390, Pathways Through an Eclectic Universe*. Astron. Soc. Pac., San Francisco, p. 472
 Molaro P., Bressan A., Barbieri M., Marigo P., Zaggia S., 2012, *Mem. Soc. Astron. Ital. Suppl.*, 22, 233
 Momany Y., Heid E. V., Saviane I., Rizzi L., 2002, *A&A*, 384, 393
 Olive K. A., Petitjean P., Vangioni E., Silk J., 2012, *MNRAS*, 426, 1427
 Piau L., Turck-Chièze S., 2002, *ApJ*, 566, 419
 Planck Collaboration XVI 2014, *A&A*, 571, 16
 Richard O., Michaud G., Richer J., Turcotte S., Turck-Chièze S., Vandenberg D. A., 2002, *ApJ*, 568, 979
 Richard O., Michaud G., Richer J., 2005, *A&A*, 430, 25
 Salaris M., Weiss A., 2001, *A&A*, 369, 955
 Sbordone L. et al., 2010, *A&A*, 522, A26
 Soderblom D., Fedele S. B., Jones B. F., 1993, *ApJ*, 106, 1080
 Somers G., Pinsonneault M. H., 2014, *ApJ*, 790, 72S
 Spezzi L., De Marchi G., Panagia N., Sicilia-Aguilar A., Ercolano B., 2012, *MNRAS*, 421, 78
 Spite F., Spite M., 1982, *A&A*, 115, 357
 Stahler S. W., 1983, *ApJ*, 274, 822
 Swenson F. J., Faulkner J., Rogers F. J., Iglesias C. A., 1994, *ApJ*, 425, 286
 Tang J., Bressan A., Rosenfield P., Slemmer A., Marigo P., Girardi L., Bianchi L., 2014, *MNRAS*, 445, 4287
 Thoul A., Bahcall J., Loeb A., 1994, *ApJ*, 421, 828
 Tognelli E., Degl'Innocenti S., Prada Moroni P. G., 2012, *A&A*, 548, A41
 Ventura P., Zeppieri A., Mazzitelli I., D'Antona F., 1998, *A&A*, 331, 1011

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