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The *Gaia*-ESO Survey: membership probabilities for stars in 32 open clusters from 3D kinematics

R. J. Jackson,¹★ R. D. Jeffries,¹★ N. J. Wright,¹ S. Randich,² G. Sacco,² E. Pancino,² T. Cantat-Gaudin,³ G. Gilmore,⁴ A. Vallenari,⁵ T. Bensby,⁶ A. Bayo,^{7,10} M. T. Costado,⁸ E. Franciosini,² A. Gonneau,⁴ A. Hourihane,⁴ J. Lewis,⁴ L. Monaco,⁹ L. Morbidelli² and C. Worley⁴

¹*Astrophysics Group, Keele University, Keele, Staffordshire ST5 5BG, UK*

²*INAF – Osservatorio Astrofisico di Arcetri, Largo E. Fermi, 5, I-50125 Firenze, Italy*

³*Institut de Ciències del Cosmos, Universitat de Barcelona (IEEC-UB), Martí i Franquès 1, E-08028 Barcelona, Spain*

⁴*Institute of Astronomy, Cambridge University, Madingley Road, Cambridge CB3 0H, UK*

⁵*INAF – Osservatorio Astronomico di Padova, Vicolo dell'Osservatorio 5, I-35122 Padova, Italy*

⁶*Lund Observatory, Department of Astronomy and Theoretical Physics, Box 43, SE-221 00 Lund, Sweden*

⁷*Instituto de Física y Astronomía, Facultad de Ciencias, Universidad de Valparaíso, Av. Gran Bretaña 1111, 5030 Casilla, Valparaíso, Chile*

⁸*Departamento de Didáctica, Universidad de Cádiz, E-11519 Puerto Real, Cádiz, Spain*

⁹*Departamento de Ciencias Físicas, Universidad Andres Bello, Fernandez Concha 700, Las Condes, Santiago, Chile*

¹⁰*Núcleo Milenio Formación Planetaria - NPF, Universidad de Valparaíso Avenida Errázuriz 1834, Valparaíso, Chile*

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ABSTRACT

The *Gaia*-ESO Survey (GES) observed many open clusters as part of its programme to spectroscopically characterize the various Milky Way populations. GES spectroscopy and *Gaia* astrometry from its second data release are used here to assign membership probabilities to targets towards 32 open clusters with ages from 1 to 3800 Myr, based on maximum likelihood modelling of the 3D kinematics of the cluster and field populations. From a parent catalogue of 14 398 individual targets, 5032 stars with uniformly determined 3D velocities, T_{eff} , $\log g$, and chemistry are assigned cluster membership with probability >0.9 , and with an average probability of 0.991. The robustness of the membership probabilities is demonstrated using independent membership criteria (lithium and parallax) in two of the youngest clusters. The addition of radial velocities improves membership discrimination over proper motion selection alone, especially in more distant clusters. The kinematically selected nature of the membership lists, independent of photometry and chemistry, makes the catalogue a valuable resource for testing stellar evolutionary models and investigating the time evolution of various parameters.

Key words: stars: evolution – stars: pre-main-sequence – open clusters and associations: general.

1 INTRODUCTION

Open star clusters and associations play a fundamental role in our understanding of stellar evolution, in the testing of stellar models and in anchoring the age scale of stars. They offer samples of stars at a range of masses and evolutionary stages, but with very similar ages and (initial) compositions. Stars in a single cluster can be used to test the mass dependence predicted by models or to use models to estimate masses; whilst the comparison of clusters across a range of ages can be used to test the time- or chemical composition-dependence predicted by models and to explore phenomena empirically that are poorly understood from a physical point of view. A non-exhaustive list of examples would include testing how models predict the positions of stars in the Hertzsprung–Russell diagram; estimating the stellar initial mass function and identifying substellar objects; calibrating white dwarf cooling models; following the spin-

down of stars and calibrating gyrochronology; and investigating the depletion of light elements in stellar interiors.

The *Gaia*-ESO Survey (GES) is a large public survey programme executed on the 8-m UT2-*Kueyen* telescope of the Very Large Telescope facility. The survey rationale, methodology, and calibration strategy are detailed in Gilmore et al. (2012), Randich, Gilmore & Gaia-ESO Consortium (2013), and Pancino et al. (2017). Over the course of about 6 yr, beginning on 2011 December 31, medium ($R \sim 17\,000$) and high ($R \sim 47\,000$) resolution multifibre spectroscopy were obtained, using FLAMES (Fiber Large Array Multi-Element Spectrograph; Pasquini et al. 2002) combined with the GIRAFFE and UVES (Ultraviolet and Visual Echelle Spectrograph; Dekker et al. 2000) spectrographs of about 10^5 and 10^4 stars in our Galaxy. The survey had the aim of understanding, through measurements of kinematics and chemical abundances, the formation and evolution of all the components of our Galaxy, and included a significant proportion (~ 40 per cent) of time devoted to studying star clusters and associations at a range of ages. At the time of writing, GES has internally delivered radial velocities and chemical

* E-mail: r.j.jackson@keele.ac.uk (RJJ); r.d.jeffries@keele.ac.uk (RDJ)

abundances for 32 clusters as part of the internal Data Release 5 (hereafter GESiDR5).

A pre-requisite for most studies using star clusters is to accurately assess which stars are actually members, in the presence of contaminating sources. Many different methods can be used to filter stars – positions, kinematics, spectroscopic parameters, abundances, photometry, but it is important that the filtering criteria are understood and do not bias any subsequent investigation of cluster properties by using those same properties to select cluster members. The addition of *Gaia* astrometric data (Gaia Collaboration et al. 2016), in the form of its first and second data releases (*Gaia* DR1 and *Gaia* DR2, Gaia Collaboration et al. 2016, 2018b), has dramatically enhanced our capability to separate cluster members from unrelated field stars using proper motion and parallax (Cantat-Gaudin et al. 2018; Gaia Collaboration et al. 2018a).

The work presented here follows on from Randich et al. (2018), where *Gaia* astrometry from *Gaia* DR1 was used in conjunction with spectroscopic parameters from GES to define samples of high-probability cluster members for eight open clusters. Here, we describe a closely related methodology that uses temperatures, gravities, and radial velocities from GESiDR5, together with astrometry from *Gaia* DR2, to define membership probabilities for sources in 32 GES clusters based on their 3D kinematics. The inclusion of the third dimension of radial velocity from GES in stars as faint as $V \sim 19$ improves our ability to separate cluster members from contaminants over studies using proper motion alone, especially in the more distant clusters. The aim is to provide rigorously determined membership lists, with quantitative membership probabilities, that can be used for a host of follow-up investigations.

2 POTENTIAL CLUSTER MEMBERS

2.1 Source data

GES Data for 32 open clusters was taken from the GESiDR5 analysis iteration in the GES archive at the Wide Field Astronomy Unit of Edinburgh University.¹ Table 1 shows a list of cluster names together with initial values of age, distance modulus and reddening reported in the literature. Also shown are the number of targets in each cluster that were observed using the GIRAFFE 665 nm filter (HR15n) and/or the UVES 520 or 580 nm filters. A summary of the target selection strategy for the GES clusters can be found in Randich et al. (2018) and summaries of the spectroscopic data and analyses can be found in Sacco et al. (2014), Damiani et al. (2014), Jeffries et al. (2014), Smiljanic et al. (2014), and Jackson et al. (2015).

The GES data are not complete in any sense. Only a (variable) fraction of members will have been observed in each cluster, either because of the inability to cover the full spatial extent of the cluster (particularly those that are nearby and of large angular extent), the inability to assign fibers to all the targets or in a few cases because the data quality were insufficient to provide the necessary parameters for further analysis (see below). Our philosophy for membership selection is therefore not to strive to be as complete as possible, but to aim to provide a secure list of kinematically selected members where any contamination is accurately accounted for by the membership probabilities.

Table 2 lists the parameters associated with the summed spectrum for a particular target observed with a particular instrumental set-up. Most targets were observed with either GIRAFFE or UVES though a single wavelength filter; the 26 000 lines in Table 2 represent 25 417

unique targets. Wherever possible, values for effective temperature (T_{eff} in K), surface gravity ($\log g$, with g in cm s^{-2}), and the gravity-sensitive spectroscopic index γ (Damiani et al. 2014) were taken from the *Recommended Astro Analysis* data base. If no value is given in that data base then the parameter is taken from the *Astro Analysis* data base. If the *Astro Analysis* data base show multiple values from different working groups then Table 2 shows their median value.

Table 2 shows the GESiDR5 values of radial velocity (RV) and uncertainty for each target/filter combination. For targets observed with the GIRAFFE 665 nm filter, an improved empirical precision, S_{RV} , is calculated from the target signal-to-noise ratio (S/N) and projected equatorial velocity ($v \sin i$), following the method described in Jackson et al. (2015), and using empirical constants determined from the analysis of GESiDR5 cluster data (see Appendix A). The values of S_{RV} are shown in Table 2 and used in the subsequent membership analysis.

GES values of RA and Dec. in Table 2 were cross-matched with the *Gaia* DR2 catalogue to obtain photometry, parallax, and proper motion data. Table 3 shows the tangential velocities in units of km s^{-1} calculated as

$$\begin{aligned} V_{\text{RA}} &= 4.74 d_{\text{c}} \text{ pm}_{\text{RA}}, \\ V_{\text{Dec}} &= 4.74 d_{\text{c}} \text{ pm}_{\text{Dec}}, \end{aligned} \quad (1)$$

where pm_{RA} and pm_{Dec} are the proper motions in units of mas yr^{-1} , and d_{c} (in pc) is the cluster distance used for the analysis of cluster membership. The implicit assumption here is that all stars are at a common distance. This will be a good approximation for cluster members, but for unassociated field stars d_{c} is simply a scaling constant that multiplies their proper motions prior to the analysis. The *Gaia* DR2 data were not filtered for possible problems with the astrometry (Lindgren et al. 2018). The issue of astrometric reliability is discussed further in Section 5.1.

2.2 Selecting potential cluster members

Potential cluster members were selected from the list of observed targets in Table 2. Targets were first selected as having reported values of 2MASS K_s magnitude, *Gaia* G magnitude and cross-match radius < 2 arcsec, T_{eff} and a spectral S/N ≥ 5 . The GES cluster targets were based on 2MASS coordinates and all have 2MASS data, but 75 were excluded here with no *Gaia* match. Fig. 1 shows Hertzsprung–Russell (HR) diagrams of targets in each cluster. The luminosity is estimated from the K_s magnitude as

$$\log L/L_{\odot} = (4.75 - M_{K_s} - BC_{K_s})/2.5, \quad (2)$$

where

$$M_{K_s} = K_s - (M - m)_0^c - R_{K_s} E(B - V)^c, \quad (3)$$

where $(M - m)_0^c$ and $E(B - V)^c$ are the adopted intrinsic distance modulus and reddening of the cluster.

The bolometric correction BC_{K_s} is estimated from the de-reddened $G - K_s$ using solar-metallicity Pisa model isochrones (Tognelli, Prada Moroni & Degl’Innocenti 2011) at the cluster age, $(M - m)_0^c$, and $E(B - V)^c$, and assuming extinctions in the G and K_s bands of $R \times E(B - V)^c$ where $R_G = 2.50$ (Chen et al. 2019) and $R_{K_s} = 0.35$ (Yuan, Liu & Xiang 2013). These coefficients are an approximate average for all spectral types and the possibility of discs and anomalous reddening in the youngest clusters is ignored. Since the only purpose (in this paper) of calculating luminosities is to estimate approximate masses to use in a small correction to the binary RV uncertainty distribution (see Section 3.1), further detail is not warranted.

¹ <http://ges/roe.ac.uk/>

Table 1. Cluster data. Columns 2–4 show ages, intrinsic distance moduli, and reddening from the literature (superscripts refer to references listed below the table). Columns 5–7 show the numbers of targets observed, the numbers of targets with a full set of the required data (see Section 2.2) and the number fitted in the membership analysis. Column 8 shows the spectral resolving power measured from arc lamp line widths (see Appendix A1), and columns 9–10 show the final mean values of distance modulus and reddening for the clusters, determined from high-probability cluster members; the first error bar on the distance modulus is a statistical uncertainty, the second is a systematic uncertainty corresponding to ± 0.1 mas in parallax (see Section 3.4).

Cluster	Age (Myr)	$(M - m)_0$ literature	$E(B - V)$ literature	Number observed	Number complete	Number fitted	Resolving power	$(M - m)_0^c$ members	$E(B - V)^c$ members
Trumpler 14	1–3 ¹⁹	12.3 ¹⁹	0.4–0.9 ¹⁹	1118	1063	729	12951	12.15 $\pm$ 0.03 $\pm$ 0.60	0.71 $\pm$ 0.14
Chamaeleon I	2 ²⁵	6.02 ⁴⁴	$\sim$ 1 ²⁵	720	649	148	13897	6.39 $\pm$ 0.01 $\pm$ 0.04	0.19 $\pm$ 0.10
NGC 6530	1–7 ³²	10.48 ³²	0.35 ³⁸	1980	1294	1075	13561	10.66 $\pm$ 0.03 $\pm$ 0.30	0.48 $\pm$ 0.10
NGC 2264	3 ⁴⁰	9.4 ³⁷	0.07 ⁴¹	1884	1738	1344	14968	9.35 $\pm$ 0.01 $\pm$ 0.16	0.03 $\pm$ 0.07
Rho Ophiuchus	3 ¹⁵	5.4 ²⁴	–	313	298	70	15854	5.72 $\pm$ 0.01 $\pm$ 0.03	0.58 $\pm$ 0.20
Lambda Ori	6 ¹⁴	7.9 ¹³	0.12 ¹¹	618	546	296	17281	8.02 $\pm$ 0.01 $\pm$ 0.09	0.06 $\pm$ 0.03
Gamma 2 Vel	5–10 ²²	7.72 ²²	0.04 ²²	1283	1283	496	14301	7.76 $\pm$ 0.01 $\pm$ 0.08	0.02 $\pm$ 0.03
NGC 2232	32 ²⁶	7.56 ¹⁶	0.03 ⁴³	1769	1764	760	12402	7.56 $\pm$ 0.01 $\pm$ 0.07	0.06 $\pm$ 0.04
NGC 2547	35 ²¹	7.97 ¹⁶	0.06 ²⁹	480	480	267	13862	7.97 $\pm$ 0.01 $\pm$ 0.09	0.08 $\pm$ 0.05
IC 4665	42 ⁵	7.69 ¹⁶	0.17 ⁵	567	567	300	15332	7.71 $\pm$ 0.02 $\pm$ 0.08	0.13 $\pm$ 0.04
IC 2602	46 ¹²	5.91 ¹⁶	0.03 ¹⁷	1861	1794	116	13542	5.91 $\pm$ 0.01 $\pm$ 0.03	0.01 $\pm$ 0.02
NGC 2451b	50 ¹⁸	7.84 ¹⁸	0.01 ³¹	1657	1655	418	13862	7.84 $\pm$ 0.02 $\pm$ 0.08	0.07 $\pm$ 0.03
IC 2391	53 ³	5.9 ¹⁶	0.01 ⁴³	438	420	67	12963	5.92 $\pm$ 0.01 $\pm$ 0.03	0.02 $\pm$ 0.02
NGC 2451a	50–80 ¹⁸	6.44 ¹⁶	0.01 ³¹	1657	1655	354	13862	6.44 $\pm$ 0.01 $\pm$ 0.04	0.02 $\pm$ 0.02
NGC 2516	125 ²⁶	8.09 ¹⁶	0.11 ³⁹	764	764	643	13440	8.08 $\pm$ 0.01 $\pm$ 0.09	0.13 $\pm$ 0.03
NGC 6067	120 ⁴³	10.76 ⁴³	0.38 ⁴³	532	530	489	17279	11.79 $\pm$ 0.02 $\pm$ 0.50	0.32 $\pm$ 0.04
Blanco 1	100–150 ²⁸	6.88 ¹⁶	0.01 ⁴³	468	468	326	17282	6.89 $\pm$ 0.01 $\pm$ 0.05	0.01 $\pm$ 0.04
NGC 6259	210 ²⁷	11.61 ⁹	0.66 ²⁷	447	447	373	17359	11.91 $\pm$ 0.03 $\pm$ 0.53	0.69 $\pm$ 0.06
NGC 6705	250 ³⁵	11.37 ²³	0.42 ³⁵	1070	1070	963	14393	11.90 $\pm$ 0.01 $\pm$ 0.53	0.36 $\pm$ 0.08
NGC 4815	500 ⁶	11.99 ⁶	0.7 ⁶	126	126	105	14012	12.87 $\pm$ 0.11 $\pm$ 0.85	0.67 $\pm$ 0.11
NGC 6633	575 ⁴²	7.99 ¹⁶	0.17 ⁴²	1600	1598	143	14532	7.97 $\pm$ 0.01 $\pm$ 0.09	0.16 $\pm$ 0.02
Trumpler 23	900 ⁴	11.71 ⁷	0.58 ⁴	89	89	77	13600	12.29 $\pm$ 0.04 $\pm$ 0.64	0.74 $\pm$ 0.04
NGC 6802	950 ²⁰	11.28 ²⁰	0.84 ²⁰	103	103	94	14309	12.72 $\pm$ 0.13 $\pm$ 0.79	0.82 $\pm$ 0.07
Berkeley 81	1000 ³³	12.39 ³³	1.0 ³³	203	203	169	13849	12.89 $\pm$ 0.10 $\pm$ 0.86	0.94 $\pm$ 0.05
Ruprecht 134	1000 ⁷	12.66 ⁷	0.5 ⁷	680	680	415	17299	12.09 $\pm$ 0.05 $\pm$ 0.58	0.47 $\pm$ 0.07
NGC 6005	1200 ³⁰	12.16 ³⁰	0.45 ³⁰	355	355	275	13771	12.38 $\pm$ 0.10 $\pm$ 0.67	0.27 $\pm$ 0.08
Pismis 18	1200 ³⁰	11.75 ³⁰	0.5 ³⁰	101	101	86	13965	12.43 $\pm$ 0.04 $\pm$ 0.69	0.69 $\pm$ 0.05
Trumpler 20	1400 ⁸	12.39 ⁸	0.35 ⁸	557	557	447	13743	13.09 $\pm$ 0.04 $\pm$ 0.96	0.38 $\pm$ 0.10
NGC 2420	2200 ³⁴	11.97 ³⁶	0.05 ²	563	563	514	12408	12.20 $\pm$ 0.02 $\pm$ 0.61	0.01 $\pm$ 0.03
Berkeley 31	2900 ¹⁰	14.4 ¹⁰	0.19 ¹⁰	616	616	454	13048	14.39 $\pm$ 0.23 $\pm$ 2.14	0.05 $\pm$ 0.12
Berkeley 44	2900 ²⁰	12.46 ²⁰	0.98 ²⁰	93	93	82	13600	12.63 $\pm$ 0.08 $\pm$ 0.76	0.87 $\pm$ 0.07
NGC 2243	3800 ¹	12.96 ¹	0.05 ¹	705	705	564	14051	13.41 $\pm$ 0.02 $\pm$ 1.14	0.01 $\pm$ 0.11

Notes. ¹Anthony-Twarog, Atwell & Twarog (2005); ²Anthony-Twarog et al. (2006); ³Barrado y Navascués, Stauffer & Patten (1999); ⁴Bonatto & Bica (2007); ⁵Cargile, James & Jeffries (2010); ⁶Carraro & Ortolani (1994); ⁷Carraro et al. (2006); ⁸Carraro, Costa & Ahumada (2010); ⁹Ciechanowska et al. (2006); ¹⁰Cignoni et al. (2011); ¹¹Diplas & Savage (1994); ¹²Dobbie, Lodieu & Sharp (2010); ¹³Dolan & Mathieu (1999); ¹⁴Dolan & Mathieu (2002); ¹⁵Erickson et al. (2011); ¹⁶Gaia Collaboration et al. (2018b); ¹⁷Hill & Perry (1969); ¹⁸Hünsch & Weidner (2003); ¹⁹Hur, Sung & Bessell (2012); ²⁰Janes & Hoq (2011); ²¹Jeffries & Oliveira (2005); ²²Jeffries et al. (2009); ²³Jeffries et al. (2017); ²⁴Loinard et al. (2008); ²⁵Luhman (2007); ²⁶Lyra et al. (2006); ²⁷Mermilliod et al. (2001); ²⁸Moraux et al. (2007); ²⁹Naylor & Jeffries (2006); ³⁰Piatti et al. (1998); ³¹Platais et al. (2001); ³²Prisinzano et al. (2005); ³³Sagar & Griffiths (1998); ³⁴Salaris, Weiss & Percival (2004); ³⁵Santos, Bonatto & Bica (2005); ³⁶Sharma et al. (2006); ³⁷Sung, Bessell & Lee (1997); ³⁸Sung, Chun & Bessell (2000); ³⁹Sung et al. (2002); ⁴⁰Sung, Bessell & Chun (2004); ⁴¹Turner (2012); ⁴²van Leeuwen (2009); ⁴³WEBDA (Dias et al. 2002); ⁴⁴Whittet et al. (1997).

Table 2. *Gaia*-ESO Survey data used to estimate the probability of cluster membership for targets observed in the 32 clusters of Table 1. Data are from the iDR5 iteration of the GES analysis, obtained at the GES archive of the Wide Field Astronomy Unit at Edinburgh University (see Section 2.1). A sample of the table is shown here. The full table is available as Supporting Information.

Target cname	Filter (nm)	Cluster DB name	RA (deg)	Dec (deg)	<i>S/N</i>	<i>T</i> _{eff} (K)	log <i>g</i>	γ	<i>K</i> _s (mag)	<i>RV</i> (km s ^{−1})	σ_{RV}	log <i>LL</i> _⊙
10532815–7710268	665	Cha.I	163.36729	−77.1741	17.140	5698	3.435	1.007	11.385	26.77	0.54	−1.20
10563044–7711393	665	Cha.I	164.12683	−77.1942	57.060	3989	4.590	0.968	8.631	16.84	0.30	−0.06
10573004–7620097	665	Cha.I	164.37517	−76.3360	197.930	4590	2.700	1.02	8.022	41.03	0.24	0.33
11022491–7733357	580	Cha.I	165.60379	−77.5599	56.500	4544	4.510	−999	8.199	15.67	0.33	0.30
11100704–7629377	580	Cha.I	167.52933	−76.4938	35.220	4267	4.550	−999	8.451	14.94	0.32	−0.02

Table 3. *Gaia* data for targets observed in the 32 clusters of Table 1 (see Section 2.1) and other calculated parameters. Parallax and proper motion data are from the *Gaia* DR2 catalogue. The RV precision, S_{RV} , is calculated from the GESiDR5 data as described in Appendix A. A *Gaia* flag of zero indicates targets with potentially unreliable astrometric data. The final two columns show the probability that the target is a member of its given cluster using the full data set, p , or the same probability computed using a data set filtered to remove any targets with a *Gaia* flag of zero (see Section 5.1). Targets with a membership probability of -1 were excluded from the membership analysis. A sample of the table is shown here. The full table is available as Supporting Information.

Target cname	Filter (nm)	Cluster DB name	PLX (mas)	σ_{PLX} (mas)	V_{RA} (km s $^{-1}$)	σ_{VRA} (km s $^{-1}$)	V_{Dec} (km s $^{-1}$)	σ_{VDec} (km s $^{-1}$)	RV (km s $^{-1}$)	S_{RV}	<i>Gaia</i> flag ^a	Membership Probability p	p_{filter}
10532815–7710268	665	Cha.I	1.89	0.08	−20.62	0.12	5.05	0.11	26.77	0.79	1	0.0665	0.0690
10563044–7711393	665	Cha.I	5.46	0.02	−21.32	0.03	2.54	0.03	16.84	0.29	1	0.9983	0.9983
10573004–7620097	665	Cha.I	1.26	0.02	−6.73	0.04	11.44	0.03	41.03	0.14	1	−1	−1
11022491–7733357	580	Cha.I	5.67	0.04	−20.87	0.07	5.46	0.06	15.67	0.33	1	0.9163	0.9164
11100704–7629377	580	Cha.I	5.57	0.28	−20.43	0.50	−3.14	0.52	14.94	0.32	0	0.9954	−1

^aA *Gaia* flag of zero indicates targets with potentially unreliable *Gaia* data – <8 visibility periods or renormalized weighted error <1.4 (see Section 5.1 and Lindegren et al. 2018).

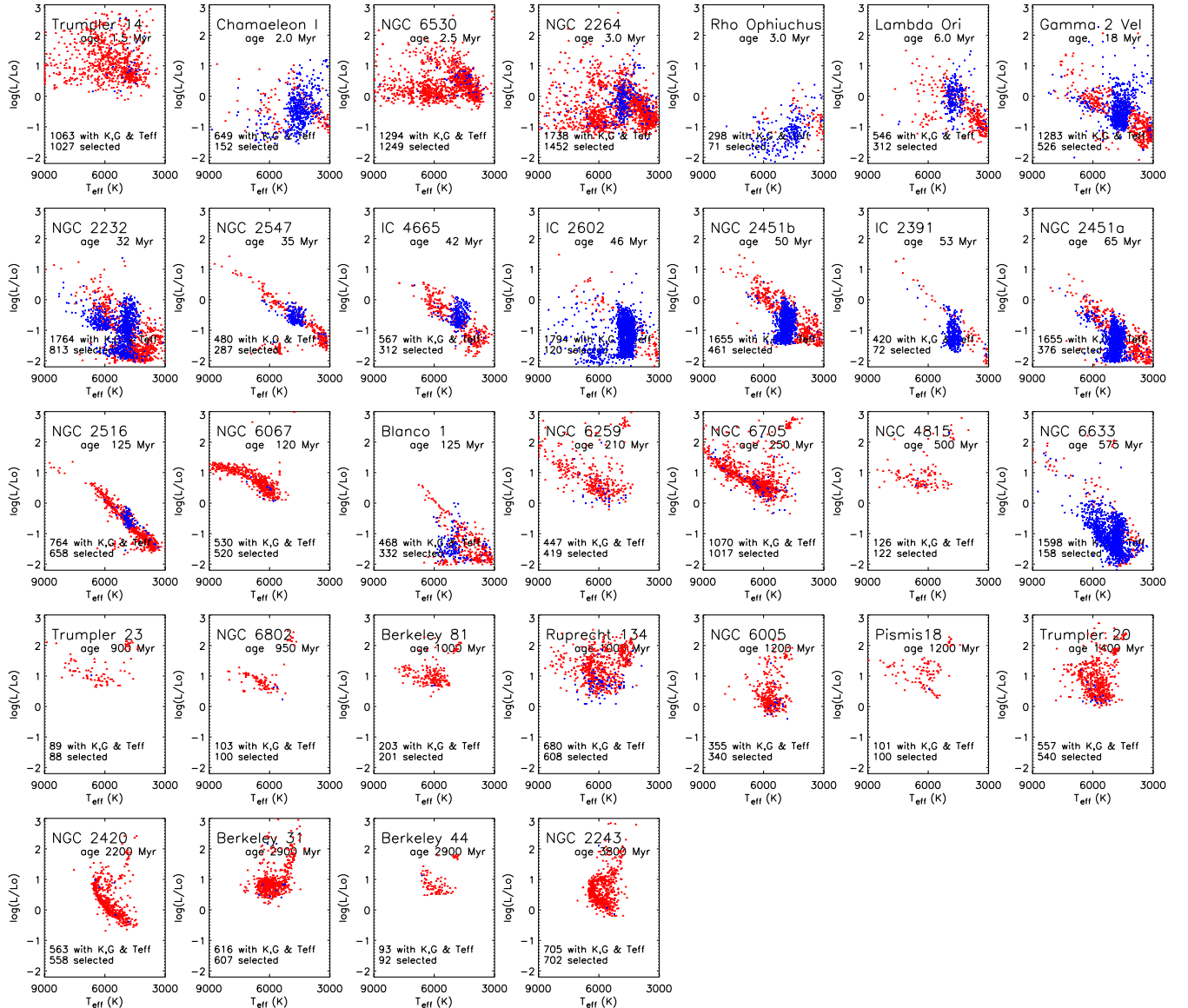


Figure 1. Cluster HR diagrams based on targets that have a 2MASS K_s magnitude, *Gaia* G magnitude, and GES T_{eff} estimate. Blue points are classified as probable field giants or other distant field stars; red points are potential cluster members selected for further analysis (see Section 2.2). Text on the plots shows the number of targets with complete data (K_s and G magnitude and T_{eff}) and the number selected as potential cluster members having no strong indication of being a background giant or other distant field star.

In the first analysis pass, the intrinsic distance modulus and reddening of each cluster were taken from the literature (columns 3–4 in Table 1). A subsequent iteration used revised values derived from an initial list of cluster members (see Section 3.4).

The next step was to screen out targets that, based on their surface gravity and distance, were highly likely to be field giants rather than cluster members. For younger clusters (<1 Gyr), field giants were identified as having $\log g \leq 3.4$, $4000 < T_{\text{eff}} < 7000$ K, and with a parallax smaller than (by at least 2σ) a value corresponding to the intrinsic distance modulus of the cluster +2 mag. For the few targets that had T_{eff} but no available $\log g$, a modified version of the γ index $\gamma' = \gamma + \tau/6$ (see Damiani et al. 2014) was used as a proxy for $\log g$ with a threshold value of $\gamma' \geq 1.335$ corresponding to $\log g \leq 3.4$ (Randich et al. 2018) and where the temperature index τ was interpolated as a cubic function of T_{eff} .

For all clusters, a further screening was then made to cut out distant targets whatever their gravity (including distant giants in the older clusters). Stars were removed if their parallax was smaller than (by at least 4σ) a value corresponding to the intrinsic distance modulus of the cluster +2 mag. Finally, a few targets were rejected that had reported values of V_{RA} , V_{Dec} , or RV outside a 150 km s^{-1} window located approximately $\pm 75 \text{ km s}^{-1}$ either side of the median velocity of the remaining targets. This effectively rejected targets with bad velocity data whilst retaining almost the entire velocity spectrum of the field population. Since the UVES observations have larger systematic RV uncertainties associated with their wavelength calibration (Sacco et al. 2014; Jackson et al. 2015), any UVES observations were discarded if a GIRAFFE observation was present for the same target, ensuring that each target was modelled only once in the membership analysis. Targets rejected for whatever reason are flagged as $\text{mem} = -1$ in Table 3. The number of targets with good photometry, temperature, no strong indication of being a background giant or other distant field star and lying within all three 150 km s^{-1} velocity windows are shown in Table 1 (as the ‘Number fitted’), whereas all targets with complete data are shown in Fig. 1.

The fraction of targets excluded prior to the membership analysis varies from about 1 per cent (e.g. NGC 2243, Berkeley 44) to 93 per cent (IC 2602). The reasons for high exclusion rates in the younger clusters are because the GES target lists were designed to be as inclusive as possible, selecting from very broad regions in colour–magnitude diagrams usually including the lower main sequence and as a result usually including large numbers of distant background giants at similar colours. Conversely, in the older clusters, the target selection was much more focused, using lists of targets that were more likely to be cluster members based on their photometry. In addition, since these older clusters tended to be more distant and usually included cluster giants as targets, there was little scope for the observed targets to be of similar brightness but much further away than the cluster.

3 PROBABILITY OF CLUSTER MEMBERSHIP

The velocity data in Table 3 were used to determine membership probabilities for individual targets, using the maximum likelihood method originally proposed by Pryor & Meylan (1993) and later updated by Cottaar, Meyer & Parker (2012) to include the effect of binarity. This technique assumes that the observed velocities are taken from an intrinsic model broadened by the measurement uncertainties and the effects of unresolved binaries. Given a model specified by a number of free parameters (see below), the best-fitting model was found by maximizing the summed logarithmic likelihood for all stars considered.

3.1 Observational uncertainty

Tangential velocities from the proper motions are assumed to be unaffected by binarity; in other words, it is assumed that the uncertainties in the tangential velocities are due only to the measurement uncertainties reported in *Gaia* DR2, and are described by Gaussian distributions that are scaled with σ_{VRA} and σ_{VDec} . It may be that there are additional systematic errors due to unresolved binarity (at separations < 0.1 arcsec, e.g. Lindegren et al. 2018) affecting the proper motion estimates. A detailed treatment of this is not possible without knowledge of the sampling and particular scan pattern for each object, but the influence of these possible additional systematic uncertainties is discussed further in Section 5.4.

The observational uncertainty in RV can be treated in a more complex way. First, the distribution of measurement uncertainty in RV is non-Gaussian, being better described by a ($\nu = 6$) Student’s t -distribution, scaled with the values of S_{RV} shown in Table 3 (see Appendix A). Secondly, for binary stars, the observational uncertainty includes the effects of RV offsets expected from a set of randomly oriented binary systems with a specified distribution of orbital periods, eccentricities, and mass ratios. The total likelihood of a target’s observed RV is then given by the sum of its likelihood if it were a single star and its likelihood if it were in an unresolved binary:

$$\mathcal{L}^{\text{RV}} = (1 - f_{\text{B}})\mathcal{L}_{\text{S}} + f_{\text{B}}\mathcal{L}_{\text{B}}, \quad (4)$$

where $\mathcal{L}_{\text{S}}$ and $\mathcal{L}_{\text{B}}$ are the likelihood of single and binary stars and f_{B} is the adopted binary fraction.

The calculation of the distribution of RV offsets followed the method described by Cottaar et al. (2012), adopting $f_{\text{B}} = 0.46$, a lognormal period distribution with a mean log period = 5.03 (in days) and dispersion 2.28 dex, and a flat mass ratio distribution for $0.1 < q < 1$ (Raghavan et al. 2010). The influence of these assumptions is discussed further in Section 5.4. For the purposes of calculating the offsets, the primary masses were estimated from the target $\log L/L_{\odot}$ using the Pisa model isochrones with the mass capped at a level equivalent to $\log L/L_{\odot} = 1$. A correction for the dilution effect due to the unresolved light from the secondary at the reflex velocity was also made, using the Pisa models to estimate the secondary contribution for a given mass ratio. Note that the amplitude of the binary-induced RV offsets is relatively insensitive to the primary mass, scaling as $(\text{mass})^{1/3}$.

3.2 Intrinsic models

The maximum likelihood calculation was made in two stages: first, the mean velocity and dispersion of the background population of field stars was characterized for each velocity component using a series of 1D maximum likelihood analyses. These results were then used in a full 3D analysis to determine the likelihood of cluster membership.

For the majority of clusters, we assumed an intrinsic model that is the sum of two Gaussian distributions (cluster plus background) of unknown central velocity and dispersion. In this case, the 1D likelihood was determined as a function of five free parameters, the intrinsic velocity and dispersion of the cluster and background populations and the overall fraction of the observed population that are cluster members. A more complex model, comprising three Gaussian distributions was used to fit five of the clusters and gave a significantly higher maximum log likelihood. In three cases, this was expected since the clusters are known to lie close in velocity space to a second association or cluster. These three clusters are Gamma

Vel (Jeffries et al. 2014), NGC 2547 (Sacco et al. 2015), and the cluster pair NGC 2451a and NGC 2451b (Hünsch & Weidner 2003). We also found that three Gaussian distributions were preferred to fit NGC 2264 and NGC 2232. For these five cases, the central velocity and broad dispersion of the field population was fixed, so that the 1D likelihood was determined as a function of six free parameters.

A 3D analysis was then made to determine the intrinsic cluster properties and the membership probabilities of individual targets. In this case, the model likelihood was determined as a function of seven free parameters; the cluster velocity and dispersion in each dimension and the fraction of the observed population that are cluster members, F , fixing the intrinsic background population velocity and dispersion components at the values found from the 1D analyses (see Appendix B).

3.3 Likelihood and membership probabilities

We compute the likelihood of a star being observed with a given velocity vector $\vec{V}_i$ and measurement uncertainty $\vec{\sigma}_i$ as

$$\mathcal{L}_{i,m,n} = F_n \mathcal{L}_{i,m}^{\mathcal{L}} + (1 - F_n) \mathcal{L}_{i,m}^{\mathcal{B}}, \quad (5)$$

where $\mathcal{L}_{i,m}^{\mathcal{L}}$ and $\mathcal{L}_{i,m}^{\mathcal{B}}$ are the likelihood of fit of the i th target to the m th cluster and background models, respectively, and we step through n values for F , the value of the fraction of stars that belong to the cluster, giving a total log likelihood for a fit of all targets to the m th model of

$$\log \mathcal{L}_m = \sum_i \sum_n \log [F_n \mathcal{L}_{i,m}^{\mathcal{L}} + (1 - F_n) \mathcal{L}_{i,m}^{\mathcal{B}}]. \quad (6)$$

The uncertainty in RV is independent of the uncertainty in proper motions allowing $\mathcal{L}_{i,m}^{\mathcal{L}}$ and $\mathcal{L}_{i,m}^{\mathcal{B}}$ to be calculated as the product of the likelihood of fits in RV, $\mathcal{L}_{i,m}^{\text{RV}}$, and fits in proper motion space, $\mathcal{L}_{i,m}^{\text{pm}}$. Calculation of $\mathcal{L}_{i,m}^{\text{RV}}$ takes account of the effects of binarity on measurement uncertainty as described in Section 3.1. Calculation of $\mathcal{L}_{i,m}^{\text{pm}}$ takes account of the correlated uncertainty between V_{RA} and V_{Dec} by using the covariance matrix elements in the *Gaia* DR2 data set (Gaia Collaboration et al. 2016).

Membership probabilities, p_i , are computed as the expectation value of an individual target being a cluster member summed over the uniform grids of component velocities and dispersions (a total of m models) and n values of fractional membership F_n . The range explored for each parameter is set to be greater than $\pm 5\sigma$ from the maximum likelihood value of that parameter, but with a minimum value of zero for the velocity dispersions. The probability of the i th target being a cluster member is given by

$$p_i = \frac{\sum_n \sum_m F_n \mathcal{L}_{i,m}^{\mathcal{L}}}{\sum_n \sum_m F_n \mathcal{L}_{i,m}^{\mathcal{L}} + \sum_n \sum_m (1 - F_n) \mathcal{L}_{i,m}^{\mathcal{B}}}. \quad (7)$$

3.4 Distance modulus and reddening

A list of probable cluster members (with $p > 0.9$) was used to re-evaluate the cluster distance modulus and reddening. The distance was estimated in two steps. First, the 3σ clipped mean and standard deviation of cluster members were used to estimate an intrinsic dispersion of cluster parallax equal to the standard deviation of cluster members less the rms parallax uncertainty (subtracted in quadrature). The quadrature sum of the estimated dispersion and the parallax uncertainties was then used as weights to calculate a weighted mean parallax and uncertainty. The uncertainty in this weighted mean was always much less than 10 per cent, so the mean parallax was inverted to yield a cluster distance. Table 1 gives the corresponding distance

modulus, $(M - m)_0^c$, where two error bars are quoted. The first is the statistical uncertainty and is larger for more distant clusters or those with few members; the second is a systematic uncertainty equivalent to 0.1 mas in parallax, which accounts for possible correlated errors in the parallax zero-point on small spatial scales (Lindgren et al. 2018), and which is generally much larger than the statistical uncertainty. Cluster reddening was estimated by comparing the measured $G - K_s$ colours with Pisa model predictions of $(G - K_s)_0$ for main-sequence stars (giant members of the older clusters were not included in the calculation) at the target luminosity and the literature age (see Table 1). $E(B - V)^c$ was taken as the median value of the 50 per cent the members with the lowest reddening, since these are more likely to be single stars. The quoted uncertainty in Table 1 is the median absolute deviation (MAD) of this subset.

It is important to note that $(M - m)_0^c$ and $E(B - V)^c$ are scaling constants in the maximum likelihood calculation. Changing them has *no direct effect on the membership probabilities*. It is however useful to determine $(M - m)_0^c$ with reasonable accuracy in order to compare the distance-dependent mean cluster velocity and dispersion in RA and Dec. with those in RV, and an estimate of $E(B - V)^c$ is required to compare luminosity and absolute photometric magnitudes with evolutionary models. Both $(M - m)_0^c$ and $E(B - V)^c$ have a weak, indirect effect on the membership probabilities, since they affect the estimated mass used to determine the RV offsets of binary stars; but even then, the binary RV offsets scale only as $(\text{mass})^{1/3}$.

Once a revised cluster distance modulus and reddening were found, the analysis steps described in Sections 2 and 3 were iterated. The final values of distance modulus and reddening are reported in Table 1.

4 RESULTS

Typical results of the maximum likelihood analysis are shown in Fig. 2 for the young cluster NGC 2264. Similar plots are available for all 32 clusters (Figs B1–B32) in Appendix B (see Supporting Information). The three upper plots in Fig. 2 show contour maps of maximum likelihood as a function of the cluster central velocity and intrinsic dispersion for each velocity component (V_{RA} , V_{Dec} , and RV) over the range of velocity and dispersion explored. The central cross marks the location of maximum likelihood. Contours mark decreasing levels of log likelihood with respect to this maximum.

The three central plots show histograms of the number of targets per 1 km s^{-1} bin for each velocity component. Text on the plots shows the maximum likelihood values of cluster velocity (U_c) and dispersion (D_c), the fraction of targets that are cluster members (F), together with the central velocity (U_b) and dispersion (D_b) of the model distribution of background stars. The black curve shows the model probability distribution based on the maximum likelihood parameter values, the median of the target measurement uncertainties and the assumed fraction of binary systems ($f_B = 0.46$).

The lower plots show, for each velocity component, the variation in log likelihood as a function of the cluster central velocity and intrinsic velocity dispersion, over the range of parameter values explored in the maximum likelihood calculation. The dotted line marks the 1σ level (a log likelihood of -0.5 relative to the maximum value). Text on the plots show the maximum likelihood value of each parameter and values at the upper and lower 1σ levels. These probability distributions were used to calculate the weighted mean and rms values of the cluster velocity components and their intrinsic dispersion that are reported in Table 4.

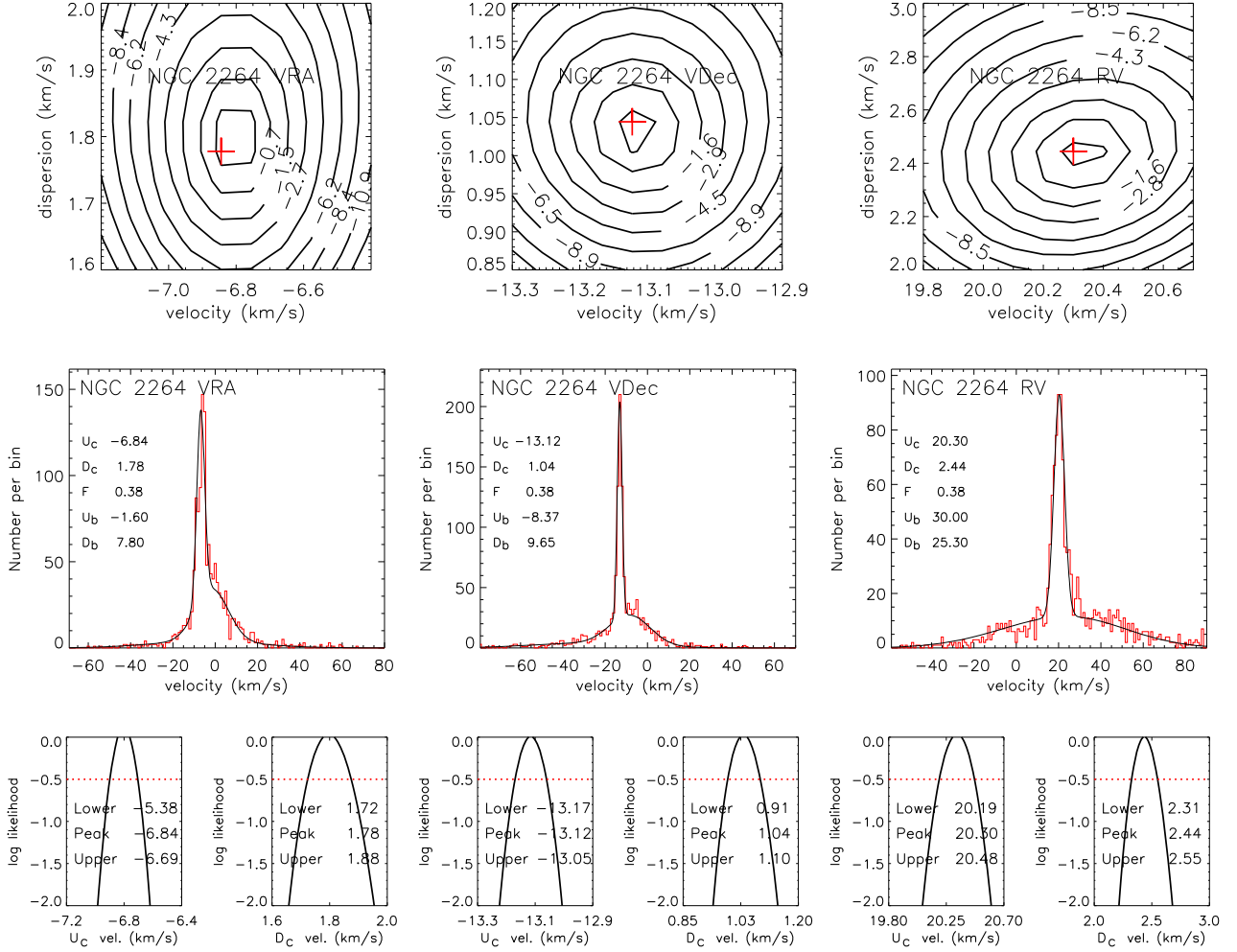


Figure 2. Results of the 3D maximum likelihood analysis of NGC 2264. The upper plots show contours of log likelihood of cluster intrinsic velocity and dispersion relative to the maximum likelihood value located the cross on each plot. The centre plots show histograms of measured velocities together with model probability distributions evaluated at the maximum likelihood values of fitted parameters and a median measurement uncertainty. The lower plots show the variation in log likelihood of each component of cluster mean velocity and intrinsic dispersion.

NGC 2264 is one of the better populated, less distant, clusters with ~ 40 per cent of targets being cluster members, producing a clear peak in number density versus all three components of velocity. This is not always the case. The large relative proper motion uncertainties for more distant clusters (with $(M - m)_0 > 10$) lead to tangential velocity uncertainties that become larger than the intrinsic cluster velocity dispersions. This leads to the cluster members ‘blending in’ to the background and causes larger uncertainties in the tangential velocity dispersions. A well populated example would be NGC 2243 (Fig. B32), where the peaks in V_{RA} and V_{Dec} are much broader than in Fig. 2, but still well defined. In other distant clusters, where fewer members are identified with $p > 0.9$ (e.g. Berkeley 31, fig. B30), the ability to select members is dominated by a narrow peak in RV, which is not affected by distance. For one cluster (NGC 6005; fig. B26), a combination of few high-probability members and a probable large intrinsic velocity dispersion made it difficult to separate the cluster from the background and it was not possible to determine a peak in likelihood for D_{RV} versus velocity. In this case, the maximum D_{RV} was artificially fixed at 6 km s^{-1} in order to determine cluster membership.

The weighted mean and rms values of cluster central velocity and the intrinsic cluster dispersion of the models used to determine cluster

membership are shown in Table 4. The membership probabilities of the individual targets in each cluster were calculated from equation (7) and are reported in Table 3. Targets with a membership probability of -1 were not included in the maximum likelihood analysis. Total numbers of targets in each cluster with a membership probability, $p > 0.9$ and $p > 0.95$, are also shown in Table 4. There are 5032 targets with $p > 0.9$.

The distribution of membership probabilities is a measure of how well the data are able to separate the cluster from the background. In a cluster like NGC 2264 (Fig. B4), this distribution is quite bimodal, with most objects either being clear cluster members with $p > 0.9$ (the average probability for this subset is $\bar{p} = 0.986$) or very unlikely to be cluster members, with $p < 0.05$, with relatively few stars in between. This is the case for most of the observed clusters (see Appendix B), illustrating the power of combining three orthogonal velocity constraints. There are some exceptions to this very sharply bimodal probability distribution (e.g. Trumpler 14, Fig. B1; NGC 6530, Fig. B3; Berkeley 31, Fig. B30).

Both Trumpler 14 and NGC 6530 show evidence that a single 3D Gaussian model is a poor representation of the cluster in one of the dimensions (RV in the case of Trumpler 14, although there is a hint of

Table 4. Results of the maximum likelihood analysis of cluster membership. Columns 2–7 show the weighted mean central velocity and intrinsic dispersion of the 3D Gaussian distribution fitted to cluster members (see Fig. 2 and Section 4). Column 8 shows the proportion of the targets analysed (the number fitted from Table 1) that are expected cluster members. Columns 9 and 10 show the number of targets with membership probability $p > 0.90$ and $p > 0.95$, respectively. Note that there are likely additional systematic uncertainties in the velocity dispersions as a result of assumptions about the sample binary properties – see Section 5.4.

Cluster	Cluster central velocity (km s ⁻¹)			Intrinsic dispersion of cluster (km s ⁻¹)			Fraction members	Number members	
	U_{RA}	U_{Dec}	U_{RV}	D_{RA}	D_{Dec}	D_{RV}		$p > 0.9$	$p > 0.95$
Trumpler 14	−85.32 ± 0.33	30.79 ± 0.31	−5.37 ± 0.80	5.62 ± 0.31	5.46 ± .27	12.92 ± 0.90	0.61 ± 0.02	350	279
Chamaeleon I	−20.03 ± 0.12	0.38 ± 0.15	15.75 ± 0.15	0.93 ± 0.10	1.21 ± 0.14	0.97 ± 0.13	0.50 ± 0.04	74	72
NGC 6530	8.34 ± 0.22	−12.90 ± 0.13	0.21 ± 0.17	3.57 ± 0.21	2.06 ± 0.12	2.45 ± 0.17	0.36 ± 0.02	327	297
NGC 2264	−6.80 ± 0.09	−13.12 ± 0.05	20.33 ± 0.13	1.80 ± 0.07	1.04 ± 0.05	2.44 ± 0.12	0.38 ± 0.01	471	439
Rho Ophiuchus	−4.43 ± 0.14	−17.26 ± 0.15	−6.33 ± 0.32	0.85 ± 0.11	0.98 ± 0.12	1.43 ± 0.34	0.59 ± 0.06	41	41
Lambda Ori	1.93 ± 0.09	−3.51 ± 0.13	26.72 ± 0.14	1.15 ± 0.08	1.50 ± 0.11	1.33 ± 0.13	0.57 ± 0.03	161	157
Gamma 2 Vel	−10.81 ± 0.06	15.63 ± 0.10	18.28 ± 0.16	0.83 ± 0.05	1.42 ± 0.08	1.66 ± 0.15	0.45 ± 0.02	206	198
NGC 2232	−7.33 ± 0.05	−2.80 ± 0.06	25.40 ± 0.06	0.36 ± 0.04	0.42 ± 0.06	0.10 ± 0.12	0.12 ± 0.01	82	80
NGC 2547	−15.98 ± 0.05	7.95 ± 0.06	12.80 ± 0.09	0.60 ± 0.05	0.71 ± 0.06	0.66 ± 0.08	0.62 ± 0.03	159	157
IC 4665	−1.58 ± 0.10	−13.93 ± 0.11	−13.75 ± 0.13	0.55 ± 0.10	0.58 ± 0.12	0.37 ± 0.17	0.14 ± 0.02	37	33
IC 2602	−12.71 ± 0.09	7.82 ± 0.10	17.58 ± 0.10	0.64 ± 0.07	0.60 ± 0.09	0.27 ± 0.15	0.46 ± 0.05	50	49
NGC 2451b	−17.15 ± 0.11	8.34 ± 0.09	15.00 ± 0.12	0.89 ± 0.08	0.61 ± 0.10	0.62 ± 0.13	0.16 ± 0.02	63	62
IC 2391	−17.77 ± 0.09	16.68 ± 0.14	14.95 ± 0.17	0.56 ± 0.07	0.82 ± 0.11	0.56 ± 0.17	0.57 ± 0.06	38	35
NGC 2451a	−19.42 ± 0.17	13.93 ± 0.08	23.41 ± 0.08	0.99 ± 0.16	0.44 ± 0.06	0.13 ± 0.13	0.11 ± 0.02	36	35
NGC 2516	−9.15 ± 0.04	21.93 ± 0.05	23.90 ± 0.06	0.91 ± 0.04	0.92 ± 0.03	0.75 ± 0.06	0.74 ± 0.02	467	459
NGC 6067	−20.68 ± 0.13	−27.93 ± 0.13	−38.20 ± 0.23	1.58 ± 0.11	1.69 ± 0.12	1.57 ± 0.24	0.40 ± 0.02	179	167
Blanco 1	21.13 ± 0.05	3.05 ± 0.04	6.02 ± 0.08	0.50 ± 0.04	0.40 ± 0.04	0.33 ± 0.08	0.40 ± 0.03	129	128
NGC 6259	−11.52 ± 0.19	−32.92 ± 0.18	−33.04 ± 0.41	1.98 ± 0.17	1.83 ± 0.17	3.21 ± 0.48	0.40 ± 0.03	137	132
NGC 6705	−17.91 ± 0.10	−47.31 ± 0.10	35.53 ± 0.16	1.88 ± 0.09	1.92 ± 0.08	2.30 ± 0.17	0.59 ± 0.02	540	524
NGC 4815	−102.12 ± 0.46	−16.85 ± 0.38	−27.22 ± 0.72	2.58 ± 0.43	2.06 ± 0.34	3.73 ± 0.74	0.51 ± 0.05	50	48
NGC 6633	2.18 ± 0.22	−3.02 ± 0.20	−28.18 ± 0.21	1.19 ± 0.20	0.99 ± 0.18	0.77 ± 0.20	0.26 ± 0.04	34	31
Trumpler 23	−57.06 ± 0.39	−64.44 ± 0.27	−61.25 ± 0.27	2.01 ± 0.43	1.43 ± 0.24	1.06 ± 0.32	0.51 ± 0.06	39	39
NGC 6802	−46.36 ± 0.28	−106.39 ± 0.40	13.29 ± 0.39	1.10 ± 0.29	1.81 ± 0.56	1.54 ± 0.43	0.59 ± 0.05	53	50
Berkeley 81	−21.70 ± 0.45	−33.03 ± 0.48	48.11 ± 0.20	1.09 ± 0.62	1.97 ± 0.67	0.59 ± 0.29	0.33 ± 0.04	50	49
Ruprecht 134	−20.54 ± 0.18	−30.21 ± 0.18	−40.91 ± 0.16	0.88 ± 0.22	1.09 ± 0.19	0.48 ± 0.20	0.17 ± 0.02	59	57
NGC 6005	−57.24 ± 0.40	−53.86 ± 0.40	−24.46 ± 0.98	2.25 ± 0.49	2.09 ± 0.55	4.37 ± 1.36	0.23 ± 0.03	46	42
Pismis18	−82.06 ± 0.31	−33.33 ± 0.37	−27.66 ± 0.50	1.22 ± 0.30	1.40 ± 0.35	1.20 ± 0.71	0.30 ± 0.05	24	22
Trumpler 20	−139.37 ± 0.20	2.96 ± 0.19	−39.75 ± 0.17	2.36 ± 0.22	2.14 ± 0.19	1.47 ± 0.24	0.41 ± 0.02	157	149
NGC 2420	−15.60 ± 0.10	−27.71 ± 0.09	74.63 ± 0.04	1.29 ± 0.10	1.15 ± 0.08	0.61 ± 0.05	0.77 ± 0.02	395	391
Berkeley 31	3.79 ± 1.49	−31.74 ± 0.85	56.94 ± 0.13	9.06 ± 1.75	4.44 ± 1.85	0.54 ± 0.25	0.25 ± 0.03	81	72
Berkeley 44	−22.00 ± 0.33	−45.28 ± 0.32	−8.60 ± 0.21	1.03 ± 0.45	0.99 ± 0.47	0.66 ± 0.25	0.54 ± 0.06	43	41
NGC 2243	−29.26 ± 0.13	124.35 ± 0.13	59.82 ± 0.05	1.77 ± 0.13	1.74 ± 0.14	0.59 ± 0.07	0.81 ± 0.02	454	449

bifurcation in proper motion space too; V_{RA} for NGC 6530). Adding a further Gaussian component for these clusters did not significantly improve the maximum likelihood of the fits, probably because the kinematic substructure is more complex than such a simple model. In these cases, the discrimination between members and non-members may not be optimal and lowers the average probability of cluster members ($p > 0.9$), but it does not invalidate the estimated membership probabilities, which simply reflect the fit of the model to the data. Kinematic substructure would be unsurprising in these very young clusters. The GES survey covers an area of Trumpler 14 known to have considerable spatial substructure (Feigelson et al. 2011; Damiani et al. 2017), and kinematic substructure in NGC 6530 has already been noted by Wright et al. (2019).

Fig. 3 plots V_{RA} versus V_{Dec} for ($p > 0.9$) members of each cluster. Ellipses represent one and two times the cluster velocity dispersions along each axis, with error bars indicating their uncertainties. For the majority of clusters, there is a clear grouping of members in proper motion space with >90 per cent located within the 2σ ellipse. For the most distant clusters (with $(M - m)_0^c > 12$), the rms uncertainties in tangential velocity are larger than the estimated values of cluster dispersion and so the velocities scatter well beyond the 2σ ellipse. These are cases in which case RV becomes the most important parameter determining the probability of cluster membership, since in principle it is distance independent. Also shown in Fig. 3 are

histograms, indicating the number of members (with $p > 0.9$) versus distance modulus over a ± 1 magnitude range relative to $(M - m)_0^c$. Nearby clusters (with $(M - m)_0^c < 10$) show a well-defined peak, as expected for cluster members. For more distant clusters, the uncertainty in the distance modulus of individual targets becomes too large to give a clear indication of cluster membership.

5 DISCUSSION

5.1 The effect of *Gaia* data quality

For the analysis of cluster membership, all reported values of *Gaia* DR2 proper motion and parallax (and their uncertainties) were accepted as valid. An alternative approach would be to filter targets with potentially unreliable *Gaia* data by requiring all sources to have ≥ 8 visibility periods (Arenou et al. 2018) and to apply the cut recommended in Lindegren et al. (2018), requiring that the ‘re-normalized unit weighted error’ was < 1.4 (see Wright et al. 2019 for further details).

To assess the effect of additional filtering of the *Gaia* astrometry, the membership analysis was repeated after removing all targets which did not pass the tests above. The objects with potentially unreliable *Gaia* data (according to these tests) are indicated in

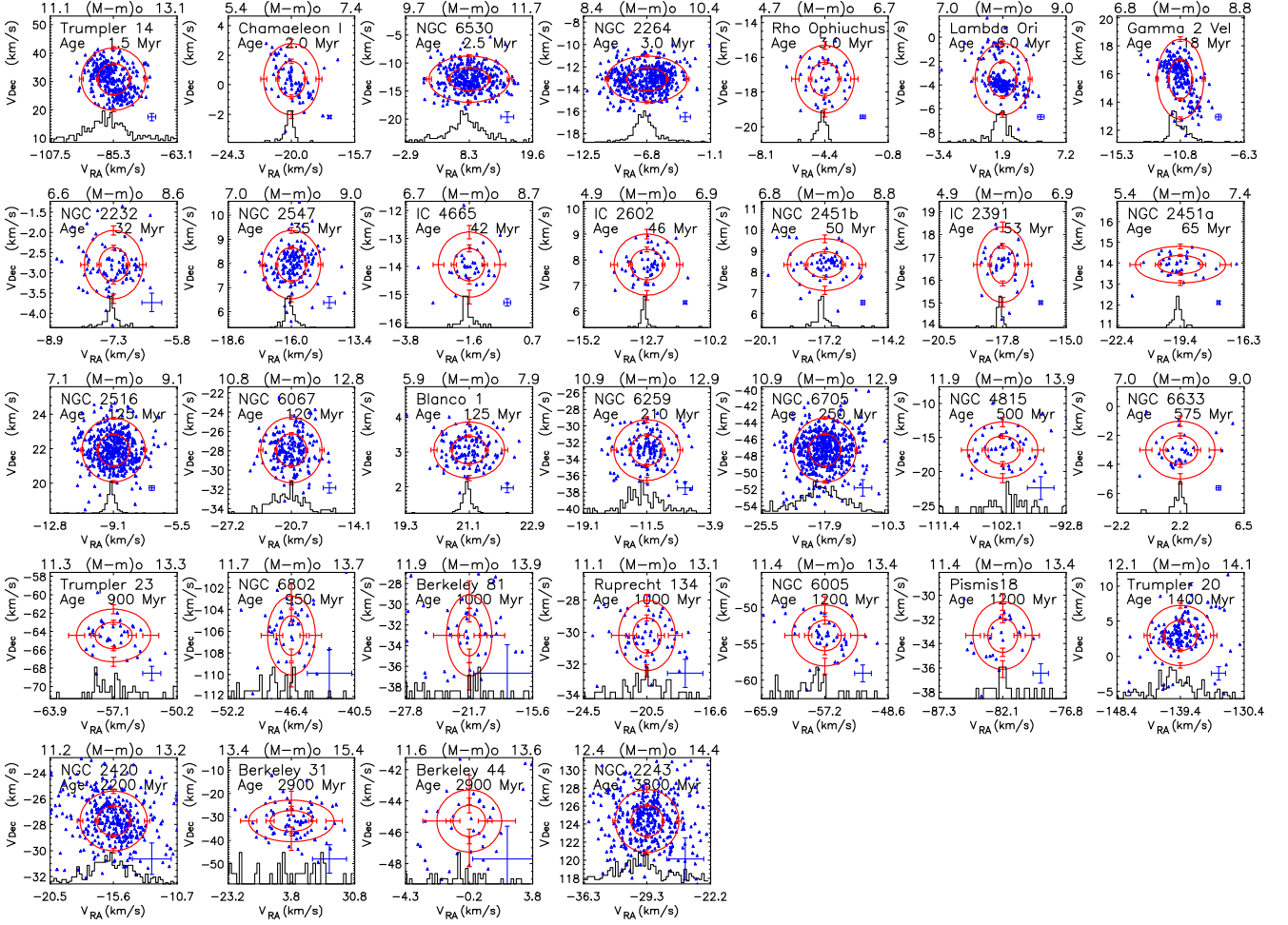


Figure 3. The tangential velocities of cluster members. Blue points show the orthogonal velocity components (in RA and Dec.) for targets with a membership probability $p > 0.9$ (see Table 3) with the blue cross in the lower right hand of the plot, indicating median measurement uncertainties. Red ellipses show the 1σ and 2σ intrinsic dispersion of cluster members (see Table 4) with error bars, indicating the rms uncertainty on the calculated dispersion in V_{RA} and V_{Dec} . The black histograms show the distribution of distance modulus for the cluster members over a ± 1 mag range relative to the cluster center (see scale above each plot).

Table 3, along with revised membership probabilities calculated using just the targets with ‘good’ *Gaia* data.

Applying the filtering reduces the number of valid targets (with all necessary data) by 10 per cent from 12659 to 11386 and the total number of $p > 0.9$ members in the 32 clusters by 7 per cent from 5032 to 4694. A total of 338 targets that were cluster members with $p > 0.9$ are rejected as having suspect *Gaia* data. Only four targets change from $p < 0.9$ to $p > 0.9$. In three cases, this is caused by an incremental (< 0.003) change in p . In the fourth case, the target (18043441–2428057) had a GIRAFFE-based RV that was highly discrepant from the cluster centroid that was rejected in favour of the UVES-derived RV of the same star. The UVES measurement was in much better accord with the cluster centroid, yielding $p > 0.9$. There is only one case of movement in the other direction, where a target with $p = 0.901$ becomes $p = 0.898$ after filtering the *Gaia* data.

Fig. 4 shows how p changes before and after filtering the *Gaia* data. For the vast majority of sources, the changes are very minor. For the subset of 4694, high-probability members that originally had $p > 0.9$ the maximum change in p is 0.011 and the rms difference is 0.0004. These extremely small changes indicate that the membership probabilities (for targets with $p >$

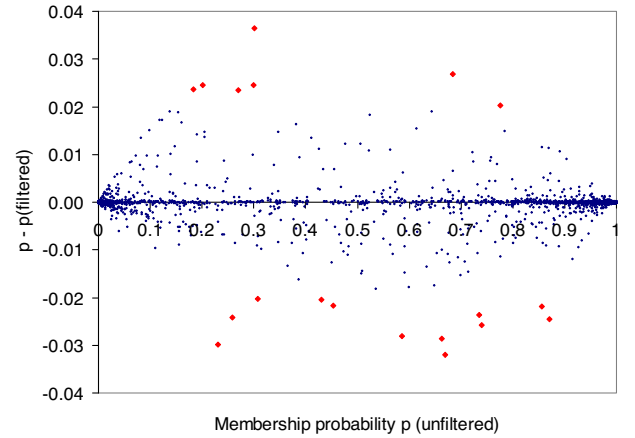


Figure 4. The effect of applying filtering to the *Gaia* DR2 on the derived membership probabilities for the 11 386 sources in the sample present before and after filtering. Large symbols highlight targets with $p - p_{\text{filtered}} > 0.02$.

0.9) evaluated using the unfiltered data set are for practical purposes the same as those calculated if excluding the suspect *Gaia* data.

The reason that the changes in p are so small is that targets that are true cluster members but have potentially unreliable measurements of V_{RA} or V_{Dec} may scatter out of the cluster and appear as background stars. The converse is generally not true. Targets that are true background stars but have an unreliable measurement of V_{RA} or V_{Dec} are very unlikely to scatter into the right cluster velocity range and appear as cluster members. So the process of membership selection using the full data set effectively filters out targets with bad V_{RA} and/or V_{Dec} , whilst retaining targets that are flagged as having suspect *Gaia* data but which actually have V_{RA} and V_{Dec} that agree well with their cluster siblings. By *not* filtering the *Gaia* data we identify 7 per cent more cluster members with $p > 0.9$, with a negligible penalty in terms of additional contamination, as demonstrated in Sections 5.2 and 5.3.

Filtering the *Gaia* data does change estimates of the mean cluster velocities and intrinsic dispersions. The average change in mean cluster velocity is 20 per cent of its uncertainty and the average change in intrinsic dispersion is 30 per cent of its measured uncertainty. The exceptions are Lambda Ori and NGC 2547 that show larger changes in intrinsic dispersion of ~ 1.5 times their estimated uncertainty.

The objective of our analysis is to identify high-probability cluster members from their velocities. It is not to characterize the cluster shapes in velocity space. The results in Table 4 can only be considered rough estimates of the true intrinsic velocity dispersion, whether or not the *Gaia* data are filtered, since the mean velocities and dispersions are the result of fitting the measured velocities in a fixed co-ordinate system to determine the probability of cluster membership. To determine the true cluster shape in velocity space may require a determination of cluster membership that is independent of kinematics and a more general model of the cluster allowing for free rotation of the cluster axes in velocity space, bulk rotation of the cluster, the finite size of the cluster and ‘perspective expansion’ (see van Leeuwen 2009; Kuhn et al. 2019), none of which are explored in this paper and are not crucial to the cluster membership calculations. In addition, the contribution of unresolved binaries will affect the inferred intrinsic velocity dispersion. This does not greatly affect cluster membership calculations, but does result in a further systematic uncertainty in the RV dispersion, where an attempt has been made to account for binarity, and a small over-estimate of the tangential velocity dispersion, where it has not (see Section 5.4).

5.2 Testing cluster membership: I. The Lambda Ori Cluster

There are a small subgroup of clusters where we can compare the kinematically determined membership probabilities with a second independent determination of cluster membership. These are nearby young clusters, aged < 10 Myr, where members will show a distance modulus close to the cluster mean and genuine low-mass members should show almost undepleted levels of lithium compared with older background field stars (e.g. Jeffries et al. 2014; Jeffries 2014). For such clusters, distance modulus and Li abundance [or equivalent width of the Li I 6708 Å feature, $\text{EW}(\text{Li})$] can be used to identify cluster members and compared with kinematically defined membership.

Fig. 5 makes this comparison for the young cluster Lambda Ori. Fig. 5(a) shows the number of targets as a function of distance modulus for targets with $p > 0.9$ and $p < 0.9$ shown separately. Note that only 5 per cent of targets populate the region $0.1 < p < 0.9$ (see

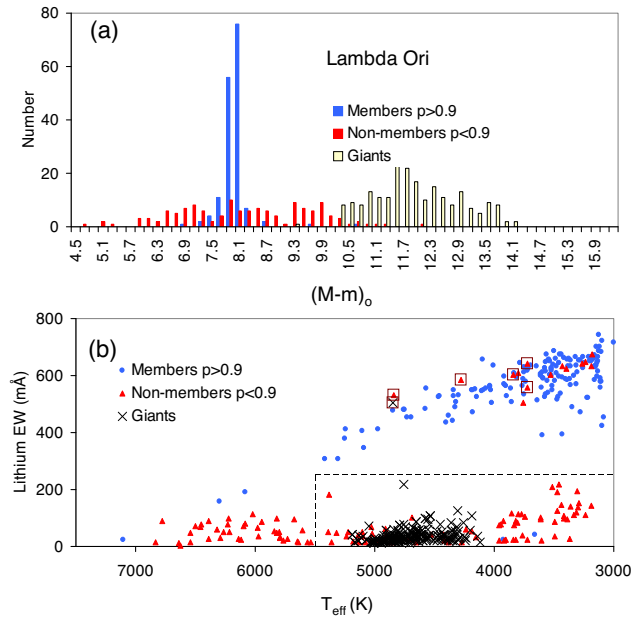


Figure 5. (a) Numbers of probable members ($p > 0.9$) and non-members ($p < 0.9$) in the Lambda Ori cluster as a function of their distance modulus. Also shown are targets identified as background giants and therefore excluded from the maximum likelihood analysis. (b) The equivalent width of the Li 6708 Å feature as a function of T_{eff} for the same targets. The dashed box and the targets marked by squares denote the region discussed in Section 5.2 and the ‘false negatives’ that remain even after filtering the *Gaia* data.

Fig. B6). The distance moduli for individual targets were calculated from the Bayesian distance determinations, given by Bailer-Jones et al. (2018) rather than inverting the individual parallaxes. Probable members are tightly grouped in distance modulus, with a dispersion of 0.4 mag. Likely non-members are much more dispersed as are targets identified as probable giants on the basis of their $\log g$ that were excluded from the likelihood analysis. It is notable that, even in this relatively nearby cluster, the use of distance as a membership constraint would not have improved our rejection of non-members significantly. There are only two $> 3\sigma$ outliers in the top panel of Fig. 5, both of which have possibly unreliable astrometry (see Section 5.1).

Fig. 5(b) shows GESiDR5 measurements of $\text{EW}(\text{Li})$ as a function of T_{eff} for targets with $p > 0.9$, $p < 0.9$, and for stars rejected as giants. Taking a clear dividing line between Li-rich members and Li-poor non-members with $T_{\text{eff}} < 5500$ K and $\text{EW}(\text{Li}) > 250$ mÅ (dashed line in Fig. 5b), there are 158 members and 91 non-members. There are two false positives – Li-poor targets with $p > 0.9$. This compares favourably with an expected number of 1 false positive based on the average probability of $\bar{p} = 0.995$ for the $p > 0.9$ cluster members. Six false negatives are also expected since $\bar{p} = 0.065$ for the $p < 0.9$ sample, but we actually see 15 (plus a Li-rich giant).

This comparison of kinematic and Li-based cluster membership indicates that the kinematic selection process is performing as expected, although the higher than expected number of false negatives might suggest either contamination by Li-rich PMS stars from adjacent, dispersed regions of recent star formation at a similar distance, like Orion OB1a (Briceño et al. 2019), or that the distribution of measurement uncertainties in V_{RA} and V_{Dec} shows a significant non-Gaussian tail. The number of Li-rich false negatives is reduced to six (highlighted with squares in Fig. 5b) by filtering for potentially unreliable *Gaia* data (see Section 5.1), but it also

reduces the number of confirmed Li-rich cluster members from 158 to 140, without changing the number of Li-poor false positives. This confirms the conclusion arrived at in Section 5.1; that the membership probabilities for objects classed as members (with high p) are reliable and that filtering the *Gaia* data merely rejects some false negatives whilst reducing the overall number of cluster members identified.

5.3 Testing cluster membership: II. NGC 2264

In general, we expect narrow distributions of V_{RA} , V_{Dec} , and RV for cluster members with much broader distributions for the background stars. This is true for the majority of clusters where the best-fitting dispersion of the apparent² velocity of the background population is between 20 and 30 km s⁻¹ (see Appendix B). The obvious exceptions are Gamma Velorum and NGC 2457, which are in a similar direction to other groups of young stars that are more spatially diffuse but still coherent in velocity (Jeffries et al. 2014; Sacco et al. 2015; Franciosi et al. 2018; Cantat-Gaudin et al. 2019); and NGC 2541a and NGC 2541b that are a pair of clusters with different distances and kinematics, but with a similar age and observed in the same direction (Hünsch & Weidner 2003). Another cluster showing a more complex distribution of background stars is NGC 2264 where the majority of non-cluster members show a narrow dispersion of ~ 9 km s⁻¹ in V_{RA} and V_{Dec} and a broader 25 km s⁻¹ in RV , with a second, less dense background population showing the usual broad velocity distribution in all three components. There are three possible explanations for the observed distribution:

- (i) Stars categorized as non-members are in fact cluster members showing a wide range of velocities in a young, unvirialized, cluster.
- (ii) NGC 2264 lies close in velocity to a second cluster or association.
- (iii) The tangential velocity distribution of background stars in the direction of NGC 2264 is less dispersed than observed for other clusters.

Since NGC 2264 is a young, not too distant, cluster Li abundance and parallax data can be used to find the likely cause. Fig. 6(a) shows a histogram of target numbers in NGC 2264 as a function of distance modulus. Stars identified as cluster members (with $p > 0.9$) show a reasonably tight distribution with a dispersion of 0.5 mag. Stars identified as non-members show a much broader distribution, suggesting they are indeed background stars, not members of NGC 2264 or some second cluster or association at a common distance. Fig. 6(b) also shows the EW(Li) of targets in NGC 2264. Stars identified as cluster members show the high levels of EW(Li) expected for a young cluster aged ~ 3 –5 Myr, whereas those identified as non-members or as giants generally show the lower EW(Li) expected of older stars, confirming that the second population are correctly identified as non-members and that the background distribution is much less dispersed in tangential velocity than for most other clusters in our sample.

For $T_{\text{eff}} < 5500$ K, Fig. 6(b) contains 455 members with $p > 0.9$ and 414 non-members. There are seven false positives – members with $\text{EW}(\text{Li}) < 250$ mÅ. This is consistent with the average membership probability $\bar{p} = 0.986$ for targets with $p > 0.9$ that suggests there should be six false positives. There are 80 false negatives identified as non-members but with $\text{EW}(\text{Li}) > 250$ mÅ. Applying the *Gaia* quality cuts reduces the number of false negatives to 45, at the expense of reducing the number of members from 455 to 430, but this is still

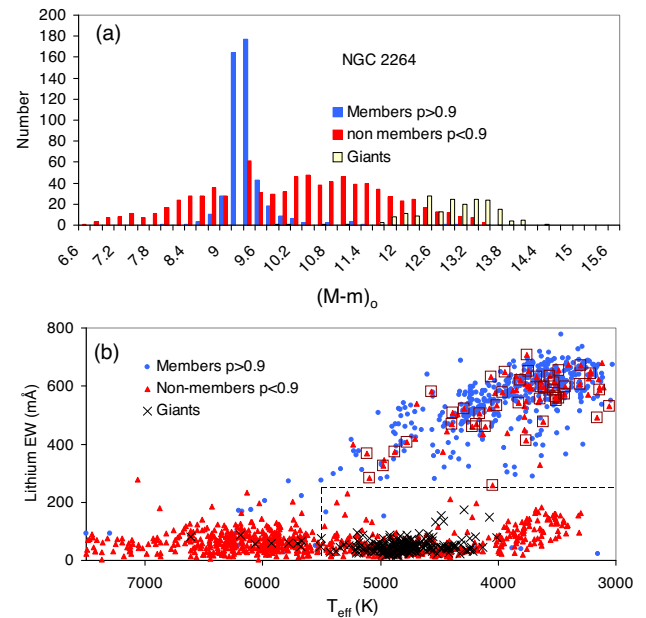


Figure 6. (a) Numbers of probable members ($p > 0.9$) and non-members ($p < 0.9$) NGC 2264 as a function of their distance modulus. Also shown are targets identified as background giants and therefore excluded from the maximum likelihood analysis. (b) The equivalent width of the Li 6708 Å feature as a function of T_{eff} for the same split of targets. The dashed box and the targets marked by squares denote the region discussed in Section 5.3 and the ‘false negatives’ that remain even after filtering the *Gaia* data.

higher than the expected number of 33 false negatives based on $\bar{p} = 0.081$ for targets with $p < 0.9$, suggesting either that we have not fully accounted for non-Gaussian errors in V_{RA} , V_{Dec} , and/or RV or that the intrinsic model for the cluster as a simple 3D Gaussian in velocity space is inadequate. That the latter is a factor is suggested by the complex spatial and kinematic structure of NGC 2264 that has already emerged from previous radial velocity studies and early work with *Gaia* DR2 (Tobin et al. 2015; Venuti et al. 2018; Buckner et al. 2020). In any case, like Lambda Ori, the tests above suggest that the membership probabilities of $p > 0.9$ members are trustworthy and accurately reflect the amount of contamination in any sample drawn from them.

5.4 Binarity

In this work, we have assumed a fixed binary fraction $f_B = 0.46$ and a separation distribution that is the average for solar-type field stars measured by Raghavan et al. (2010). This binary fraction is also similar to that estimated by modelling the photometric deviations from single-star isochrones in a number of these clusters (e.g. NGC 2547 and NGC 2516, Jeffries, Thurston & Hambly 2001; Jeffries et al. 2004). However, both the binary fraction and separation distribution (and possibly the mass ratio distribution) are likely to be mass dependent (Duchêne et al. 2013) and may also be different in clusters as a result of mass segregation or dynamical evolution (e.g. Geller, Latham & Mathieu 2015; Leiner et al. 2015). In principle, the binary properties should be treated as a free parameter in the model RV distribution and constrained by the data, since the RV offsets caused by binarity add to the observed RV dispersion of the cluster, and the shape of the RV offset distribution due to binarity is different to the distribution of RV uncertainties (Cottaar et al. 2012). However, in our case, there are insufficient data to estimate the binary fraction

²The conversion from proper motion to tangential velocity assumes all stars are at the distance of the cluster.

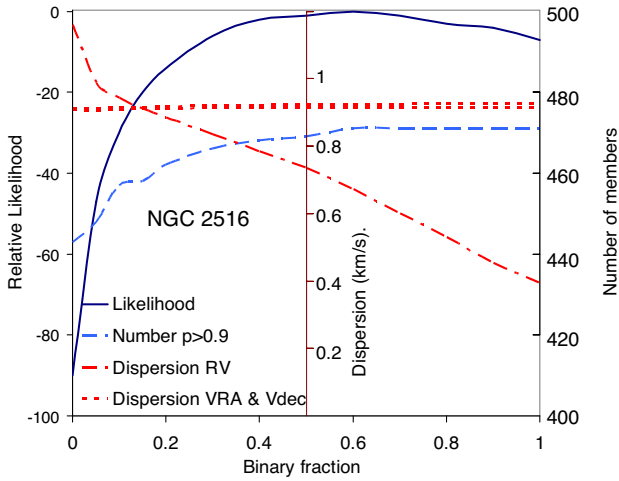


Figure 7. Variation of key parameters with the assumed binary fraction for open cluster NGC 2516. The black solid line shows the maximum likelihood relative to the peak value, the red dot–dashed line shows the intrinsic cluster RV dispersion, the red dotted line shows that dispersion in V_{RA} and V_{Dec} , and the blue dashed line shows the number of targets with a membership probability $p > 0.9$.

with any precision (see below) and an accurate estimation would require an exceedingly precise knowledge of the shape, particularly the tails, of the RV uncertainty distribution.

To assess the effect of fixing f_B , Fig. 7 shows what happens to key parameters from the maximum likelihood analysis if f_B is given different values, using NGC 2516 as an example. This is the best defined of our clusters with a high number of cluster members and a low background population. The maximum likelihood shows a broad peak centred at $f_B \sim 0.5$, but any value between about 0.3 and 0.7 would fit the data equally well. The number of cluster members with $p > 0.9$ changes by only ± 0.5 per cent for $0.3 < f_B < 0.7$. However, changing f_B has a stronger effect on estimates of the intrinsic cluster RV dispersion, since the distribution in offsets of RV due to binarity is convolved with D_{RV} in the calculation of maximum likelihood. If broadening due to binarity is increased by increasing f_B then less of the observed dispersion needs to be explained by intrinsic dispersion in the cluster. This negative correlation between the intrinsic RV dispersion and the assumed value of f_B is clearly shown in Fig. 7.

If $0.3 < f_B < 0.7$, then the best-fitting intrinsic RV dispersion changes by $\sim \pm 0.1 \text{ km s}^{-1}$, which is larger than the formal statistical uncertainties on the fit. For this reason, there is a significant additional systematic uncertainty in the values of D_{RV} shown in Table 4. The size of this uncertainty will have a greater or lesser effect on each cluster depending on the size of the intrinsic RV dispersion of the cluster and RV uncertainties compared with the dispersion introduced by binaries. In many clusters (those with velocity dispersions $> 1 \text{ km s}^{-1}$ or those with large uncertainties in velocity dispersion), it is unimportant. A rule of thumb would be to add or subtract about 0.4 km s^{-1} in quadrature to the quoted value of intrinsic RV dispersion to simulate changing the binary fraction between $0.3 < f_B < 0.7$.

Fig. 7 shows the importance of modelling of the offsets in measured RVs of binary stars relative to the barycentre. A similar, but smaller, effect is expected in the tangential velocities due to offsets in proper motion induced by the motion of the photo-centre of binary systems averaged over the observing period of *Gaia* (e.g. Lindegren et al. 2018). To test the significance of this, a simplistic

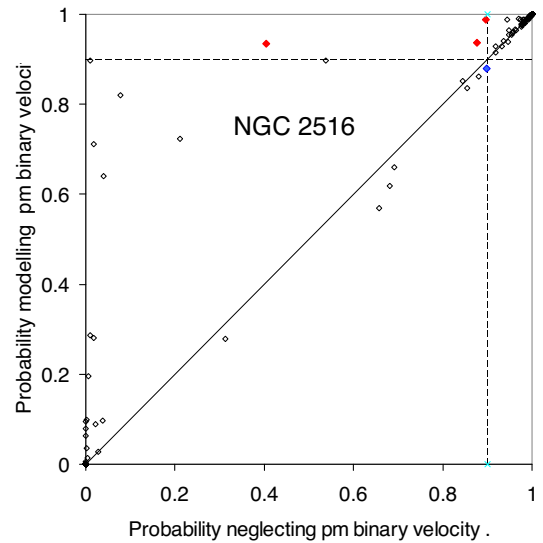


Figure 8. Membership probabilities in NGC 2516 taking approximate account of the variable offset in proper motion caused by binarity, averaged over the period of *Gaia* observations, compared with those (in Table 4) calculated neglecting this effect.

model was developed that includes the additional uncertainty due to offsets in proper motions assuming uniform sampling over the 22 months of *Gaia* observations but neglecting any covariance between the astrometric parameters. As in Section 3.1, the binary velocity offsets, V_b , were calculated for a set of randomly orientated binary systems, with a correction based on the mass ratio of the system to model the influence of the secondary on the motion of the photocentre. We assume that the effects of the averaging diminish the observed velocity offset roughly as $V_b^{pm} = V_b \text{sinc}(\pi T/P)$, where P is the binary period and T is the total observation time. The net effect on a simulated population is similar to that of binarity on the observed RV distribution, but the effect is smaller, and the tails of the distribution are suppressed, because of the strong averaging effect for shorter period binary systems.

This model was applied to the membership probability calculation for NGC 2516, using the same distribution of orbital parameters as before and with $f_B = 0.46$. Fig. 8 shows the comparison of the membership probabilities calculated with and without the effect of binarity on the tangential velocities. Only 6 per cent of stars have a change in membership probability of > 0.01 ; these are targets that were outliers in V_{RA} or V_{Dec} in the original calculation. The general trend is to slightly increase membership probabilities; three stars (highlighted in red) have their membership probability increased to > 0.9 , whilst only one star (in blue) moves in the opposite direction.

This simple treatment suggests that the effect of neglecting binary motion on V_{RA} and V_{Dec} may produce a small (< 1 per cent) underestimate in the number of cluster members with $p > 0.9$ and is much less important than including the effects of binarity on RV (c.f. the blue dashed line in Fig. 8). The effect of binarity on the proper motions also has a weaker effect on the inferred tangential velocity dispersions. The inclusion of the effect for NGC 2516 predicts reduces the estimated intrinsic dispersions, D_{RA} and D_{Dec} , from $0.91\text{--}0.92 \text{ km s}^{-1}$ (see Table 4) to $0.78 \pm 0.04 \text{ km s}^{-1}$ (in better agreement with D_{RV}).

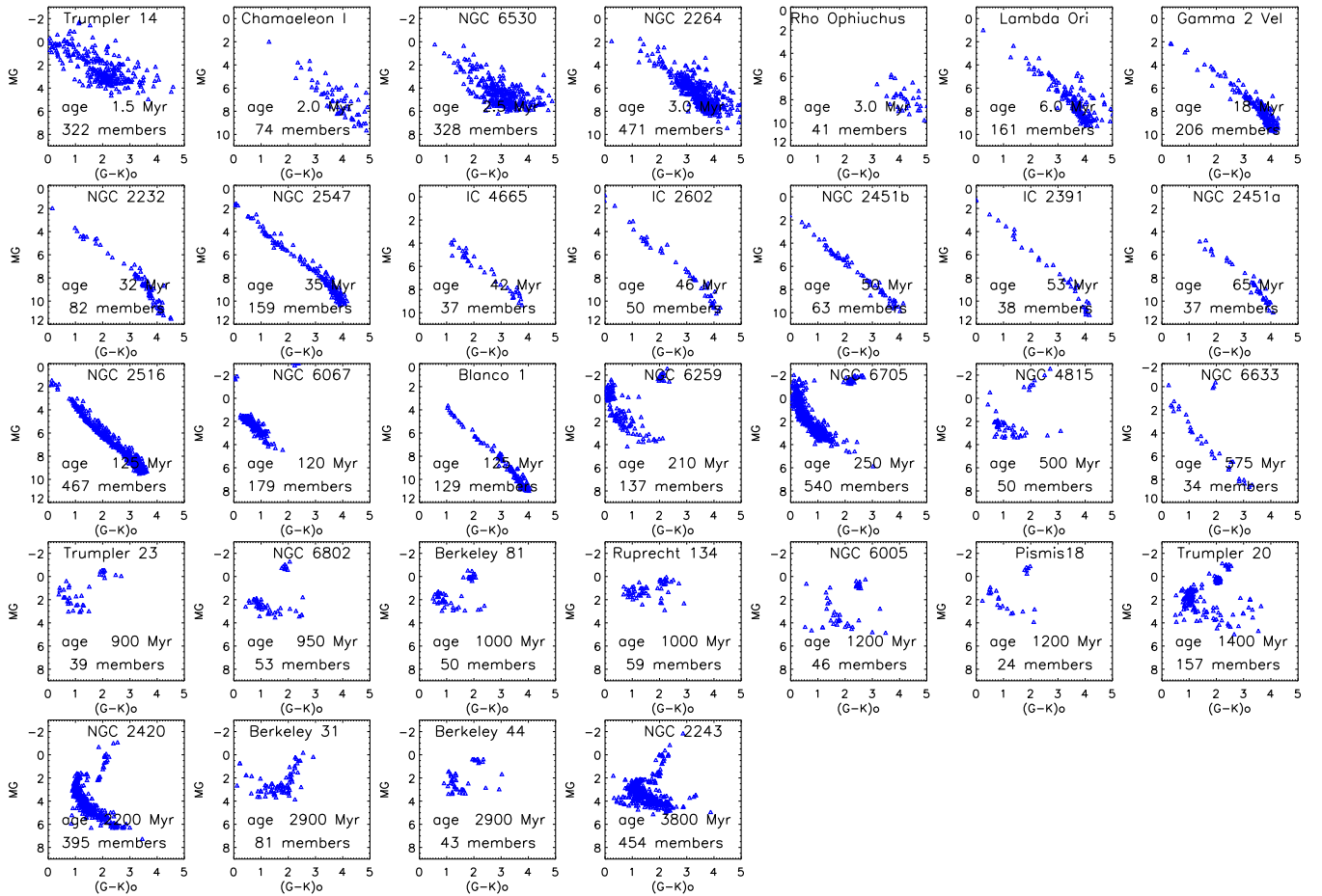


Figure 9. Cluster colour–magnitude diagrams, showing targets with a membership probability $p > 0.9$.

5.5 The advantages of kinematic membership selection

Whilst there are still caveats in examining the detailed kinematics of the clusters using the analysis presented here (see Sections 5.1 and 5.4), the fact that the membership probabilities are almost exclusively based on stellar kinematics makes our membership lists a valuable resource for investigating other properties of clusters and the stars within them. The targets in GES clusters were selected mainly³ on their photometric properties, but with a broad selection in colour–magnitude diagrams that should easily encompass the entire cluster population within the GES magnitude limits. Then, since the membership probabilities here are largely independent of photometry or estimates of stellar chemistry, then the results of this work can serve as inputs to investigate the HR diagram, cluster chemistry, rotation, magnetic activity, light element depletion etc. (e.g. Spina et al. 2017; Randich et al. 2018), but without the concern that results could be biased by using these properties to select members in the first place.

Pursuing these projects is beyond the scope of this paper, but as an illustration Fig. 9 shows the absolute G versus $(G - K_s)_0$ colour–magnitude plots of cluster members (with $p > 0.9$), using the calculated values of distance modulus and reddening shown in Table 1. A comparison with Fig. 1 shows the extent to which our membership selection has ‘cleaned’ these diagrams. Many of

the clusters now clearly follow a single, age-dependent, isochrone with many less outliers than were seen in HR diagrams of the GES targets in Fig. 1. The form of the isochrones are best seen in nearby intermediate age clusters. The scatter is greater on young clusters (< 10 Myr) where the photometry is likely affected by differential reddening.

5.6 The advantage of 3D over 2D kinematic selection

The methodology described in Section 3.3 permits a comparison of how well selection using proper motion alone performs compared with proper motion plus the additional constraints provided by RV. Fig. 10 illustrates the distribution of p obtained for two clusters in the sample (similar plots are available for all clusters in Appendix B) – NGC 2264 a young cluster (~ 5 Myr) at intermediate distance (~ 750 pc) and NGC 6005 an older (~ 1.2 Gyr) and more distant (~ 3 kpc) cluster. The solid lines indicate the final results of the 3D kinematic selection, whilst the dashed lines indicate 2D selection using tangential velocities alone.

In almost all the clusters considered here, the addition of RV steepens the transition between objects with high and low p . That is, it reduces the number of objects with intermediate values of p and improves the contrast between members and non-members. Another general feature is that the average value of p for those objects considered to be likely cluster members ($p > 0.9$) increases. These increases are small, but highly significant if the aim is to provide secure samples with minimal contamination, since it is this latter

³Many of the bright targets observed by UVES were selected as likely cluster members from previous studies.

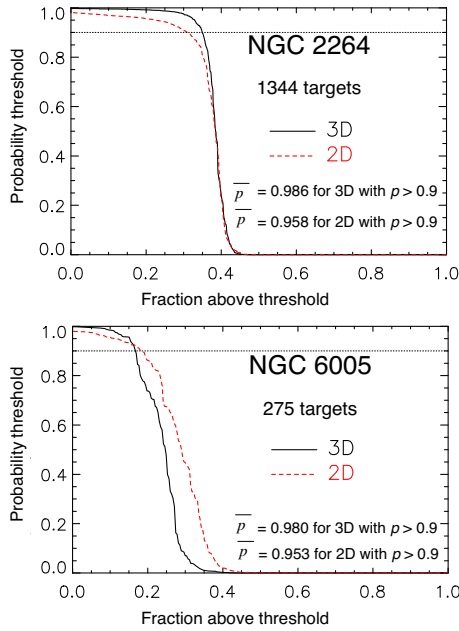


Figure 10. The fraction of targets with membership probability, p , above a threshold level versus the threshold level for NGC 2264 and NGC 6530. The solid line shows the results for the standard 3D (proper motions and radial velocity) analysis, whilst the red dashed lines show the results for a 2D analysis that only uses proper motion. Similar plots are available for all clusters in Appendix B.

statistic that determines the estimated numbers of false positives in the sample. For example in NGC 2264, a sample with $p > 0.9$ selected from 2D velocity data would have 20 false positives, whereas adding RV selection reduces this number to 7.

The magnitude of these improvements depends on the size of the background population and the overlap between the cluster kinematics and that of the background. The latter is increased if the peak defined by the cluster in tangential velocity is blurred by the increased uncertainties that accompany greater distance. In contrast, the resolving power of the RV measurements is not directly distance dependent.

Thus in nearby clusters or where there is relatively little background contamination, the improvements of 3D over 2D selection are very small (e.g. Rho Oph, Fig. B5; NGC 2516, Fig. B15; NGC 2243, Fig. B32). However, when the background is significant and the tangential velocity of the cluster is not distinct from that background, especially in more distant clusters, the improvement in the fidelity of membership selection when adding RV is considerable (e.g. NGC 6530, Fig. B3; Pismis 18, Fig. B27; Berkeley 31, Fig. B30).

6 SUMMARY

In this paper, we have set out a methodology designed to give a secure, rather than complete, set of members for 32 open clusters observed as part of the *Gaia*-ESO Survey. After filtering the observed targets to exclude those without the necessary data or which are obvious background stars or giants, membership is assessed solely using the 3D kinematics of the stars. Using a maximum likelihood technique, robust membership probabilities have been computed and in all of the clusters there is a clear separation between the population of high probability members and objects which are most likely to be unrelated to the cluster. The addition of radial velocities improves the

ability to separate cluster and background populations over proper motion data alone. This is especially important for the distant clusters where the uncertainties in proper motion are larger than the intrinsic dispersions within the cluster or in clusters, where there is significant kinematic overlap between the cluster and background populations.

Tests using independent membership criteria in young clusters suggest that the derived membership probabilities give an accurate indication of the contamination remaining in any sample of high-probability cluster members. However, it appears likely that the membership probabilities of some genuine members, which are not classified as such, may be underestimated. The explanation for this may lie in an imperfect understanding of the reliability of some of the kinematic data or of the tails of the radial velocity and proper motion uncertainty distributions. Alternatively, it is probable that the simple Gaussian models we have used for the intrinsic velocity distributions are too simplistic to fully reflect the kinematics of the young clusters where we have been able to do these tests.

The results of our investigation are presented in the form of a catalogue of compiled data that includes the membership probability of each star observed towards each GES cluster. There are 5032 high probability ($p > 0.9$) members of the 32 clusters, with an average probability of $\bar{p} = 0.991$. We also show the RA, Dec, and RV components of mean velocity and intrinsic dispersion of the cluster model used to determine membership probabilities. We caution that these latter results do not fully characterize the shape of clusters in velocity space, since any cluster asymmetry does not necessarily align with the chosen axes.

Since the membership criteria are almost purely kinematic, and independent of stellar photometry and chemistry, then the catalogue will be valuable for investigating other non-kinematic stellar and cluster properties available from the GES data, without having to compromise the investigation by using those properties as membership criteria. Examples include testing stellar evolutionary models using HR and colour–magnitude diagrams or following the evolution of magnetic activity, rotation, and light element depletion.

With the final data release of GES due towards the end of 2020 and the improvements expected in *Gaia* DR3, it is anticipated that the cluster membership catalogue will be updated in the future to include the full set of clusters observed as part of GES and more reliable and precise astrometric data.

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DATA AVAILABILITY STATEMENT

The data underlying this article are available in the GES archive at the Wide Field Astronomy Unit at Edinburgh University at <http://ge.roe.ac.uk/>

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SUPPORTING INFORMATION

Supplementary data are available at [MNRAS](https://www.mnras.org) online.

Appendix B. Results for individual clusters.

Table S2. *Gaia*-ESO Survey data used to estimate the probability of cluster membership for targets observed in the 32 clusters of Table 1.

Table S3. *Gaia* data for targets observed in the 32 clusters of Table 1 (see Section 2.1) and other calculated parameters.

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APPENDIX A: RADIAL VELOCITY MEASUREMENT PRECISION

An empirical estimate of measurement precision is used for GIRAFFE measurements of RV for the analysis of cluster membership. As described in Jackson et al. (2015), the empirical measurement precision, E_{RV} , is characterized as a Student's t -distribution scaled by an empirical uncertainty S_{RV} that varies with signal-to-noise ratio (S/N) and projected equatorial velocity ($v \sin i$).

The scaling constant of uncertainty for short term repeats (e.g. spectra taken consecutively), using the same instrument set-up and wavelength calibration, is given by

$$S_{RV,0} = B \frac{(1 + ([v \sin i]/C)^2)^{3/4}}{(S/N)}, \quad (A1)$$

where B is an empirically determined parameter that depends on the intrinsic stellar spectrum (largely characterized by the effective temperature) and C is a function of the spectrograph resolving power.

For long-term repeats (e.g. spectra taken on different nights), there is an additional contribution to the measurement uncertainty due to variations in instrument set-up and wavelength calibration, A , which adds in quadrature to the short-term uncertainty, such that the distribution of E_{RV} for long-term repeats is characterized by

$$S_{RV} = \sqrt{A^2 + S_{RV,0}^2}, \quad (A2)$$

Jackson et al. (2015) used data for nine clusters reported in the *Gaia*-ESO Survey (GES) data release iDR2/3 to determine empirical values for A , B , and C . In this case, both A and C were treated as constants over the whole analysis.

The analysis of Jackson et al. (2015) has been repeated here using data for the 32 clusters from GESiDR5 to determine appropriate expressions for A , B , and C . This has required two modifications to the analysis: first, the use of a reduced value of $v \sin i$ to account for changes in instrument resolving power over time; and secondly, the scaling of constant A with S/N as $A = A_0 + A_1/(S/N)$ in order to

fit data from more distant clusters that show lower average levels of S/N.

A1 Calculation of the reduced projected equatorial velocity

The GES pipeline used to estimate $v \sin i$ for GESiDR5 data assumes a fixed spectral resolving power, $R = 17\,000$ for filter HR15n. In practice, the effective resolution of spectra observed using the HR15n filter, measured from the line width of arc-lamp spectra has varied with time over the period of the GES observations, falling from $R \sim 15\,000$ in 2012 January to $R \sim 13\,000$ in 2015 February after which a new focusing procedure for the instrument produced a consistent level of $R \sim 17\,000$. As a result the pipeline values of projected equatorial velocity (V_{ROT}) are higher than the true value of $v \sin i$ for observations made before 2015 February. The effect is most pronounced for the slowest rotating stars where a V_{ROT} of $\sim 12 \text{ km s}^{-1}$ is reported. To correct for the reduction in R below the expected level, a reduced value of $v \sin i$ is used to determine the effect of rotational velocity on measurement precision of RV:

$$v \sin i = \sqrt{V_{ROT}^2 - V_{cor}^2} \text{ for } V_{ROT} > V_{cor}, \quad (A3)$$

where

$$V_{cor} = 0.895c \sqrt{\frac{1}{R^2} - \frac{1}{17\,000^2}}, \quad (A4)$$

c is the speed of light and R is the resolving power over the period when the cluster was observed (values R are shown in Table 1).

A2 Fitted parameters

Data for 30 000 short term repeats and 4400 long-term repeats were analysed to determine empirical values of A , B , and C in equations (A1) and (A2) giving (in units of km s^{-1}):

$$A = 0.09 + 10.0/(S/N);$$

$$B = 4.12 + 242.6 \log(T_{\text{eff}}/3600 \text{ K});$$

$$C = 0.895c/R. \quad (A5)$$

APPENDIX B: RESULTS FOR INDIVIDUAL CLUSTERS

Figs B1–B32 (see Supporting Information) graphically show the results of the maximum likelihood analysis procedure, along with the distribution of membership probabilities (described in Sections 3.2 and 3.3, for each of the 32 GES clusters considered in this paper.

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