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AT 2017be – a new member of the class of intermediate-luminosity red transients

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ABSTRACT

We report the results of our spectrophotometric monitoring campaign for AT 2017be in NGC 2537. Its light curve reveals a fast rise to an optical maximum, followed by a plateau lasting about 30 d, and finally a fast decline. Its absolute peak magnitude ($M_r \simeq -12$ mag) is fainter than that of core-collapse supernovae, and is consistent with those of supernova impostors and other intermediate-luminosity optical transients. The quasi-bolometric light-curve peaks at $\sim 2 \times 10^{40}$ erg s⁻¹, and the late-time photometry allows us to constrain an ejected ⁵⁶Ni mass of $\sim 8 \times 10^{-4} M_{\odot}$. The spectra of AT 2017 be show minor evolution over the observational period, a relatively blue continuum showing at early phases, which becomes redder with time. A prominent H α emission line always dominates over other Balmer lines. Weak Fe II features, Ca II H&K, and the Ca II NIR triplet are also visible, while P-Cygni absorption troughs are found in a high-resolution spectrum. In addition, the [Ca II] $\lambda\lambda 7291, 7324$ doublet is visible in all spectra. This feature is typical of intermediate-luminosity red transients (ILRTs), similar to SN 2008S. The relatively shallow archival *Spitzer* data are not particularly constraining. On the other hand, a non-detection in deeper near-infrared *HST* images disfavors a massive Luminous Blue Variable eruption as the origin for AT 2017be. As has been suggested for other ILRTs, we propose that AT 2017be is a candidate for a weak electron-capture supernova explosion of a superasymptotic giant branch star, still embedded in a thick dusty envelope.

Key words: stars: AGB and post-AGB – stars: mass-loss – supernovae: general – galaxies: individual: NGC 2537 – supernovae: individual: AT 2017be.

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1 INTRODUCTION

In recent years, the proliferation of wide-field sky surveys has led to the discovery of a large variety of transient events of varying luminosities. In particular, an increasing number of events show luminosities intermediate between those of core-collapse (CC) supernovae (SNe) and classical novae (see e.g. Kasliwal 2012, for a short review on the subject). These events are generally called intermediate-luminosity optical transients (ILOTs), and are challenging our observational and theoretical understanding of standard scenarios of stellar evolution.

A sub-set of ILOTs are designated ‘SN impostors’ (Van Dyk et al. 2000), and mimic the observables of real SN explosions, although their progenitor stars survive the outburst. In many cases their peculiarity is not immediately recognized and in fact, several impostors have been labelled in the past with supernova designation. SN impostors have absolute magnitudes fainter than -15 mag, and their spectra are dominated by prominent hydrogen lines in emission, which usually lack high-velocity components in their profile (Smith et al. 2011; Tartaglia et al. 2016). Various scenarios have been proposed to explain SN impostors, including outbursts triggered by binary interaction, or violent mass-loss events during the late evolution of single massive stars. Stars producing SN impostors range from the most massive luminous blue variables (LBVs; e.g. Wolf 1989; Humphreys & Davidson 1994; Smith, Vink & de Koter 2004) to potentially lower mass superasymptotic giant branch (S-AGB) stars (see e.g. Miyaji et al. 1980; Nomoto 1984; Miyaji & Nomoto 1987; Hashimoto, Iwamoto & Nomoto 1993; Kitaura, Janka & Hillebrandt 2006; Poelarends et al. 2008).

LBVs form a class of massive stars with high luminosity, blue colours, large variability, enormous mass-loss rates, and H-rich spectra. The LBV phase is believed to be a short-duration stage in the evolution of massive stars. Humphreys & Davidson (1994) proposed that the transition of LBVs from quiescence to the eruptive phase is characterized by an increase in the optical luminosity accompanied by a modest colour evolution (towards redder colours). The best-known examples of LBV eruptions are the ‘giant eruption’ of the Galactic LBV η Carinae in the mid-19th Century (Humphreys & Davidson 1994) and P Cygni in the 17th Century (de Groot 1969, 1988). In addition to these eruptions in the Milky Way, similar non-terminal outbursts are also occasionally discovered in extra-Galactic environments (Smith et al. 2011), where they are the most frequently observed SN impostors.

An alternative scenario to explain some putative SN impostors is the explosion of an S-AGB star as an electron capture supernova (ECSN) (see e.g. Prieto et al. 2008, 2009; Thompson et al. 2009). These stars are usually surrounded by a thin helium shell and an extended hydrogen envelope (Nomoto 1987), and may experience major mass-loss accompanied by significant outbursts. It has been proposed that when a star’s initial mass is in the range $\sim 8\text{--}12 M_{\odot}$, it can die as an ECSN (see e.g. Miyaji et al. 1980; Nomoto 1984, 1987; Miyaji & Nomoto 1987; Hashimoto et al. 1993; Kitaura et al. 2006; Poelarends et al. 2008; Wanajo, Janka & Müller 2011). The Ne–O cores of these H-exhausted stars are very close to the boundary for an SN explosion. Electron-capture reactions by ^{24}Mg and ^{20}Ne nuclei reduce electron pressure, triggering the core to collapse before Ne burning ignites, and giving rise to an ECSN. We note however, that the precise mass range for ECSN progenitors is controversial and is likely limited to $\mu M \sim 1 M_{\odot}$ or less (Siess 2007; Doherty et al. 2015, 2017).

Recently, Kankare et al. (2012) proposed that CSM interaction following an ECSN can explain the Type IIn SN 2009kn (peak

$M_B \approx -18$ mag), as did Smith (2013) and Mauerhan et al. (2013) for the Type IIn SN 2011ht. However, ECSNe are expected to be fairly faint events (e.g. Tominaga, Blinnikov & Nomoto 2013). Therefore, candidates are expected to have modest luminosities and type IIn SN like spectra (e.g. Schlegel 1990) if a mass-loss event happened a short time before the SN explosion. The faint SN 2008S is an appealing ECSN candidate (see e.g. Prieto et al. 2008; Welch, Clayton & Sugerman 2008; Wesson et al. 2008; Yee, Eastman & Prieto 2008; Botticella et al. 2009; Smith et al. 2009; Wesson et al. 2010; Kochanek 2011; Szczygieł et al. 2012; Adams et al. 2016; Sugerman et al. 2016). In fact, a possible S-AGB progenitor has been proposed by several authors (Botticella et al. 2009; Szczygieł et al. 2012; Adams et al. 2016), based on pre-outburst archival Spitzer images which reveal a candidate in the correct mass range, and which was embedded in a dusty cocoon. SN 2002bu (Puckett & Gauthier 2002; Szczygieł, Kochanek & Dai 2012), NGC 300 OT2008-1 (e.g. Berger et al. 2009; Bond et al. 2009; Prieto et al. 2009; Kashi, Frankowski & Soker 2010; Patat et al. 2010; Humphreys et al. 2011), M85 OT2006-1 (see e.g. Kulkarni et al. 2007; Pastorello et al. 2007; Rau et al. 2007), PTF10fq (Kasliwal et al. 2011), NGC 5775 OT2012-1 (Barsukova et al. 2012; Berger et al. 2012) are also very likely ILOTs arising from ECSNe. These transients are collectively referred to as intermediate-luminosity red transients (ILRTs).

In this article, we present a study of a newly discovered ILRT, AT 2017be (also known as KAIT-17A or iPTF17be). AT 2017be was first discovered on 2017 January 06.508 (UT will be used hereafter) by the Lick Observatory Supernova Search (LOSS), at an unfiltered magnitude of 18.5 (Stephens, Zheng & Filippenko 2017). We note that Hosseinzadeh et al. (2017) first classified AT 2017be as an LBV, while Adams et al. (2018) recently presented some follow-up observations of AT 2017be, and also proposed an LBV eruption scenario for this event. This work presents a more extensive dataset, with higher cadence spectro-photometric coverage, and a monitoring window covering all phases of the transients evolution, from which we reach a somewhat different conclusion as to the nature of AT 2017be.

This paper is organized as follows: in Section 2, we characterize the host galaxy NGC 2537. In Section 3, we describe the photometric data reduction techniques and present the light curves of AT 2017be. In Section 4, the spectroscopic data are shown and analysed. In Section 5, we examine the pre-explosion frames, in particular *Hubble Space Telescope* (HST) archival images, and characterize the quiescent progenitor star a few years before the AT 2017be event. A discussion on the nature of AT 2017be and a short summary are presented in Section 6.

2 THE HOST GALAXY

AT 2017be was discovered in NGC 2537 at RA = $08^{\text{h}}13^{\text{m}}13^{\text{s}}.38$, Dec. = $+45^{\circ}59'29''.54$ [J2000] (see Fig. 1). The host galaxy is also known as Mrk 86, Arp 6, or the Bear’s Paw Galaxy. According to the standard morphological classification scheme (Hubble 1926; de Vaucouleurs 1959), NGC 2537 is a blue SB(rs) galaxy. Detailed information on the host galaxy taken from the NASA/IPAC Extragalactic database¹ (NED) is reported in Table 1.

The distance to NGC 2537 has been estimated with multiple methods, including using the brightest stars in the galaxy (Sharina, Karachentsev & Tikhonov 1999) and the Tully–Fisher technique

¹<http://nedwww.ipac.caltech.edu/>

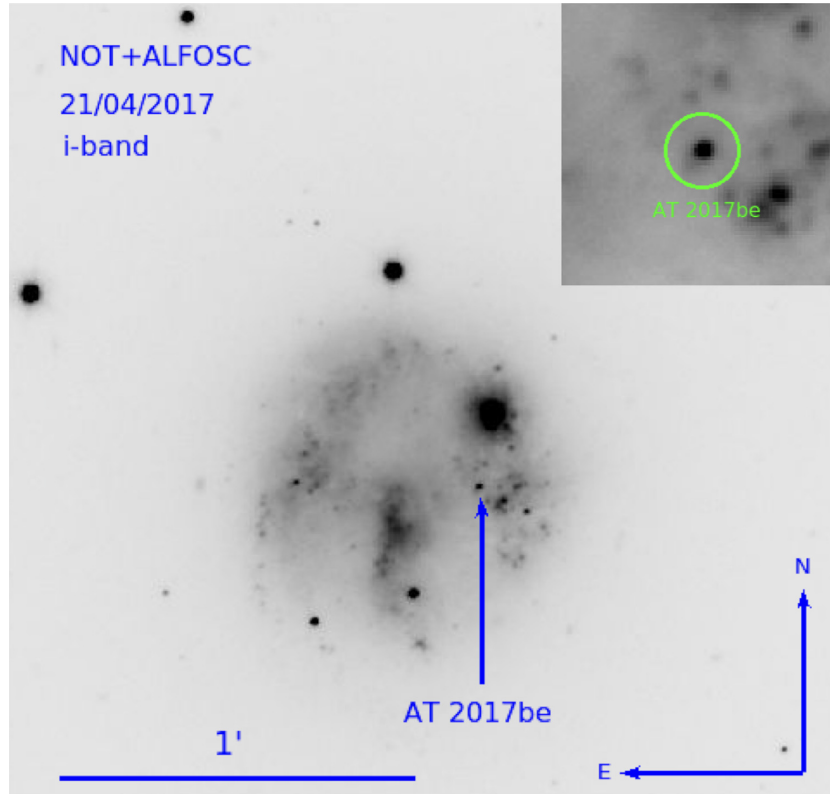


Figure 1. The field of AT 2017be. A zoom in on the location of AT 2017be is shown in the inset.

Table 1. Basic parameters of NGC 2537.

α (J2000)	$08^h 13^m 14^s.64$
δ (J2000)	$45^\circ 59' 23''.3$
Major axis (arcmin)	1.7
Minor axis (arcmin)	1.5
Morphological type	SB(rs)dm
Apparent magnitude (V)	12.32 mag
Redshift	0.001438 ± 0.000003
$v_{\text{heliocentric}}$	$431 \pm 1 \text{ km s}^{-1}$
Adopted distance	$7.82 \pm 0.54 \text{ Mpc}$
Galactic absorption	$A_B = 0.195 \text{ mag}$

(Bottinelli et al. 1984; Tully 1988). However, the reported estimates are largely discrepant. In particular, different estimates based on the Tully–Fisher method give a wide range of values (~ 9 – 20 Mpc). From the radial velocity (cz) corrected for local group infall into Virgo ($V_{\text{vir}} = 571 \pm 4 \text{ km s}^{-1}$) from LEDA² (Makarov et al. 2014) and adopting standard cosmological parameters ($H_0 = 73 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_M = 0.27$, $\Omega_\Lambda = 0.73$), we derive a luminosity distance $d_0 = 7.82 \pm 0.54 \text{ Mpc}$.³ The resulting distance modulus $\mu = 29.47 \pm 0.15 \text{ mag}$ will be used throughout this paper. We note that this kinematic distance is in good agreement with that inferred from the absolute magnitude of the brightest stars in NGC 2537 (Sharina et al. 1999).

We adopt a Milky Way reddening $E(B - V)_{\text{Gal}} = 0.048 \text{ mag}$ (from Schlafly & Finkbeiner 2011) towards AT 2017be. A narrow

Na I D absorption is visible in our spectra at the redshift of the host galaxy NGC 2537, but with low signal-to-noise (S/N) ratio. We increased the S/N by averaging the early spectra, and measured for the blended doublet an equivalent width (EW) of $\simeq 0.57 \pm 0.17 \text{ \AA}$. Following the relation in Turatto, Benetti & Cappellaro 2003 [i.e. $E(B - V) = 0.16 \times \text{EW}(\text{Na I D})$], and assuming $R_V = 3.1$ (Cardelli, Clayton & Mathis 1989), we obtain a host galaxy reddening value of $E(B - V)_{\text{host}} = 0.04 \pm 0.02 \text{ mag}$ and a total line-of-sight colour excess $E(B - V)_{\text{total}} = 0.09 \pm 0.03 \text{ mag}$.

Our highest resolution spectrum (+30.2 d from the r -band maximum, see Section 4) allows us to deblend the narrow Galactic and host galaxy Na I D absorption line components, which appear to have similar strengths. This supports our claim of a similar reddening contribution from the Milky Way and NGC 2537. A final check has been performed by computing the Balmer decrement (which is the ratio of the $H\alpha$ and $H\beta$ line intensities) in our spectra. We obtained decrements roughly in agreement with the classical value of $H\alpha/H\beta = 2.86$ for typical gas conditions (Osterbrock & Bochkarev 1989) adopting an $E(B - V)_{\text{total}} = 0.09 \text{ mag}$. This total reddening estimate will be used throughout the paper.

3 PHOTOMETRY

3.1 Observations and data reduction

As mentioned previously, the transient AT 2017be was first found by the Lick Observatory Supernova Search (LOSS; reported in TNS⁴). An early-phase photometric measurement was also obtained by the

²<http://leda.univ-lyon1.fr/>

³This value is obtained using Ned Wright’s Cosmological Calculator <http://www.astro.ucla.edu/~wright/CosmoCalc.html>; see Wright (2006).

⁴<https://wis-tns.weizmann.ac.il/object/2017be>

Intermediate Palomar Transient Factory (iPTF⁵). We included these data in our light curve. Our multiband monitoring campaign started on 2017 January 6, soon after the classification, when we triggered systematic follow-up observations of the target, with a more aggressive strategy during the first month, when the evolution of the transient was more rapid. Subsequently, we slightly relaxed the cadence of the campaign, which went on to span a period of approximately 1 yr. Detailed information on the observations including the photometric measurements are provided in Table A1.

The optical follow-up campaign of AT 2017be was performed using the following instruments: the 1-m telescope of the Las Cumbres Observatory (LCO⁶) at McDonald Observatory (Texas, USA) equipped with a Sinistro (labelled in this paper as f05) camera; the 2.56-m Nordic Optical Telescope (NOT⁷) equipped with ALFOSC and the 10.4-m Gran Telescopio Canarias (GTC⁸) with OSIRIS, both located at Roque de los Muchachos Observatory (La Palma, Canary Islands, Spain); and the 1.82-m Copernico Telescope of the INAF – Padova Observatory at Mt. Ekar (Asiago, Italy)⁹ with the AFOSC camera. A complementary near-infrared (NIR) photometric campaign started on 2017 February 7 (approximately 1 month after the discovery) and ended on 2018 January 4. The NIR data were obtained using the NOT equipped with NOTCam.

All optical imaging frames were first pre-processed using standard tasks in IRAF¹⁰ for overscan, bias, and flat-field corrections, and trimming of unexposed frame regions. The steps necessary to obtain the SN magnitude were performed using the dedicated pipeline SNOOPY¹¹ developed by Cappellaro (2014). SNOOPY consists of a collection of PYTHON-based scripts calling standard IRAF tasks within PYRAF. The package also makes use of other public tools, including SEXTRACTOR¹² for source extraction from images, DAOPHOT¹³ (e.g. Stetson 1987) for instrumental magnitude measurements using the point-spread function (PSF) fitting technique, and HOTPANTS¹⁴ (e.g. Becker 2015) for PSF matching and template subtraction.

We used the PSF-fitting technique to measure instrumental magnitudes because, with a complex background, it provides more reliable results than aperture photometry. First, we subtracted the underlying background as calculated via a low-order polynomial fit. Then, using the profiles of isolated stars in the field of AT 2017be, we obtained the PSF model to be used for extracting the instrumental magnitude of the transient. The measured source is subtracted from the image and a new estimate for local background is obtained. Then the fitting procedure is iterated to improve the results. Error estimates are performed via artificial star experiments, in which fake stars with similar magnitude as the object are positioned very close to the location of AT 2017be. The simulated sources are then measured by the same PSF-fitting techniques, and the dispersion of these measurements provides an estimate of the recovered magnitude errors. However at late phases, when the transient became too faint to obtain reliable measurements with the PSF-fitting technique, we used template subtraction to properly remove the flux contamination of the host galaxy. The template images of the host galaxy were

observed as part of the Sloan Digital Sky Survey (SDSS¹⁵) on 2000 May 4, and were subtracted from the late-time *griz* images. This procedure was needed in particular for the LCO images, because the seeing was typically poor and hence background contamination more severe.

We determined the colour terms for all filters using observations of standard fields. Specifically, we used standard photometric fields from Landolt (1992) for calibrating the Johnson–Cousins data, and fields covered by the SDSS DR 13 catalogue for the Sloan-filter observations. The photometric zero-point of each image was calibrated with reference to a subset of stars in the field of the transient from the SDSS DR 13 catalogue. To determine the zero-point of the Johnson–Cousins filter images, the Sloan-filter magnitudes of SDSS sources in the field were converted to Johnson–Cousins filters using the relations of Chonis & Gaskell (2008). The final optical magnitudes of AT 2017be are reported in Table A1. We note that very few observations were obtained using the Johnson–Cousins *R* and *I* filters, hence these data were converted to Sloan-*r* and *i* SDSS systems following again the relations given in Chonis & Gaskell (2008).

NIR reductions include differential flat-field correction, sky subtraction, and distortion correction. The reductions were based on a modified version of the external IRAF package NOTCAM version 2.5.¹⁶ To improve the S/N, for each filter we combined multiple individual dithered exposures. NIR instrumental magnitudes of AT 2017be were measured with SNOOPY, and then calibrated using sources in the field from the Two Micron All Sky Survey (2MASS¹⁷) catalogue. The resulting NIR magnitudes (in the Vegamag system) for AT 2017be are reported in Table A2.

3.2 Light curves

The last pre-discovery non-detection reported by the KAIT is from 2016 December 29.479 UT (MJD = 57751.479), while the discovery epoch is 2017 January 6.508 UT (MJD = 57759.508). Therefore, it is reasonable to assume that the explosion of AT 2017be occurred between these two epochs. In the absence of more stringent indications, we adopt an explosion epoch of MJD = 57755 ± 4. Soon after the discovery of AT 2017be, we initiated a photometric follow-up campaign which lasted for around 1 yr. The resulting multiband light curves of AT 2017be are shown in Fig. 2. In the forthcoming discussion, we will adopt the epoch of the *r*-band peak magnitude (MJD = 57769.8 ± 0.1) as the reference time for the light and colour curves (see Section 3.3). This was obtained by fitting a third-order polynomial to the early *r*-band light curve.

The apparent magnitudes of AT 2017be in the different bands at maximum are listed in Table 2, while the decline rates in different phases of the evolution obtained through linear fits are reported in Table 3. Similar to other ILRTs, AT 2017be rises rapidly and reaches the peak magnitude in all bands in less than 2 weeks. The blue-band light curves reach their luminosity peaks a few days earlier than the red-band ones (see Fig. 2). After maximum, the luminosity of AT 2017be decays quite rapidly in all bands until day ~30, with the blue band (e.g. *B*gV) light curves usually fading more rapidly than the red bands (e.g. *ri*). Subsequently, the luminosity settles on to a plateau phase lasting approximately 30 d. The averaged

⁵<https://www.ptf.caltech.edu/iptf/>

⁶LCO is a global robotic network of telescopes; see <http://lco.global/>

⁷<http://www.not.iac.es/>

⁸<http://www.gtc.iac.es/>

⁹<http://www.oapd.inaf.it/>

¹⁰<http://iraf.noao.edu/>

¹¹<http://sngroup.oapd.inaf.it/snoopy.html/>

¹²www.astromatic.net/software/sextractor/

¹³<http://www.star.bris.ac.uk/~mbt/daophot/>

¹⁴<http://www.astro.washington.edu/users/becker/v2.0/hotpants.html/>

¹⁵<http://www.sdss.org/>

¹⁶<http://www.not.iac.es/instruments/notcam/guide/observe.html#reduction>

¹⁷<http://irsa.ipac.caltech.edu/Missions/2mass.html/>

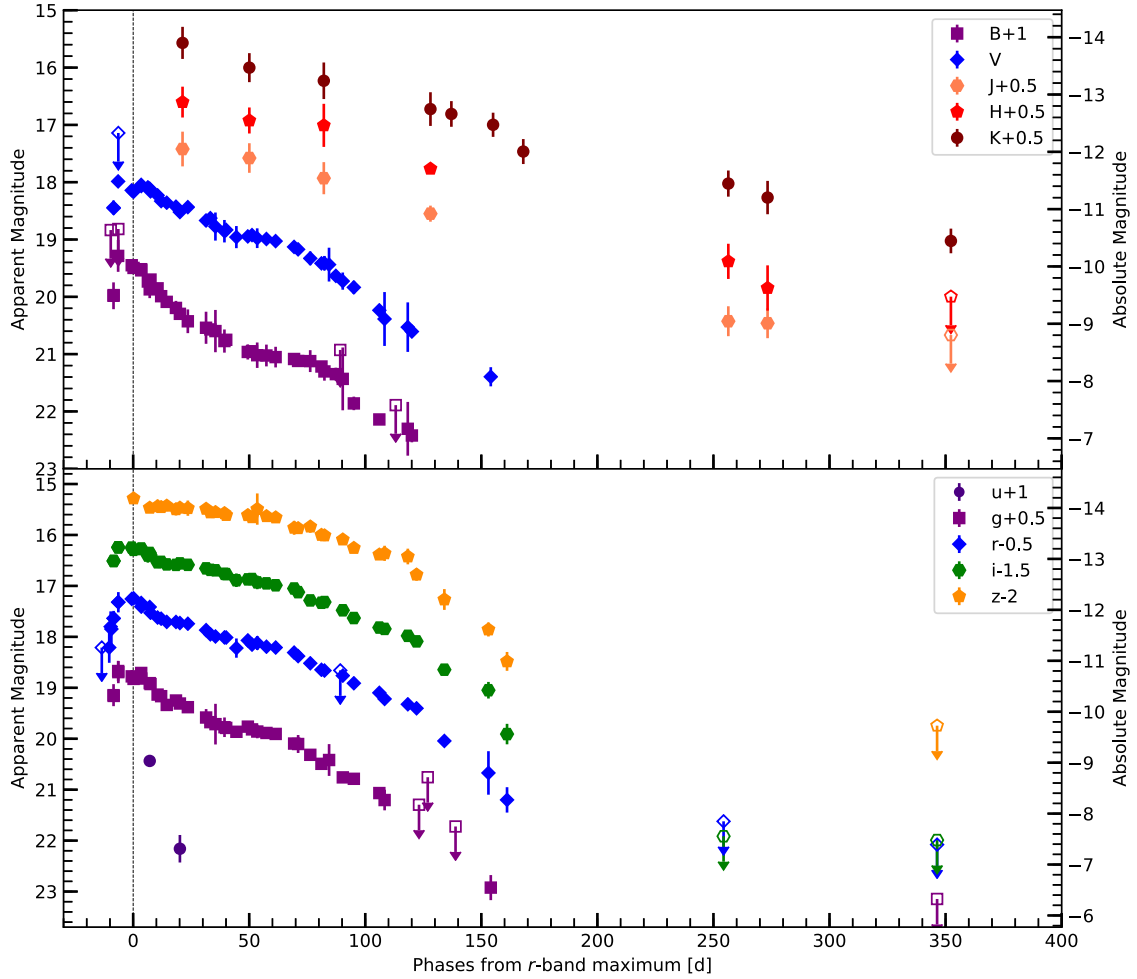


Figure 2. Multiband light curves of the transient AT 2017be. Both the optical and NIR observations are shown in this figure. Note that the *ugriz* magnitudes are in the AB magnitude scale, while *BV* and *NIR* magnitudes are calibrated to the Vega magnitude scale. Some points are taken from published values (as reported in Table A1). Detection limits at early and late phases are indicated with an empty symbol with a down-arrow. The errors for most data points are smaller than the marked symbols.

Table 2. Parameters for the peak and the plateau.

Filter	m_{peak} (mag)	m_{plateau} (mag)	μ_m (mag) ^a
<i>B</i>	18.46 ± 0.07	20.07 ± 0.06	1.61 ± 0.09
<i>V</i>	18.05 ± 0.05	19.03 ± 0.05	0.98 ± 0.05
<i>g</i>	18.21 ± 0.05	19.33 ± 0.04	1.12 ± 0.06
<i>r</i>	17.70 ± 0.04	18.63 ± 0.04	0.93 ± 0.04
<i>i</i>	17.73 ± 0.03	18.38 ± 0.03	0.65 ± 0.03
<i>z</i>	17.45 ± 0.35	18.12 ± 0.03	0.67 ± 0.35

^a μ_m is the magnitude difference between the averaged plateau and the peak.

Table 3. Decline rates in mag/100 d for the light curves.

Filter	γ_1 (mag/100 d)	γ_2 (mag/100 d)	γ_3 (mag/100 d)
<i>B</i>	3.22 ± 0.04	0.61 ± 0.06	2.80 ± 0.07
<i>V</i>	1.90 ± 0.03	0.74 ± 0.05	2.84 ± 0.05
<i>g</i>	2.67 ± 0.03	0.40 ± 0.04	2.77 ± 0.07
<i>r</i>	1.81 ± 0.02	0.60 ± 0.04	2.29 ± 0.03
<i>i</i>	1.18 ± 0.02	0.89 ± 0.03	2.12 ± 0.03
<i>z</i>	0.54 ± 0.03	0.61 ± 0.03	2.31 ± 0.04
<i>J</i>	–	0.84 ± 0.14	1.35 ± 0.28
<i>H</i>	–	0.66 ± 0.14	1.64 ± 0.27
<i>K</i>	–	1.08 ± 0.15	1.31 ± 0.12

plateau apparent magnitudes are also reported in Table 2. After the plateau phase, all light curves fade again more rapidly. From a morphological point of view, we can hence split the light-curve evolution in three phases: from the peak to ~ 30 d there is an initial fast decline ($\gamma_1 \approx 1\text{--}3$ mag/100 d, see Table 3); from ~ 30 to ~ 70 d a plateau ($\gamma_2 \approx 0.6\text{--}0.9$ mag/100 d), and from ~ 70 d to the latest detections before solar conjunction a faster late decline ($\gamma_3 \approx 1.3\text{--}2.8$ mag/100 d). The transient was followed in photometry up to late phases, but only upper limits were obtained in all optical bands after solar conjunction (which lasted ~ 100 d). Complementary NIR photometry is shown in Fig. 2 and the data are reported in Table 3. Unfortunately, we did not follow the transient in the NIR during the first month, hence the light-curve rise and the subsequent NIR peak have not been monitored, while the initial decline, and the late decline phases have been observed. The NIR decline rates γ_2 (optical plateau-like period) and γ_3 are reported in Table A2. We included the earliest photometric points in our measurements of γ_2 , and this may explain the relatively large decline rates found in the NIR light curves during the intermediate plateau phase (up to ~ 80 d). We also note that the measured *K*-band decline rate at very late phases in the *K* band is around 1.02 mag/100 d, which is consistent with the theoretical ^{56}Co decay of 0.98 mag/100 d assuming complete gamma-ray and positron trapping. However,

this is not a strong argument to support the ^{56}Co decay powering mechanism at late times, as the spectral energy distribution (SED; see Section 3.5) suggests the presence of a clear near-IR excess at these epochs probably due to emission from pre-existing dust in the CSM of AT 2017be. For this reason, we believe that the observed K -band decline match with the ^{56}Co decay rate is likely a coincidence.

3.3 Colour and absolute light curves

Fig. 3 shows the $B - V$ (upper panel), $R - I/r - i$ (middle panel), $J - K$ (lower panel) colour evolution of AT 2017be along with that of other ILRTs, including SN 2008S (Botticella et al. 2009), NGC 300 OT2008-1 (Humphreys et al. 2011), M85 OT2006-1 (Kulkarni et al. 2007), and PTF10fq (Kasliwal et al. 2011).

For SN 2008S, we adopted a distance modulus $\mu = 28.78 \pm 0.08$ mag, a Galactic reddening $E(B - V)_{\text{Gal}} = 0.36$ mag, and a host galaxy reddening $E(B - V)_{\text{host}} = 0.32$ mag (Botticella et al. 2009). For NGC 300 OT2008-1, we adopted a distance modulus of $\mu = 26.37 \pm 0.14$ mag (e.g. Bono et al. 2010; Tully et al. 2013; Bhardwaj et al. 2016, averaging several values available in the literature with Cepheid estimations), a Galactic reddening $E(B - V)_{\text{Gal}} = 0.01$ mag (Schlafly & Finkbeiner 2011), and a host galaxy reddening $E(B - V)_{\text{host}} = 0.25$ mag (Na I D measurement), which agree with the estimates of Humphreys et al. (2011). For M85 OT2006-1, Kulkarni et al. (2007) estimated a distance modulus of $\mu = 31.33 \pm 0.14$ mag and a total reddening $E(B - V)_{\text{total}} = 0.14$ mag. Here we updated the distance modulus of $\mu = 31.00 \pm 0.12$ mag (Tully et al. 2013) and used the same reddening estimate. Finally, for PTF10fq, Kasliwal et al. (2011) adopted Galactic reddening $E(B - V)_{\text{Gal}} = 0.04$ mag only. Note that a non-negligible contribution from host galaxy reddening cannot be ruled out, although the host galaxy is an S0-type, as there is evidence of low-level star formation (Pastorello et al. 2007). Here we adopt the Kasliwal et al. (2011) extinction and update the distance modulus to $\mu = 30.77 \pm 0.52$ mag (Springob et al. 2009).

At early phases, the transients included in our sample show a homogeneous colour evolution, which can be indicative of similar temperatures. The $B - V$ colour of AT 2017be evolves to progressively redder colours reaching 0.9 mag at about 60 d after maximum, which suggests a significant decrease in the ejecta temperature. This is confirmed through the inspection of the spectral blackbody temperature (Section 4). Later on, the $B - V$ index moves back to a bluer colour and remains between 0.5 and 0.9 mag at later phases. This evolution of AT 2017be is similar to that of NGC 300 OT2008-1. Fig. 3 (middle) shows the $R - I/r - i$ colour evolution of the five ILRTs. The $r - i$ colour of AT 2017be has very little evolution to redder colours (reaching 0.3 mag) until about day 30, while $B - V \sim 0.7$ mag at the same epoch (see the upper panel of Fig. 3). The $r - i$ colour of AT 2017be experiences a monotonic rise from 0.1 to 0.5 mag during the entire follow-up campaign. We note that a significant dispersion is observed in the colour evolution of AT 2017be in the data presented by Adams et al. (2018, see their fig. 1), while our data define a clearer colour evolution. However, the maximum colour of ~ 0.5 mag in our $r - i$ colour curve is in agreement with their measurement. A comparison at later phases is more difficult due to the larger dispersion in the Adams et al. (2018) data. The $J - K$ colour curves for our ILRT sample are shown in Fig. 3 (lower panel). We note that AT 2017be shows a NIR colour minimum of 1.5 mag at the phase of around 50 d, then it becomes progressively redder with time, reaching 2.2 mag. The overall trend with $J - K$

colours becoming redder with time seems to be a common property for all ILRTs.

Fig. 4 shows the absolute r -band light curve of AT 2017be, compared with those of the ILRTs mentioned above. AT 2017be reached an absolute peak magnitude of $M_r = -11.98 \pm 0.09$ mag as measured using a third-order polynomial fit to the light curve. A multi-band absolute light curve of AT 2017be was also shown in Adams et al. (2018, see their fig. 15). Their measurement of the Sloan r -band peak is around -11.4 mag. This difference of over half a magnitude is mainly due to different adopted distances and (more marginally) to the different extinction estimates. The Adams et al.'s apparent magnitude light curves are in good agreement with our data. The above value is intermediate between those of luminous novae and classical CC SNe (i.e. $-10 > M_V > -15$ mag), and is consistent with that expected for an ILOT (Smith et al. 2011; Tartaglia et al. 2016). The r -band light curve of AT 2017be has an asymmetric light-curve peak, with a rapid rise but a modest post-peak decline, with the light-curve settling on to a sort of plateau. AT 2017be is fainter than SN 2008S ($M_R = -14.24 \pm 0.03$ mag) and NGC 300 OT2008-1 ($M_R = -13.23 \pm 0.05$ mag), and nearly the same luminosity as M85 OT2006-1 ($M_R = -12.64 \pm 0.50$ mag) and PTF10fq (absolute peak magnitude $M_R = -11.56 \pm 0.33$ mag; see the upper panel of Fig. 4). Fig. 4 (lower panel) shows the R/r -band absolute magnitude evolution of AT 2017be and PTF10fq in the ABmag scale (Oke & Gunn 1983). The comparison shows the striking similarity in the early light-curve evolution for these two ILRTs, which reach approximately the same maximum magnitudes.

3.4 Quasi-bolometric light curve

Quasi-bolometric luminosities are obtained by converting the available magnitudes to flux densities, then correcting these fluxes with the adopted extinction coefficients, and finally integrating the SED over the entire wavelength range. We assume that the flux at the integration limits (at effective wavelengths shorter than the u band and longer than the K band) is zero. At some epochs, in a given filter, in which the photometric data are not available, we estimate the flux contribution of the missing bands by interpolating between epochs with available data or through an extrapolation from the earlier or later available epochs, assuming a constant colour evolution. We calculated the pseudo-bolometric light curve of AT 2017be and other ILRTs using only the $BVRi$ bands. This allows for meaningful comparisons among ILOTs, for which often u , z , or NIR bands were not available. In the case of AT 2017be we also computed a more complete quasi-bolometric light curve including the contribution to the total luminosity of the u and the NIR (z , J , H , K) bands. The resulting quasi-bolometric light curves are shown in Fig. 5.

The quasi-bolometric light-curve peak of AT 2017be is $\sim 8 \times 10^{39}$ erg s $^{-1}$, including the optical bands only; this value increases to $\sim 1.8 \times 10^{40}$ erg s $^{-1}$ including also the u and NIR bands contributions (~ 55.6 per cent). The peak luminosities of SN 2008S, NGC 300 OT2008-1, M85 OT2006-1, and PTF10fq (considering optical bands only) are about 6×10^{40} , 1.8×10^{40} , 10^{40} , and 6×10^{39} erg s $^{-1}$, respectively. The light-curve shape of AT 2017be is different from those of SN 2008S and NGC 300 OT2008-1, having a somewhat narrower peak and a post-peak plateau, and is very similar to those of M85 OT2006-1 and PTF10fq, in overall shape, peak luminosity, and post-peak light-curve flattening. The main difference is in the duration of the plateau. While the post-plateau decline in AT 2017be begins about 70 d after maximum light, and at a similar phase in PTF10fq, for M85 OT2006-1 the

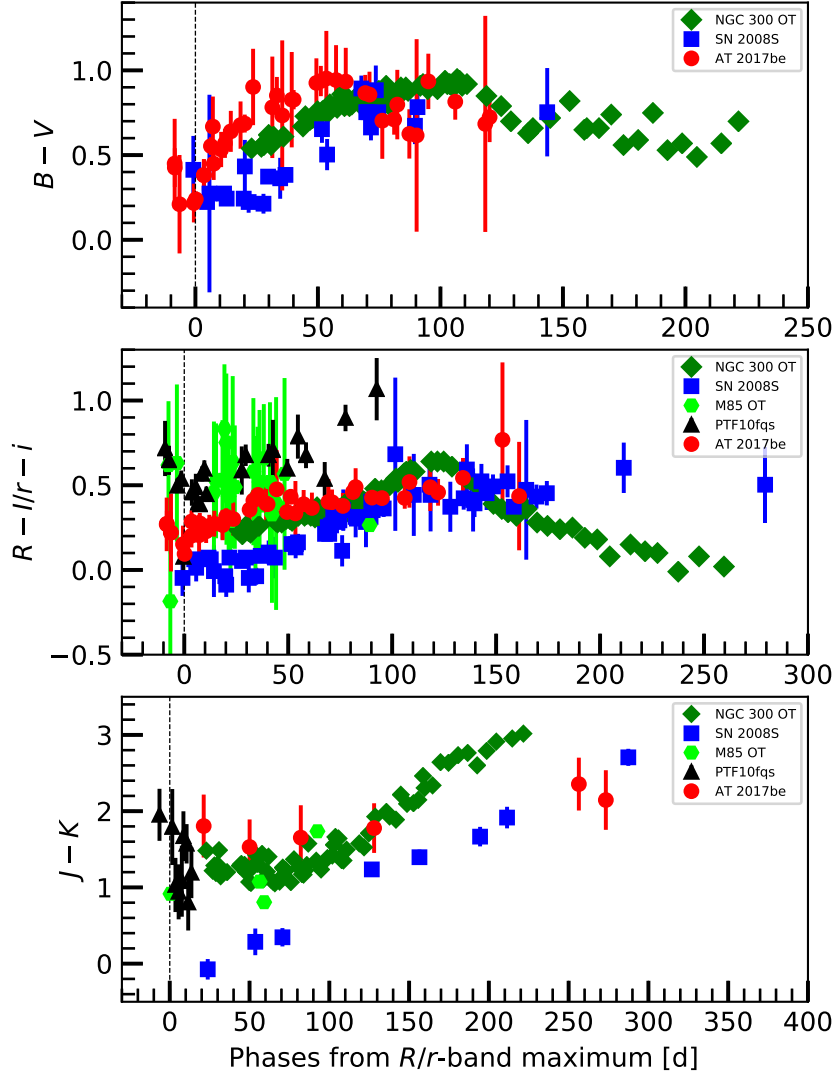


Figure 3. Comparison of $B - V$, $R - I/r - i$, and $J - K$ colour curves of AT 2017be, SN 2008S, NGC 300 OT2008-1, M85 OT2006-1, and PTF10fqs. The data of SN 2008S are from Botticella et al. (2009), those of NGC 300 OT2008-1, M85 OT2006-1, and PTF10fqs are from Humphreys et al. (2011), Kulkarni et al. (2007), and Kasliwal et al. (2011), respectively. The adopted reddening values are given in the main text. The $r - i$ colours shown here are obtained from Sloan magnitudes, converted to the Vega system. The errors on $J - K$ colour were not included in the NIR dataset for NGC 300 OT2008-1.

plateau only lasts for around 30–35 d. The post-plateau decline of AT 2017be is very rapid, with a rate that is faster than the decay of ^{56}Co (see Fig. 5). Unfortunately, we could not follow the object to observe an optical flattening on to the radioactive tail, because the transient disappeared behind the sun. However, we observed again the field a few months later, when it became visible again before the sunrise, and noticed that AT 2017be was no longer visible in the optical bands, while it was clearly detected in the NIR bands.

As we have a limited light-curve information for AT 2017be at late phases, we cannot securely estimate its bolometric light curve, which is essential to constrain the ejected ^{56}Ni mass. In addition, AT 2017be has no mid-IR monitoring, and we know that similar objects like SN 2008S and NGC 300 OT2008-1 have a non-negligible mid-IR contribution at late phases, with the peak of their SEDs shifting from the optical to the mid-IR (Kochanek 2011). Given that SN 2008S and NGC 300 OT2008-1 have relatively complete datasets spanning from the optical to mid-IR and extended to late phases, we use their SEDs to estimate the relative flux contribution of the different wavelength regions. Assuming that the SED evolution of

AT 2017be is similar to that of the two template objects, we can in this way constrain the contribution of the missing bands and estimate a reliable bolometric light curve also for the former object.

We first used SN 2008S as template for AT 2017be, in view of the similar $J - K$ colour evolution. We estimated the NIR luminosity of SN 2008S from its light curve, and adopted its model bolometric light curve obtained by Kochanek (2011) through the relation: $L_{\text{bol}} = L_0 \exp(-t/t_0) + L_1$, with $L_0 \simeq 10^{7.3} L_{\odot}$, $L_1 \simeq 10^{5.8} L_{\odot}$, and $t_0 \simeq 48$ d. Hence we computed the $L_{\text{NIR}}/L_{\text{bol}}$ ratio for SN 2008S to be ≈ 0.31 – 0.36 , from ~ 245 to 345 d after maximum. Finally, we obtained the bolometric light curve of AT 2017be using its L_{NIR} values obtained from the light curve at late phases assuming the same $L_{\text{NIR}}/L_{\text{bol}}$ ratio as SN 2008S. The resulting bolometric light curve for AT 2017be at late phases is also shown in Fig. 5 with a solid line.

Assuming that the energy from γ rays and positrons is fully trapped and thermalized, we constrain the ^{56}Ni mass deposited by AT 2017be by comparing its late-time bolometric luminosity with that of SN 1987A at the same epochs (~ 260 – 360 d after their ex-

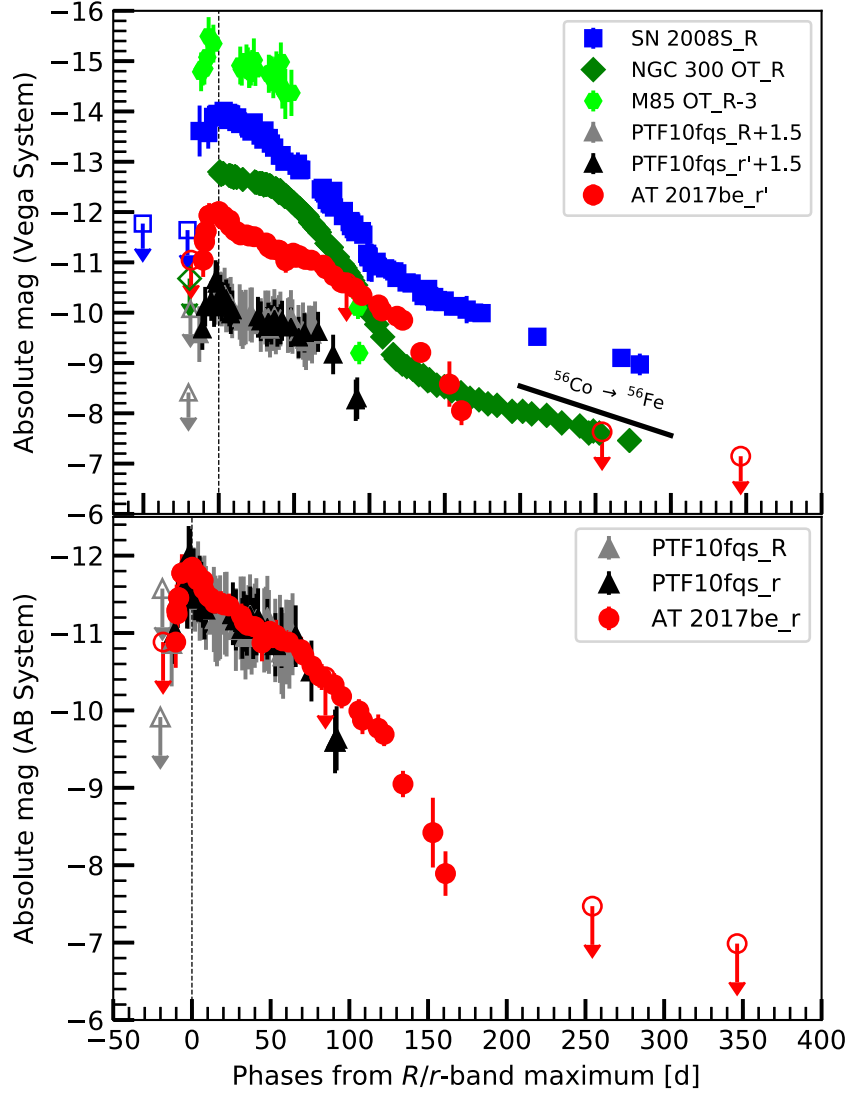


Figure 4. *R/r* band absolute light curves of AT 2017be and other ILRTs, including SN 2008S (Botticella et al. 2009), NGC 300 OT2008-1 (Humphreys et al. 2011), M85 OT2006-1 (Kulkarni et al. 2007), and PTF10fqs (Kasliwal et al. 2011). In order to improve visualization, arbitrary shifts are applied to M85 OT2006-1 and PTF10fqs as per the legend. We adopt the time of the *R/r*-band light-curve peaks as our reference epoch.

plosion). The luminosity of SN 1987A was calculated by Whitelock et al. (1988), including the *U*- to *M*-band contribution (we will assume that the contribution to the total luminosity from the regions outside this range is negligible). The ^{56}Ni mass synthesized by AT 2017be ($\sim 7 \times 10^{-4} M_{\odot}$) is obtained through the relation:

$$M(^{56}\text{Ni})_{\text{AT 2017be}} = 0.073 M_{\odot} \times \left(\frac{L_{\text{AT 2017be}}(t)}{L_{\text{SN 1987A}}(t)} \right). \quad (1)$$

As a further test, we can use the SED of NGC 300 OT2008-1 as a template, adopting the same algorithm as SN 2008S, and we find $M(^{56}\text{Ni}) \sim 9 \times 10^{-4} M_{\odot}$. Averaging the above estimates, we infer $M(^{56}\text{Ni}) = 8 (\pm 1) \times 10^{-4} M_{\odot}$ for AT 2017be. However, we remark that additional contributions to the late bolometric light curve from interaction between ejecta and CSM, or even IR echoes cannot be definitely ruled out (see the further discussion in Section 3.5). In that case, the above $M(^{56}\text{Ni})$ value should be regarded as an upper limit.

Integrating the bolometric light curve over the entire observational campaign (~ 370 d) of AT 2017be, we can estimate a re-

liable value for the total radiated energy. We have a good early time coverage in the optical and NIR bands at early phases. Hence, we can estimate the contribution of the missing bands using again both SN 2008S and NGC 300 OT2008-1 templates. In the two cases, we obtain $\sim 1.19 \times 10^{47}$ and $\sim 1.31 \times 10^{47}$ erg as energy radiated by AT 2017be, with an average of $1.25 (\pm 0.06) \times 10^{47}$ erg.

Although ECSNe are expected to arise from stars in a very narrow range of masses, and the mass of ^{56}Ni only marginally affect the observed luminosity (in particular, at early phases), the quasi-bolometric light curves of ILRTs are expected to be reasonably similar (e.g. Kitaura et al. 2006; Pumo et al. 2009; Moriya & Eldridge 2016). Indeed, as shown in Fig. 5, the light curves of all our ILRTs are indeed broadly similar. The differences between these events can potentially be explained by a number of factors, including the mass-loss history (hence the geometric distribution and density profile of the CSM), the H recombination timescale, and the influence of ^{56}Ni distribution inside the ejecta.

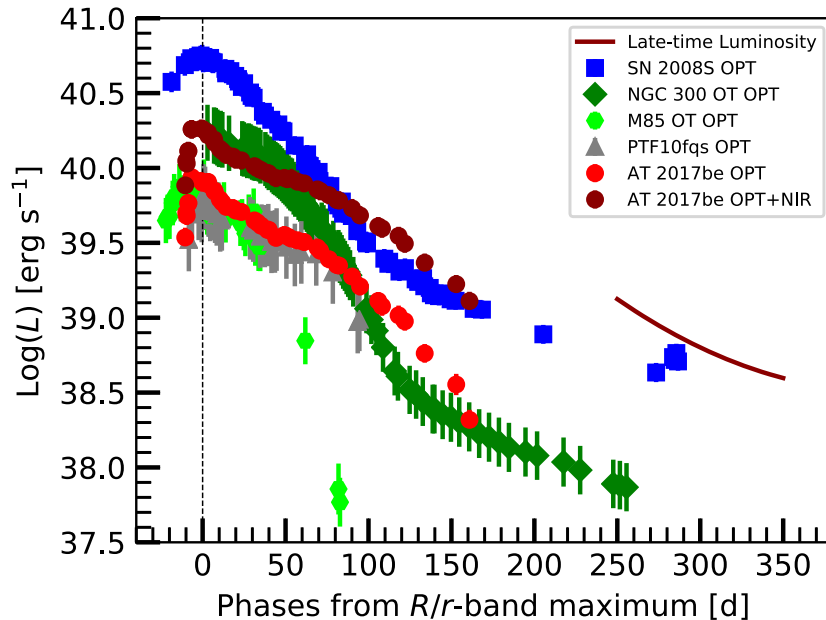


Figure 5. Quasi-bolometric light curves of AT 2017be, SN 2008S (Botticella et al. 2009), NGC 300 OT2008-1 (Humphreys et al. 2011), M85 OT2006-1 (Kulkarni et al. 2007), and PTF10fqs (Kasliwal et al. 2011). For the purpose of comparison, only the Johnson/Cousins *BVR* and SDSS *gri* filters were used. Note that for AT 2017be we also computed the *u*-to-*K*-band light curve (shown with dark red dots). Errors shown account for the uncertainties on photometric measurements, distance and reddening estimates, and SED fitting.

3.5 Spectral energy distribution analysis

The temporal evolution of the SED of AT 2017be is shown in Fig. 6. We did not collect NIR observations until about 3 weeks after maximum, so our SED at early epochs is based solely on optical data. These early epochs are well represented by a single blackbody, which is fitted to the optical data, whose temperature declines rapidly from ~ 6500 to ~ 5500 K by $\sim +20$ d (see the upper panel of Fig. 7). The first optical+NIR SED is obtained at +21 d from maximum, which reveals a notable NIR excess over a single black body (see inset in the upper-right corner of Fig. 6). To investigate the nature of the NIR excess, we carried out a simultaneous two-component fit to the optical+NIR SEDs consisting of two (hot plus warm components) black-body functions. Occasionally, for some epochs fluxes for unobserved bands were estimated through an interpolation procedure. At these later epochs (after +20 d), the temperature of the hot component slowly declines from about 5500 to 4300 K at about 160 d. As expected, the above temperature estimates are in agreement with those obtained from the blackbody fit to the continuum of the spectra obtained at the similar epochs (see Section 4.1, and Fig. 7, upper panel). Furthermore, the derived temperature is consistent with the hydrogen recombination temperature. The temperature of the warm blackbody is apparently constant over the 3 months (from ~ 20 to 150 d) of monitoring, within substantial errors, $T = 1200 \pm 200$ K, indicative that this component is very likely due to thermal emission from surrounding dust.

From luminosity and temperature we derive the radius of AT 2017be using the Stefan–Boltzmann law. The evolution of the wavelength-integrated luminosity of the hot component is shown in the middle panel of Fig. 7. It decreases from $\sim 1.7 \times 10^{40}$ erg s $^{-1}$ at -6.5 d to $\sim 4.5 \times 10^{38}$ erg s $^{-1}$ at about 3 months from maximum. The radius of the hot component (see the lower panel of Fig. 7) slowly increases from 1.1×10^{14} cm near maximum light to 1.3×10^{14} cm at about 1 month after peak luminosity (see lower panel in Fig. 7). After this until about four months past-maximum

the radius decreases from about 1.4×10^{14} to 1.2×10^{14} cm, before dropping more rapidly to about 4×10^{13} cm at 161 d. The time evolution of the luminosity as well as the blackbody radius and hot temperature are reported in Table 4, and shown in Fig. 7.

NIR excess emission has been observed in several SNe and is often attributed to the thermal emission from warm dust heated by the SN. On the other hand, Kochanek (2011) investigated the evolving SEDs of SN 2008S and NGC 300 OT2008-1, extending to mid-IR wavelengths, and found that the transient SEDs are largely determined by the dust reformation beyond the shock front.

Assuming the near-infrared emission is indeed due to dust, we fit a modified blackbody function (Hildebrand 1983) to the optical+NIR SEDs. We fit for the dust temperature, T_d , and dust mass, M_d , adopting the formalism as described in Gall et al. (2017), which assumes that the dust mass absorption coefficient, $\kappa_{\text{abs}}(\nu, a)$ (in units of cm 2 g $^{-1}$) behaves as a λ^{-x} power law in the wavelength range 0.9–2.5 μm , with x as the power-law slope. Here we perform fits for either $x = 1.2$ or 1.5 and adopt $\kappa_{\text{abs}}(\lambda = 1 \mu\text{m}) = 1.0 \times 10^4$ cm 2 g $^{-1}$ [appropriate for carbonaceous dust, e.g. Rouleau & Martin (1991)].

The results for the temperature and radius of the blackbody function for the supernova are identical to the results obtained from the pure hot and warm blackbody fits. However, here we obtain a slightly lower temperature (~ 1000 K) for the dust component which also shows no evolution over time. The amount of dust decreases over time from a few times $10^{-5} M_{\odot}$ to a few times $10^{-6} M_{\odot}$ in the case of carbonaceous dust. However, the low dust temperatures do not rule out other dust species such as silicates. In the case of silicate dust the estimate of the dust mass would be about an order of magnitude larger.

Unfortunately, in the absence of longer wavelength observations to better constrain the dust component we can only derive an upper limit to the dust temperature and a lower limit on the dust mass. From the analysis of the SED of AT 2017be SED, we infer that pre-existing dust from the progenitor likely survived in the circumstellar

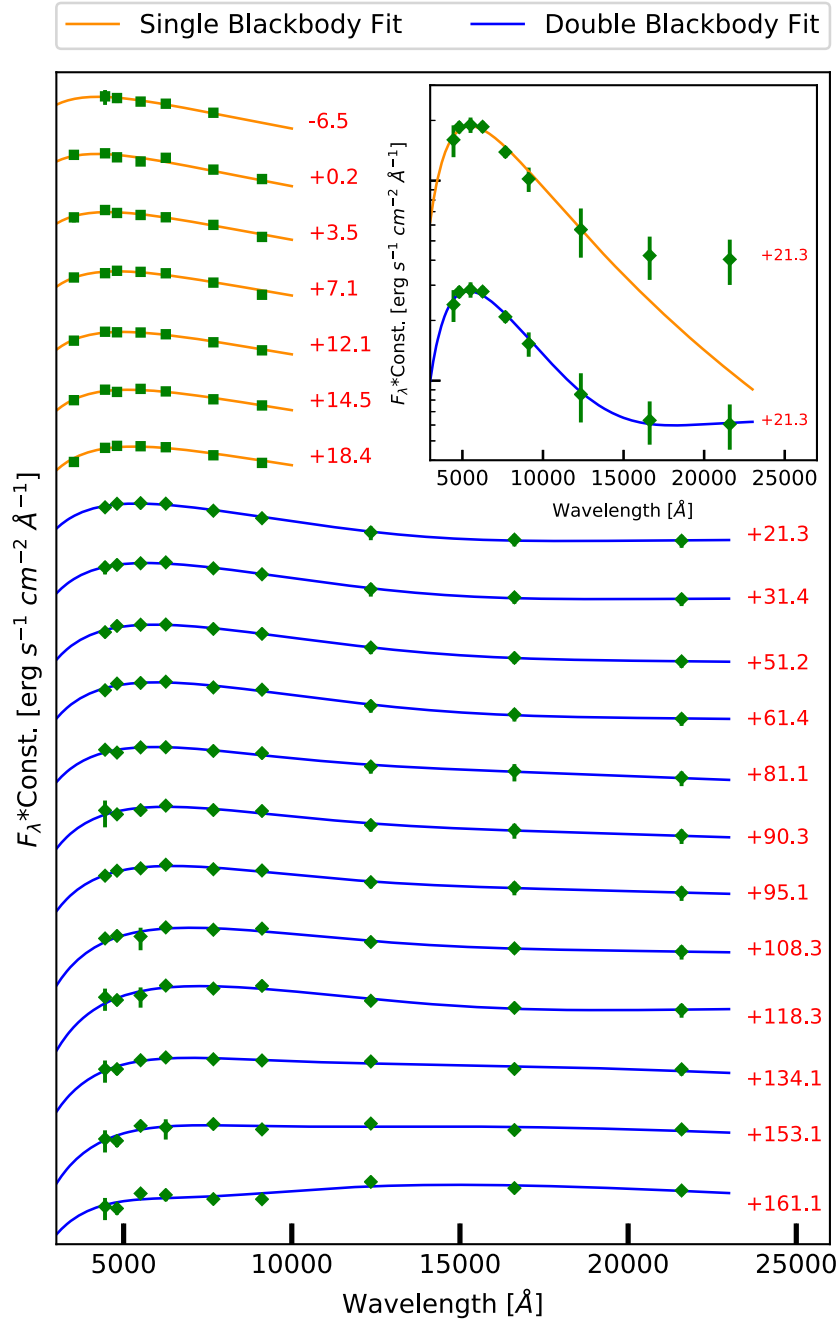


Figure 6. The observed SED evolution of AT 2017be. The fluxes have been corrected for reddening. We adopt the value of $\text{MJD} = 57769.8 \pm 0.1$ (Sloan r -band maximum) as our reference epoch.

environment, and was probably heated by the SN. This dust has to be distributed far from the progenitor and such that it does not contribute to additional extinction along the line of sight towards the SN (see the extinction discussion in Section 2). This implies that either the dust mass is only slightly above our lower limit, or the dust distribution of AT 2017be may be similar to that of V838 Mon (e.g. Tylenda 2004; Tylenda & Kamiński 2012), in form of clumpy, complex structure.

4 SPECTROSCOPY

Our spectroscopic follow-up campaign started on 2017 January 10, and lasted until 2017 April 18, hence covering almost 100 d. Spectroscopic observations were performed using the 2-m Faulkes North Telescope of LCO with FLOYDS, the 2.56-m NOT with ALFOSC, and the 10.4-m Gran Telescopio Canarias (GTC) with OSIRIS. Details on the instrumental configurations and information on our spectra of AT 2017be are reported in Table 5.

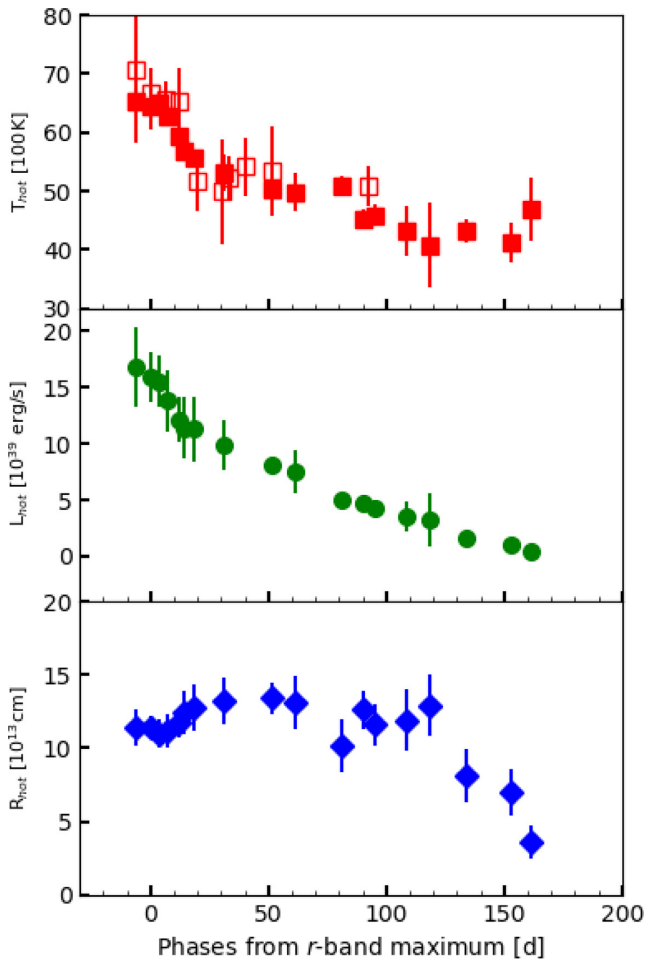


Figure 7. Temporal evolution of parameters determined from the SED of AT 2017be. In the upper panel, empty squares indicate the temperature evolution deduced from the blackbody fit to the spectral continuum. The uncertainty in the inferred blackbody temperature has been estimated based on the uncertainties in flux in each bandpass. We adopt $\text{MJD} = 57769.8 \pm 0.1$ (Sloan *r*-band maximum) as our reference epoch.

The spectra of AT 2017be were reduced using standard tasks in IRAF.¹⁸ Overscan, bias, and flat-field corrections were applied using a similar procedure as for the photometric images. From the two-dimensional frames, we performed an optimal extraction of one-dimensional spectra using the package APALL. The one-dimensional spectrum was wavelength-calibrated using spectra of comparison lamps obtained during the same night and with the same instrumental setup. The accuracy of the wavelength calibration was checked by measuring the observed wavelengths of a few night-sky lines, and (if necessary) a constant shift was applied to match the expected line wavelengths. Spectra of selected spectro-photometric standard stars were obtained during the same night as the SN observation to allow flux calibration of the spectra of AT 2017be. The precision of the flux calibration was verified using the available coeval photometric data. When some flux discrepancy was found, a scaling factor was applied to the spectrum to obtain a good match with our broad-band photometry. Finally, the strongest telluric absorption

bands of O₂ and H₂O in the spectra of AT 2017be were removed using the spectra of early-type standard stars.

In order to properly analyse our spectra of AT 2017be, we applied a correction to account for the line-of-sight reddening contribution (see Section 2), and a Doppler shift to place the spectra in the host galaxy rest frame. The spectral resolutions, reported in Table 5, were obtained by measuring the average widths of night sky lines.

4.1 Spectral evolution

The spectra of AT 2017be show little evolution over the entire monitoring period (from 7 d to almost 100 d after the explosion, see Fig. 8). All spectra consist of a nearly featureless continuum with strong Balmer emission lines. The [Ca II] doublet at 7291, 7324 Å and the Ca II NIR triplet are also detected. The P-Cygni absorption troughs, which usually characterize SN explosions or the presence of CSM material if narrow, are observed in our highest resolution GTC (R2500R) spectra. The H α line possibly shows a multicomponent profile, although in most spectra of AT 2017be the line is well reproduced by fitting a single Lorentzian profile (see Section 4.3).

The spectra are characterized by a significant contribution of the continuum to the total flux, hence we estimated the temperature of the photosphere through a blackbody fit. The blackbody temperatures from spectra and imaging are consistent within the errors (see the upper panel of Fig. 7). We note that this temperature evolution is consistent with that of hydrogen recombination in normal Type II CC SNe (e.g. Filippenko 1997). The spectral blackbody temperature values are reported in Table 6 and are plotted in Fig. 7 (upper panel).

4.2 Line identification

The line identification was performed on two high-S/N spectra: an early epoch (+0.2 d) ALFOSC spectrum, and a late-time (92.1 d) GTC spectrum. The identification is shown in Fig. 9. H α and other Balmer series emission lines are clearly present. Paschen series lines are also visible at the reddest wavelengths in our highest S/N spectra. Additionally, weak lines of Fe II, Na I D, O I, Ca II H&K, and the Ca II NIR triplet are identified at their rest wavelengths. A possible Fe II line is visible from first spectrum up to 3 weeks after maximum light (see the question mark in the upper panel of Fig. 9).

We also identify prominent [Ca II] doublet lines at $\lambda\lambda 7291, 7324$, which are visible in all spectra of AT 2017be (see Fig. 8), and which are seen in all the ILRTs. In Fig. 10, we compare spectra of AT 2017be and other four ILRTs at similar phases (marked with red labels in the figure). The striking spectral match of AT 2017be with the SN 2008S-like family is evident not only from the detection of the [Ca II] doublet (see right-hand panel in Fig. 10), but also from the presence of other shared spectral features (both in absorption and emission), providing further observational evidence to support our classification of AT 2017be as an ILRT.

4.3 Evolution of the line profiles

In order to search for signatures of ejecta–CSM interaction we analysed the evolution of the line profiles, along with the velocity and the luminosity of H α at different phases. Multiple line components in spectra are normally associated with interaction between supernova ejecta and circumstellar wind, with the different components resulting from multiple emitting gas regions (Turatto et al. 1993). The classical interpretation is that the narrow component of H α is

¹⁸NOT+ALFOSC spectra were reduced using a dedicated graphical user interface (ALFOSC GUI) developed by E. Cappellaro (<http://sngroup.oapd.inaf.it/alfoscgui.html>).

Table 4. Properties of the hot blackbody component fit to the uBVgrizJHK bands of AT 2017be. The warm component does not show a significant evolution within its uncertainties. Uncertainties are given in parentheses.

Date	MJD	Phase ^a (d)	Luminosity (hot) (10 ³⁹ erg s ⁻¹)	Temperature (hot) (K)	Radius (hot) (10 ¹³ cm)
20170110 ^b	57763.3	-6.5	16.8 (4.0)	6530 (190)	11.4 (1.0)
20170116 ^b	57770.0	+0.2	15.9 (2.0)	6430 (390)	11.4 (1.0)
20170120 ^b	57773.3	+3.5	15.5 (2.0)	6510 (170)	11.0 (1.0)
20170123 ^b	57777.0	+7.1	13.8 (3.0)	6280 (170)	11.2 (1.0)
20170128 ^b	57781.9	+12.1	12.1 (2.0)	5930 (70)	11.7 (1.0)
20170131 ^b	57784.3	+14.5	11.4 (3.0)	5680 (130)	12.4 (2.0)
20170204 ^b	57788.3	+18.4	11.3 (3.0)	5580 (140)	12.8 (2.0)
20170207	57791.1	+21.3	10.9 (3.0)	5540 (220)	12.8 (1.0)
20170217	57801.2	+31.4	9.9 (2.0)	5310 (300)	13.2 (2.0)
20170309	57821.0	+51.2	8.1 (1.0)	5020 (100)	13.4 (1.0)
20170319	57831.2	+61.4	7.6 (2.0)	4990 (320)	13.1 (2.0)
20170407	57850.9	+81.1	5.0 (1.0)	5100 (170)	10.2 (2.0)
20170417	57860.1	+90.3	4.8 (1.0)	4530 (160)	12.6 (1.0)
20170421	57864.9	+95.1	4.3 (1.0)	4590 (200)	11.6 (1.0)
20170505	57878.1	+108.3	3.5 (1.0)	4320 (420)	11.9 (2.0)
20170515	57888.1	+118.3	3.3 (2.0)	4070 (720)	12.9 (2.0)
20170530	57903.9	+134.1	1.6 (0.3)	4310 (200)	8.1 (2.0)
20170618	57922.9	+153.1	1.0 (0.3)	4130 (330)	7.0 (2.0)
20170626	57930.9	+161.1	0.4 (0.2)	4690 (540)	3.6 (1.0)

^aPhases are relative to *r*-band maximum light curve, on MJD = 57769.8 ± 0.1.^bData are obtained by fitting a single blackbody function, as we have no contemporaneous NIR observations.**Table 5.** Log of the spectroscopic observations of AT 2017be.

Date		Phase ^a (d)	Telescope + instrument	Grism	Spectral range (Å)	Resolution (Å)	Exp. time (s)
20170110	57763.4	−6.4	LCO + FLOYDS	Channels red + blue	3600–10 000	15.5	3600
20170117	57770.0	+0.2	NOT + ALFOSC		gm4	3400–9700	15
20170123	57776.2	+6.4	NOT + ALFOSC	gm4	3390–9650	18	3512
20170128	57781.9	+12.1	NOT + ALFOSC	gm4	3500–9590	15	3600
20170205	57790.0	+20.1	NOT + ALFOSC	gm4	3400–9650	18	3600
20170216	57800.0	+30.2	GTC + OSIRIS	R2500R	5580–7680	3.5	1800
20170219	57803.0	+33.2	NOT + ALFOSC	gm4	3400–9700	14	3600
20170225	57809.9	+40.1	NOT + ALFOSC	gm4	3700–9650	19	3600
20170308	57821.0	+51.2	NOT + ALFOSC	gm4	3400–9650	18	3600
20170418	57861.9	+92.1	GTC + OSIRIS	R1000R	5100–10 400	8	2 × 900

^aPhases are relative to *r*-band maximum light, on MJD = 57769.8 ± 0.1.

emission from unshocked CSM, while broader line components are signatures of shocked material and/or freely expanding SN ejecta.

(i) H α line

The profile of H α is characterized by a narrow component which dominates in all our spectra of AT 2017be. Its profile is usually well reproduced by a simple Lorentzian fit. We first fitted the H α profile with a single Lorentzian function using standard IRAF tasks, and measured the full width at half-maximum (FWHM) and flux values. In order to account for the instrumental limitations, we corrected the FWHM values for the spectral resolution (width = $\sqrt{\text{FWHM}^2 - \text{resolution}^2}$), and finally we computed the H α FWHM velocities (v_{FWHM}). The resulting velocities and luminosities are reported in Table 6, and the H α luminosity evolution is plotted in Fig. 11.

The velocity ranges from $v_{\text{FWHM}} \sim 300$ to ~ 850 km s⁻¹ in a rather erratic way. However, we should note that in most spectra the FWHM of H α is close to the instrumental resolution, hence these values are unreliable for determining the velocity of the emitting H-rich material. From our highest quality and best resolution GTC+OSIRIS spectra, we infer $v_{\text{FWHM}} \sim 820$ km s⁻¹ and $v_{\text{FWHM}} \sim 730$ km s⁻¹ at phases +30.2 and +92.1 d, respectively, which we regard as our most reliable values for the line velocity. During the first ~ 30 d, the H α luminosity decreases from $\sim 13 \times 10^{37}$ to $\sim 3 \times 10^{37}$ erg s⁻¹. Later on, it settles on to slightly higher values, around $4\text{--}6 \times 10^{37}$ erg s⁻¹. It is unclear if this apparent H α luminosity increase at later phases is real or is rather due to systematic errors (such as slit losses). However, as the H α luminosity increases by nearly a factor of two, we tentatively attribute it to a progressive strengthening of the interaction of ejecta with CSM.

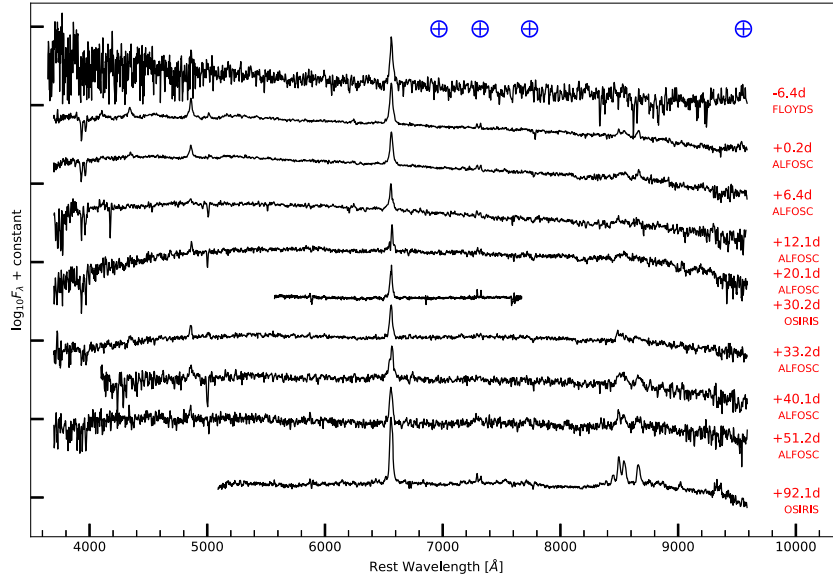


Figure 8. Spectroscopic evolution of AT 2017be from 2017 January 10 (MJD = 57763.38) to April 18 (MJD = 57861.93). The epochs reported to the right of each spectrum are relative to the Sloan *r*-band maximum. The positions of the strongest telluric absorption bands are marked with the $\oplus$ symbol; all spectra have been redshift-corrected.

Table 6. Main parameters inferred from the spectra of AT 2017be.

Date	MJD	Phase ^a (d)	FWHM H α (Å)	Velocity H α (km s ⁻¹)	Luminosity H α ^d (10 ³⁷ erg s ⁻¹)	Temperature (K)
20170110 ^c	57763.4	-6.4	19.9	570	13.0	7650 $\pm$ 1250
20170117 ^c	57770.0	+0.2	19.8	590	10.2	6650 $\pm$ 400
20170123 ^c	57776.2	+6.4	20.8	480	7.3	6550 $\pm$ 300
20170128 ^c	57781.9	+12.1	18.6	500	4.1	6550 $\pm$ 550
20170205 ^c	57790.0	+20.1	19.2 ^b	310	3.7	5150 $\pm$ 500
20170216	57800.0	+30.2	18.4	820	3.0	5000 $\pm$ 900
20170219 ^c	57803.0	+33.2	19.3	610	3.8	5250 $\pm$ 350
20170225 ^c	57809.9	+40.1	24.9	740	5.3	5400 $\pm$ 500
20170308 ^c	57821.0	+51.2	22.0	580	4.3	5650 $\pm$ 750
20170418	57861.9	+92.1	17.9	730	5.9	5100 $\pm$ 350

^aPhases are relative to *r*-band maximum light, on MJD = 57769.8 $\pm$ 0.1.

^bThere is a hot pixel close to the H α position, affecting the recovered line profile and making the above measurement uncertain.

^cThese measurements are close to the instrumental resolution, hence after correcting for instrumental resolution the resulting velocity has a large uncertainty.

^dWe assume conservative errors on the luminosity of about 20 per cent.

Unfortunately, the modest resolution of our spectra is insufficient to reveal very narrow components in the H α profile. The only exception is the spectrum at +30.2 d, that can resolve features down to a limit of 160 km s⁻¹. This spectrum will be discussed separately later on (Section 4.4). The evolution of the H α profile during the 3 months of spectral coverage is shown in Fig. 12 (left-hand panel). We note that there is no significant change in the wavelength position of the H α emission peak.

(ii) Ca II lines

The evolution of the Ca II NIR triplet line profiles in the wavelength window 8300–8800 Å for AT 2017be is shown in Fig. 12 (right-hand panel). In all spectra we mark the individual positions of Ca II triplet lines at 8498, 8542, and 8662 Å. These are produced through radiative de-excitation from the upper 4p²P_{1/2,3/2} levels to the lower metastable 3d²D_{3/2,5/2} levels (Mallik 1997). This triplet is a very common feature observed in the spectra of many types of transients. In our low-resolution spectra, the 8498 and 8542 Å

members of the Ca II triplet are usually blended. In order to estimate the FWHM velocity of the Ca II lines, we considered the two spectra with the highest S/N: the NOT (phases +0.2 d) and GTC (phases +92.1 d) spectra. These spectra exhibit narrow features with velocity of 700–800 km s⁻¹, which are consistent with the values obtained for H α . This is consistent with the Ca II triplet arising from the same region as the main H α emission. On the other hand, the Ca II H&K absorptions are also clearly discernible in some epochs (the spectra at +0.2, +6.4, +12.1, +20.1, and 33.2 d in Fig. 8), with a velocity of $\sim$ 180 km s⁻¹ (as measured from two minima of the Ca II H&K lines).

The [Ca II] doublet lines at 7291, 7324 Å are always visible in our spectra, and appear to be slightly narrower than H α , hence in most cases these lines are unresolved. The [Ca II] doublet is unblended in all our spectra, although the two components show different relative intensities and widths. The intensity ratio of the two [Ca II] components (λ 7291 versus λ 7328) slightly increases from $\sim$ 0.9 in our early spectra to $\sim$ 1.3 in our late spectra. However, we remark

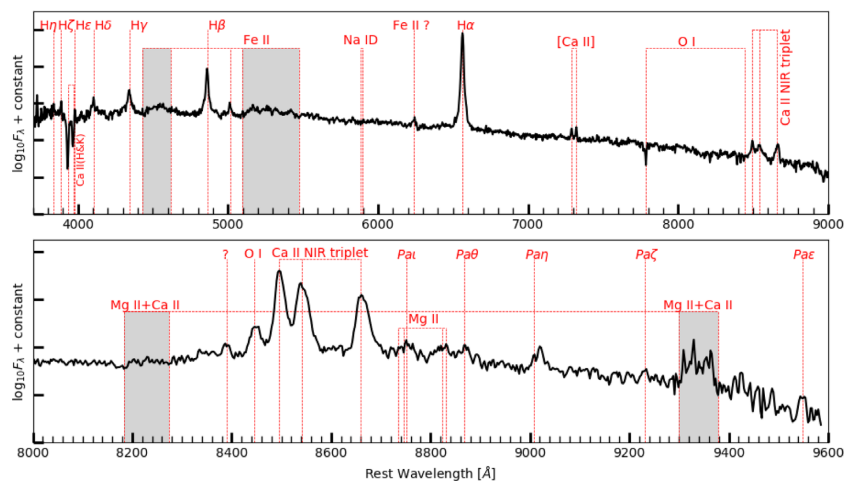


Figure 9. Upper panel: Identification of the strongest emission lines in the early-time (~ 1 d) NOT + ALFOSC spectrum of AT 2017be. Lower panel: Identification of the most prominent emission lines in the late-time (~ 90 d) GTC + OSIRIS spectrum of AT 2017be. Both the spectra have been redshift-corrected.

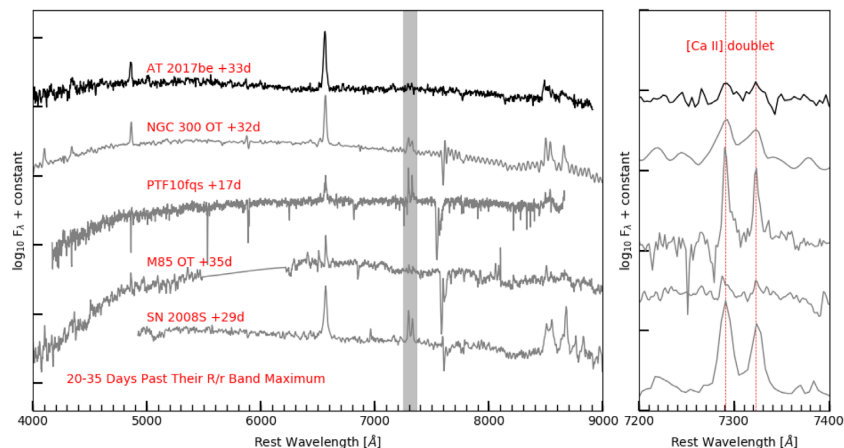


Figure 10. Comparison of spectra of those of ILRTs obtained around the same epoch after their R/r -band maximum. All the spectra are taken from Cai et al. (in preparation), Kasliwal et al. (2011), Kulkarni et al. (2007), and Botticella et al. (2009). The [Ca II] doublet lines are marked in the grey shaded area in the left-hand panel. The right-hand panel shows the detailed profiles. The spectra are all redshift-corrected and extinction-corrected.

that the uncertainty in the flux ratio measurements is of the order of 20 per cent.

As in most spectra the [Ca II] lines are unresolved, we consider only the two moderate resolution GTC spectra (at phases +30.2 and +92.1 d, respectively). Also in these spectra, however, the FWHMs of the [Ca II] doublet are close to the spectral resolution limit, therefore any velocity measurement is affected by a significant uncertainty. For the two spectra, we obtain ~ 115 and ~ 250 km s^{-1} for the individual components of the doublet. The [Ca II] doublet lines are produced in a relatively low-density gas, from radiative de-excitation related to the metastable $3d^2D$ level (Chevalier & Fransson 1994). A major implication of the above analysis is that the [Ca II] and Ca II H&K lines form in a different region as the Ca II NIR triplet (and $H\alpha$). In particular, [Ca II] and Ca II H&K lines are likely forming in a slowly expanding (100 – 200 km s^{-1}) circumstellar shell, while the bulk of the Ca II NIR triplet and $H\alpha$ emission comes from faster-moving ejected material.

4.4 Constraining the CSM using the +30 d spectrum

The GTC+OSIRIS spectrum at phase +30.2 d (resolution $v_{\text{FWHM}} \approx 160$ km s^{-1}) reveals that the $H\alpha$ line exhibits a more complex profile than what appears from the inspection of other lower resolution spectra shown in Fig. 12. Note, however, that some bumps likely due to blends of Fe II lines (see Fig. 13), are visible slightly above the flux level of the continuum. Most of these features were observed in SN 2008S (Botticella et al. 2009, see their fig. 9) and classified as Fe II lines (multiplets 40, 74, 92). Hence, they are not weak $H\alpha$ components. In particular the bump peaking (in the Doppler-corrected spectrum) at 6518 Å ($v = 2100$ km s^{-1}) is inconsistent with a blue-shifted tail of $H\alpha$, whose profile is, in fact, well-fitted by a single Lorentzian emission component with $v_{\text{FWHM}} \sim 820$ km s^{-1} . Superimposed on the Lorentzian component, we note a very narrow P-Cygni feature (see Fig. 13), which is marginally resolved, and that can be attributed to a slow-moving circumstellar shell. From the position of the narrow P-Cygni absorption, we estimate an upper

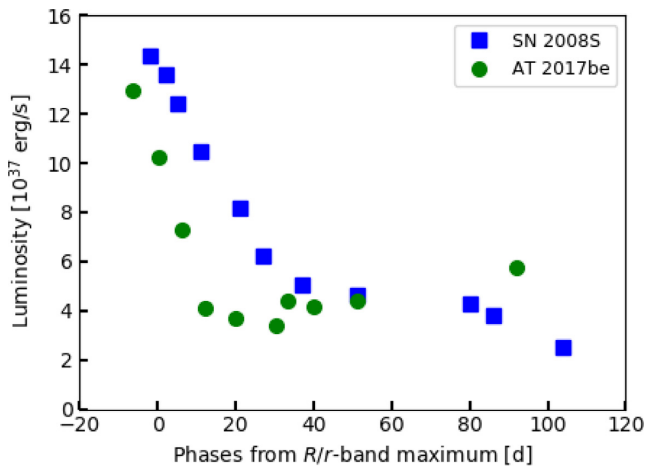


Figure 11. Comparison between $H\alpha$ luminosity evolution in AT 2017be and SN 2008S. Values for SN 2008S are taken from the literature (Botticella et al. 2009, see their table 10). We conservatively estimate the uncertainties to be around 20 per cent.

limit velocity of $\sim 70 \text{ km s}^{-1}$ for this H-rich shell. We note that this velocity is barely consistent with the v_{FWHM} that we infer for the $[\text{Ca II}]$ doublet lines in the same spectrum. In fact, the average velocity obtained after instrumental resolution correction for the two lines of the doublet is $v_{\text{FWHM}} \sim 115 \pm 8 \text{ km s}^{-1}$ (the error on the line fit is only accounted for here). It likely suggests that the narrow $H\alpha$ P-Cygni and the $[\text{Ca II}]$ doublet are both produced in slow-moving material ($\leq 70\text{--}115 \text{ km s}^{-1}$) expelled by the progenitor star years before the explosion. This modest wind velocity allows us to disfavour an LBV-like outburst as the origin for this mass-loss event, and is more consistent with CSM produced by a RSG or a super-AGB star.

Berger et al. (2009) investigated the environmental properties of the archetypal ILRT NGC 300 OT2008-1 through high-resolution spectroscopy. They discussed the complex profile of prominent spectral features such as the Balmer lines, the $[\text{Ca II}]$ doublet and Ca II NIR triplet (visible in emission) and $\text{Ca II H\&K}$ (in absorption). These high-quality spectra of NGC 300 OT2008-1 showed also weaker features of He I , $[\text{O I}] \lambda 6300$, $\lambda 6364$, $\text{O I } \lambda 7774$, and $\lambda 8446$, along with Fe II multiplets. The $H\alpha$ emission line was symmetric, but accompanied by narrow absorption features superimposed on the main emission, one (weaker) blue-shifted by -130 km s^{-1} , and another more prominent red-shifted by 30 km s^{-1} . A similar narrow P-Cygni absorption is visible in our GTC + OSIRIS spectrum (blue-shifted by -70 km s^{-1}). As in the case of NGC 300 OT2008-1, the $[\text{Ca II}]$ doublet has a strongly asymmetric profile without a blue wing, while the Ca II NIR triplet has a broad absorption feature, supporting the existence of multiple line components. The features of Berger et al.’s high-resolution spectra suggest that NGC 300 OT2008-1 has a complex circumstellar environment, marked by outflows and inflows from different gas components, including the wind from a possible companion star (a Wolf–Rayet star or a blue supergiant).

5 PROGENITOR ANALYSIS

In the attempt to detect (or place stringent detection limits) to the quiescent progenitor star, we searched pre-explosion images of the field of AT 2017be in ground-based and space telescope archives. Unfortunately while the host galaxy is relatively nearby, very few

deep images are available in the archives. The most constraining image is a *HST* + *NICMOS* frame taken on 2002 October 14, using the NIC3 camera and the *F160W* filter ($\sim H$ band). A narrow-band NIR image taken at the same epoch does not place useful magnitude limits to the progenitor.

To identify the position of AT 2017be in the NICMOS data, we aligned the image to good seeing images taken at the NOT. We aligned the NICMOS data to both an *i*-band image taken on 2017 April 21, and a *H*-band image taken on April 8, using 10 sources in the former and 7 in the latter that were also present in the NICMOS frame. Rotation, translation, and a scale factor between each pair of images were fitted for, and the resulting transformation had an RMS scatter of 69 and 53 mas when using either *i* band or *H* band, respectively, as our reference image. In both cases, we obtained an almost identical position for AT 2017be on the NICMOS image (see the left-hand panel of Fig. 14).

To determine a limiting magnitude for the progenitor of AT 2017be, we performed artificial star tests on the NICMOS data (see the middle and right-hand panels of Fig. 14). A model point spread function was created from sources in the field, and then scaled in flux and inserted at the SN position. The 3σ limiting magnitude was taken to be the value at which the source was recovered with a photometric uncertainty of 0.2 mag. We determine a flux limit for the progenitor with this technique of $F < 2.04 \times 10^{-18} \text{ erg s}^{-1} \text{ cm}^{-2} \text{ \AA}^{-1}$, or a limiting magnitude in the ABmag system of $F160W > 20.8 \text{ mag}$.

We used the same techniques to determine limiting magnitudes for the Spitzer IRAC images (2004-04-01.5 UT, see Fig. 15) as were used for the NICMOS data. Limits of $M_{3.6\mu} > 15.80 \text{ mag}$ and $M_{4.5\mu} > 15.82 \text{ mag}$ (Vega mag) were found in IRAC CH1 and CH2, respectively.

The limits we can place on the progenitor of AT 2017be are unfortunately not particularly constraining. The implied absolute magnitude of the progenitor given a distance modulus of $\mu = 29.47 \text{ mag}$ is fainter than $H = -10.1 \text{ mag}$ (converting to Vegamag). For comparison, the progenitor of SN 2008S, which was detected in the MIR with *Spitzer*, had a *K*-band limit of $> -10.8 \text{ mag}$ (Botticella et al. 2009).

Adams et al. (2018) performed an independent analysis on pre-explosion archive images, although they did not describe in depth the methods used to infer their limits. While the *H*-band limit inferred by Adams et al. from the *HST* + *NICMOS* data, is similar to ours, their IRAC limits are significantly deeper (Adams et al. 2018, almost 1.5 mag discrepancy for CH1 and CH2, respectively; see their table 4). We note that the putative progenitor was blended with an IR-luminous extended source, and this can possibly explain the discrepancy. However, as it is unclear from Adams et al. exactly what data were used in their analysis, and how their limiting magnitudes were determined, we cannot investigate this discrepancy further.

6 DISCUSSION AND SUMMARY

6.1 Comparison with other ILRTs

In recent years, the number of discovered ILRTs has grown significantly, and includes such events as SN 2008S (Botticella et al. 2009), NGC 300 OT2008-1 (Humphreys et al. 2011), M85 OT2006-1 (Kulkarni et al. 2007), and PTF10fq (Kasliwal et al. 2011). The latest addition is AT 2017be, whose absolute magnitude at peak lies in the same range ($-10 \text{ mag} > M_V > -15 \text{ mag}$) as several putative classes of ILOTs. More specifically, AT 2017be exhibits the characteristic properties an ILRT, having a relatively slow spectro-

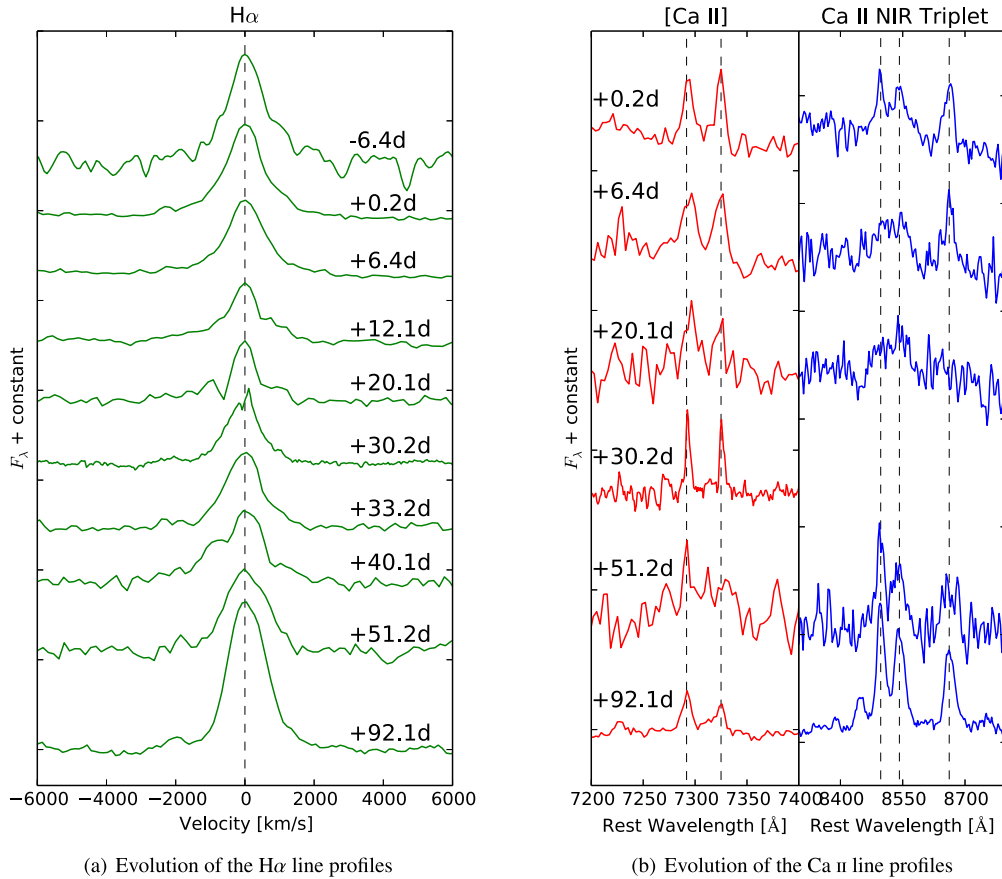


Figure 12. Evolution of the H α (left-hand panel), [Ca II] doublet, and Ca II NIR triplet (right-hand panel) line profiles. The vertical black dashed lines mark the rest-wavelength position of these emission lines.

photometric evolution, along with a light curve that resembles those of Type II-P or Type II-L SNe. In addition, the spectra of AT 2017be are dominated by prominent and narrow H emission lines ($v_{\text{FWHM}} < 1000 \text{ km s}^{-1}$), weak Fe II features, Ca II H&K, and the Ca II NIR triplet, similar to other ILRTs. Most importantly, the [Ca II] doublet (see Fig. 10, right), which is a typical marker of ILRTs, is visible in all our spectra.

The light curve of AT 2017be is very similar to that of PTF10fq, and shares some similarity with M85 OT2006-1. Their light curves show a sort of short-duration plateau. The colour evolution of AT 2017be in the optical bands matches that of NGC 300 OT2008-1, especially the $B - V$ and $R - I/r - i$ colours. The $J - K$ colour of AT 2017be, SN 2008S, NGC 300 OT2008-1, and M85 OT2006-1 becomes redder with time. The Sloan r -band absolute magnitude at maximum of AT 2017be is $M_r = -11.98 \pm 0.09 \text{ mag}$, and is almost equal to that of PTF10fq. The constraint on the ^{56}Ni mass inferred from the late-phase quasi-bolometric light curve, implies that very little ^{56}Ni was ejected in the explosion ($\sim 8 \times 10^{-4} M_\odot$ or less), which is less than that estimated by Botticella et al. for SN 2008S.

A common characteristic of ILRTs is their MIR luminous pre-explosion progenitors, with the quiescent star being embedded in a dusty CSM (Botticella et al. 2009, see their fig. 17). The best-studied cases indicate moderate-mass ($8\text{--}12 M_\odot$) progenitors, although in some cases the studies point to slightly lower mass stars (e.g. M85 OT2006-1, Ofek et al. 2008; Pastorello et al. 2007; Rau et al. 2007). Unfortunately, we cannot set a stringent constraints on the AT 2017be progenitor star in pre-explosion archive images.

6.2 Plausible scenarios for ILRTs and conclusions

The nature of ILRTs is debated. Several scenarios have been proposed in the attempt to interpret their peculiar observables. For SN 2008S, a scenario invoking a weak ECSN explosion of a super-AGB star embedded in an optical thick dust envelope has been proposed (Prieto et al. 2008; Botticella et al. 2009; Thompson et al. 2009). Alternatively, the LBV-like super-Eddington outburst of a highly obscured, massive progenitor ($\sim 20 M_\odot$) has been suggested by Smith et al. (2009). Another well-studied object is NGC 300 OT2008-1, and which is still consistent with both an ECSN (Prieto et al. 2008; Thompson et al. 2009) or an LBV-like eruption (Smith et al. 2011).

Berger et al. (2009) investigated NGC 300 OT2008-1 with high-resolution spectroscopy, and proposed that the observed outburst was produced by a $\sim 10\text{--}20 M_\odot$, relatively compact progenitor (a blue supergiant or a Wolf-Rayet star), obscured by a dusty cocoon. An alternative explanation involved interaction in a binary system. Bond et al. (2009) proposed NGC 300 OT2008-1 to be an eruption of a dust-enshrouded, IR-luminous, and relatively massive star ($\sim 10\text{--}15 M_\odot$), likely an OH/IR source, which began its evolution on a blue loop toward higher temperatures. Both SN 2008S and NGC 300 OT2008-1 are slightly more luminous than other ILRTs, and have light curves similar to those of Type II-L SNe. From an observational point of view, ILRTs appear to be scaled-down analogues of Type II-P or Type II-L SNe.

Therefore, we now discuss three possible scenarios (stellar mergers, LBV-like eruption, and EC SN explosions) for ILRTs.

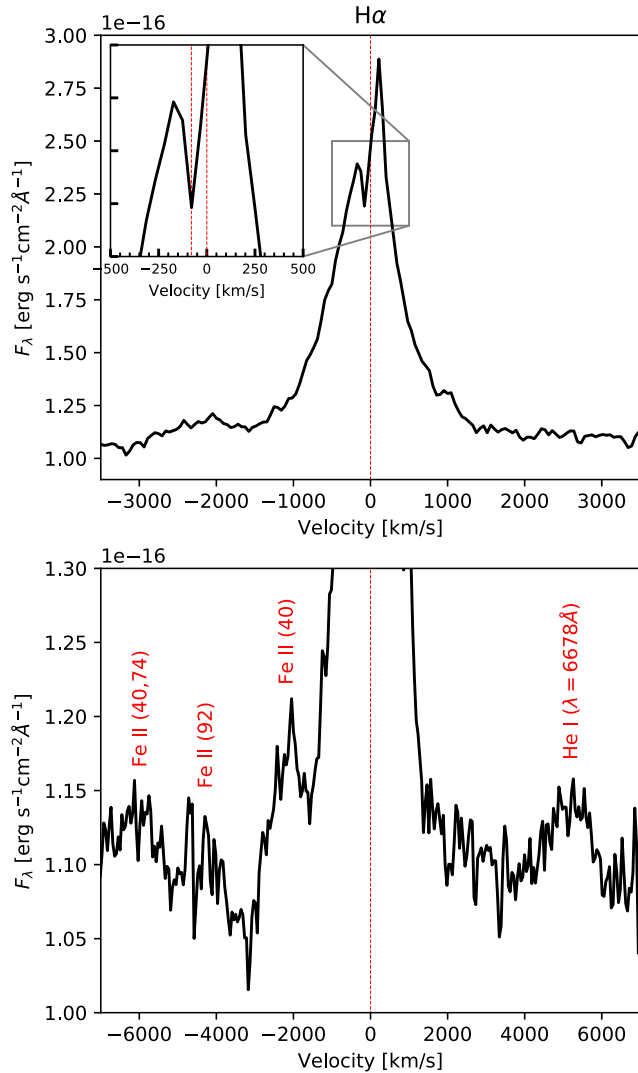


Figure 13. The profile of H α in our highest resolution spectrum (GTC, phase 30.2 d) in velocity space. The vertical red dashed line marks the rest-velocity position of H α . Lower contrast Fe II lines and He I λ 6678 are also identified in the lower panel.

(i) The most extreme outcome of binary interaction is a stellar merger. The discovery of V1309 Sco (Barsukova, Goranskij & Valeev 2017) provided strong support to the idea that luminous red novae (LRNe) likely result from a merging event (see e.g. Mason et al. 2010; Tytenda et al. 2011). The typical observational features of LRNe such as V8383 Mon, V1309 Sco, and V4332 Sgr are multi-peaked light curves, low ejecta velocities (~ 100 – 300 km s⁻¹), no evidence of [Ca II] features, and very late-time spectra dominated by molecular bands (TiO, VO) in the optical domain (Martini et al. 1999; Crause et al. 2003; Rushton et al. 2003; Banerjee & Ashok 2004; Barsukova, Goranskij & Valeev 2013; McCollum et al. 2014; Kamiński et al. 2015; Loebman et al. 2015; Tytenda et al. 2015). All of this makes AT 2017be quite different from LRNe, hence we agree with Adams et al. in ruling out a stellar merger scenario for this transient.

(ii) LBVs are erupting massive stars that may experience severe mass-loss. Multiple outbursts can be observed during a single eruptive phase that may last even decades (like the Great Eruption of η Carinae). During these outburst, no ⁵⁶Ni is expected to be ejected.

In addition, the ejected gas expands at velocities of a few hundreds km s⁻¹, occasionally exceeding 10^4 km s⁻¹ during a giant eruption, which are inconsistent with our spectroscopic observations (e.g. ~ 820 and 730 km s⁻¹ at epoch of +30.2 and +92.1 d, respectively). An LBV outburst scenario for AT 2017be has been recently proposed by Adams et al. (2018). However, we do not find robust arguments to support this, as we did not detect any luminous progenitor in pre-explosion data (in contrast with expectations from a massive LBV). A caveat here is that lower luminosity LBVs may have initial masses down to about $20 M_\odot$, as claimed by Smith et al. (2010, 2016, 2018). In particular, UGC 2773-OT is a moderate-mass LBV of $\sim 20 M_\odot$, sharing some observational properties with ILRTs (in particular, the overall spectral evolution and the detection of the [Ca II] doublet). However, the light curve was totally different, showing a rise phase lasting over a decade, which is a clear signature of a long-lasting eruptive phase. AT 2017be and ILRTs have short-duration light curves, resembling that of some CC SNe, and the best-followed objects (SN 2008S and NGC 300 OT2008-1) show a late-time decline rate consistent with the ⁵⁶Co decay into ⁵⁶Fe, suggesting a SN explosion (see below).

(iii) Pumo et al. (2009) described the expected properties of EC SNe from super-AGB progenitors combining theoretical models with observational evidence. This type of CC SNe are triggered by electron-capture in the core of a progenitor which has evolved to reach the super-AGB phase. The threshold to trigger an EC SN explosion is determined by the competition between core growth and mass-loss (see e.g. Woosley, Heger & Weaver 2002; Herwig 2005). If the core grows to reach the Chandrasekhar limit (i.e. $M_{\text{CH}} \sim 1.37 M_\odot$; Nomoto 1984), electron-capture reactions are triggered, allowing the super-AGB star to explode as faint (⁵⁶Ni-poor) ECSN. In contrast, if envelope mass-loss is large enough, electron-capture reactions fail, and the super-AGB star evolves to become an ONe white dwarf (WD). Although the ECSN mechanism in super-AGB is a plausible explanation, several uncertainties still exist. None the less, as a major support to the ECSN scenario for ILRTs, Adams et al. (2016) obtained late-time *Hubble* and *Spitzer Space Telescope* imaging of SN 2008S and NGC 300 OT2008-1, and detected both transients at a mid-IR luminosity fainter than those of their quiescent progenitors, without a detection in the optical and NIR bands.

Unfortunately, in contrast with SN 2008S and NGC 300 OT2008-1, a tight constraint on the progenitor of AT 2017be does not exist. However, we favour an ECSN scenario on the basis of the following supporting observational arguments:

(i) The light curves of ILRTs resemble sub-luminous counterparts of Type II-P or II-L SNe.

(ii) The expansion velocity of the ejecta inferred from broader resolved H and Ca II lines of AT 2017be is significantly lower than those of canonical CC SNe, but only slightly slower than those observed in some faint Type II SNe (Pastorello et al. 2004; Spiro et al. 2014). The wind velocity expected in the super-AGB stars is low, ~ 100 km s⁻¹ (Weaver & Woosley 1979), or even less. This agrees with the velocities inferred for the slow-moving CSM in the highest resolution spectra of NGC 300 OT2008-1 (Berger et al. 2009), and the GTC spectrum of AT 2017be at +30.2 d.

(iii) AT 2017be was photometrically monitored until very late epochs. Although already below the detection threshold in the other monitoring bands, in the *K* band it was visible up to ~ 350 d from maximum. The late time photometric data allow us to constrain to about $8 \times 10^{-4} M_\odot$ for the ejected ⁵⁶Ni mass. This ⁵⁶Ni limit is over a factor of one lower than the amount estimated for other ILRTs

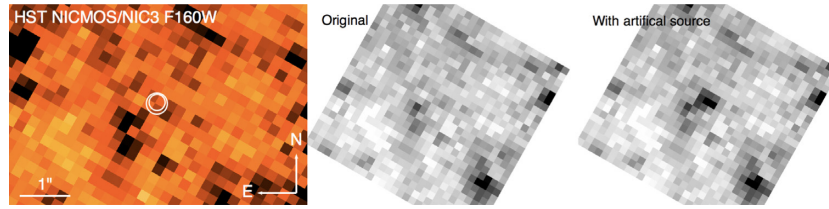


Figure 14. A section of the archival *HST* + NICMOS *F160W* image of the site of AT 2017be. The circles mark the location of AT 2017be as determined from *i*- and *H*-band images of the SN; the radii of the circles correspond to 3x the RMS scatter in the astrometric transformation (left). The *HST* + NICMOS *F160W* image of the site of AT 2017be (middle), and the same region with an artificial source at 3 σ significance (right).

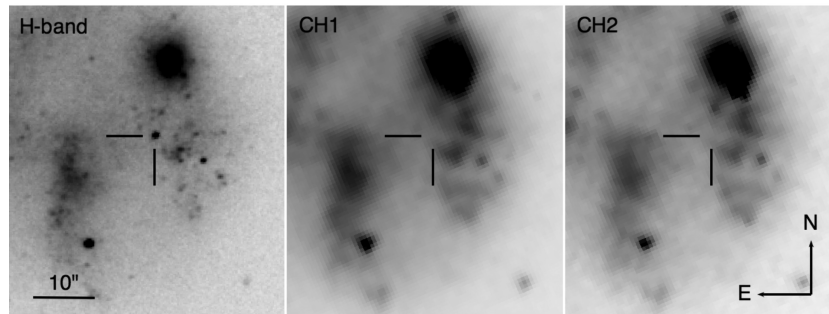


Figure 15. *H*-band post-explosion and *Spitzer* IRAC CH1 (3.6 μ m) and CH2 (4.5 μ m) pre-explosion images of the site of AT 2017be. The position of the transient is indicated with tick marks in all panels.

($\lesssim 10^{-3} M_{\odot}$). The above limit is still consistent with the predictions for an EC SN explosion (Botticella et al. 2009; Pumo et al. 2009). However, on the other hand, we cannot exclude a totally ^{56}Ni -free outburst.

(iv) For a few ILRTs, the progenitor stars were recovered in pre-explosion archive images, and in all cases they were moderate mass stars (in the mass range $\sim 8\text{--}12 M_{\odot}$) embedded in a dusty environment, consistent with super-AGB star scenario.

(v) Adams et al. (2016) reported the mid-IR fading of some ILRTs at very late phases to be more than 15 times fainter than the luminosity of the quiescent progenitor detection. This is a crucial argument to support a terminal SN explosion of ILRT progenitor stars. A deep IR limit with James Webb Space Telescope (JWST) can determine if there is a surviving star following AT 2017be.

It is evident that high-resolution spectroscopy at early phases, the availability of deep pre-discovery images in the optical and IR archives (ideally with high spatial resolution), and deep late-time observations in the IR domain are crucial for better constraining the explosion scenario of AT 2017be and other ILRTs. Future surveys in the optical bands (e.g. with the Large Synoptic Survey Telescope; LSST) will discover a larger number of new ILRTs, enabling us to unequivocally unveil the nature of their stellar progenitors and their explosion mechanism.

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(iii) The Gran Telescopio Canarias (GTC) operated on the island of La Palma at the Spanish Observatorio del Roque de los Muchachos of the Instituto de Astrofísica de Canarias.

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REFERENCES

- Adams S. M., Kochanek C. S., Prieto J. L., Dai X., Shappee B. J., Stanek K. Z., 2016, *MNRAS*, 460, 1645
- Adams S. M. et al., 2018, *PASP*, 130, 034202
- Banerjee D. P. K., Ashok N. M., 2004, *ApJ*, 604, L57
- Barsukova E. A., Fabrika S. N., Goranskij V., Valeev A., 2012, *Peremennye Zvezdy*, 32
- Barsukova E. A., Goranskij V. P., Valeev A. F., 2013, *Cent. Eur. Astrophys. Bull.*, 37, 325
- Barsukova E. A., Goranskij V. P., Valeev A. F., 2017, in Balega Yu. Yu., Kudryavtsev D. O., Romanyuk I. I., Yakunin I. A., eds, *ASP Conf. Ser. Vol. 510, Stars: From Collapse to Collapse*. Astron. Soc. Pac., San Francisco, p. 401
- Becker A., 2015, *Astrophysics Source Code Library*, record ascl:1504.004
- Berger E. et al., 2009, *ApJ*, 699, 1850
- Berger E., Marion G. H., Hsiao E., Foley R., Chornock R., 2012, *Astron. Telegram*, 4009, 1
- Bhardwaj A., Kanbur S. M., Macri L. M., Singh H. P., Ngeow C. C., Wagner-Kaiser R., Sarajedini A., 2016, *AJ*, 151, 88
- Bond H. E., Bedin L. R., Bonanos A. Z., Humphreys R. M., Monard L. A. G. B., Prieto J. L., Walter F. M., 2009, *ApJ*, 695, L154
- Bono G., Caputo F., Marconi M., Musella I., 2010, *ApJ*, 715, 277
- Botticella M. T. et al., 2009, *MNRAS*, 398, 1041
- Bottinelli L., Gouguenheim L., Paturel G., de Vaucouleurs G., 1984, *A&AS*, 56, 381
- Cappellaro E., 2014, *SNOoPy: a package for SN photometry*, <http://sngroup.oapd.inaf.it/snoopy.html>
- Cardelli J. A., Clayton G. C., Mathis J. S., 1989, *ApJ*, 345, 245
- Chevalier R. A., Fransson C., 1994, *ApJ*, 420, 268
- Chonis T. S., Gaskell C. M., 2008, *AJ*, 135, 264
- Crause L. A., Lawson W. A., Kilkenny D., van Wyk F., Marang F., Jones A. F., 2003, *MNRAS*, 341, 785
- de Groot M., 1969, *BAN*, 20, 225
- de Groot M., 1988, *IRAJ*, 18, 163
- de Vaucouleurs G., 1959, *Handbuch der Physik*, 53, 275
- Doherty C. L., Gil-Pons P., Siess L., Lattanzio J. C., Lau H. H. B., 2015, *MNRAS*, 446, 2599
- Doherty C. L., Gil-Pons P., Siess L., Lattanzio J. C., 2017, *PASA*, 34, e056
- Filippenko A. V., 1997, *ARA&A*, 35, 309
- Gall C., Hjorth J., Rosswog S., Tanvir N. R., Levan A. J., 2017, *ApJ*, 849, L19
- Grebel E. K., 2001, *Astrophys. Space Sci. Suppl.*, 277, 231
- Hashimoto M., Iwamoto K., Nomoto K., 1993, *ApJ*, 414, L105
- Herwig F., 2005, *ARA&A*, 43, 435
- Hildebrand R. H., 1983, *QJRAS*, 24, 267
- Hosseinzadeh G., McCully C., Arcavi I., Howell D. A., Valenti S., 2017, *Transient Name Server Classification Report*, 43, 1
- Hubble E., 1926, *Contributions from the Mount Wilson Observatory / Carnegie Institution of Washington*, 324, 1
- Humphreys R. M., Davidson K., 1994, *PASP*, 106, 1025
- Humphreys R. M., Bond H. E., Bedin L. R., Bonanos A. Z., Davidson K., Berto Monard L. A. G., Prieto J. L., Walter F. M., 2011, *ApJ*, 743, 118
- Kamiński T., Mason E., Tyndla R., Schmidt M. R., 2015, *A&A*, 580, A34
- Kankare E. et al., 2012, *MNRAS*, 424, 855
- Kashi A., Frankowski A., Soker N., 2010, *ApJ*, 709, L11
- Kasliwal M. M., 2012, *PASA*, 29, 482
- Kasliwal M. M. et al., 2011, *ApJ*, 730, 134
- Kitaura F. S., Janka H.-T., Hillebrandt W., 2006, *A&A*, 450, 345
- Kochanek C. S., 2011, *ApJ*, 741, 37
- Kulkarni S. R. et al., 2007, *Nature*, 447, 458
- Landolt A. U., 1992, *AJ*, 104, 340
- Loebman S. R. et al., 2015, *AJ*, 149, 17
- Makarov D., Prugniel P., Terekhova N., Courtois H., Vauglin I., 2014, *A&A*, 570, A13
- Mallik S. V., 1997, *A&AS*, 124, 359
- Martini P., Wagner R. M., Tomaney A., Rich R. M., della Valle M., Hauschildt P. H., 1999, *AJ*, 118, 1034
- Mason E., Diaz M., Williams R. E., Preston G., Bensby T., 2010, *A&A*, 516, A108
- Mauerhan J. C. et al., 2013, *MNRAS*, 431, 2599
- McCollum B. et al., 2014, *AJ*, 147, 11
- Miyaji S., Nomoto K., 1987, *ApJ*, 318, 307
- Miyaji S., Nomoto K., Yokoi K., Sugimoto D., 1980, *PASJ*, 32, 303
- Moriya T. J., Eldridge J. J., 2016, *MNRAS*, 461, 2155
- Nomoto K., 1984, *ApJ*, 277, 791
- Nomoto K., 1987, *ApJ*, 322, 206
- Ofek E. O. et al., 2008, *ApJ*, 674, 447
- Oke J. B., Gunn J. E., 1983, *ApJ*, 266, 713
- Osterbrock D. E., Bochkarev N. G., 1989, *SvA*, 33, 694
- Pastorello A. et al., 2004, *MNRAS*, 347, 74
- Pastorello A. et al., 2007, *Nature*, 449, 1
- Patat F., Maund J. R., Benetti S., Botticella M. T., Cappellaro E., Harutyunyan A., Turatto M., 2010, *A&A*, 510, A108
- Poelarends A. J. T., Herwig F., Langer N., Heger A., 2008, *ApJ*, 675, 614
- Prieto J. L. et al., 2008, *ApJ*, 681, L9
- Prieto J. L., Sellgren K., Thompson T. A., Kochanek C. S., 2009, *ApJ*, 705, 1425
- Puckett T., Gauthier S., 2002, *IAU Circ.*, 7863, 1
- Pumo M. L. et al., 2009, *ApJ*, 705, L138
- Rau A., Kulkarni S. R., Ofek E. O., Yan L., 2007, *ApJ*, 659, 1536
- Rouleau F., Martin P. G., 1991, *ApJ*, 377, 526
- Rushton M. T. et al., 2003, *A&A*, 412, 767
- Schlaflly E. F., Finkbeiner D. P., 2011, *ApJ*, 737, 103
- Schlegel E. M., 1990, *MNRAS*, 244, 269
- Sharina M. E., Karachentsev I. D., Tikhonov N. A., 1999, *Astron. Lett.*, 25, 322
- Siess L., 2007, *A&A*, 476, 893
- Smith N., 2013, *MNRAS*, 434, 102
- Smith N., Vink J. S., de Koter A., 2004, *ApJ*, 615, 475
- Smith N. et al., 2009, *ApJ*, 697, L49
- Smith N. et al., 2010, *AJ*, 139, 1451
- Smith N., Li W., Silverman J. M., Ganeshalingam M., Filippenko A. V., 2011, *MNRAS*, 415, 773
- Smith N., Andrews J. E., Mauerhan J. C., Zheng W., Filippenko A. V., Graham M. L., Milne P., 2016, *MNRAS*, 455, 3546
- Smith N. et al., 2018, preprint ([arXiv:1805.03298](https://arxiv.org/abs/1805.03298))
- Spiro S. et al., 2014, *MNRAS*, 439, 2873
- Springob C. M., Masters K. L., Haynes M. P., Giovanelli R., Marinoni C., 2009, *ApJS*, 182, 474
- Stephens H. F., Zheng W., Filippenko A. V., 2017, *Transient Name Server Discovery Report*, p. 33
- Stetson P. B., 1987, *PASP*, 99, 191
- Sugerman B., Bengt A., Cosgrove A., Snyder K., 2016, *American Astronomical Society Meeting Abstracts*, 227, p. 237.06

- Szczygieł D. M., Prieto J. L., Kochanek C. S., Stanek K. Z., Thompson T. A., Beacom J. F., Garnavich P. M., Woodward C. E., 2012, *ApJ*, 750, 77
- Szczygieł D. M., Kochanek C. S., Dai X., 2012, *ApJ*, 760, 20
- Tartaglia L. et al., 2016, *ApJ*, 823, L23
- Thompson T. A., Prieto J. L., Stanek K. Z., Kistler M. D., Beacom J. F., Kochanek C. S., 2009, *ApJ*, 705, 1364
- Tominaga N., Blinnikov S. I., Nomoto K., 2013, *ApJ*, 771, L12
- Tully R. B., 1988, NBGC.C....0000T
- Tully R. B. et al., 2013, *AJ*, 146, 86
- Turatto M., Cappellaro E., Danziger I. J., Benetti S., Gouiffes C., della Valle M., 1993, *MNRAS*, 262, 128
- Turatto M., Benetti S., Cappellaro E., 2003, in Hillebrandt W., Leibundgut B. eds, *From Twilight to Highlight: The Physics of Supernovae*, Springer-Verlag, New York, p. 200
- Tylenda R., 2004, *A&A*, 414, 223
- Tylenda R., Kamiński T., 2012, *A&A*, 548, A23
- Tylenda R. et al., 2011, *A&A*, 528, A114
- Tylenda R., Górny S. K., Kamiński T., Schmidt M., 2015, *A&A*, 578, A75
- Van Dyk S. D., Peng C. Y., King J. Y., Filippenko A. V., Treffers R. R., Li W., Richmond M. W., 2000, *PASP*, 112, 1532
- Wanajo S., Janka H.-T., Müller B., 2011, *ApJ*, 726, L15
- Weaver T. A., Woosley S. E., 1979, *BAAS*, 11, 724
- Welch D. L., Clayton G. C., Sugerman B., 2008, *Central Bureau Electron. Telegrams*, 1330, 1
- Wesson R., Fabbri J., Barlow M., Meixner M., 2008, *Central Bureau Electron. Telegrams*, 1381, 1
- Wesson R. et al., 2010, *MNRAS*, 403, 47
- Whitelock P. A. et al., 1988, *MNRAS*, 234, 5P
- Wolf B., 1989, *A&A*, 217, 87
- Woosley S. E., Heger A., Weaver T. A., 2002, *Rev. Mod. Phys.*, 74, 1015
- Wright E. L., 2006, *PASP*, 118, 1711
- Yee J. C., Eastman J. D., Prieto J. L., 2008, *Central Bureau Electron. Telegrams*, 1340, 1

APPENDIX A: LIGHT CURVES OF AT 2017BE

Table A1. Optical (*BVgriz*) photometric data for AT 2017be. Early unfiltered observations are also reported.

Date	MJD	Phase ^a	<i>B</i> (err)	<i>V</i> (err)	<i>g</i> (err)	<i>r</i> (err)	<i>i</i> (err)	<i>z</i> (err)	Instrument
20161229	57751.48	−18.34	—	—	—	>18.7 ^b	—	—	LOSS ^e
20170106	57759.51	−10.31	—	—	—	18.71 ^b	—	—	LOSS ^e
20170107	57760.12	−9.70	—	—	—	18.30	—	—	iPTF ^f
20170107	57760.16	−9.66	>17.8	—	—	—	—	—	f05
20170107	57760.43	−9.39	—	—	—	18.347	—	—	iPTF ^g
20170108	57761.36	−8.47	18.982 (0.234)	18.442 (0.122)	18.645 (0.216)	18.135 (0.144)	18.010 (0.065)	—	f05
20170108	57761.42	−8.40	18.970 (0.097)	18.454 (0.061)	18.660 (0.118)	18.140 (0.060)	18.011 (0.047)	—	f05
20170110	57763.34	−6.48	18.286 (0.278)	17.986 (0.084)	18.171 (0.200)	17.818 (0.091)	17.742 (0.083)	—	f05
20170110	57763.42	−6.40	>17.8	>17.141	18.193 (0.217)	17.821 (0.199)	17.744 (0.115)	—	f05
20170116	57769.20	−0.63	18.449 (0.088)	18.142 (0.073)	18.276 (0.123)	17.756 (0.065)	17.748 (0.088)	—	f05
20170116	57769.98	+0.16	18.496 (0.031)	18.167 (0.034)	18.326 (0.027)	17.743 (0.026)	17.792 (0.031)	17.783 (0.068)	ALFOSC
20170120	57773.27	+3.45	18.518 (0.075)	18.046 (0.055)	18.209 (0.059)	17.835 (0.050)	17.777 (0.048)	—	f05
20170120	57773.31	+3.49	18.538 (0.083)	18.067 (0.046)	18.319 (0.048)	17.908 (0.040)	17.766 (0.053)	—	f05
20170123	57776.15	+6.33	18.733 (0.023)	18.090 (0.047)	—	—	17.905 (0.139)	—	ALFOSC
20170123 ^c	57776.95	+7.13	18.870 (0.153)	18.112 (0.090)	18.409 (0.096)	17.914 (0.083)	17.853 (0.088)	17.967 (0.075)	AFOSC
20170124	57777.28	+7.46	18.699 (0.077)	18.160 (0.057)	18.427 (0.051)	18.028 (0.042)	17.897 (0.045)	—	f05
20170127	57780.33	+10.51	18.852 (0.087)	18.229 (0.057)	18.634 (0.061)	18.121 (0.045)	18.037 (0.052)	17.938 (0.129)	f05
20170128	57781.87	+12.05	18.989 (0.074)	18.330 (0.043)	18.665 (0.045)	18.144 (0.041)	18.028 (0.051)	17.954 (0.057)	ALFOSC
20170131	57784.27	+14.45	19.088 (0.102)	18.358 (0.065)	18.837 (0.066)	18.207 (0.056)	18.078 (0.052)	17.927 (0.106)	f05
20170204	57788.25	+18.43	19.194 (0.123)	18.427 (0.066)	18.750 (0.071)	18.208 (0.058)	18.087 (0.068)	17.996 (0.122)	f05
20170205 ^{c, d}	57789.95	+20.13	19.297 (0.043)	18.521 (0.031)	18.809 (0.043)	18.229 (0.038)	18.058 (0.031)	17.966 (0.075)	ALFOSC
20170209	57793.37	+23.55	19.427 (0.207)	18.435 (0.089)	18.881 (0.069)	18.243 (0.067)	18.086 (0.071)	17.974 (0.150)	f05
20170217	57801.22	+31.40	19.540 (0.280)	18.668 (0.106)	19.084 (0.166)	18.368 (0.058)	18.156 (0.041)	17.988 (0.084)	f05
20170219	57803.07	+33.25	19.569 (0.046)	18.627 (0.099)	19.175 (0.032)	18.451 (0.021)	18.185 (0.036)	18.058 (0.040)	ALFOSC
20170221	57805.26	+35.44	19.598 (0.369)	18.774 (0.245)	19.214 (0.399)	18.493 (0.062)	18.194 (0.038)	18.052 (0.064)	f05
20170225	57809.14	+39.32	19.772 (0.202)	18.855 (0.196)	19.275 (0.192)	18.511 (0.121)	18.263 (0.064)	18.073 (0.108)	f05
20170225	57809.91	+40.09	19.750 (0.047)	18.833 (0.032)	19.288 (0.031)	18.513 (0.027)	18.269 (0.016)	18.109 (0.042)	ALFOSC
20170302	57814.28	+44.46	—	18.958 (0.193)	19.364 (0.107)	18.721 (0.189)	18.389 (0.133)	—	f05
20170307	57819.20	+49.38	19.962 (0.132)	18.946 (0.060)	19.263 (0.068)	18.570 (0.045)	18.372 (0.043)	18.112 (0.085)	f05
20170309	57821.02	+51.20	19.950 (0.076)	18.923 (0.047)	19.317 (0.045)	18.653 (0.028)	18.365 (0.020)	18.153 (0.040)	ALFOSC
20170311	57823.26	+53.44	20.019 (0.222)	18.977 (0.172)	19.359 (0.132)	18.621 (0.136)	18.431 (0.134)	17.992 (0.307)	f05
20170315	57827.22	+57.40	20.024 (0.190)	18.991 (0.079)	19.387 (0.058)	18.694 (0.065)	18.450 (0.053)	18.132 (0.112)	f05
20170319	57831.19	+61.37	20.052 (0.180)	19.028 (0.085)	19.406 (0.084)	18.710 (0.056)	18.488 (0.071)	18.157 (0.083)	f05
20170327	57839.14	+69.32	20.085 (0.092)	19.129 (0.055)	19.595 (0.055)	18.809 (0.085)	18.551 (0.104)	18.360 (0.136)	f05
20170328	57840.87	+71.05	20.119 (0.093)	19.172 (0.098)	19.604 (0.176)	18.879 (0.059)	18.623 (0.046)	18.368 (0.082)	AFOSC
20170403	57846.10	+76.28	20.122 (0.191)	19.329 (0.119)	19.816 (0.077)	19.019 (0.057)	18.785 (0.051)	18.336 (0.121)	f05
20170407	57850.94	+81.12	20.216 (0.079)	19.416 (0.042)	19.993 (0.084)	19.147 (0.052)	18.831 (0.040)	18.500 (0.038)	ALFOSC
20170409	57852.22	+82.40	20.302 (0.161)	19.414 (0.127)	—	19.162 (0.085)	18.817 (0.072)	18.510 (0.112)	f05
20170411	57854.21	+84.39	>19.9	19.438 (0.295)	19.918 (0.312)	>19.2	—	—	f05
20170414	57857.14	+87.32	20.349 (0.112)	19.634 (0.093)	—	—	—	—	f05
20170417	57860.11	+90.29	20.432 (0.548)	19.726 (0.150)	20.259 (0.065)	19.262 (0.046)	18.979 (0.043)	18.590 (0.095)	f05
20170421	57864.89	+95.07	20.860 (0.118)	19.835 (0.113)	20.287 (0.051)	19.411 (0.059)	19.131 (0.078)	18.756 (0.064)	ALFOSC
20170502	57875.87	+106.05	21.140 (0.090)	20.235 (0.056)	20.567 (0.067)	19.599 (0.045)	19.318 (0.045)	18.886 (0.054)	ALFOSC
20170505	57878.12	+108.30	>20.9	20.387 (0.470)	20.704 (0.202)	19.719 (0.109)	19.344 (0.105)	18.861 (0.149)	f05
20170515	57888.13	+118.31	21.303 (0.469)	20.529 (0.432)	>20.8	19.824 (0.106)	19.480 (0.092)	18.923 (0.154)	f05
20170516	57889.89	+120.07	21.420 (0.119)	20.606 (0.087)	—	—	—	—	ALFOSC
20170518	57891.15	+121.33	—	—	>20.3	19.903 (0.047)	19.589 (0.063)	19.279 (0.120)	ALFOSC
20170530	57903.89	+134.07	—	—	>21.2	20.544 (0.093)	20.145 (0.075)	19.767 (0.204)	ALFOSC
20170618	57922.88	+153.06	—	—	—	21.172 (0.427)	20.548 (0.163)	20.356 (0.137)	ALFOSC
20170619	57923.89	+154.07	—	21.395 (0.167)	22.423 (0.246)	—	—	—	ALFOSC
20170626	57930.89	+161.07	—	—	—	21.701 (0.249)	21.409 (0.201)	20.982 (0.185)	ALFOSC
20170922	58018.19	+248.37	—	—	—	>22.1	—	—	ALFOSC
20170928	58024.18	+254.36	—	—	—	—	>22.4	—	ALFOSC
20171229	58116.17	+346.35	—	—	>22.6	>22.6	>22.5	>21.2	OSIRIS

^aPhases are relative to *r*-band maximum light, MJD = 57769.8 ± 0.1.^bThe original unfiltered magnitudes reported on the Vegamag system have been rescaled to ABmags.^cThe *u*-band magnitudes for these two epochs are reported as 19.436 (0.098) and 21.160 (0.270) mag, respectively.^dThe *R*- and *I*-band magnitudes for this epoch are reported as 18.037 (0.040) and 17.583 (0.035), respectively.^eLick Observatory Supernova Search (LOSS) observation (using the Katzman Automatic Imaging Telescope (KAIT; <http://w.astro.berkeley.edu/bait/kait.html>), Lick Observatory, California, USA) obtained through the Transient Name Server (TNS; <https://wis-tns.weizmann.ac.il/object/2017be>).^fIntermediate Palomar Transient Factory (iPTF; <https://www.ptf.caltech.edu/iptf>) observation using Sloan-band filters, see Hosseinzadeh et al. (2017).^giPTF observation (R-PTF filter, ABmag) obtained via the TNS.

Table A2. *JHK* photometric data for AT2017be.

Date	MJD	Phase ^a	<i>J</i> (err)	<i>H</i> (err)	<i>K</i> (err)	Instrument
20170207	57791.09	+21.27	16.921 (0.303)	16.102 (0.269)	15.069 (0.280)	NOTCam
20170307	57819.87	+50.05	17.078 (0.258)	16.423 (0.227)	15.501 (0.253)	NOTCam
20170408	57851.95	+82.13	17.430 (0.280)	16.508 (0.376)	15.730 (0.319)	NOTCam
20170524	57897.90	+128.08	18.049 (0.140)	17.265 (0.089)	16.224 (0.294)	NOTCam
20170602	57906.91	+137.09	–	–	16.309 (0.224)	NOTCam
20170620	57924.88	+155.06	–	–	16.498 (0.211)	NOTCam
20170703	57937.89	+168.07	–	–	16.965 (0.218)	NOTCam
20170930	58026.23	+256.41	19.926 (0.261)	18.881 (0.307)	17.524 (0.228)	NOTCam
20171017	58043.14	+273.32	19.961 (0.261)	19.346 (0.395)	17.768 (0.291)	NOTCam
20180104	58122.12	+352.30	>20.2	>19.5	18.525 (0.215)	NOTCam

^aPhases are relative to *r*-band maximum light, MJD = 57769.8 ± 0.1.

NOTCam: the 2.56-m Nordic Optical Telescope (NOT) with NOTCam.

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