



Publication Year	2017
Acceptance in OA	2020-09-04T10:02:04Z
Title	Radar sounding of Lucus Planum, Mars, by MARSIS
Authors	OROSEI, ROBERTO, Rossi, Angelo Pio, Cantini, Federico, Caprarelli, Graziella, Carter, Lynn M., Papiano, Irene, CARTACCI, MARCO, CICHETTI, ANDREA, NOSCHESE, RAFFAELLA
Publisher's version (DOI)	10.1002/2016JE005232
Handle	http://hdl.handle.net/20.500.12386/27133
Journal	JOURNAL OF GEOPHYSICAL RESEARCH (PLANETS)
Volume	122

RESEARCH ARTICLE

10.1002/2016JE005232

Key Points:

- Radar sounding data over Lucus Planum were analyzed to derive bulk dielectric properties
- Results support the interpretation of a volcanic origin for those areas in which echoes from the base of Lucus Planum could be detected
- Lucus Planum is laterally inhomogeneous, implying that it originated from multiple events and/or processes

Correspondence to:

R. Orosei,
roberto.oroisei@inaf.it

Citation:

Orosei, R., A. P. Rossi, F. Cantini, G. Caprarelli, L. M. Carter, I. Papiano, M. Cartacci, A. Cicchetti, and R. Noschese (2017), Radar sounding of Lucus Planum, Mars, by MARSIS, *J. Geophys. Res. Planets*, 122, doi:10.1002/2016JE005232.

Received 28 NOV 2016

Accepted 24 MAY 2017

Accepted article online 30 MAY 2017

Radar sounding of Lucus Planum, Mars, by MARSIS

Roberto Orosei¹, Angelo Pio Rossi², Federico Cantini³, Graziella Caprarelli^{4,5}, Lynn M. Carter⁶, Irene Papiano⁷, Marco Cartacci⁸, Andrea Cicchetti⁸, and Raffaella Noschese⁸
¹Istituto di Radioastronomia, Istituto Nazionale di Astrofisica, Bologna, Italy, ²Department of Physics and Earth Sciences, Jacobs University Bremen, Bremen, Germany, ³Space Engineering Center, Ecole Polytechnique Federale de Lausanne, Lausanne, Switzerland, ⁴Division ITEE, University of South Australia, Adelaide, South Australia, Australia, ⁵International Research School of Planetary Sciences, Pescara, Italy, ⁶Lunar and Planetary Laboratory, University of Arizona, Tucson, Arizona, USA, ⁷Liceo Scientifico Augusto Righi, Bologna, Italy, ⁸Istituto di Astrofisica e Planetologia Spaziali, Istituto Nazionale di Astrofisica, Rome, Italy

Abstract Lucus Planum, extending for a radius of approximately 500 km around 181°E, 5°S, is part of the Medusae Fossae Formation (MFF), a set of several discontinuous deposits of fine-grained, friable material straddling across the Martian highland-lowland boundary. The MFF has been variously hypothesized to consist of pyroclastic flows, pyroclastic airfall, paleopolar deposits, or atmospherically deposited icy dust driven by climate cycles. Mars Advanced Radar for Subsurface and Ionosphere Sounding (MARSIS), a low-frequency subsurface-sounding radar carried by European Space Agency's Mars Express, acquired 238 radar swaths across Lucus Planum, providing sufficient coverage for the study of its internal structure and dielectric properties. Subsurface reflections were found only in three areas, marked by a distinctive surface morphology, while the central part of Lucus Planum appears to be made of radar-attenuating material preventing the detection of basal echoes. The bulk dielectric properties of these areas were estimated and compared with those of volcanic rocks and ice-dust mixtures. Previous interpretations that east Lucus Planum and the deposits on the northwestern flanks of Apollinaris Patera consist of high-porosity pyroclastic material are strongly supported by the new results. The northwestern part of Lucus Planum is likely to be much less porous, although interpretations about the nature of the subsurface materials are not conclusive. The exact origin of the deposits cannot be constrained by radar data alone, but our results for east Lucus Planum are consistent with an overall pyroclastic origin, likely linked to Tharsis Hesperian and Amazonian activity.

Plain Language Summary Lobe-shaped thick deposits, collectively known as the Medusae Fossae Formation, are found west of the Olympus Mons volcano on Mars. Visual observations of these smooth and relatively unremarkably looking materials have not definitively determined how they formed with hypotheses ranging from volcanic ash to remnants of a materials deposited at a previous location of the north pole, to accumulation of atmospheric dust. In this study we used the ground penetrating radar on board Mars Express to see through these deposits to derive information about Lucus Planum, the central lobe of the Medusae Fossae Formation. Through our analysis of the way the radar waves were reflected by subsurface layering, we concluded that the materials forming Lucus Planum are spatially variable: the east and west portions of the deposits are highly porous and probably composed of ashes and rocks from nearby volcanoes. In the north-west the deposits are much denser, but we could not unequivocally define their nature. Finally, our instrument could not detect signals from the central part of Lucus Planum, which suggests yet a different component in the deposits. This diversity points to a dynamic geological history in this unique region of Mars.

1. Introduction

Lucus Planum, extending for a radius of approximately 500 km around 181°E, 5°S, is part of the Medusae Fossae Formation (MFF), a set of several discontinuous deposits of fine-grained, friable material straddling across the Martian highland-lowland boundary [e.g., Carter *et al.*, 2009].

The MFF covers an extensive area, spanning latitudinally more than 1000 km and longitudinally some 6000 km. It is separated into several discontinuous lobes (Figure 1). The lobe that occupies the central part

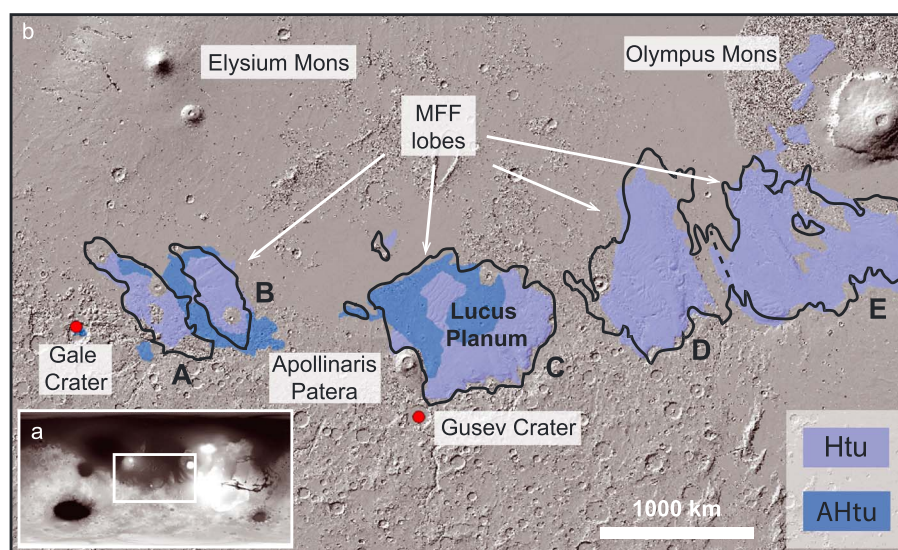


Figure 1. Geomorphologic and geologic setting for Lucus Planum: (a) Location of the greater Medusa Fossae Formation (MFF) deposits, Figure 1b is outlined in white. (b) Extent of the Medusa Fossae Formation as outlined, in black, by Kerber *et al.* [2011]. The MFF lobes are also indicated with letters, according to Harrison *et al.* [2010]. Lucus Planum constitutes the center lobe of the MFF. Units from the revised global geological map of Mars [Tanaka *et al.*, 2014] are indicated *Htu* and *AHtu*. The location of the NASA MER Spirit rover landing site in Gusev crater and NASA Mars Science Laboratory Curiosity in Gale crater is indicated. Figure 1a is a global map in equirectangular projection, centered at longitude 180. Figure 1b is a map centered on Lucus Planum and covering the MFF, in equirectangular projection, and centered at 180°E. The longitudinal extension is about 5000 km. Latitude range is from -30° to 30° , and longitude ranges from 130° E to 230° E.

of the Medusae Fossae Formation is known as Lucus Planum [e.g., Kerber *et al.*, 2011] (or alternatively, lobe B [Harrison *et al.*, 2010]). In the recently revised global geologic map of Mars [Tanaka *et al.*, 2014] two units make up Lucus Planum, namely the Hesperian and Amazonian-Hesperian transitional units (respectively, *Htu* and *AHtu*) (Figure 1) [Tanaka *et al.*, 2014].

The MFF has been variously hypothesized to consist of pyroclastic flows [Scott and Tanaka, 1982; Mandt *et al.*, 2008; Bradley *et al.*, 2002], pyroclastic airfall [Tanaka, 2000; Hynek *et al.*, 2003; Kerber *et al.*, 2011], paleopolar deposits [Schultz and Lutz, 1988], or atmospherically deposited icy dust driven by climate cycles [Head III and Kreslavsky, 2004]. A branching positive relief system within Lucus Planum was interpreted by Harrison *et al.* [2013] as an ancient fluvial system originating from seepage sapping, implying that Lucus Planum was volatile rich. The MFF shows evidence of a complex history of deposition, erosion, and exhumation of both landforms and deposits [e.g., Kerber and Head, 2012]. Both erosional and depositional landforms are visible at different stratigraphic levels, resulting in complex morphologies.

Two sounding radars have been flown on Martian missions: Mars Advanced Radar for Subsurface and Ionosphere Sounding (MARSIS) [Picardi *et al.*, 2005] and SHallow RADar (SHARAD) [Seu *et al.*, 2007]. Both instruments are synthetic aperture, low-frequency radars carried by European Space Agency's (ESA) Mars Express and NASA's Mars Reconnaissance Orbiter, respectively. They transmit low-frequency radar pulses that penetrate below the surface and are reflected by dielectric discontinuities in the subsurface. MARSIS is optimized for deep penetration, with a free-space range resolution of approximately 150 m, a footprint size of 10–20 km across track, and 5–10 km along track. SHARAD has tenfold better resolution, at the cost of reduced penetration. Parts of the MFF have been probed by both of these sounding radars [Watters *et al.*, 2007; Carter *et al.*, 2009], revealing a dielectric permittivity of the MFF material that is consistent with either a substantial component of water ice or a low-density, ice-poor material. While the work by Watters *et al.* [2007] was focused on Lucus Planum, estimates of dielectric properties by Carter *et al.* [2009] were based on observations over Zephyria Planum, in the westernmost part of the Medusae Fossae Formation, and the area between Gordii Dorsum and Amazonis Mensa, at the Eastern end of the MFF.

The dielectric permittivity of the MFF material [Watters *et al.*, 2007; Carter *et al.*, 2009] is consistent with either a substantial component of water ice or a low-density, ice-poor material. There is no evidence for internal

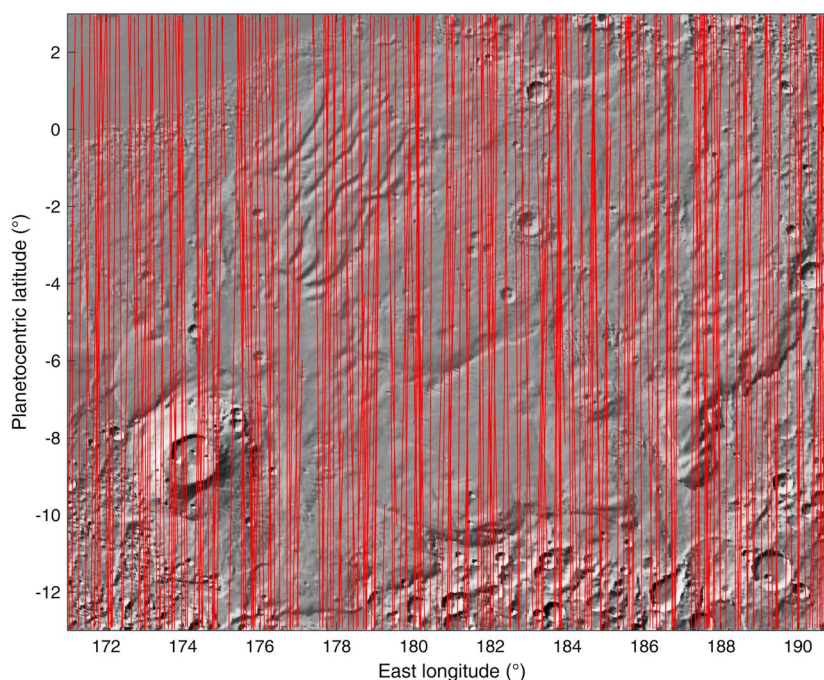


Figure 2. Shaded relief map of the Lucus Planum area showing MARSIS ground tracks as red lines. Lines are thinner than the MARSIS swath width for legibility.

layering from SHARAD data [Carter *et al.*, 2009], despite the fact that layering at scales of tens of meters has been reported in many parts of the MFF [Kerber, 2014]. This lack of detection can be the result of one or more factors, such as high interface roughness, low dielectric contrast between materials, or discontinuity of the layers.

2. Method

Operating since mid-2005, MARSIS has acquired 238 swaths of echoes across Lucus Planum, shown in Figure 2. Each swath consists of a few hundred observations, for a total of over 38,000 echoes. Data are affected by the dispersion and attenuation of the radar signal caused by ionospheric plasma, but a number of methods have been developed over the years to attenuate or compensate such effects [Picardi and Sorge, 2000; Armand *et al.*, 2003; Safaeinili *et al.*, 2007; Mougnot *et al.*, 2008; Zhang *et al.*, 2009; Smirnov and Yushkova, 2013; Campbell and Watters, 2016]. Data used in this work have been processed using the methodology described by Cartacci *et al.* [2013], which consists in the maximization of the signal power in an interval centered around the strongest echo through the differential variation of the phase of the components of the Fourier signal spectrum.

MARSIS data acquired continuously during the movement of the spacecraft are usually displayed in the form of radargrams, grey-scale images in which the horizontal dimension is distance along the ground track, the vertical one is the round trip time of the echo, and the brightness of the pixel is a function of the strength of the echo (refer to example in Figure 3). The first step in data analysis consisted in the visual inspection of radargrams to determine their quality. Observations were discarded if the ionospheric distortion compensation algorithm had failed, if spurious signals from the electronics of the spacecraft were present, or if exceptional ionosphere conditions resulted in a severe attenuation or absence of the signal. This reduced the number of radargrams suitable for further analysis by approximately 25%.

The next step consisted in the identification of subsurface echoes in radargrams, which is complicated by the so called “clutter,” that is by echoes coming from off-nadir surface features, such as craters or mountains, and reaching the radar after the nadir surface echo. As clutter can dwarf subsurface echoes, numerical electromagnetic models of surface scattering have been developed [see e.g., Nouvel *et al.*, 2004; Spagnuolo *et al.*, 2011] to validate the detection of subsurface interfaces in MARSIS data. They are used to produce simulations of surface echoes, which are then compared to the ones detected by the radar: any secondary echo visible in radargrams but not in simulations is interpreted as caused by subsurface reflectors (Figure 3).

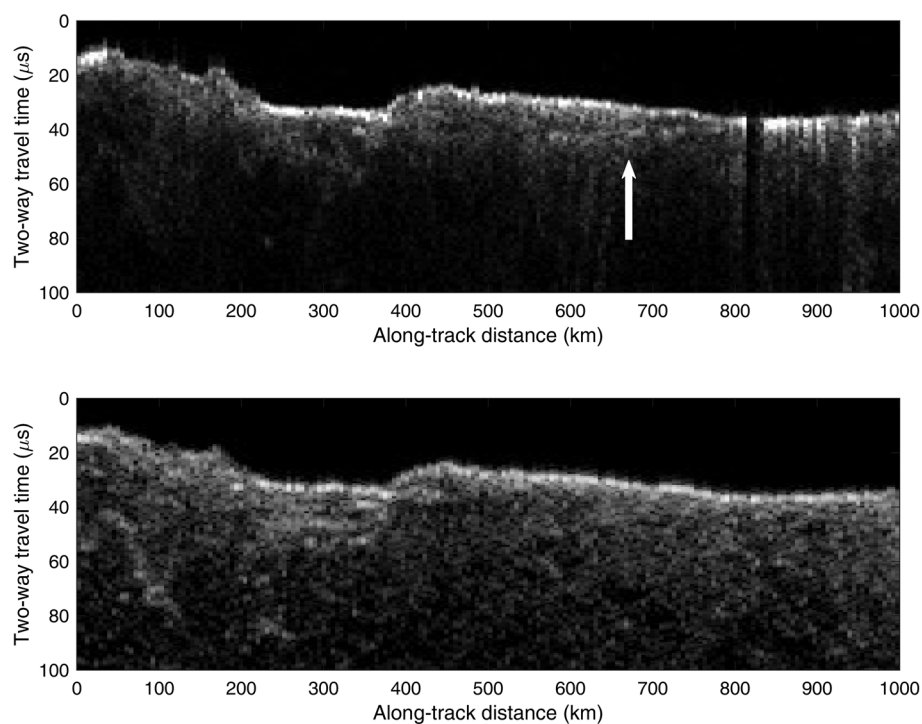


Figure 3. Comparison between (top) real and (bottom) simulated radargrams for orbit 4011. The simulation reproduces echoes from topography only, while real data contain both surface and subsurface echoes. The arrow highlights a weak subsurface reflection.

To analyze clutter, a code for the simulation of radar wave surface scattering was developed, based on the algorithm of *Nouvel et al.* [2004]. The Mars Orbiter Laser Altimeter (MOLA) topographic data set [*Smith et al.*, 2001] was used to represent the Martian surface as a collection of flat plates called facets. Radar echoes were computed as the coherent sum of reflections from all facets illuminated by the radar. The computational burden of simulations required the use of the SuperMUC supercomputer at the Leibniz-Rechenzentrum, Garching, Germany.

Subsurface reflections in Lucus Planum are usually weak and often have a diffuse appearance (Figure 3). Several methods were attempted to automatically identify such reflections in radargrams, but eventually a supervised procedure was used, in which an operator manually selects a few points marking the position of the interface in a radargram, and then the procedure itself outlines the interface and records of its aerocentric coordinates, its time delay from the surface echo, and its reflected power. The confidence in the retrieved coordinates is based on the accuracy of the reconstructed Mars Express trajectory, which is estimated to be a fraction of the MARSIS footprint size. To better determine the position and power of subsurface echoes, radar signals have been interpolated with the Fourier interpolation method to reduce the sampling interval to 0.1 μ s. The precision in the determination of the time delay is assumed to be the one-way delay resolution (or 0.5 μ s, corresponding to 150 m free space), while the uncertainty in echo power is considered to be below 0.5 dB because of the interpolation.

3. Results

A total of 97 subsurface reflectors were identified, extending along track over distances up to 500 km. Their distribution across Lucus Planum is shown in Figure 4. In spite of several high-quality radargrams crossing the central part of Lucus Planum, only a handful of subsurface interfaces could be detected there, most of which are shallow, often associated with pedestal craters. Reflectors concentrate in specific areas: the deposits on the northwestern flanks of Apollinaris Patera, the rugged terrain north of Tartarus Scopulus, and the large lobe located north-east of Memnonia Sulci. The contours of these areas follow closely morphologically distinct provinces within Lucus Planum, which suggests that variations in surface morphology could be tied to

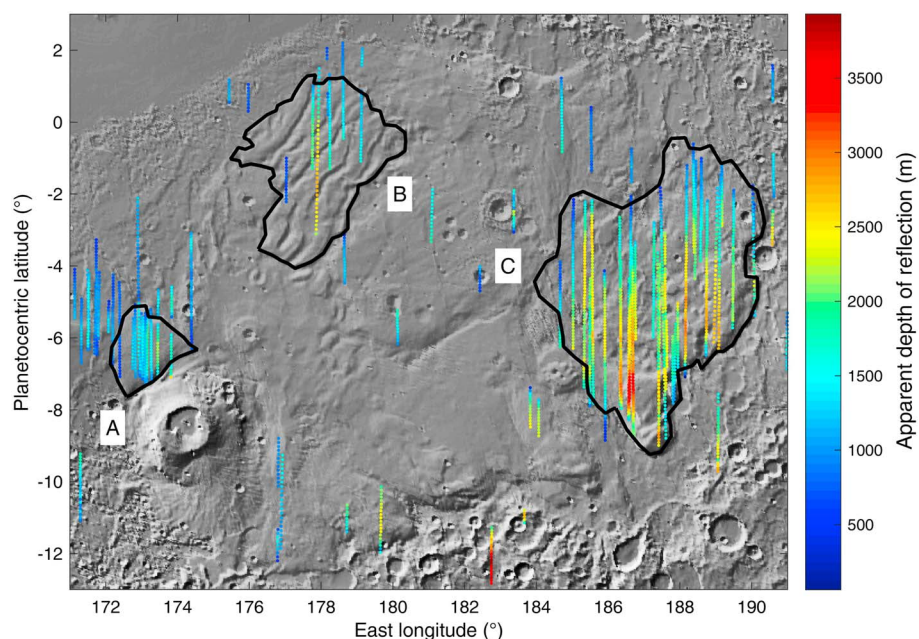


Figure 4. Map of the location and apparent depth of subsurface echoes in Lucus Planum. Apparent depth is coded according to the color scale on the right and differs from real depth by a factor $\sqrt{\epsilon}$ (see equation (1)). Black contours outline what appear to be different morphologic provinces within Lucus Planum.

changes in the material forming Lucus Planum. These areas are outlined in Figure 4 and labeled “A,” “B,” and “C,” respectively.

Figure 4 shows the apparent depth of reflectors, estimated from the measured round-trip time delay between surface and subsurface echo by

$$z = \frac{c \tau}{2\sqrt{\epsilon}} \quad (1)$$

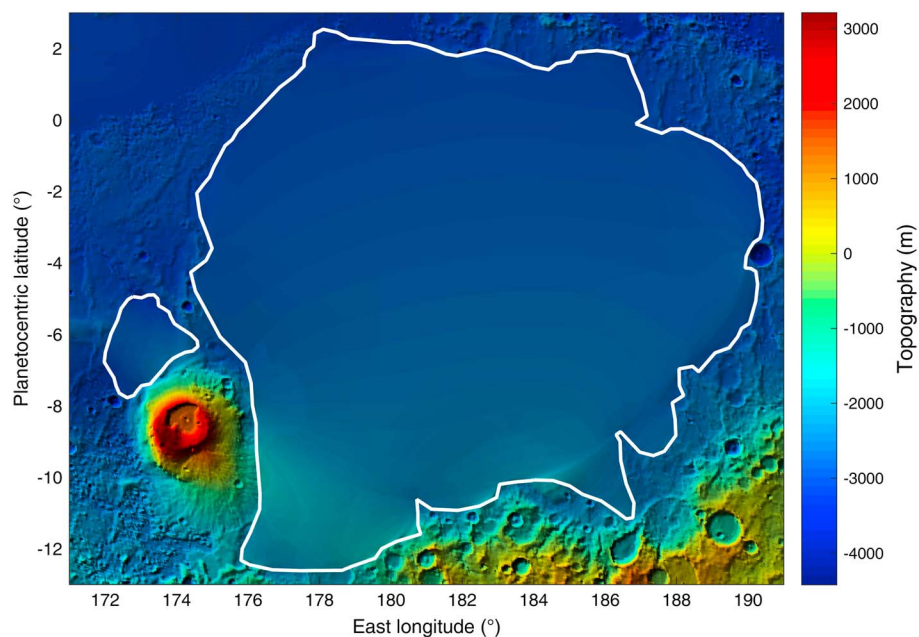


Figure 5. Topographic map of the Lucus Planum area, based on the MOLA data set. The white contours identify the areas in which MOLA topography was removed, and then the base of the Lucus Planum and Apollinaris Patera deposits was interpolated from the remaining topographic information through the natural neighbor method.

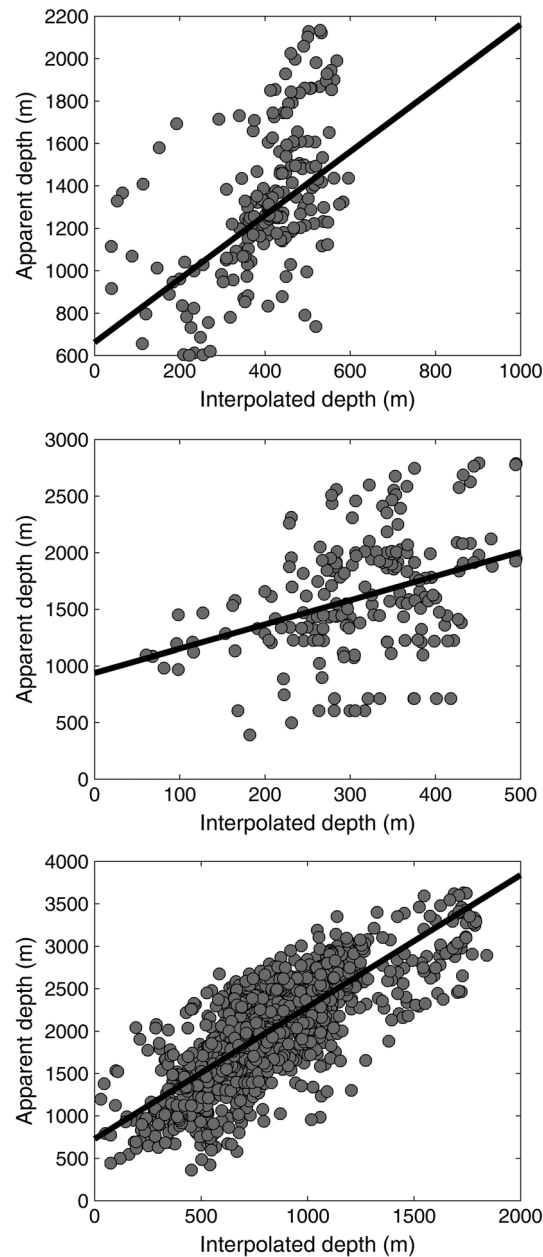


Figure 6. Plot of the thickness of Lucas Planum, estimated by interpolating the surrounding topography, versus the apparent thickness derived from the delay time of subsurface radar echoes. (top to bottom) Graphs refer to areas A, B, and C (see Figure 4). The best fit line has been computed using the least absolute deviation method.

from which we see that the slope of the best fit line in a plot of interpolated versus apparent depth provides an estimate of $\sqrt{\epsilon}$. The resulting plots for areas A, B, and C are shown in Figure 6. Because of the large dispersion of data points in some of the plots, the best fit line was computed using the least absolute deviations method [Bloomfield and Steiger, 1983], which is less sensitive to outliers than the least squares method.

The slopes and constant terms of the best fit lines in Figure 6 are reported in Table 1. For each value, the corresponding 95% confidence interval of the estimate is listed, providing some insight on the quality of the data fit. Table 1 reports also estimates of ϵ , obtained from the values of slopes and their corresponding 95% confidence bounds through equation (3).

where z is depth, c the speed of light *in vacuo*, τ the round-trip time delay between surface and subsurface echo, and ϵ is the real part of the relative complex permittivity (also called dielectric constant) of the Lucas Planum material. The apparent depth z_a was computed assuming that ϵ is equal to 1, corresponding to the permittivity of free space:

$$z_a = \frac{c \tau}{2} \quad (2)$$

Apparent depths overestimate the thickness of Lucas Planum by a factor comprised between $\sqrt{3}$ and 3, depending on the nature of the material through which the wave propagates [Ulaby *et al.*, 1986, Appendix E].

Estimates of permittivity for the different regions of Lucas Planum provide some insight on their nature and a more precise evaluation of their thickness. Following the approach first presented in Picardi *et al.* [2005] and used also in Watters *et al.* [2007], we produced an independent estimate of the thickness of Lucas Planum assuming that the deposits rest on a surface in lateral continuity with the surrounding topography, and that MARSIS echoes come from such surface. The white contours in Figure 5 encompass those areas in which MOLA topography was removed and then interpolated from the remaining topographic information through the natural neighbor method [Sibson, 1981].

The difference between the actual topography and the interpolated basal topography of Lucas Planum provides an estimate of the depth of the base of Lucas Planum, z_i . By inserting z_i in equation (1), solving equation (2) by $c\tau$, and rearranging and simplifying equal terms, we obtain

$$z_a = \sqrt{\epsilon} z_i \quad (3)$$

Table 1. Coefficients of the Best Fit Lines in Figure 6, Together With Their 95% Confidence Bounds^a

Area	Slope	95% Confidence Interval
A	1.50	[1.17, 1.82]
B	2.13	[1.34, 2.93]
C	1.56	[1.49, 1.62]
Area	Constant Term	95% Confidence Interval
A	661.2	[527.9, 794.5]
B	937.6	[679.1, 1196.1]
C	727.1	[671.4, 782.8]
Area	ϵ	95% Confidence Interval
A	2.25	[1.38, 3.33]
B	4.56	[1.80, 8.57]
C	2.42	[2.23, 2.62]

^aThe table reports also the corresponding estimates of ϵ and of its confidence interval, derived from equation (3).

From equation (3), the value of the constant term in best fit lines should be zero, different from what is reported in Table 1. The presence of a constant term indicates a systematic error in the evaluation of z_a , z_i , or both. Because the range resolution of MARSIS is about 150 m in free space [Picardi *et al.*, 2005], the constant terms in Table 1 correspond to a few range resolution cells. One possible explanation is that the interpolation method failed to provide a correct estimate of the basal topography: because Lucus Planum straddles the dichotomy boundary, the topography beneath it is expected to be complex, affecting the precision of results. Another possibility is a systematic overestimation of the time delay of subsurface echoes in radargrams, perhaps because subsurface reflections are less sharp than surface ones, and the manual determination of their exact position introduces additional uncertainties.

Permittivity is a complex quantity: its real part affects the velocity of an electromagnetic wave, while its imaginary part is related to the dissipation (or loss) of energy within the medium. The ratio between the imaginary and the real part of the complex permittivity is called the loss tangent. Estimating the loss tangent of the material within Lucus Planum provides an additional constraint on its nature and can be used by way of checking on the significance of the values of ϵ in Table 1.

The loss tangent over parts of the Medusae Fossae Formation was estimated from the rate of decay of the subsurface echo power as a function of time delay by Watters *et al.* [2007]. Following a similar approach, we assumed that the surface and the subsurface interfaces over Lucus Planum are smooth at MARSIS frequencies, meaning that the RMS height of topography is a fraction of the wavelength, and that Lucus Planum consists only of nonmagnetic, low loss material. While a higher roughness would cause only a fluctuation of surface and subsurface power without affecting the mean rate of subsurface power decay with depth, the assumption that Lucus Planum consists of a low loss, nonmagnetic material is validated by previous results [Picardi *et al.*, 2005] and would result in little or no subsurface interface detections if violated. Under these assumptions, following Porcello *et al.* [1974], the surface echo power P_s can be written as follows:

$$P_s = P_t \cdot \left(\frac{G\lambda}{8\pi H} \right)^2 \cdot |R_s|^2 \quad (4)$$

with P_t the transmitted power, G the antenna gain, λ the wavelength, H the spacecraft altitude, and R_s the surface Fresnel reflection coefficient at normal incidence. Analogously, the subsurface echo power P_{ss} can be computed through the following expression:

$$P_{ss} = P_t \cdot \left(\frac{G\lambda}{8\pi(H+z)} \right)^2 \cdot (1 - |R_s|^2)^2 \cdot |R_{ss}|^2 \cdot \exp(-2\pi f \tan \delta \tau) \quad (5)$$

where R_{ss} is the subsurface Fresnel reflection coefficient at normal incidence, f the radar frequency, and $\tan \delta$ the loss tangent of the Lucus Planum material, here assumed to be constant through its entire thickness, while z and τ have been defined in equation (1).

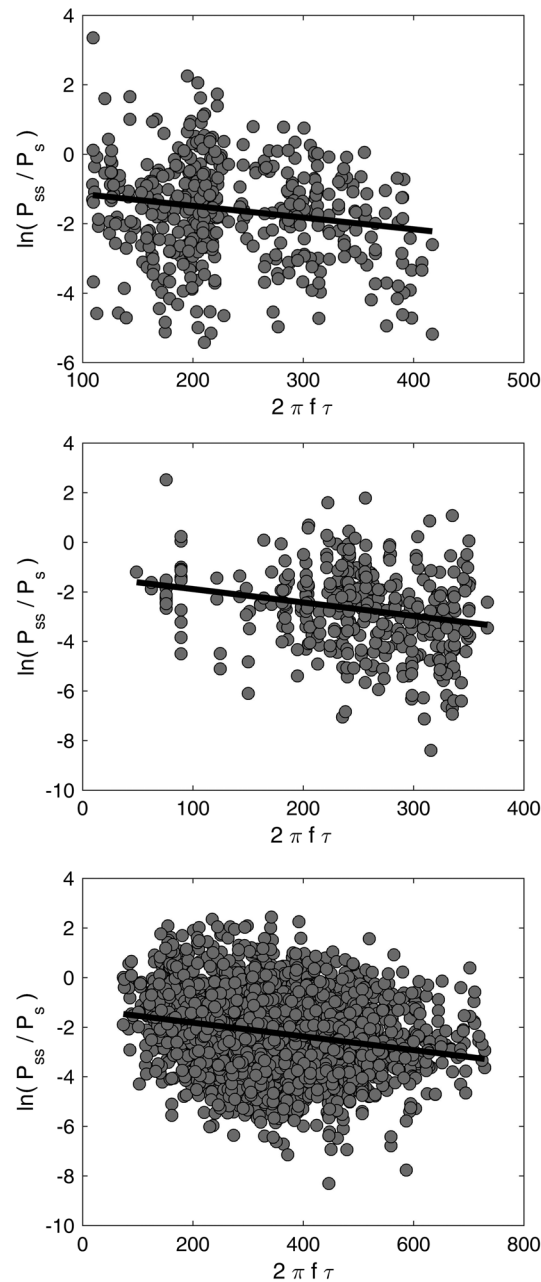


Figure 7. Plot of the number of cycles completed by the radar wave within Lucus Planum versus the natural logarithm of the subsurface to surface echo power ratio. (top to bottom) Graphs refer to areas A, B, and C (see Figure 4). The best fit line has been computed using the least absolute deviation method.

to produce surface echoes whose power is dominated by the dielectric permittivity of the layers closest to the surface, as discussed in *Grima et al.* [2014]. This effect would alter P_{ss}/P_s , but it would not change the rate at which this quantity decreases with depth, that is the first term of the right side of equation (6). Random inhomogeneities within the deposits, whose characteristic size is comparable to the MARSIS wavelength, would result in volume scattering, that is in the diffusion of electromagnetic radiation within the deposits away from the direction of propagation.

The diffuse, weak echoes between surface and basal reflections visible in Figure 2 could be caused by volume scattering, although they could also originate from surface roughness. Volume scattering cannot be

By dividing equation (5) and equation (4) and then taking the natural logarithm of the result, the following expression is obtained:

$$\ln\left(\frac{P_{ss}}{P_s}\right) = -2\pi f \tan \delta \tau + K \quad (6)$$

where K is a term depending on R_s and R_{ss} . The topography of Lucus Planum is characterized by a roughness that is not negligible compared to the MARSIS wavelength [see, *Kreslavsky and Head*, 2000; *Neumann et al.*, 2003]. This implies that P_s and P_{ss} fluctuate around a mean value that is a function of statistical parameters characterizing the topography, such as RMS height and RMS slope [see, for example, *Ogilvy*, 1991]. Under the assumption that such parameters do not vary significantly within each of the three areas A, B, and C, then roughness will cause only a variation of the value of parameter K and the addition of a random noise to $\ln(P_{ss}/P_s)$ in equation (6).

Other factors connected to the internal structure of the Lucus Planum and Apollinaris Patera deposits are unlikely to affect equation (6) significantly. A surface layer thinner than the vertical resolution of the radar can generate interferences so as to drastically reduce surface reflectivity, as in the case of the CO_2 layer over the South Polar Layered Deposits (SPLD) identified by *Mouginot et al.* [2009]. However, such coherent effects require a very smooth surface and are strongly frequency dependent. Both the rougher surface of Lucus Planum and Apollinaris Patera [Kreslavsky and Head, 2000; Neumann et al., 2003], and the fact that such dependence on frequency was not found in the data seem to rule out the presence of such a layer.

Other material inhomogeneities in the dielectric properties at depths below the vertical resolution of MARSIS would tend

Table 2. Coefficients of the Best Fit Lines in Figure 7, Together With Their 95% Confidence Bounds^a

Area	Slope	95% Confidence Interval
A	−0.0034	[−0.0052, −0.0016]
B	−0.0054	[−0.0077, −0.0031]
C	−0.0028	[−0.0032, −0.0023]
Area	Constant Term	95% Confidence Interval
A	−0.809	[−1.255, −0.363]
B	−1.349	[−1.933, −0.764]
C	−1.270	[−1.436, −1.104]
Area	$\tan \delta$	95% Confidence Interval
A	0.0034	[0.0016, 0.0052]
B	0.0054	[0.0031, 0.0077]
C	0.0028	[0.0023, 0.0032]

^aThe table reports also the corresponding estimates of $\tan \delta$ and of its confidence interval, derived from equation (6).

easily characterized from the measure of backscattered radiation, but it would attenuate the subsurface radar echo. This effect cannot be separated from dielectric attenuation, and it would thus lead to a systematic overestimate of $\tan \delta$ from equation (6), which thus constitutes an upper bound for the true dielectric attenuation.

With these caveats, the slope of the best fit line in a plot of $2\pi f\tau$ (that is the number of cycles completed by the radar wave within Lucus Planum) versus the natural logarithm of the subsurface to surface echo power ratio will provide an estimate of $\tan \delta$. Such plots for areas A, B, and C, with the corresponding best fit lines, are shown in Figure 7. In analogy with the estimation of ϵ , the best fit line was computed using the least absolute deviations method.

The slopes and constant terms of the best fit lines in Figure 7 are reported in Table 2. For each value we report the corresponding 95% confidence interval of the estimate, to provide some insight on the quality of the data fit. Table 2 reports also the corresponding estimates of $\tan \delta$ with their 95% confidence bounds.

4. Discussion

The lack of subsurface reflections in the central part of Lucus Planum can be the result of several factors, some of which depend on surface properties. A high topographic roughness at scales comparable to the radar wavelength causes scattering of the incident pulse, resulting in weaker surface and subsurface echoes. However, RMS heights estimated from MOLA data both over baselines of a few to several kilometers [Kreslavsky and Head, 2000] and within the MOLA footprint [Neumann *et al.*, 2003] are higher in area C, where subsurface detections are frequent, than in the central part of Lucus Planum. Another possibility is that the basal roughness is higher in its central part. Because subsurface echoes appear to be associated with areas of distinct surface morphology, a third possibility is that the central part of Lucus Planum consists of denser, more radar-attenuating material.

Values of ϵ in areas A and C are similar to those found by Watters *et al.* [2007] and Carter *et al.* [2009], while those in area B appears to be higher, although the estimate is affected by a larger uncertainty. The same trend, both in values and confidence intervals, is observed also for $\tan \delta$. It can also be seen in Figure 4 that the spatial density of subsurface interface detections is much higher in areas A and C than in area B, in spite of a comparable density of coverage (see Figure 2). Surface roughness at kilometer scale in area B is similar to that of the central part of Lucus Planum and smaller than that of area C [Kreslavsky and Head, 2000], in spite of the different surface morphology, while roughness in area B at hundred meters scale is comparable to that of area C [Neumann *et al.*, 2003]. Because the dearth and weakness of subsurface echoes in area B do not correlate with a higher surface roughness compared to areas A and C, we favor the interpretation that, in spite of the large uncertainties, the higher value of the complex relative permittivity in area B is an indication of a change in bulk dielectric properties with respect to areas A and C.

The relative dielectric constant of volcanic rocks such as those thought to constitute the Martian crust is variable, ranging between 2.5 for pumice to about 10 or even higher for dense basalts [Ulaby *et al.*, 1986, Appendix E]. Following a search in the literature and a set of new measurements between 0.01 and 10 MHz, Rust *et al.* [1999] concluded that the main factor in determining the value of ϵ is porosity, finding the following empirical relation for dacitic rocks:

$$(\epsilon)^{0.96} = \Phi + 6.51 (1 - \Phi) \quad (7)$$

where Φ is porosity. Such relation is similar to estimates for other nonbasaltic rocks, and holds also for basalts, although with a greater variability due perhaps to Fe-Ti oxide mineral content [Rust *et al.*, 1999]. Modeling the dependence of ϵ on porosity through equation (7), and inverting such equation to obtain estimates of Φ , it is found that the values of ϵ in Table 1 are consistent with a porosity between 0.6 and 0.9 for area A, up to 0.85 for area B, and between 0.7 and 0.8 for area C. Such high values are typical of volcanic rocks extruded through explosive, rather than effusive, processes. In their study of the correlation between the distribution of porosity in pyroclasts and eruption styles, Mueller *et al.* [2011] found that porosity values above 0.5 are characteristic of the product of explosive basaltic eruptions (eruptions of gas-rich, low-viscosity magma [e.g., Wilson and Head, 1994]) or even highly explosive subplinian, plinian, or ultraplinian eruptions, whose deposits derive from fallout or pyroclastic density currents.

Potential sources of Martian pyroclastic deposits have been discussed in the literature [e.g., Hynek *et al.*, 2003; Carter *et al.*, 2009] and might be largely related to Tharsis. Also Apollinaris Patera is at close reach for Lucus Planum [Kerber *et al.*, 2011]. The role of possibly buried volcanic edifices has been suggested by Carter *et al.* [2009]: evidence of such edifices has not been found so far in the area.

Tridymite has recently been discovered by the Curiosity rover within lacustrine sediments in Gale crater [Morris *et al.*, 2016], suggesting the presence of silica-rich volcanics within the crater's watershed. If the Medusae Fossae Formation is composed of explosive volcanic material, the upper portion of Mount Sharp, which is thought to be a part of the MFF [Thomson *et al.*, 2011; Tanaka *et al.*, 2014], could be a source for that material.

Another plausible explanation of the nature of Lucus Planum materials and the Medusae Fossae Formation in general is that they might consist of ice-rich dust or ice-laden porous rock, although previous estimates of dielectric properties based on radar data proved inconclusive [Watters *et al.*, 2007; Carter *et al.*, 2009].

The permittivity of a mixture of ice and dust can be estimated using a mixing formula. Because of the lack of knowledge about the size and shape of pores or ice inclusions in the rock, in the following analysis we selected the general Polder-van Santen model [Polder and van Santen, 1946]. This formula is one of the simplest and yet more widely used, and it has the special property that treats the inclusions and the hosting material symmetrically; it balances both mixing components with respect to the unknown effective medium, using the volume fraction of each component as a weight:

$$(1 - f) \frac{\epsilon_h - \epsilon_{\text{eff}}}{\epsilon_h + 2\epsilon_{\text{eff}}} + f \frac{\epsilon_i - \epsilon_{\text{eff}}}{\epsilon_i + 2\epsilon_{\text{eff}}} = 0 \quad (8)$$

where f is the volume fraction of inclusions in the mixture, ϵ_h is the permittivity of the host material, ϵ_i is the permittivity of the inclusions, and ϵ_{eff} is the effective permittivity of the mixture.

Water ice has a relative dielectric constant well within the range of values typical of porous rocks (2–6), while its loss tangent can vary by orders of magnitude as a function of temperature in the range 100–270 K, which is applicable to Martian conditions. Using the empirical formulas presented in Mätzler [1998] and a mean surface temperature of 210 K, typical for the latitudes of Lucus Planum according to Mellon *et al.* [2004], it is found that the real part of the permittivity of water ice is ≈ 3.1 , and the loss tangent is $\approx 5 \cdot 10^{-5}$. We hypothesized that the relative dielectric constant of the rocky component in the Lucus Planum material could range from 7 to 15 [Rust *et al.*, 1999], and that its loss tangent could independently vary between 10^{-3} and 10^{-1} [Ulaby *et al.*, 1986, Appendix E]. The Polder–van Santen mixing rule was then used to model the effective permittivity of all possible combinations of relative dielectric constant, loss tangent, and porosity, similar to the method described in Alberti *et al.* [2012].

Comparing the results with the estimated values in areas A, B, and C, we found that no mixture of rock and ice could produce a complex permittivity compatible with that of areas A and C. It is possible to obtain compatible permittivity values for these two areas using a three-component mixture, that is rock, ice, and void,

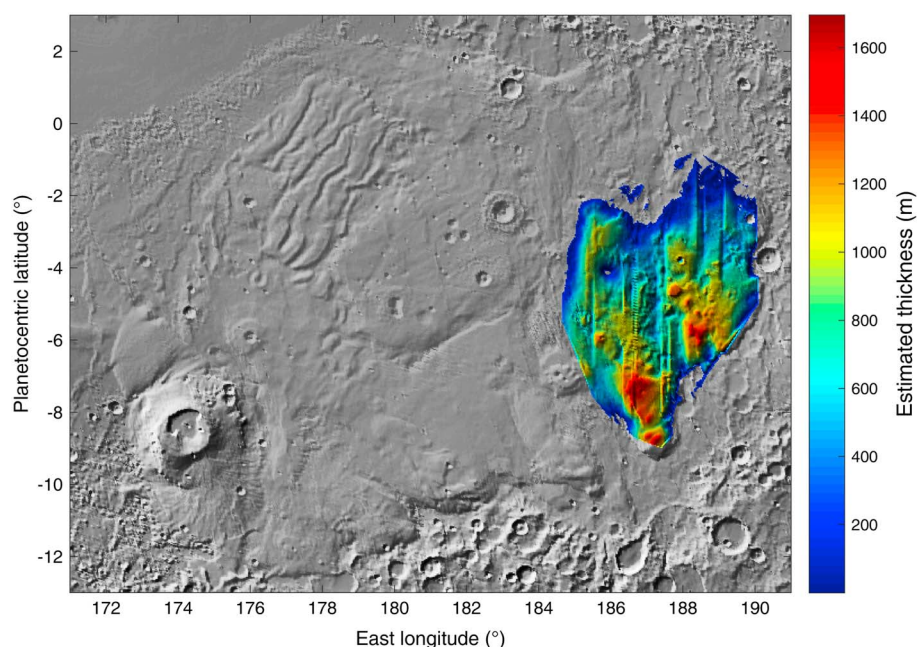


Figure 8. Map of the thickness of the eastern part of Lucus Planum, superimposed to a shaded relief map covering the same area shown in Figure 4. Thickness has been determined using the estimated relative dielectric constant in Table 1 and interpolated through the natural neighbor method.

but the significance of this result, given the weakly constrained multidimensional parameter space, is difficult to assess. In the map of water-equivalent hydrogen content for the Martian soil produced by *Feldman et al.* [2004], the Lucus Planum area appears to be relatively water rich, with a water-equivalent hydrogen content estimated at around 8%. This value however is referred to the first meter of depth, while the dielectric permittivity derived from MARSIS data is an average over the whole thickness of the Lucus Planum and Apollinaris Patera deposits.

For area B, mixtures with an ice volume fraction between 0.3 and 0.9 and a loss tangent for the rocky material comprised between $3 \cdot 10^{-3}$ and $3 \cdot 10^{-2}$ could return a range of permittivity values consistent with estimates. To determine the significance of this result, we also computed the effective permittivity of a mixture of rock and void (empty pores) over the same parameter space. We found that values consistent with those of area B could be obtained for a range of porosity and loss tangent values similar to that of the mixture of rock and ice. We thus conclude that the nature of the bulk material in area B cannot be reliably determined using only the data provided by this analysis.

Area C presents the highest number and density of subsurface detections and the smallest uncertainty in the estimates of dielectric properties. We therefore inserted in equation (1) the value of ϵ from Table 1 to estimate the thickness of the Lucus Planum deposits in such area and then interpolated this quantity over area C through the natural neighbor method. The result is included in Figure 8, in which the color-coded thickness is layered on a shaded relief map of Lucus Planum. The deposits are several hundred meters thick on average, locally reaching a thickness up to 1.5 km, for a total volume of $\approx 6.8 \cdot 10^4 \text{ km}^3$. The deposit thickness varies positively with regional elevations, being higher in the south and lower in the north.

Overall, Lucus Planum subsurface as sounded by MARSIS appears to be locally to regionally inhomogeneous. This can be interpreted in terms of complex, multiprocess components of the deposits constituting Lucus Planum and possibly the Medusa Fossae Formation as a whole [e.g., *Kerber and Head, 2010; Kerber, 2014*]. An interplay between a possibly dominating volcano-sedimentary component and local, possibly late-stage erosional and partially depositional episodes could be envisaged. Such episodes on or in the vicinity of Lucus Planum likely occurred in the relatively recent past [*Harrison et al., 2013*], leading to extensive resedimentation of Lucus Planum materials [*Kerber and Head, 2012*].

The dielectric properties of the northwestern part of Lucus Planum, implying a higher density compared to the other radar-transparent areas, and the inferred strong attenuation of the radar signal in its central part

could be interpreted as due to the presence of indurated sedimentary deposits. Their existence within Lucus Planum is consistent with extensive reworking of those deposits through time [Harrison *et al.*, 2013; Kerber and Head, 2012; Kerber, 2014]. Although such deposits could be compositionally similar to the overall MFF materials [Harrison *et al.*, 2013], they could be locally remobilized, thus changing their architecture, structure, and texture, including their degree of cementation [Kerber and Head, 2012]. Their Hesperian-Amazonian age [Kerber and Head, 2010; Tanaka *et al.*, 2014] would match both late stage valley network as well as vigorous Tharsis activity. MARSIS data cannot shed much light on small- to medium-scale lateral and vertical variations, which will require additional work at an appropriate scale.

The possibility of sampling with Mars Science Laboratory on its way uphill on Mount Sharp in Gale Crater some material, even not necessarily in situ but made available through mass wasting and resedimentation, would allow for some indirect ground truth: pyroclastic, possibly acidic, volcanic material of an age comparable with that of Lucus Planum (Hesperian to Amazonian) would offer support to a pyroclastic origin of the MFF. On the other hand, resedimented material so far from the MFF main bodies would not allow for volatiles to be embedded and preserved.

5. Conclusions

MARSIS acquired 238 radar swaths across Lucus Planum, providing sufficient coverage for the study of the internal structure and dielectric properties of this part of the MFF. Subsurface reflections were found only in three areas, marked by a distinctive surface morphology, while the central part of Lucus Planum appears to be made of radar-attenuating material preventing the detection of basal echoes. The bulk dielectric constant of these areas was estimated by comparing their apparent thickness from radar data with their basal topography, extrapolated from the surrounding terrains. The complex part of the dielectric permittivity was derived from the weakening of basal echoes as a function of apparent depth, yielding results that are consistent with the estimated dielectric constant. The inferred bulk properties were compared with known materials such as volcanic rocks and ice-dust mixtures. The interpretation that the eastern area of Lucus Planum and the deposits on the northwestern flanks of Apollinaris Patera consist of high-porosity pyroclastic material is strongly supported by results, while northwestern Lucus Planum is likely to be much less porous. No conclusion could be drawn about the presence of pore ice.

All evidence points to Lucus Planum being highly inhomogeneous. The exact origin of the deposits cannot be constrained by radar data alone, but our results are consistent with an overall pyroclastic origin as suggested by Carter *et al.* [2009] and Watters *et al.* [2007] for the MFF. The geological complexity of the subsurface revealed through MARSIS data is consistent with a combination of processes acting through space and time, including fluvial (possibly outflow-related) activity occurred during the emplacement of the *Htu* and *AHtu* units [Tanaka *et al.*, 2014], as well as eolian deposition [Kerber and Head, 2012]. The overall surface textural and topographical heterogeneity might be linked to postemplacement erosional processes related to regional wind dynamics [Kerber and Head, 2012] on a variably indurated substrate. The evidence in this work was not sufficient to demonstrate the presence of an ice-related component in the central part of Lucus Planum, although this cannot be conclusively excluded. A full understanding of such a complex geological history will require the integration of several data sets at different scales and with different resolutions.

References

- Alberti, G., L. Castaldo, R. Orosei, A. Frigeri, and G. Cirillo (2012), Permittivity estimation over Mars by using SHARAD data: The Cerberus Palus area, *J. Geophys. Res.*, **117**, E09008, doi:10.1029/2012JE004047.
- Armand, N. A., V. M. Smirnov, and T. Hagfors (2003), Distortion of radar pulses by the Martian ionosphere, *Radio Sci.*, **38**, 1090, doi:10.1029/2002RS002849.
- Bloomfield, P., and W. L. Steiger (1983), *Least Absolute Deviations, Progress in Probability and Statistics*, vol. 6, Birkhäuser Boston, Boston, Mass.
- Bradley, B. A., S. E. H. Sakimoto, H. Frey, and J. R. Zimbleman (2002), Medusae Fossae Formation: New perspectives from Mars Global Surveyor, *J. Geophys. Res.*, **107**, 5058, doi:10.1029/2001JE001537.
- Campbell, B. A., and T. R. Watters (2016), Phase compensation of MARSIS subsurface sounding data and estimation of ionospheric properties: New insights from SHARAD results, *J. Geophys. Res. Planets*, **121**, 180–193, doi:10.1002/2015JE004917.
- Cartacci, M., E. Amata, A. Cicchetti, R. Nosciese, S. Giuppi, B. Langlais, A. Frigeri, R. Orosei, and G. Picardi (2013), Mars ionosphere total electron content analysis from MARSIS subsurface data, *Icarus*, **223**, 423–437, doi:10.1016/j.icarus.2012.12.011.
- Carter, L. M., et al. (2009), Shallow radar (SHARAD) sounding observations of the Medusae Fossae Formation, Mars, *Icarus*, **199**, 295–302, doi:10.1016/j.icarus.2008.10.007.
- Feldman, W. C., et al. (2004), Global distribution of near-surface hydrogen on Mars, *J. Geophys. Res.*, **109**, E09006, doi:10.1029/2003JE002160.
- Grima, C., D. D. Blankenship, D. A. Young, and D. M. Schroeder (2014), Surface slope control on firn density at Thwaites Glacier, West Antarctica: Results from airborne radar sounding, *Geophys. Res. Lett.*, **41**, 6787–6794, doi:10.1002/2014GL061635.

Acknowledgments

This work was supported by the Italian Space Agency (ASI) through contract I/032/12/1. The numerical code for the simulation of surface scattering was developed at the Consorzio Interuniversitario per il Calcolo Automatico dell'Italia Nord-Orientale (CINECA) in Bologna, Italy. Simulations were produced thanks to the Partnership for Advanced Computing in Europe (PRACE), awarding us access to the SuperMUC computer at the Leibniz-Rechenzentrum, Garching, Germany, through project 2013091832. Test simulations were run at Jacobs University CLAMV HPC cluster, and we are grateful to Achim Gelessus for his support. This research has made use of NASA's Astrophysics Data System. A.P.R. has been supported by the European Union FP7 and Horizon 2020 research and innovation programmes under grant agreements 654367 (EarthServer-2) and 283610 (EarthServer). Data used in this analysis have been taken from the MARSIS public data archive, which is currently undergoing validation before publication on the Planetary Science Archive of the European Space Agency (<http://www.cosmos.esa.int/web/psa/mars-express>) and mirroring on NASA's Planetary Data System Geosciences Node (http://pds-geosciences.wustl.edu/missions/mars_express/marsis.htm). Simulations used in the identification of subsurface interfaces will be published in the same archives at a yet undefined date, but they are already available in the following public repository: <https://doi.org/10.5281/zenodo.582651>.

- Harrison, S., M. Balme, A. Hagermann, J. Murray, and J.-P. Muller (2010), Mapping Medusae Fossae Formation materials in the southern highlands of Mars, *Icarus*, 209(2), 405–415.
- Harrison, S. K., M. R. Balme, A. Hagermann, J. B. Murray, J.-P. Muller, and A. Wilson (2013), A branching, positive relief network in the middle member of the Medusae Fossae Formation, equatorial Mars—Evidence for sapping?, *Planet. Space Sci.*, 85, 142–163, doi:10.1016/j.pss.2013.06.004.
- Head III, J. W., and M. Kreslavsky (2004), Medusae Fossae Formation: Ice-Rich Airborne Dust Deposited During Periods of High Obliquity?, paper presented at the 35th Lunar and Planetary Science Conference, Lunar and Planet. Inst., Houston, Tex.
- Hynek, B. M., R. J. Phillips, and R. E. Arvidson (2003), Explosive volcanism in the Tharsis region: Global evidence in the Martian geologic record, *J. Geophys. Res.*, 108, 5111, doi:10.1029/2003JE002062.
- Kerber, L. (2014), The Distribution and Diversity of Layering Within the Medusae Fossae Formation, paper presented at the 45th Lunar and Planetary Science Conference, Lunar and Planet. Inst., The Woodlands, Tex.
- Kerber, L., and J. W. Head (2010), The age of the Medusae Fossae Formation: Evidence of Hesperian emplacement from crater morphology, stratigraphy, and ancient lava contacts, *Icarus*, 206, 669–684, doi:10.1016/j.icarus.2009.10.001.
- Kerber, L., and J. W. Head (2012), A progression of induration in Medusae Fossae Formation transverse aeolian ridges: Evidence for ancient aeolian bedforms and extensive reworking, *Earth Surf. Processes Landforms*, 37(4), 422–433.
- Kerber, L., J. W. Head, J.-B. Madeleine, F. Forget, and L. Wilson (2011), The dispersal of pyroclasts from Apollinaris Patera, Mars: Implications for the origin of the Medusae Fossae Formation, *Icarus*, 216, 212–220, doi:10.1016/j.icarus.2011.07.035.
- Kreslavsky, M. A., and J. W. Head (2000), Kilometer-scale roughness of Mars: Results from MOLA data analysis, *J. Geophys. Res.*, 105, 26,695–26,712, doi:10.1029/2000JE001259.
- Mandt, K. E., S. L. de Silva, J. R. Zimbelman, and D. A. Crown (2008), Origin of the Medusae Fossae Formation, Mars: Insights from a synoptic approach, *J. Geophys. Res.*, 113, E12011, doi:10.1029/2008JE003076.
- Mätzler, C. (1998), Microwave properties of ice and snow, in *Solar System Ices, Astrophysics and Space Science Library*, vol. 227, edited by B. Schmitt, C. de Bergh, and M. Festou, pp. 241–257, Kluwer, Acad., Dordrecht, Netherlands, doi:10.1007/978-94-011-5252-5_10.
- Mellon, M. T., W. C. Feldman, and T. H. Prettyman (2004), The presence and stability of ground ice in the southern hemisphere of Mars, *Icarus*, 169, 324–340, doi:10.1016/j.icarus.2003.10.022.
- Morris, R. V., et al. (2016), Silicic volcanism on Mars evidenced by tridymite in high-SiO₂ sedimentary rock at Gale crater, *Proc. Natl. Acad. Sci.*, 113(26), 7071–7076, doi:10.1073/pnas.1607098113.
- Mouginot, J., W. Kofman, A. Safaeinili, and A. Herique (2008), Correction of the ionospheric distortion on the MARSIS surface sounding echoes, *Planet. Space Sci.*, 56, 917–926, doi:10.1016/j.pss.2008.01.010.
- Mouginot, J., W. Kofman, A. Safaeinili, C. Grima, A. Herique, and J. J. Plaut (2009), MARSIS surface reflectivity of the south residual cap of Mars, *Icarus*, 201, 454–459, doi:10.1016/j.icarus.2009.01.009.
- Mueller, S., B. Scheu, U. Kueppers, O. Spieler, D. Richard, and D. B. Dingwell (2011), The porosity of pyroclasts as an indicator of volcanic explosivity, *J. Volcanol. Geotherm. Res.*, 203, 168–174, doi:10.1016/j.jvolgeores.2011.04.006.
- Neumann, G. A., J. B. Abshire, O. Aharonson, J. B. Garvin, X. Sun, and M. T. Zuber (2003), Mars Orbiter Laser Altimeter pulse width measurements and footprint-scale roughness, *Geophys. Res. Lett.*, 30, 1561, doi:10.1029/2003GL017048.
- Nouvel, J.-F., A. Herique, W. Kofman, and A. Safaeinili (2004), Radar signal simulation: Surface modeling with the Facet Method, *Radio Sci.*, 39, RS1013, doi:10.1029/2003RS002903.
- Ogilvy, J. A. (1991), *Theory of Wave Scattering From Random Rough Surfaces*, IOP Publ., Bristol.
- Picardi, G., and S. Sorge (2000), Adaptive compensation of ionosphere dispersion to improve subsurface detection capabilities in low-frequency radar systems, in *Proceedings of 8th International Conference on Ground Penetrating Radar, SPIE*, vol. 4084, edited by D. A. Noon, G. F. Stickley, and D. Longstaff, pp. 624–629, SPIE - The Int. Soc. for Opt. Eng., Bellingham, Wash.
- Picardi, G., et al. (2005), Radar soundings of the subsurface of Mars, *Science*, 310, 1925–1928, doi:10.1126/science.1122165.
- Polder, D., and J. H. van Santen (1946), The effective permeability of mixtures of solids, *Physica*, 12, 257–271.
- Porcello, L. J., R. L. Jordan, J. S. Zelenka, G. F. Adams, R. J. Phillips, W. E. Brown, S. H. Ward Jr., and P. L. Jackson (1974), The Apollo lunar sounder radar system, *Proc. IEEE*, 62, 769–783.
- Rust, A. C., J. K. Russell, and R. J. Knight (1999), Dielectric constant as a predictor of porosity in dry volcanic rocks, *J. Volcanol. Geotherm. Res.*, 91, 79–96, doi:10.1016/S0377-0273(99)00055-4.
- Safaeinili, A., W. Kofman, J. Mouginot, Y. Gim, A. Herique, A. B. Ivanov, J. J. Plaut, and G. Picardi (2007), Estimation of the total electron content of the Martian ionosphere using radar sounder surface echoes, *Geophys. Res. Lett.*, 34, L23204, doi:10.1029/2007GL032154.
- Schultz, P., and A. B. Lutz (1988), Polar wandering of Mars, *Icarus*, 73, 91–141, doi:10.1016/0019-1035(88)90087-5.
- Scott, D. H., and K. L. Tanaka (1982), Ignimbrites of Amazonis Planitia region of Mars, *J. Geophys. Res.*, 87, 1179–1190, doi:10.1029/JB087iB02p01179.
- Seu, R. et al. (2007), SHARAD sounding radar on the Mars Reconnaissance Orbiter, *J. Geophys. Res.*, 112, E05S05, doi:10.1029/2006JE002745.
- Sibson, R. (1981), A brief description of natural neighbor interpolation, in *Interpolating Multivariate Data*, edited by V. Barnett, pp. 21–36, John Wiley, New York.
- Smirnov, V. M., and O. V. Yushkova (2013), The influence of the ionosphere in subsurface Martian soil sounding experiments and a method of its correction, *Sol. Syst. Res.*, 47, 430–436, doi:10.1134/S0038094613060099.
- Smith, D. E., et al. (2001), Mars Orbiter Laser Altimeter: Experiment summary after the first year of global mapping of Mars, *J. Geophys. Res.*, 106, 23,689–23,722, doi:10.1029/2000JE001364.
- Spagnuolo, M. G., F. Grings, P. Perna, M. Franco, H. Karszenbaum, and V. A. Ramos (2011), Multilayer simulations for accurate geological interpretations of SHARAD radargrams, *Planet. Space Sci.*, 59, 1222–1230, doi:10.1016/j.pss.2010.10.013.
- Tanaka, K. L. (2000), Dust and ice deposition in the Martian geologic record, *Icarus*, 144, 254–266, doi:10.1006/icar.1999.6297.
- Tanaka, K. L., J. A. Skinner, J. M. Dohm, R. P. Irwin III, E. J. Kolb, C. M. Fortezzo, T. Platz, G. G. Michael, and T. Hare (2014), *Geologic Map of Mars, U.S. Geol. Surv. Sci. Invest. Map*, 3292, U.S. Dep. of the Inter., Reston, Va.
- Thomson, B. J., N. T. Bridges, R. Milliken, A. Baldridge, S. J. Hook, J. K. Crowley, G. M. Marion, C. R. de Souza Filho, A. J. Brown, and C. M. Weitz (2011), Constraints on the origin and evolution of the layered mound in Gale crater, Mars using Mars Reconnaissance Orbiter data, *Icarus*, 214, 413–432, doi:10.1016/j.icarus.2011.05.002.
- Ulaby, F. T., R. K. Moore, and A. K. Fung (1986), *Microwave Remote Sensing: Active and Passive*, Artech House microwave library, vol. 3, Addison-Wesley Publ. Co., Adv. Book Program/ World Sci. Div., Norwood, Mass.

- Watters, T. R., et al. (2007), Radar sounding of the Medusae fossae formation Mars: Equatorial ice or dry, low-density deposits?, *Science*, *318*, 1125–1128, doi:10.1126/science.1148112.
- Wilson, L., and J. W. Head (1994), Mars: Review and analysis of volcanic eruption theory and relationships to observed landforms, *Rev. Geophys.*, *32*(3), 221–263.
- Zhang, Z., E. Nielsen, J. J. Plaut, R. Orosei, and G. Picardi (2009), Ionospheric corrections of MARSIS subsurface sounding signals with filters including collision frequency, *Planet. Space Sci.*, *57*, 393–403, doi:10.1016/j.pss.2008.11.016.