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First Measurement of the $^{14}\text{N}/^{15}\text{N}$ Ratio in the Analog of the Sun Progenitor OMC-2 FIR4

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Abstract

We present a complete census of the $^{14}\text{N}/^{15}\text{N}$ isotopic ratio in the most abundant N-bearing molecules toward the cold envelope of the protocluster OMC-2 FIR4, the best known Sun progenitor. To this scope, we analyzed the unbiased spectral survey obtained with the IRAM 30 m telescope at 3, 2, and 1 mm. We detected several lines of CN, HCN, HNC, HC_3N , N_2H^+ , and their respective ^{13}C and ^{15}N isotopologues. The lines' relative fluxes are compatible with LTE conditions, and moderate line opacities have been corrected via a population diagram method or theoretical relative intensity ratios of the hyperfine structures. The five species lead to very similar $^{14}\text{N}/^{15}\text{N}$ isotopic ratios, without any systematic difference between amine- and nitrile-bearing species as previously found in other protostellar sources. The weighted average of the $^{14}\text{N}/^{15}\text{N}$ isotopic ratio is 270 ± 30 . This $^{14}\text{N}/^{15}\text{N}$ value is remarkably consistent with the [250–350] range measured for the local galactic ratio but significantly differs from the ratio measured in comets (around 140). High-angular resolution observations are needed to examine whether this discrepancy is maintained at smaller scales. In addition, using the CN, HCN, and HC_3N lines, we derived a $^{12}\text{C}/^{13}\text{C}$ isotopic ratio of 50 ± 5 .

Key words: astrochemistry – ISM: abundances – ISM: individual objects (OMC-2) – stars: formation – stars: low-mass – stars: protostars

1. Introduction

The solar system is the result of a long and complex process, several aspects of which still remain a mystery. One of these is the so-called “anomalous” $^{14}\text{N}/^{15}\text{N}$ value in the objects of the solar system (Caselli & Ceccarelli 2012; Hily-Blant et al. 2013). Based on the solar wind particles (Marty 2012), the Solar Nebula value is 441 ± 6 . However, $^{14}\text{N}/^{15}\text{N}$ is 272 in the Earth atmosphere (Marty 2012), around 140 in comets (Manfroid et al. 2009; Mumma & Charnley 2011; Rousselot et al. 2014; Shinnaka et al. 2014), and between 5 and 300 in meteorites (Busemann et al. 2006; Aléon 2010; Bonal et al. 2010). Consequently, solar system primitive objects as well as the terrestrial atmosphere are enriched with ^{15}N with respect to the presumed initial value. It has been long known that, similarly to the ^{15}N enrichment, the D/H ratio in terrestrial water is about 10 times larger than in the Solar Nebula, and this is very likely due to the conditions in the earliest phases of the solar system (see e.g., the reviews by Ceccarelli et al. 2014a and Cleaves et al. 2015). Because of this, the reason for the ^{15}N enrichment has been sought after in the chemical evolution of matter during the first steps of the solar system's formation (e.g., Terzieva & Herbst 2000; Rodgers & Charnley 2008; Wiström et al. 2012; Hily-Blant et al. 2013).

Several observations in Solar-like star-forming regions have been reported in the literature. In prestellar cores, $^{14}\text{N}/^{15}\text{N}$ varies between 70 and more than 1000 (Ikeda et al. 2002; Gerin et al. 2009; Milam & Charnley 2012; Bizzocchi et al. 2013; Daniel et al. 2013; Hily-Blant et al. 2013, 2017; Taniguchi & Saito 2017), between 150 and 600 in Solar-like Class 0 protostars (Gerin et al. 2009; Wampfler et al. 2014), and

between 80 and 160 in protoplanetary disks (Guzmán et al. 2015, 2017).

Whereas the $^{14}\text{N}/^{15}\text{N}$ values reported in the literature for prestellar cores, protostars, disks, and comets have been derived from the observations of half a dozen different species (CN, HCN, HNC, NH_3 , N_2H^+ , cyanopolyynes), it should be noted that, for each of these objects, only a few species were used each time.

Nonetheless, one has to consider that, rather than to the Solar Nebula $^{14}\text{N}/^{15}\text{N}$ value, these measurements should be compared to the current local interstellar $^{14}\text{N}/^{15}\text{N}$ ratio of ~ 300 , which results from cosmic evolution in the solar neighborhood (Hily-Blant et al. 2017; Romano et al. 2017 and references therein) and is, apparently by coincidence, very close to the terrestrial atmosphere value.

In order to understand the origin of the solar system ^{15}N enrichment, we need to measure it in objects that are as similar as possible to the Sun progenitor. To date, the best known analog of the Sun progenitor is represented by the source OMC-2 FIR4, in the Orion Molecular Complex at a distance of 420 pc (Hirota et al. 2007; Menten et al. 2007), north of the famous KL object. Several recent observations show that FIR4 is a young protocluster containing several protostars, some of which will eventually become suns (Shimajiri et al. 2008; López-Sepulcre et al. 2013; Furlan et al. 2014). In addition, OMC-2 FIR4 shows signs of the presence of one or more sources of energetic ≥ 10 MeV particles, the dose of which is similar to that measured in meteoritic material (Ceccarelli et al. 2014b; Fontani et al. 2017). In this article, we report the first measure of the $^{14}\text{N}/^{15}\text{N}$ ratio in OMC-2 FIR4, using different molecules: HC_3N , HCN, HNC, CN, and N_2H^+ .

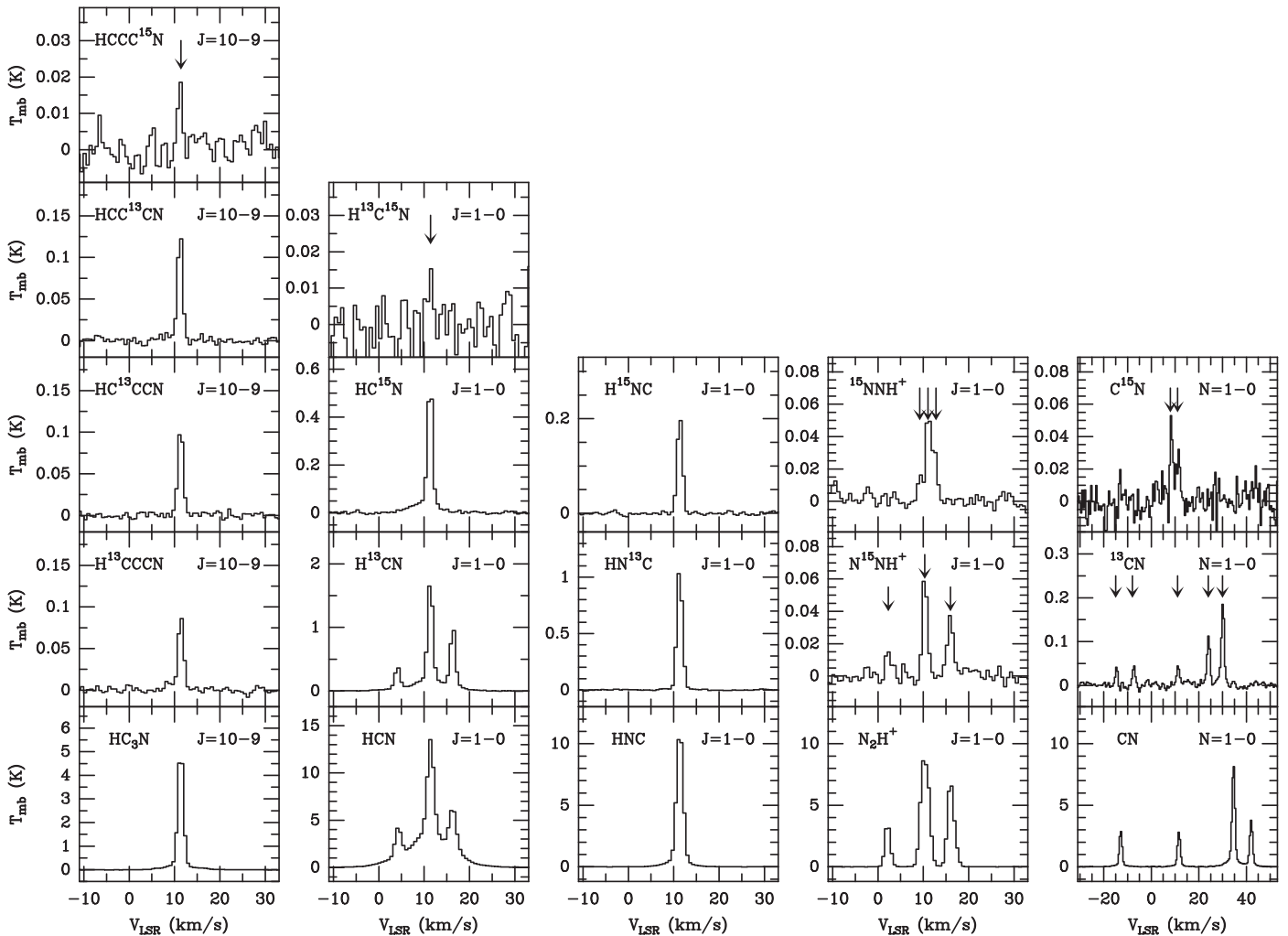


Figure 1. Representative sample of the observed spectra for the HC_3N , HCN , HNC , N_2H^+ , and CN families. The temperature scale is main beam temperature. Weak lines are indicated by an arrow.

2. Observations

We carried out an unbiased spectral survey of OMC-2 FIR4 (PI: Ana López-Sepulcre) in the 1, 2, and 3 mm bands with the IRAM 30 m telescope. The 3 mm (80.5–116.0 GHz) and 2 mm (129.2–158.8 GHz) bands were observed between 2011 August 31 and September 5, and on 2013 June 24. The 1 mm (202.5–266.0 GHz) range was observed on 2012 March 10–12 and on 2014 February 7. The Eight Mixer Receiver (EMIR) has been used, connected to the 195 kHz resolution Fourier Transform Spectrometer (FTS) units. The observations were conducted in wobbler switch mode, with a throw of $120''$. Pointing and focus measurements were performed regularly. The telescope half-power beam width (HPBW) is $21''$ – $30''$, $15''$ – $19''$, and $9''$ – $12''$ in the 1, 2, and 3 mm bands respectively. The package CLASS90 of the GILDAS software collection⁶ was used to reduce the data. The uncertainties of calibration are estimated to be lower than 10% at 3 mm and 20% at 2 and 1 mm. After subtraction of the continuum emission via first-order polynomial fitting, a final spectrum was obtained by stitching the spectra from each scan and frequency setting. The intensity was converted from antenna temperature (T_{ant}^*) to main beam temperature (T_{mb}) using the beam

efficiencies provided at the IRAM web site. The typical rms noise, expressed in T_{mb} unit, is 4–7 mK in the 3 mm band, 8–10 mK in the 2 mm band, and 15–25 mK in the 1 mm band.

3. Results

In order to derive the $^{14}\text{N}/^{15}\text{N}$ ratio in OMC-2 FIR4, we have looked for the ^{15}N -bearing substitutes of all the abundant N-bearing species present in the survey. We have clearly detected and identified the ^{15}N isotopologues of five species: HC_3N , HCN , HNC , CN , and N_2H^+ , most with more than one line, and tentatively detected one line of $\text{H}^{13}\text{C}^{15}\text{N}$. In addition, we have also included in our analysis the ^{13}C -bearing isotopes of HC_3N , HCN , HNC , and CN . A representative sample of these lines is shown in Figure 1, and all the observed lines are plotted in Appendix A in Figures 4–12.

The lines analysis and modeling presented here make use of several tools of the CASSIS package.⁷ Gaussian fits have been used to derive the lines' integrated intensities (called fluxes in the following) and kinematical properties. All lines are well fitted with narrow Gaussian components showing low

⁶ <http://www.iram.fr/IRAMFR/GILDAS/>

⁷ (Centre d'Analyse Scientifique de Spectres Instrumentaux et Synthétiques (CASSIS) is a line analysis and modeling software developed by IRAPUPS/CNRS (<http://cassis.irap.omp.eu>).

dispersions in central velocities and linewidths ($V_{\text{LSR}} = 11.3$ (0.1) km s⁻¹, FWHM = 1.4 (0.2) km s⁻¹). In addition, the strongest lines also show a slightly displaced broad component ($V_{\text{LSR}} = 10.8$ (0.3) km s⁻¹, FWHM = 6.4 (0.4) km s⁻¹ from

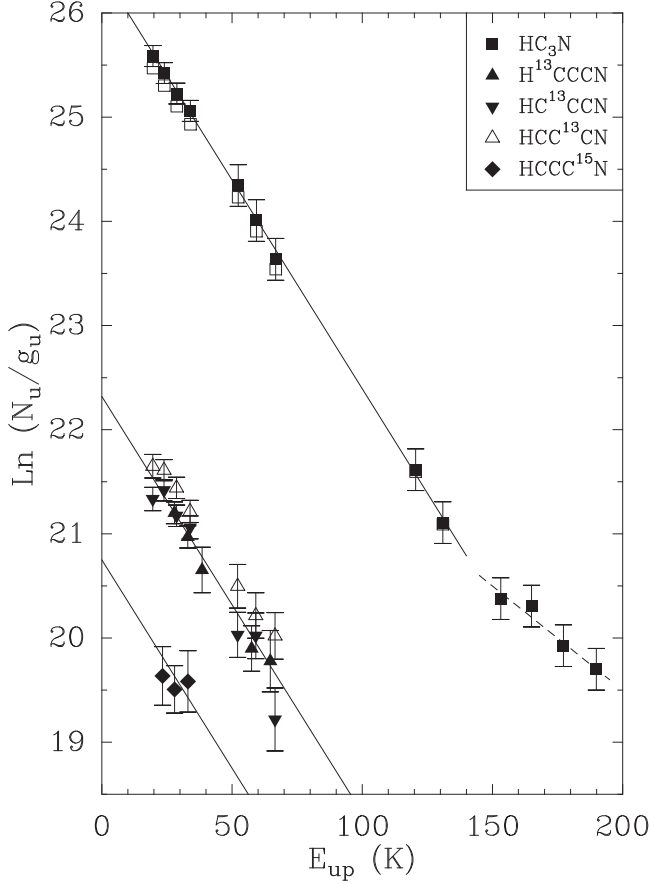


Figure 2. Population diagrams for HC₃N and its isotopologues. No beam dilution correction has been applied. The HC₃N data points are shown as black squares when corrected for opacity effects, and as empty squares without the opacity correction. The three solid lines correspond to a rotational temperature of 25 K; the dashed line corresponds to an excitation temperature of 50 K.

Gaussian fits), which is not detected in the ¹⁵N-bearing species, except in HC¹⁵N. Therefore, we have focused our work on the narrow component. Figures 4–12 of Appendix A show the Gaussian profiles superimposed to the observed lines. The fluxes reported in Tables 1–4 correspond to the narrow Gaussian components only, and the 1 sigma error bars include the fit and the calibration uncertainties.

For HCN, HNC, and N₂H⁺, our survey covers only the ($J = 1-0$) transitions of the ¹⁵N-bearing isotopologues. Thus, for these species, to derive of the ¹⁴N/¹⁵N ratio we used the “flux ratio method” applied to the ($J = 1-0$) transitions of the observed isotopologues. It leads to reliable abundance ratios provided that (i) the ($1-0$) transitions of the various isotopologues correspond to the same excitation temperature, (ii) the lines are not significantly affected by or can be corrected for opacity effects, and (iii) the emission size is the same for the various isotopologues.

For HC₃N and CN, because more than one rotational transition is observed for the ¹⁵N-bearing isotopologues, we could perform a local thermal equilibrium (LTE) modeling to derive the ¹⁴N/¹⁵N abundance ratios, as discussed for each species below.

Whenever possible, we have obtained direct ¹⁴N/¹⁵N measurements. In two cases, HCN and HNC, we have obtained indirect ¹⁴N/¹⁵N derivations from the less abundant isotopologues H¹³CN and HN¹³C, assuming a ¹²C/¹³C ratio.

In all cases, the isotopic ratios that we derive are beam-averaged values, at the scale of the largest HPBW of our observations ($\sim 30''$).

3.1. HC₃N

To rely on a coherent set of lines likely to sample the same gas, we have restricted our analysis to the narrow HC₃N emission, because broad emission from the ¹⁵N and the ¹³C isotopes of HC₃N is not detected in our survey.

To derive the column densities of the five species, we have performed a Population Diagram analysis of the lines reported in Table 1, which corrects iteratively for the opacity effects. Large-scale maps obtained recently with the IRAM 30 m

Table 1
Integrated Intensities from Gaussian Fits for HC₃N and Its Isotopologues

HC ₃ N				H ¹³ CCCN	HC ¹³ CCN	HCC ¹³ CN	HCCC ¹⁵ N
Transition ($J'-J$)	Frequency ^a (MHz)	E_{up} ^a (K)	$\int T_{\text{mb}} dv$ (K km s ⁻¹)	$\int T_{\text{mb}} dv$ (K km s ⁻¹)	$\int T_{\text{mb}} dv$ (K km s ⁻¹)	$\int T_{\text{mb}} dv$ (K km s ⁻¹)	$\int T_{\text{mb}} dv$ (K km s ⁻¹)
9–8	81881.5	19.6	7.1(0.7)	...	0.11(0.01)	0.16(0.02)	...
10–9	90979.0	24.0	7.4(0.7)	0.14(0.01)	0.16(0.02)	0.18(0.02)	0.024(0.007)
11–10	100076.4	28.8	7.3(0.7)	0.15(0.02)	0.16(0.02)	0.19(0.02)	0.028(0.006)
12–11	109173.6	34.0	7.4(0.7)	0.14(0.01)	0.17(0.02)	0.18(0.02)	0.030(0.010)
13–12	114615.0 ^b	38.5 ^b	...	0.11(0.01)	...	...	...
15–14	136464.4	52.4	5.7(1.1)	...	0.10(0.02)	0.13(0.03)	...
16–15	145561.0	59.4	4.7(0.9)	0.08(0.02)	0.10(0.02)	0.12(0.02)	...
17–16	154657.3	66.8	3.7(0.7)	0.08(0.02)	0.06(0.01)	0.13(0.03)	...
24–23	218324.7	130.9	0.6(0.1)	...	...	...	...
26–25	236512.8	153.2	0.36(0.07)	...	...	...	...
27–26	245606.3	165.0	0.37(0.07)	...	...	...	...
29–28	263792.3	189.9	0.22(0.04)	...	...	...	...

Notes. The integrated intensities correspond to the narrow Gaussian components only, and the 1σ error bars include the fit and the calibration uncertainties.

^a We report in this table only the main isotopologue frequencies and upper-level energies, except for the (13–12) transition, which falls outside the observed frequency range.

^b These values correspond to the H¹³CCCN isotopologue, the only one for which the (13–12) transition is covered by our observations.

Table 2
Integrated Intensities from Gaussian Fits for HCN and Its Isotopologues

Transition ($J'_F - J_F$) or ($J' - J$) ^a	HCN				H ¹³ CN	HC ¹⁵ N	H ¹³ C ¹⁵ N
	Frequency ^b (MHz)	E_{up} ^b (K)	Component ^c	$\int T_{\text{mb}} dv$ ^d (K km s ⁻¹)	$\int T_{\text{mb}} dv$ ^d (K km s ⁻¹)	$\int T_{\text{mb}} dv$ ^d (K km s ⁻¹)	$\int T_{\text{mb}} dv$ ^d (K km s ⁻¹)
1 ₁ –0 ₁	88630.4	4.3	N	8.4(0.9)	1.3(0.1)	...	...
1 ₂ –0 ₁ or 1–0	88631.9	4.3	N	13(2)	2.2(0.2)	0.70(0.07)	0.020(0.008)
			B	50 (5)	2.0 (0.2)	0.34 (0.03)	...
1 ₀ –0 ₁	88633.9	4.3	N	5.2(0.5)	0.44(0.06)	...	...
3–2	265886.4	25.2	N	24(1)	1.3(0.3)	...	...
			B	85(18)	7.2(1.5)	...	...

Notes.

^a For the ¹⁵N-bearing isotopologues and the broad Gaussian component, the reported transition is the ($J = 1-0$) one.

^b We report in this table only the main isotopologue frequencies and upper-level energies.

^c Narrow (N) and broad (B) Gaussian components fitted to the observed spectra, with no attempt to distinguish the hyperfine components in the broad emission.

^d The 1σ error bars include the fit and the calibration uncertainties.

telescope (J. Al-Edhari et al. 2017, in preparation) show that the cyanopolyne emission is extended. Thus, no beam dilution correction has been applied. The results of the analysis are shown in Figure 2.

The HC₃N lines diagram suggests the existence of two components responsible for the narrow emission. After correction of their (moderate) opacity (0.20–0.25), the HC₃N lines with upper-level energy E_{up} lower than 130 K are compatible with an LTE excitation at a rotational temperature of 25 K, in agreement with the values derived from the ¹³C isotopes, and an HC₃N column density of $3.3 \times 10^{13} \text{ cm}^{-2}$. Assuming the same excitation temperature for the HCCC¹⁵N lines, we obtain a column density of $1.2 \times 10^{11} \text{ cm}^{-2}$ and a direct determination of the ¹⁴N/¹⁵N abundance ratio of 275 ± 65 . For the three ¹³C-bearing isotopologues of HC₃N, we derive, in the same way, ¹²C/¹³C abundance ratios of 57 ± 7 , 59 ± 11 , and 46 ± 6 . We have checked that an LTE model based on these parameters is perfectly coherent with the non-detection of the isotopologues lines at $E_{\text{up}} > 70 \text{ K}$.

The HC₃N lines with $E_{\text{up}} > 150 \text{ K}$ suggest the existence of a warmer component, but its analysis is out of the scope of this article. An extended and detailed modeling of the cyanopolyynes emission in OMC-2 FIR4, relying on a more complete set of data and including the broad emission, is presented in another study (J. Al-Edhari et al. 2017, in preparation).

3.2. HCN

The lines of HCN and of its ¹³C- and ¹⁵N- bearing isotopologues show a narrow and a broad emission. We have fitted them by Gaussian profiles (see Table 2 and Figure 9).

3.2.1. Narrow Emission

The ($J = 1-0$) spectra of HCN and H¹³CN split into three hyperfine components, which provides a measure of the line opacity when comparing the relative intensities of the hyperfine components. Under LTE optically thin conditions, this ratio is 1:3:5. Relative to the weakest component, the observed flux ratios obtained for HCN and H¹³CN are, respectively, $1:(1.6 \pm 0.2)$; (2.6 ± 0.5) and $1:(2.9 \pm 0.5)$; (5.1 ± 0.9) (see Table 5). It suggests that the H¹³CN hyperfine components have the same excitation

temperature and are optically thin, whereas the HCN lines suffer from significant opacity or/and anomalous excitation effects, which will prevent a direct determination of the ¹⁴N/¹⁵N ratio. Neglecting the weak differences of line frequencies, the double isotopic ratio ¹³C¹⁴N/¹²C¹⁵N is then simply equal to the ratio of the total fluxes, obtained by adding the hyperfine component contributions: 5.6 ± 0.8 . This may lead to an indirect determination of the ¹⁴N/¹⁵N ratio, assuming a ¹²C/¹³C ratio. In addition, the H¹³C¹⁵N (1–0) line being tentatively detected (at 2.5 sigmas, see Figures 1 and 9), the comparison with the H¹³C¹⁴N (1–0) total flux provides a direct measurement of the ¹⁴N/¹⁵N ratio, equal to 200 ± 85 , whereas the comparison with the HC¹⁵N (1–0) flux gives a ¹²C/¹³C ratio of 36 ± 15 . It should be noted that these ratios may be somewhat underestimated, the H¹³C¹⁵N (1–0) line being only tentatively detected, which tends to an overestimation of its flux.

In addition, the Population Diagram built with the $J = (1-0)$ and $(2-1)$ fluxes of H¹³CN leads to an excitation temperature of 6 K, if the emission is assumed to be more extended than the largest beam (30''), a reasonable assumption for such cold gas.

3.2.2. Broad Emission

The broad emission is particularly evident in the ($J = 2-1$) spectra of HCN and H¹³CN, but it is also detected in the ($J = 1-0$) emission, even for HC¹⁵N. We have estimated the corresponding ¹³C¹⁴N/¹²C¹⁵N double isotopic ratio from the ($J = 1-0$) line flux ratio. This value, 6.0 ± 1.5 , shows a quite large uncertainty but is fully compatible with the ratio derived from the narrow emission. A Population Diagram applied to the H¹³CN ($J = 1-0$) and ($J = 2-1$) fluxes, with no beam dilution, leads to an excitation temperature of 22 K, which suggests that this broad emission could come from a warmer component than the narrow one.

3.3. HNC

The broad emission, which is marginally visible in the HNC (1–0) and HN¹³C (2–1) spectra only, has not been included in our analysis (see Table 3 and Figure 10). The HNC hyperfine structure is too narrow to be spectrally resolved in our data. However, a Population Diagram analysis applied to the

Table 3
Integrated Intensities from Gaussian Fits for HNC, N₂H⁺ and Their Isotopologues

HNC				HN ¹³ CN		H ¹⁵ NC	
Transition ($J'_F - J_F$)	E_{up} (K)	Frequency (MHz)	$\int T_{\text{mb}} dv$ (K km s ⁻¹)	Frequency (MHz)	$\int T_{\text{mb}} dv$ (K km s ⁻¹)	Frequency (MHz)	$\int T_{\text{mb}} dv$ (K km s ⁻¹)
1–0	4.4	90663.6	21(2)	87090.9	1.7(0.2)	88865.7	0.31(0.01)
3–2	25.1	...	...	261263.3	1.09(0.04)	...	...

N ₂ H ⁺				¹⁵ NNH ⁺		N ¹⁵ NH ⁺	
Transition ($J'_F - J_F$)	E_{up} (K)	Frequency (MHz)	$\int T_{\text{mb}} dv$ (K km s ⁻¹)	Frequency ^a (MHz)	$\int T_{\text{mb}} dv$ (K km s ⁻¹)	Frequency ^a (MHz)	$\int T_{\text{mb}} dv$ (K km s ⁻¹)
1 ₁ –0 ₁	4.5	93171.9	13.1(1.5)	90263.5	0.08(0.01)	91204.3	0.059(0.008)
1 ₂ –0 ₁	4.5	93173.7	21.3(2.5)	90263.9	0.036(0.005)	91206.0	0.09(0.02)
1 ₀ –0 ₁	4.5	93176.1	5.3(0.5)	90264.5	0.015(0.004)	91208.5	0.025(0.006)

Note. The integrated intensities correspond to the narrow Gaussian components only, and the 1 σ error bars include the fit and the calibration uncertainties.

^a Frequencies from Dore et al. (2009).

($J = 1-0$) and ($2-1$) lines of HN¹³C, assuming extended emission, indicates moderate opacities (~ 0.2) and leads to an excitation temperature of 7–8 K. Such a low excitation temperature appears coherent with the assumption of extended emission.

As the HNC line is certainly more severely affected by opacity effect, a direct derivation of the ¹⁴N/¹⁵N ratio is not possible. From the flux ratio of ¹³C- and ¹⁵N- bearing species, the double isotopic ratio ¹³C¹⁴N/¹²C¹⁵N equal to 5.4 ± 0.8 , becomes 6.0 ± 0.8 when the opacity correction is applied to the HN¹³C line flux.

3.4. N₂H⁺

The main isotope and each of the ¹⁵N-bearing substitutes of N₂H⁺ show three hyperfine components, with relative intensities of 1:3:5 if LTE optically thin emission applies (see Table 5). The frequencies of the hyperfine lines of ¹⁵NNH⁺ and N¹⁵NH⁺ have been presented by Dore et al. (2009). The observed hyperfine flux ratios are 1:(2.5 $\pm$ 0.4):(4.1 $\pm$ 0.6) for the main isotopologue, 1:(2.4 $\pm$ 0.7):(5.0 $\pm$ 1.5) for ¹⁵NNH⁺, and 1:(2.4 $\pm$ 0.6):(3.8 $\pm$ 1.2) for N¹⁵NH⁺. We conclude that, as expected, the emission of the ¹⁵N-bearing species is optically thin and that the line opacity is very moderate for the main isotopologue components. Assuming that the weakest line of N₂H⁺ (1–0) is optically thin, we can estimate the opacity-corrected fluxes of the two others. With such a correction, the ¹⁴N/¹⁵N ratios derived from the total fluxes are 320 ± 60 from ¹⁵NNH⁺ and 240 ± 50 from N¹⁵NH⁺.

3.5. CN

The CN family members present extremely rich rotational spectra, combining fine and hyperfine structure interactions and our survey covers both the ($N = 1-0$) and ($N = 2-1$) transitions for the three isotopologues (see Figures 11 and 12).

In addition, the main isotopologue shows a broad emission, more visible on the ($N = 2-1$) transitions than on the ($N = 1-0$) ones.

Most of the CN ($N = 1-0$) and ($N = 2-1$) hyperfine components reported in the Cologne Database for Molecular Spectroscopy (CDMS) and the NASA Jet Propulsion Laboratory (JPL) database are easily detected. We have compared the observed flux ratios with the theoretical ratios (proportional to the g_{up} , A_{ij} ratios, the slight frequency differences being neglected). The results (see Table 5) suggest that the hyperfine components follow an intensity distribution very close to LTE and that the line opacities are moderate.

The same analysis shows that the ¹³CN and C¹⁵N ($N = 1-0$) and ($N = 2-1$), which are clearly detected and do not suffer from blending, follow an intensity distribution very close to LTE and that the lines are optically thin.

We have thus performed a simultaneous LTE modeling of the ($N = 1-0$) and ($N = 2-1$) transitions for the three isotopologues. For ¹³CN and C¹⁵N we have assumed that the emission comes from a single extended component with kinematical properties that are derived from the Gaussian fits ($V_{\text{LSR}} = 11.4$ km s⁻¹, FWHM = 1.3 km s⁻¹). For CN, to account for the broad emission, we have added a second component ($V_{\text{LSR}} = 11.1$ km s⁻¹, FWHM = 6.9 km s⁻¹). The free parameters of our modeling, performed with a Markov Chain Monte-Carlo (MCMC) minimization to obtain the best fit to the lines, were the excitation temperature T_{ex} and the column densities of the three isotopologues for the narrow emission component, the source size, the CN column density, and the excitation temperature for the broad CN emission component.

For the first component, the best fit was obtained with the following parameters: $T_{\text{ex}} = 8 \pm 1$ K, $N(\text{CN}) = (3.5 \pm 0.5) \times 10^{14}$ cm⁻², $N(^{13}\text{CN}) = (8 \pm 1) \times 10^{12}$ cm⁻², $N(\text{C}^{15}\text{N}) = (1.3 \pm 0.2) \times 10^{12}$ cm⁻². It corresponds to the following isotopic ratios: ¹³C¹⁴N/¹²C¹⁵N = 6.2 ± 1.3 , ¹²C/¹³C = 44 ± 8 , and ¹⁴N/¹⁵N = 270 ± 60 .

For the broad component, there are three free parameters and the best fit solution is degenerate. However, the excitation temperature depends only weakly on the assumed size and is between 50 and 60 K.

The calculated profiles are superimposed to the observed ones in Figures 11 and 12.

Table 4
Integrated Intensities from Gaussian Fits for CN and Its Isotopologues

Species	Transition ^a $N_{J'F'1F'} - N_{JF1F}$ or $N_{J'F'} - N_{JF}$	Frequency (MHz)	E_{up} (K)	V_{LSR}^b (km s ⁻¹)	FWHM ^b (km s ⁻¹)	$\int T_{\text{mb}} dv^b$ (K km s ⁻¹)
CN	$1_{01/21/2} - 0_{01/21/2}$	113123.37	5.43	11.6(0.3)	1.4(0.3)	0.9(0.1)
	$1_{01/21/2} - 0_{01/23/2}$	113144.19	5.43	11.7(0.3)	1.5(0.3)	4.9(0.5)
	$1_{01/23/2} - 0_{01/21/2}$	113170.54	5.43	11.7(0.3)	1.5(0.3)	5.3(0.5)
	$1_{01/23/2} - 0_{01/23/2}$	113191.33	5.43	11.7(0.3)	1.5(0.3)	5.6(0.6)
	$1_{03/23/2} - 0_{01/21/2}$	113488.14	5.45	11.7(0.3)	1.5(0.3)	6.0(0.7)
	$1_{03/25/2} - 0_{01/23/2}$	113490.99	5.45	11.7(0.3)	1.6(0.3)	14.0(1.6)
	$1_{03/21/2} - 0_{01/21/2}$	113499.64	5.45	11.6(0.3)	1.4(0.3)	4.3(0.4)
	$1_{01/21/2} - 0_{01/23/2}$	113508.93	5.45	11.7(0.3)	1.4(0.3)	4.5(0.5)
	$1_{03/21/2} - 0_{01/23/2}$	113520.42	5.45	11.6(0.3)	1.3(0.3)	0.8(0.1)
	$2_{03/21/2} - 1_{03/21/2}$	226287.43	16.31	11.3(0.1)	1.1(0.1)	0.5(0.1)
	$2_{03/21/2} - 1_{03/23/2}$	226298.92	16.31	11.0(0.2)	1.6(0.2)	0.6(0.2)
	$2_{03/23/2} - 1_{03/21/2}$	226303.08	16.31	11.3(0.1)	1.2(0.1)	0.5(0.1)
	$2_{03/23/2} - 1_{03/23/2}$	226314.54	16.31	11.2(0.1)	1.1(0.1)	1.0(0.2)
	$2_{03/23/2} - 1_{01/25/2}$	226332.54	16.31	11.3(0.1)	1.1(0.1)	0.5(0.1)
	$2_{03/25/2} - 1_{03/23/2}$	226341.93	16.31	11.3(0.1)	1.1(0.1)	0.7(0.1)
	$2_{03/25/2} - 1_{03/25/2}$	226359.87	16.31	11.3(0.1)	1.2(0.1)	2.6(0.6)
	$2_{03/21/2} - 1_{01/23/2}$	226616.56	16.31	11.3(0.1)	1.1(0.1)	0.5(0.1)
	$2_{03/23/2} - 1_{01/23/2}$	226632.19	16.31	11.3(0.1)	1.3(0.1)	3.2(0.7)
	$2_{03/25/2} - 1_{01/23/2}$	226659.58	16.31	11.4(0.1)	1.6(0.1)	8.3(1.7)
	$2_{03/21/2} - 1_{01/21/2}$	226663.70	16.31	11.4(0.1)	1.4(0.1)	3.2(0.7)
	$2_{03/23/2} - 1_{01/21/2}$	226679.38	16.31	11.4(0.1)	1.3(0.1)	3.6(0.8)
	$2_{05/25/2} - 1_{03/25/2}$	226892.12	16.34	11.3(0.1)	1.2(0.1)	3.2(0.7)
	$2_{05/23/2} - 1_{03/25/2}$	226905.38	16.34	11.3(0.2)	0.9(0.2)	0.10(0.03)
¹³ CN	$1_{1/201} - 0_{1/211}$	108412.86	5.23	11.0(0.4)	1.1(0.4)	0.03(0.01)
	$1_{1/201} - 0_{1/212}$	108426.89	5.23	11.4(0.3)	1.3(0.3)	0.07(0.01)
	$1_{3/210} - 0_{1/201}$	108631.12	5.21	11.7(0.4)	1.3(0.4)	0.04(0.02)
	$1_{3/211} - 0_{1/201}$	108636.92	5.21	11.3(0.3)	1.4(0.3)	0.11(0.02)
	$1_{3/212} - 0_{1/201}$	108651.30	5.21	11.5(0.3)	1.4(0.3)	0.18(0.03)
	$1_{1/212} - 0_{1/212}$	108657.65	5.24	11.4(0.3)	1.3(0.3)	0.13(0.02)
	$1_{3/223} - 0_{1/212}$	108780.20	5.25	11.4(0.3)	1.5(0.3)	0.27(0.05)
	$1_{3/222} - 0_{1/211}$	108782.37	5.25	11.4(0.3)	1.5(0.3)	0.16(0.03)
	$1_{3/221} - 0_{1/210}$	108786.98	5.25	11.3(0.3)	1.5(0.3)	0.06(0.01)
	$1_{3/221} - 0_{1/211}$	108793.75	5.25	11.3(0.3)	1.4(0.3)	0.06(0.01)
	$1_{3/222} - 0_{1/212}$	108796.40	5.25	11.4(0.4)	1.1(0.4)	0.05(0.02)
	$2_{5/233} - 1_{3/222}$	217467.15	15.69	11.3(0.1)	1.3(0.1)	0.36(0.08)
	$2_{5/232} - 1_{3/221}$	217469.15	15.69	11.2(0.1)	1.8(0.1)	0.23(0.05)
C ¹⁵ N	$1_{1/21} - 0_{1/21}$	109689.61	5.27	11.1(0.4)	2.1(0.4)	0.05(0.02)
	$1_{3/21} - 0_{1/20}$	110023.54	5.28	11.5(0.3)	1.3(0.3)	0.05(0.02)
	$1_{3/22} - 0_{1/21}$	110024.59	5.28	11.4(0.4)	1.3(0.4)	0.08(0.02)
	$2_{3/22} - 1_{1/21}$	219722.49	15.81	11.3(0.2)	1.3(0.2)	0.08(0.02)
	$2_{5/22} - 1_{3/21}$	219934.04	15.84	11.4(0.1)	1.7(0.1)	0.10(0.03)

Notes.

^a For CN and ¹³CN, according to the CDMS convention, the quantum numbers are N, J, F_1, F with $F_1 = J + I_1$, $F = F_1 + I_2$ where I_1 is the ¹²C or ¹³C nuclear spin and I_2 that of ¹⁴N. For C¹⁵N, the quantum numbers are N, J, F with $J = N + S$ and $F = J + I$, where S and I are, respectively, the electronic spin and the nuclear spin of ¹⁵N.

^b The kinematical parameters and the integrated intensities correspond to the narrow Gaussian components only, and the 1σ error bars include the fit and the calibration uncertainties.

4. Discussion and Conclusions

4.1. The ¹⁴N/¹⁵N Ratio toward OMC-2 FIR4

We have reported here a complete census of the ¹⁴N/¹⁵N ratio in the most abundant N-bearing species toward the protocluster OMC-2 FIR4. The five ¹⁴N/¹⁵N ratios derived directly from the line fluxes (HC₃N, CN, N₂H⁺, H¹³CN)

appear very similar. Their weighted average is 260 ± 30 . Two indirect ¹⁴N/¹⁵N ratio derivations, obtained with HCN and HNC, can be made from the double isotopic ratios ¹³C¹⁴N/¹²C¹⁵N. Discussing the ¹²C/¹³C isotopic ratio is outside of the scope of this study. However, it is notable that one of the ¹³C-bearing isotopologues of HC₃N seems to show systematically stronger lines than the two others (see Figure 2);

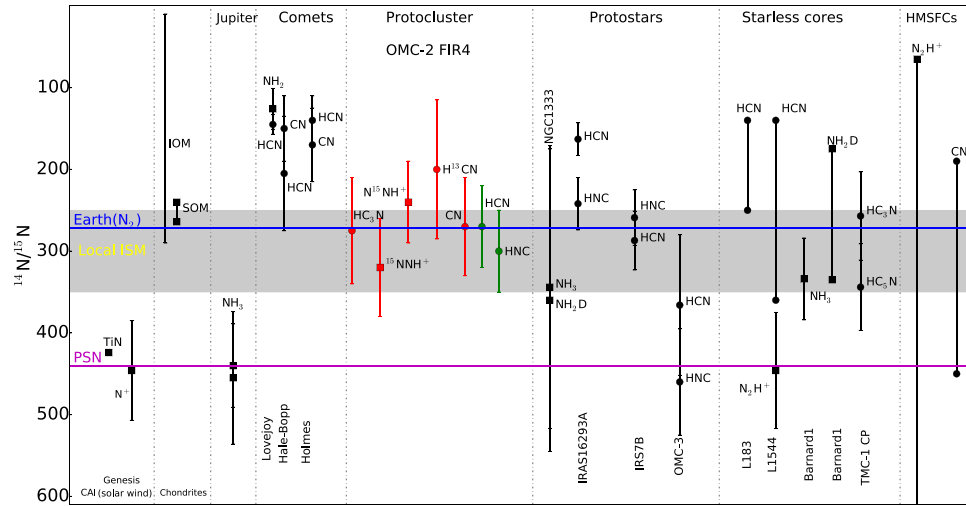


Figure 3. Nitrogen isotopic composition of solar system objects compared with different sources (figure adapted from Hily-Blant et al. 2013). The blue horizontal line represents the nitrogen isotopic composition of the terrestrial atmosphere ($^{14}\text{N}/^{15}\text{N} = 272$), while the magenta horizontal line indicates the protosolar nebula (PSN) value of 441 ± 6 Marty (2012). Square and circle symbols correspond to amine and nitrile functional groups, respectively. IOM stands for “insoluble organic matter,” SOM for “soluble organic matter,” and CAI for “Calcium-, Aluminum-rich inclusions.” Red and green symbols represent, respectively, direct and indirect measurements of $^{14}\text{N}/^{15}\text{N}$ in OMC-2 FIR4, assuming $^{12}\text{C}/^{13}\text{C} = 50 \pm 5$ (this work). Protostars data are from Wampfler et al. (2014), starless cores data from Hily-Blant et al. (2013) and Taniguchi & Saito (2017), and high-mass star-forming cores data from Fontani et al. (2015; where HMSFC stands for “high-mass starless core”). The local galactic ratio range, shown as a gray zone, takes into account the measurements reported in Adande & Ziurys (2012) and Hily-Blant et al. (2017).

however, this requires complementary observations to be confirmed and discussed. Although far from complete in terms of $^{12}\text{C}/^{13}\text{C}$ measurements, our data allow us to derive five direct estimates of this isotopic ratio and its weighted average value is 50 ± 5 . The resulting indirect $^{14}\text{N}/^{15}\text{N}$ ratios are, respectively, 270 ± 50 and 290 ± 50 for HCN and HNC. These values are in remarkable agreement with the direct determinations. The total weighted average, including both direct and indirect measures, is 270 ± 30 . All of the $^{14}\text{N}/^{15}\text{N}$ ratios derived in OMC-2 FIR4 are plotted in Figure 3.

As discussed for each species, the narrow emission of HCN, HNC, and CN, from which the $^{14}\text{N}/^{15}\text{N}$ ratio is measured, shows a very low excitation temperature (6–9 K). Without large-scale emission maps for these species, we cannot firmly establish that they trace extended parent gas of the protocluster, but it seems a reasonable interpretation. HC_3N narrow emission traces at least two components: a relatively cold gas ($T_{\text{ex}} = 25$ K) that our recent observations obtained with the 30 m telescope show to be extended, and a warmer component ($T_{\text{ex}} \sim 50$ K). A more sophisticated analysis of the HC_3N emission, also including the broad emission, will be presented in a forthcoming article (J. Al-Edhari et al. 2017, in preparation). In addition, we also tentatively measured with the HCN isotopologues the $^{14}\text{N}/^{15}\text{N}$ ratio in the broad emission, detected for the main isotopologue of all the species studied here, except N_2H^+ . This ratio appears perfectly compatible with the value derived from the narrow emission. Our data suggest that this broad emission is warmer than the narrow one, but do not allow for an estimation of the emission size. We hope that the interferometric (ALMA and NOEMA) data that we will soon obtain toward this source will allow to understand the nature of this emission.

4.2. Comparison with Other Galactic Sources

Measurements of the $^{14}\text{N}/^{15}\text{N}$ ratio in starless cores (e.g., Hily-Blant et al. 2013) and in protostars (e.g., Wampfler et al. 2014) seem to indicate (see Figure 3) that the ratios

derived from molecules carrying the amine functional group (NH_3 , N_2H^+) are larger than the ratios derived from molecules carrying the nitrile functional group (CN, HCN); a chemical origin is proposed for this effect. However, none of the studied sources show a complete set of measurements from different tracers, so it is very difficult to distinguish between variations from source to source and from molecule to molecule. On the other hand, our results, which rely on a set of five different species that trace the same (cold extended) gas, and that belong to the nitrile and amine families, do not show any significant difference. In contrast, they are very similar, and they agree remarkably with the present local $^{14}\text{N}/^{15}\text{N}$ galactic ratio of ~ 300 as derived from observations (Adande & Ziurys 2012; Hily-Blant et al. 2017 and references therein) and predicted by models of galactic CNO evolution (e.g., Romano et al. 2017).

Our observations show that in OMC-2 FIR4, which is the best analog of the Sun progenitor, there is no ^{15}N fractionation, compared to the current value at the same (8 kpc) galactic center distance. This is in agreement with the recent model predictions by Roueff et al. (2015).

In conclusion, the presented measurements of $^{14}\text{N}/^{15}\text{N}$ seem to be at odds with the previous measurements in prestellar cores and protostars (see Figure 3), which, depending on the species used, suggest ^{15}N enrichment or deficiency with respect to the local ISM value. It is possible that this discrepancy is due to the different spatial scale probed by our and other observations. Specifically, while the cold and large-scale gas might not be ^{15}N enriched, local smaller-scale clumps might present this enrichment (or deficiency). Ongoing interferometric observations toward OMC-2 FIR4 will verify this possibility. If this is the case, the enrichment of the solar system bodies could find an explanation in the ISM chemistry. If, on the contrary, the new observations would confirm the absence of ^{15}N enrichment also at small scales, then the ^{15}N enrichment in solar system bodies must have another nature.

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Software: CASSIS (Vastel et al. 2015), GILDAS (Pety 2005).

Appendix A Observed Lines

A representative sample of these lines has been shown in Figure 1. All of the observed lines are plotted in Figures 4–12. The temperature scale is main beam temperature.

A. OBSERVED LINES

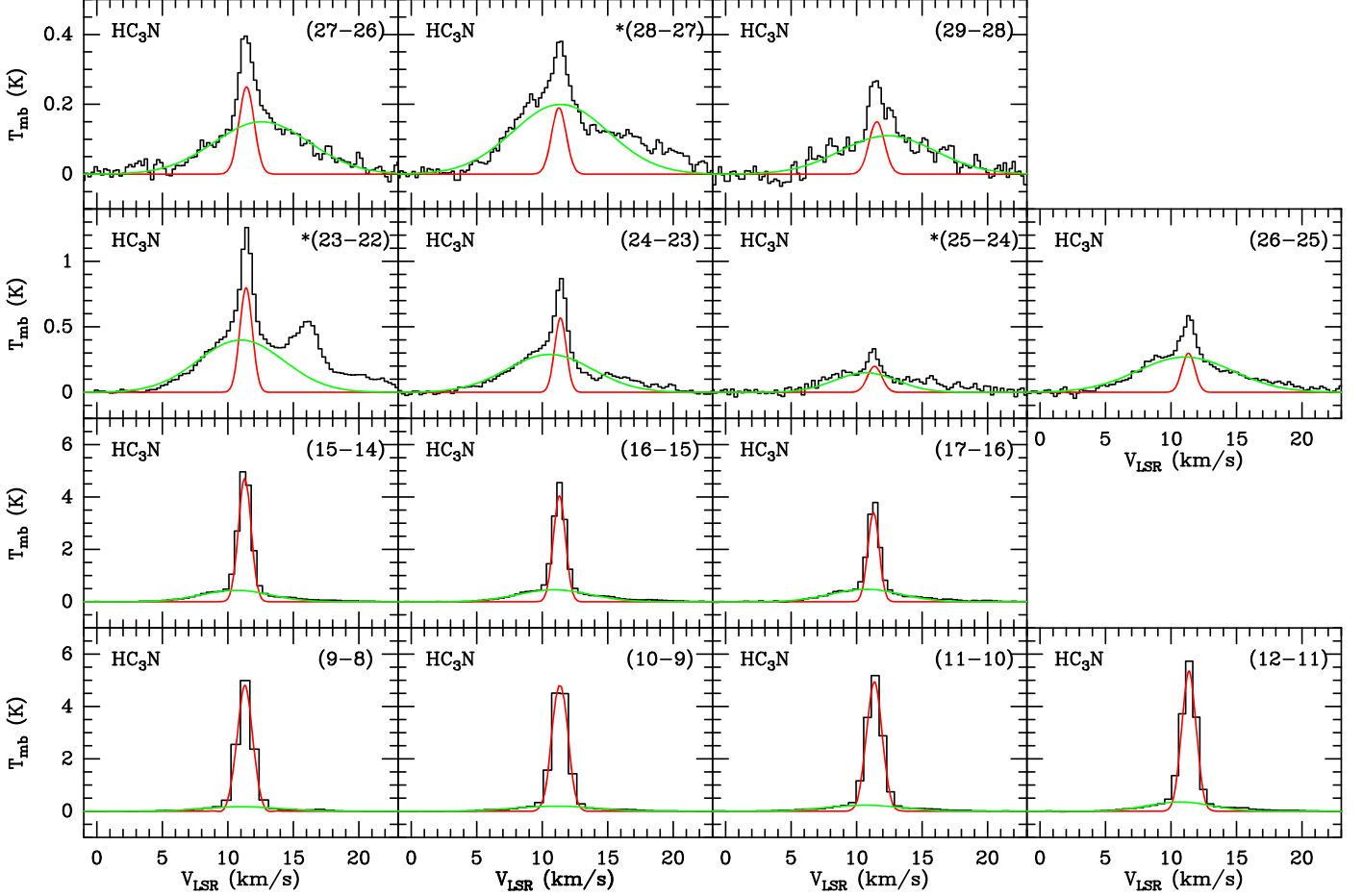


Figure 4. HC_3N observed spectra and two components' Gaussian fits to the lines. The temperature scale is main beam temperature. Due to significant overlapping with other emissions or a calibration problem, the lines indicated by an asterisk (*) have neither been included in our analysis nor reported in Table 1.

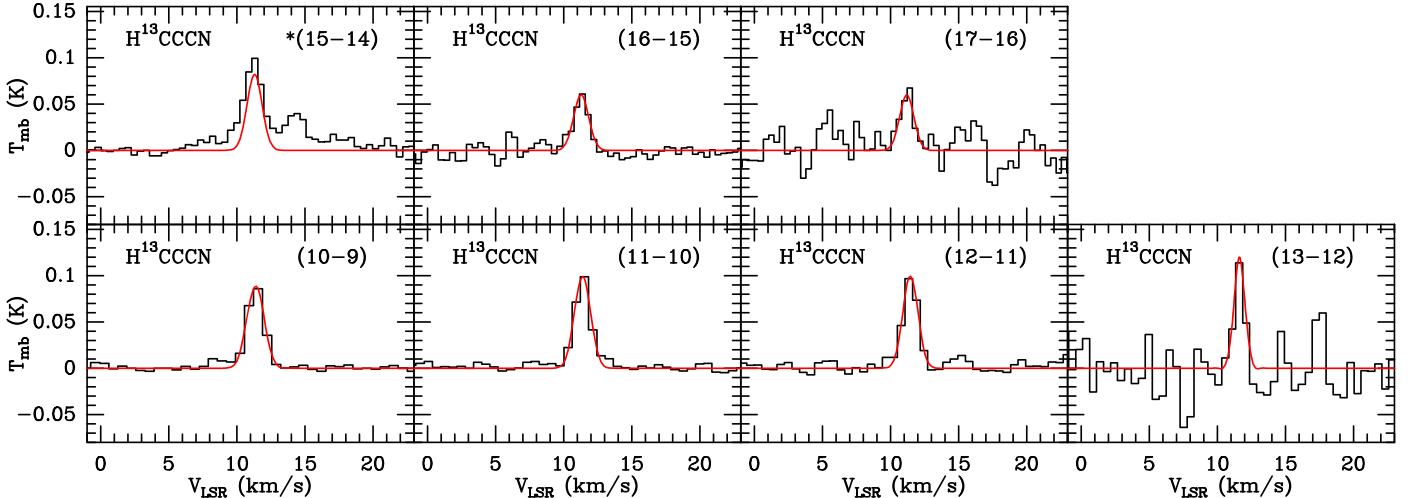


Figure 5. H^{13}CCCN observed spectra and Gaussian fit to the lines. The temperature scale is main beam temperature. Due to significant overlapping with other emissions, the line indicated by an asterisk (*) has neither been included in our analysis nor reported in Table 1.

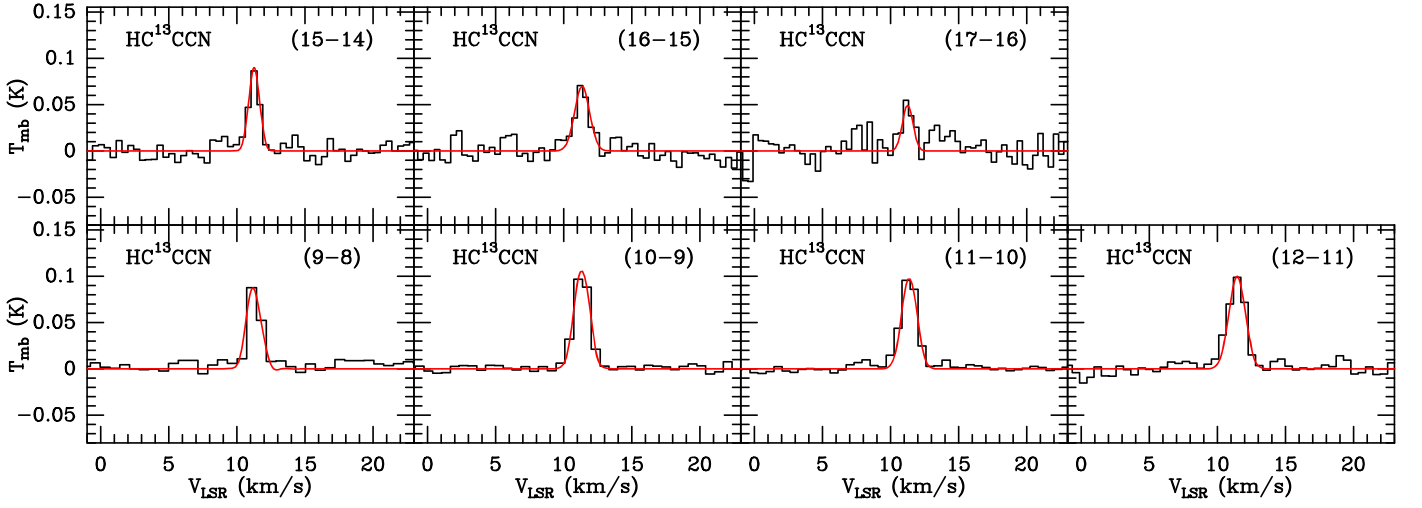


Figure 6. HC^{13}CCN observed spectra and Gaussian fit to the lines. The temperature scale is main beam temperature.

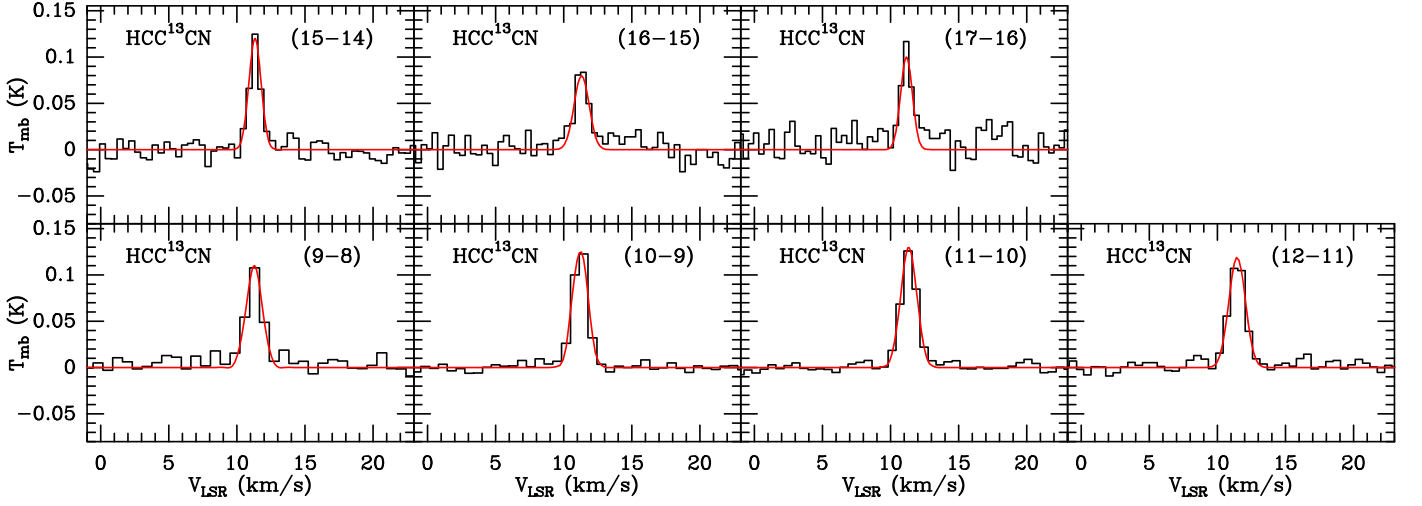


Figure 7. HCC^{13}CN observed spectra and Gaussian fit to the lines. The temperature scale is main beam temperature.

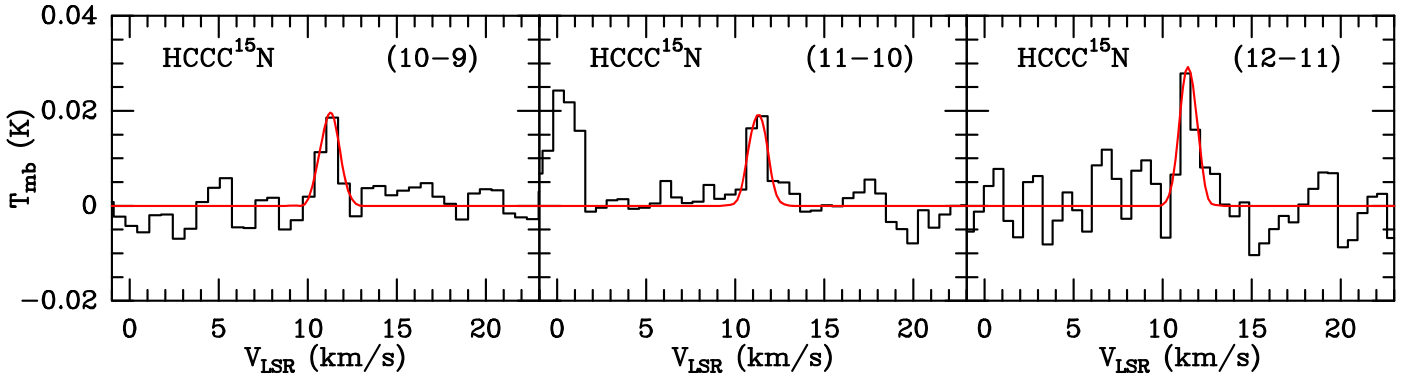


Figure 8. HCCC^{15}N observed spectra and Gaussian fit to the lines. The temperature scale is main beam temperature.

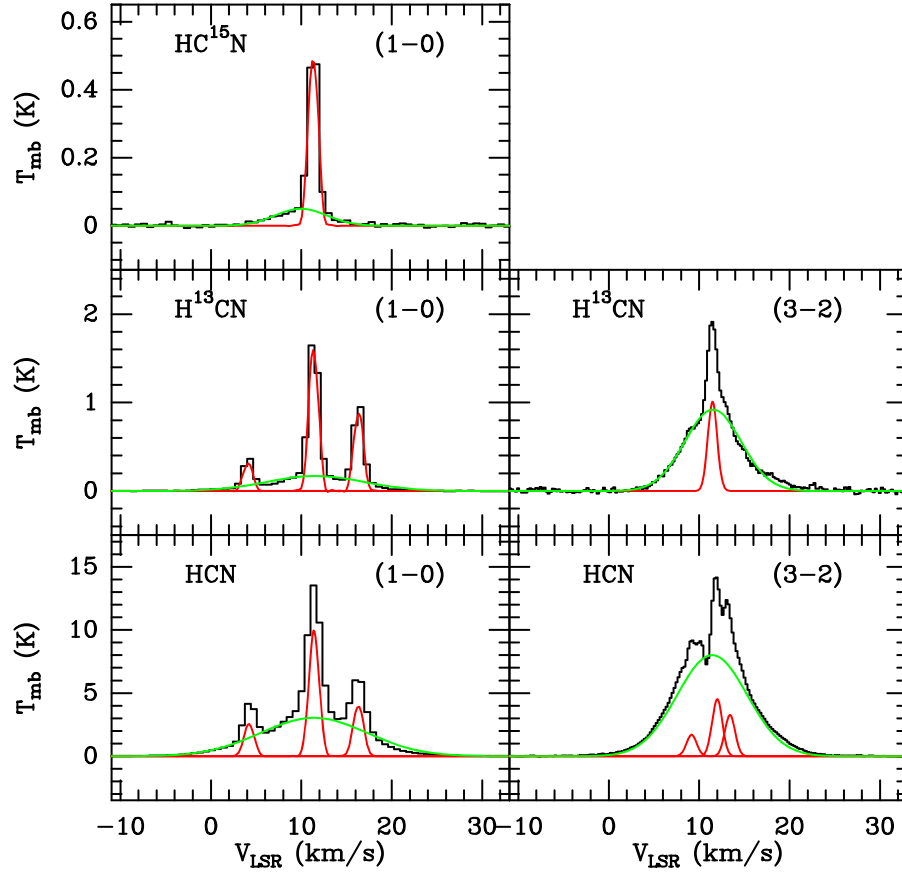


Figure 9. HCN and its isotopologues observed spectra. Two components' Gaussian fits to the lines are superimposed onto the spectra. The temperature scale is main beam temperature.

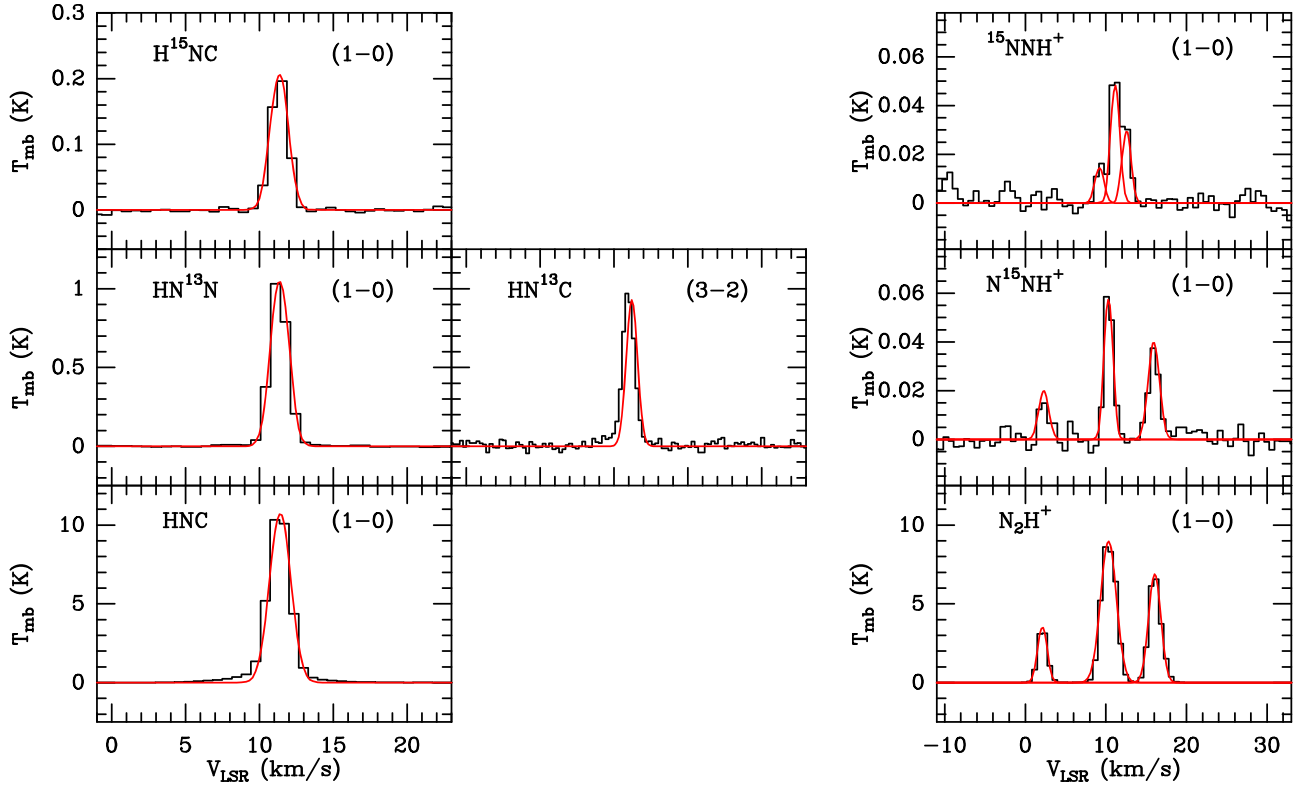


Figure 10. HNC, N_2H^+ and their isotopologues observed spectra. Gaussian fits to the lines are superimposed onto the spectra. The temperature scale is main beam temperature.

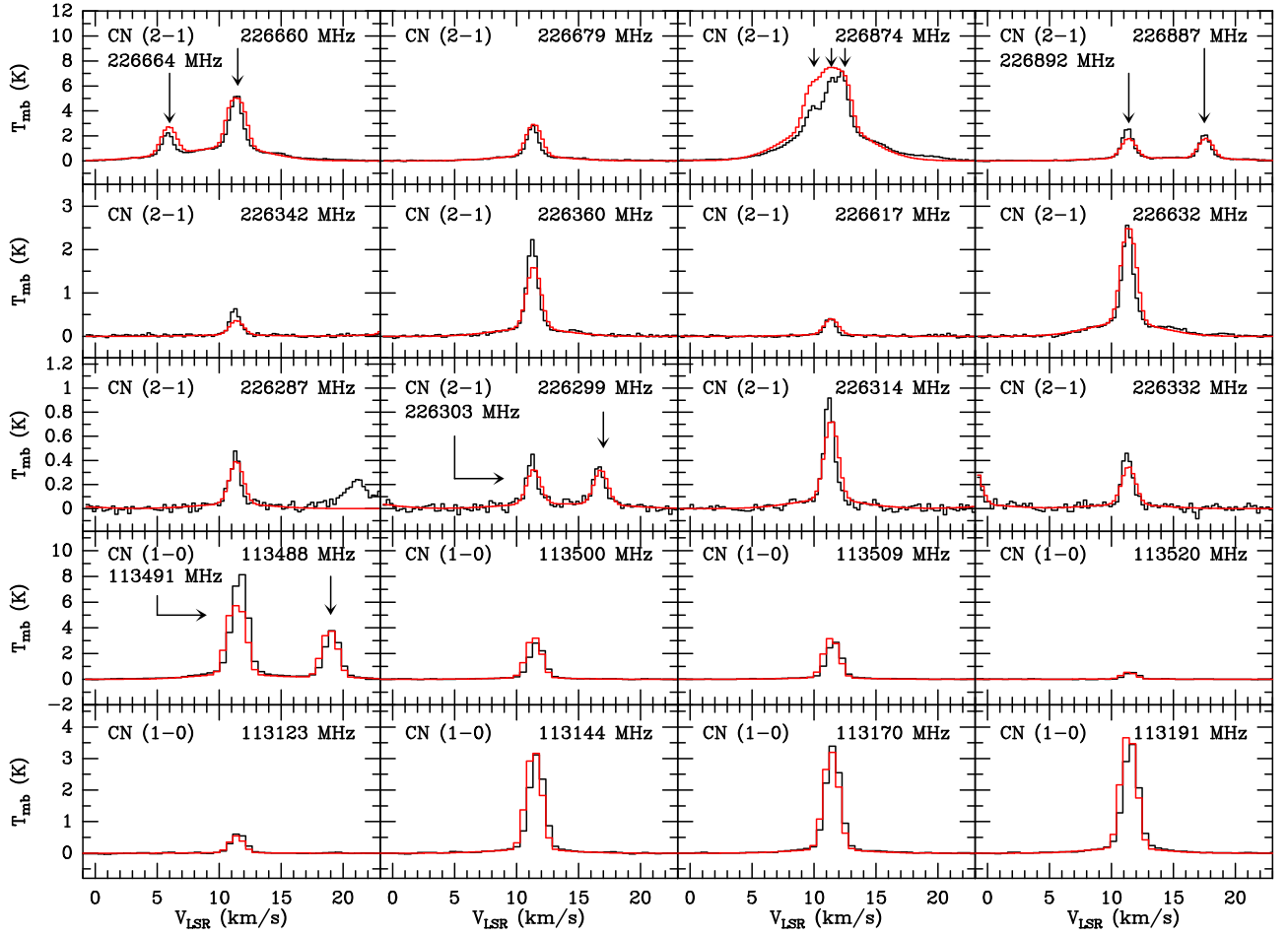


Figure 11. CN observed spectra. The profiles superimposed onto the spectra have been calculated using a two-component LTE modeling (see the text). When several close hyperfine components are present, the frequency of the main one is given. The temperature scale is main beam temperature.

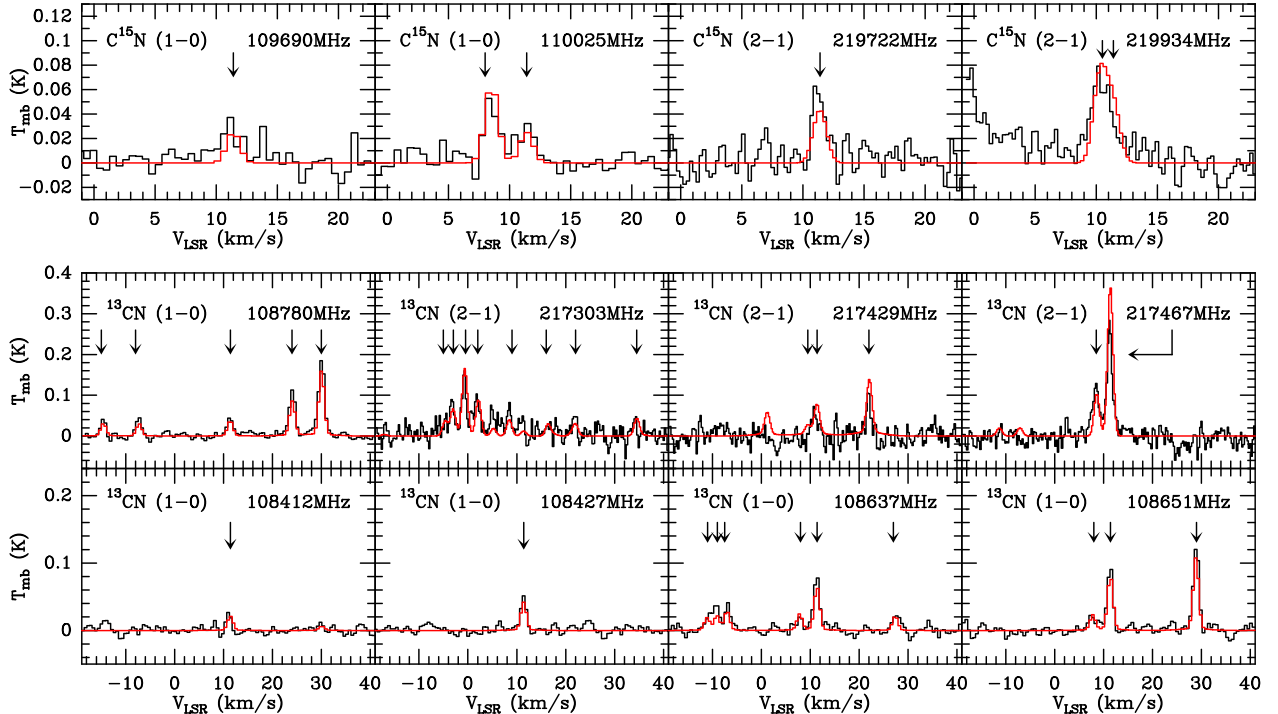


Figure 12. ^{13}CN and C^{15}N observed spectra. The profiles superimposed onto the spectra have been calculated using an LTE modeling (see the text). When several close hyperfine components are present, the frequency of the main one is given. The temperature scale is main beam temperature.

Appendix B

Opacity Checks of Hyperfine Components

We have compared the observed flux ratios with the theoretical ratios (proportional to the g_{up} ratios, the slight

frequency differences being neglected). The results (see Table 5) suggest that the hyperfine components follow an intensity distribution very close to LTE and that the line opacities are moderate.

Table 5
Relative Fluxes of the Hyperfine Components

Species	Transition	Frequency (MHz)	E_{up} (K)	g_{up}	A_{ij} (s^{-1})	R_{theo}^b	$\int T_{\text{mb}} dv^c$ (K km s^{-1})	$R_{\text{obs.}}^d$
HCN	1_1-0_1	88630.42	4.25	3	2.43×10^{-5}	3.0	8.4(0.9)	1.6(0.2)
	1_2-0_1	88631.85	4.25	5	2.43×10^{-5}	5.0	13.5(2.0)	2.6(0.5)
	$1_0-0_1^a$	88633.94	4.25	1	2.43×10^{-5}	1.0	5.2(0.5)	1.0(0.1)
H^{13}CN	1_1-0_1	86338.77	4.14	3	2.22×10^{-5}	3.0	1.3(0.1)	2.9(0.5)
	1_2-0_1	86340.18	4.14	5	2.22×10^{-5}	5.0	2.2(0.2)	5.1(0.9)
	$1_0-0_1^a$	86342.27	4.14	1	2.22×10^{-5}	1.0	0.44(0.06)	1.0(0.2)
NNH^+	1_1-0_1	93171.88	4.47	9	3.63×10^{-5}	3.0	13.1(1.5)	2.5(0.4)
	1_2-0_1	93173.70	4.47	15	3.63×10^{-5}	5.0	21.3(2.5)	4.1(0.6)
	$1_0-0_1^a$	93176.13	4.47	3	3.63×10^{-5}	1.0	5.3(0.5)	1.0(0.1)
$^{15}\text{NNH}^+$	1_2-0_1	90263.91	4.38	3	3.30×10^{-5}	3.0	0.036(0.005)	2.4(0.7)
	1_1-0_1	90263.49	4.38	5	3.30×10^{-5}	5.0	0.08(0.01)	5.0(1.5)
	$1_0-0_1^a$	90264.50	4.38	1	3.30×10^{-5}	1.0	0.015(0.004)	1.0(0.4)
N^{15}NH^+	1_1-0_1	91204.26	4.33	3	3.40×10^{-5}	3.0	0.059(0.008)	2.4(0.6)
	1_2-0_1	91205.99	4.33	5	3.40×10^{-5}	5.0	0.09(0.02)	3.8(1.2)
	$1_0-0_1^a$	91208.52	4.33	1	3.40×10^{-5}	1.0	0.025(0.006)	1.0(0.3)
CN	$1_0 1/2 1/2-0_0 1/2 1/2^a$	113123.37	5.43	2	1.29×10^{-6}	1.0	0.9(0.1)	1.1(0.2)
	$1_0 1/2 1/2-0_0 1/2 3/2$	113144.19	5.43	2	1.05×10^{-5}	8.1	4.9(0.5)	5.8(0.8)
	$1_0 1/2 3/2-0_0 1/2 1/2$	113170.54	5.43	4	5.14×10^{-6}	7.9	5.3(0.5)	6.2(0.9)
	$1_0 1/2 3/2-0_0 1/2 3/2$	113191.33	5.43	4	6.68×10^{-6}	10.3	5.6(0.6)	6.6(1.0)
	$1_0 3/2 3/2-0_0 1/2 1/2$	113488.14	5.45	4	6.73×10^{-6}	10.4	6.0(0.7)	7.0(1.1)
	$1_0 3/2 5/2-0_0 1/2 3/2$	113490.99	5.45	6	1.19×10^{-5}	27.5	14.0(1.6)	16.4(2.5)
	$1_0 3/2 1/2-0_0 1/2 1/2$	113499.64	5.45	2	1.06×10^{-5}	8.2	4.3(0.4)	5.1(0.7)
	$1_0 3/2 3/2-0_0 1/2 3/2$	113508.93	5.45	4	5.19×10^{-6}	8.0	4.5(0.5)	5.3(0.8)
	$1_0 3/2 1/2-0_0 1/2 3/2^a$	113520.42	5.45	2	1.30×10^{-6}	1.0	0.76(0.08)	0.9(0.1)
	$2_0 3/2 1/2-1_0 3/2 1/2$	226287.43	16.31	2	1.03×10^{-5}	4.6	0.5(0.1)	5.1(1.1)
	$2_0 3/2 1/2-1_0 3/2 3/2$	226298.92	16.31	2	8.23×10^{-6}	3.7	0.6(0.2)	5.8(1.7)
	$2_0 3/2 3/2-1_0 3/2 1/2$	226303.08	16.31	4	4.17×10^{-6}	3.7	0.5(0.1)	4.9(1.2)
	$2_0 3/2 3/2-1_0 3/2 3/2$	226314.54	16.31	4	9.91×10^{-6}	8.8	1.0(0.2)	10.1(2.1)
	$2_0 3/2 3/2-1_0 1/2 5/2$	226332.54	16.31	4	4.55×10^{-6}	4.0	0.5(0.1)	5.0 (1.1)
	$2_0 3/2 5/2-1_0 3/2 3/2$	226341.93	16.31	6	3.16×10^{-6}	4.2	0.7(0.1)	6.6(1.4)
	$2_0 3/2 5/2-1_0 3/2 5/2$	226359.87	16.31	6	1.61×10^{-5}	21.4	2.6(0.6)	25.4(5.3)
	$2_0 3/2 1/2-1_0 1/2 3/2$	226616.56	16.31	2	1.07×10^{-5}	4.8	0.5(0.1)	4.6(0.9)
	$2_0 3/2 3/2-1_0 1/2 3/2$	226632.19	16.31	4	4.26×10^{-5}	37.9	3.2(0.7)	31.2(6.8)
	$2_0 3/2 5/2-1_0 1/2 3/2$	226659.58	16.31	6	9.47×10^{-5}	126.2	8.3(1.7)	80(17)
	$2_0 3/2 1/2-1_0 1/2 1/2$	226663.70	16.31	2	8.46×10^{-5}	37.6	3.2(0.7)	30.9(6.6)
	$2_0 3/2 3/2-1_0 1/2 1/2$	226679.38	16.31	4	5.27×10^{-5}	46.8	3.6(0.8)	35.0 (7.3)
	$2_0 5/2 5/2-1_0 3/2 5/2$	226892.12	16.34	6	1.81×10^{-5}	24.1	3.2(0.7)	30.9 (6.5)
	$2_0 5/2 3/2-1_0 3/2 5/2^a$	226905.38	16.34	4	1.13×10^{-6}	1.0	0.10(0.03)	1.0(0.3)
^{13}CN	$1_1 2 0 1-0_1 2 1 1$	108412.86	5.23	3	3.1×10^{-6}	1.1	0.03(0.01)	1.1(0.5)
	$1_1 2 0 1-0_1 2 1 2$	108426.89	5.23	3	6.3×10^{-6}	2.3	0.069(0.008)	2.1(0.7)
	$1_1 2 1 0-0_1 2 0 1$	108631.12	5.21	1	9.6×10^{-6}	1.2	0.04(0.02)	1.1(0.7)
	$1_1 2 1 1-0_1 2 0 1$	108636.92	5.21	3	9.6×10^{-6}	3.5	0.11(0.02)	3.4(1.1)
	$1_1 2 1 2-0_1 2 0 1$	108651.30	5.21	5	9.8×10^{-6}	5.9	0.18(0.03)	5.4(1.9)
	$1_3 2 1 2-0_1 2 1 2$	108657.65	5.24	5	7.2×10^{-6}	4.4	0.13(0.02)	4.0(1.3)
	$1_3 2 2 3-0_1 2 1 2$	108780.20	5.25	7	1.1×10^{-5}	8.9	0.27(0.05)	8.2(2.8)
	$1_3 2 2 2-0_1 2 1 1$	108782.37	5.25	5	7.8×10^{-6}	4.7	0.16(0.03)	4.9(1.6)
	$1_3 2 2 1-0_1 2 1 0$	108786.98	5.25	3	5.7×10^{-6}	2.1	0.06(0.01)	2.0(0.7)

Table 5
(Continued)

Species	Transition	Frequency (MHz)	E_{up} (K)	g_{up}	A_{ji} (s^{-1})	R_{theo}^b	$\int T_{\text{mb}} dv^c$ (K km s $^{-1}$)	R_{obs}^d
	$1_{3/2,2,1}-0_{1/2,1,1}$	108793.75	5.25	3	4.5×10^{-6}	1.6	0.06(0.01)	1.9(0.6)
	$1_{3/2,2,2}-0_{1/2,1,2}$	108796.40	5.25	5	2.8×10^{-6}	1.7	0.05(0.02)	1.7(0.8)
	$1_{3/2,1,2}-0_{1/2,1,1}$	108643.59	5.24	5	2.6×10^{-6}	1.6	0.06(0.01)	1.7(0.6)
	$1_{3/2,1,0}-0_{1/2,1,1}$	108644.35	5.24	1	9.6×10^{-6}	1.2	0.06(0.02)	1.8(0.7)
	$1_{3/2,1,1}-0_{1/2,1,1}^a$	108645.06	5.24	3	2.7×10^{-6}	1.0	0.03(0.01)	1.0(0.4)
	$1_{3/2,1,1}-0_{1/2,1,2}$	108658.95	5.24	3	3.3×10^{-6}	1.2	0.03(0.01)	1.0(0.5)
	$2_{5/2,3,3}-1_{3/2,2,2}$	217467.15	15.69	7	8.92×10^{-5}	0.6	0.36(0.08)	0.6(0.1)
	$2_{5/2,3,2}-1_{3/2,2,1}$	217469.15	15.69	5	8.43×10^{-5}	0.4	0.23(0.05)	0.4(0.1)
C^{15}N	$1_{1/2,1,1}-0_{1/2,1,1}^a$	109689.61	5.27	3	7.10×10^{-6}	1.0	0.05(0.02)	1.0(0.5)
	$1_{3/2,1,1}-0_{1/2,1,1}$	110023.54	5.28	3	7.16×10^{-6}	1.0	0.05(0.02)	1.0(0.4)
	$1_{3/2,2,1}-0_{1/2,1,1}$	110024.59	5.28	5	1.09×10^{-5}	2.6	0.08(0.02)	1.7(0.7)
	$2_{3/2,2,1}-1_{1/2,1,1}$	219722.49	15.81	5	8.67×10^{-5}	0.3	0.08(0.02)	0.4(0.1)
	$2_{5/2,2,1}-1_{3/2,1,1}$	219934.04	15.84	5	9.36×10^{-5}	0.7	0.10(0.03)	0.6(0.1)

Notes.^a Weakest detected line of the hyperfine structure. This is the reference to compute R_{theo} and R_{obs} .^b Ratio of the line $G_{\text{up}} A_{ji}$ product relative to the reference line; equal to the fluxes ratio in LTE optically thin conditions, neglecting the frequency differences.^c Observed line fluxes derived from Gaussian fits; the errors include fit and calibration uncertainties.^d Ratio of observed fluxes relative to the reference line. For CN ($N = 1-0$), the reference flux is averaged over the two reference lines.**ORCID iDs**Claudine Kahane  <https://orcid.org/0000-0003-2718-7197>Ali Jaber Al-Edhari  <https://orcid.org/0000-0003-4089-841X>Cecilia Ceccarelli  <https://orcid.org/0000-0001-9664-6292>Ana López-Sepulcre  <https://orcid.org/0000-0002-6729-3640>Francesco Fontani  <https://orcid.org/0000-0003-0348-3418>**References**

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