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Galaxy evolution studies with the SPace IR telescope for Cosmology and Astrophysics (SPICA): the power of IR spectroscopy

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Abstract

IR spectroscopy in the range 12–230 μm with the SPace IR telescope for Cosmology and Astrophysics (SPICA) will reveal the physical processes that govern the formation and evolution of galaxies and black holes through cosmic time, bridging the gap between the James Webb Space Telescope (JWST) and the new generation of Extremely Large Telescopes (ELTs) at shorter wavelengths and the Atacama Large Millimeter Array (ALMA) at longer wavelengths. SPICA, with its 2.5-m telescope actively-cooled to below 8 K, will obtain the first spectroscopic determination, in the mid-IR rest-frame, of both the star-formation rate and black hole accretion rate histories of galaxies, reaching lookback times of 12 Gyr, for large statistically significant samples. Densities, temperatures, radiation fields and gas-phase metallicities will be measured in dust-obscured galaxies and active galactic nuclei (AGN), sampling a large range in mass and luminosity, from faint local dwarf galaxies to luminous quasars in the distant Universe. AGN and starburst feedback and feeding mechanisms in distant galaxies will be uncovered through detailed measurements of molecular and atomic line profiles. SPICA's large-area deep spectrophotometric surveys will provide mid-IR spectra and continuum fluxes for unbiased samples of tens of thousands of galaxies, out to redshifts of $z \sim 6$. Furthermore, SPICA spectroscopy will uncover the most luminous galaxies in the first few hundred million years of the Universe, through their characteristic dust and molecular hydrogen features.

Keywords: galaxies: evolution – galaxies: active – galaxies: starburst – infrared: galaxies – techniques: IR spectroscopy

1 INTRODUCTION

Over the past three decades, we have learned that at least half of the energy ever emitted by stars and accreting black holes in galaxies is absorbed by dust, and re-radiated in the infrared (e.g. Hauser & Dwek 2001; Lagache et al. 2005; Franceschini et al. 2008). We now know that the peak in the growth of galaxies occurs at redshifts of $z \sim 1\text{--}3$ (Lilly et al. 1996; Franceschini et al. 1999; Madau & Dickinson 2014; Rowan-Robinson

et al. 2016), when the Universe was roughly 3 Gyr old — a result achieved primarily through deep and wide-field observations with previous IR space observatories, namely the IR Astronomical Satellite (IRAS, Neugebauer et al. 1984), the IR Space Observatory (ISO, Kessler et al. 1996), *Spitzer* (Werner et al. 2004), AKARI (Murakami et al. 2007), *Herschel* (Pilbratt et al. 2010) and the Wide-field IR Survey Explorer (WISE, Wright et al. 2010). Despite their successes, these observatories had either small cold telescopes, or large, warm mirrors, ultimately limiting their ability

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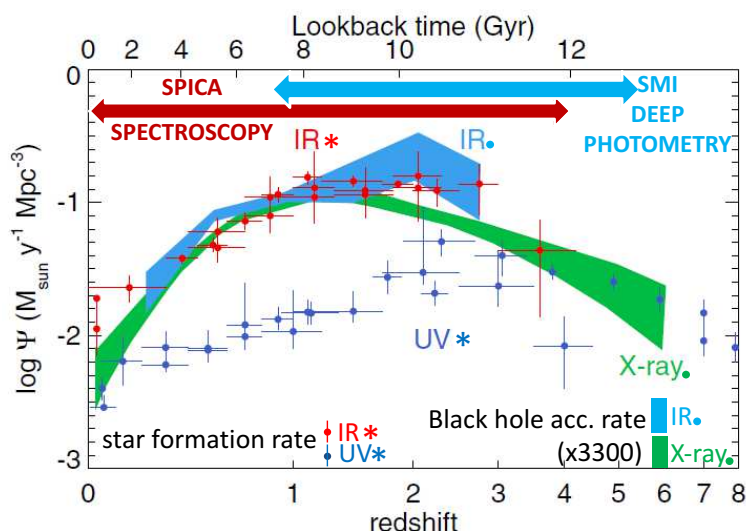


Figure 1. Estimated star-formation rate densities from the far-ultraviolet (FUV, blue points) and far-IR (FIR, red points) photometric surveys (adapted from Madau & Dickinson 2014). The estimated black hole accretion rate density, scaled up by a factor of 3300, is shown for comparison (in green shading from X-rays and light blue from the IR). For redshifts above $z=2$, very little IR data on the SFR density exist, making its determination rather poor. The redshift ranges probed with SPICA spectroscopy (up to $z=4$) and photometry (up to $z=6$) are also shown.

to probe the physics, through spectroscopy and deep photometry, of the faintest and most distant obscured sources in our Universe.

Due to the progress in detector performance and cryogenic cooling technologies, great advances in our ability to study the hidden, dusty Universe can be made through observations in the thermal infrared. The SPace IR telescope for Cosmology and Astrophysics (SPICA, Swinyard et al. 2009; Nakagawa et al. 2014) will achieve a gain of over two orders of magnitude in spectroscopic sensitivity in the mid/far-IR compared to *Herschel* and *Spitzer*. SPICA will provide access to wavelengths well beyond those reachable with the James Webb Space Telescope (JWST, Gardner et al. 2006) and the new generation of extremely large telescopes (ELTs), and at wavelengths shortward of those accessible by the Atacama Large Millimeter/Submillimeter Array (ALMA, Wootten & Thompson 2009). It will enable the discovery and detailed study of normal galaxies across their key phases of evolution, as well as probing the earliest forming galaxies and super-massive black holes.

A description of the SPICA mission was originally presented in Swinyard et al. (2009) and recent updates can be found in Nakagawa et al. (2014); Sibthorpe et al. (2016), and in Roelfsema et al. (2017). SPICA will consist of a 2.5-m primary mirror, actively cooled to below 8K. There are two primary instruments sharing the focal plane: the Far-IR Instrument (SAFARI) and the Mid-IR Instrument (SMI). SAFARI includes a grating spectrometer covering the 34–230 μm spectral range simultaneously at a resolution $R \sim 300$. The combination of the grating with a Martin-Puplett interferometer allows for observations at higher spectral resolution ($1500 < R < 11000$, depending on wavelength) over the

same spectral range. SAFARI also includes the imaging polarimeter POL, with a field of view of $80'' \times 80''$ in three spectral bands centred at 110, 220 and 350 μm . A description of the SAFARI optical system architecture and design concept is given in Pastor et al. (2016). SMI covers the wavelength range of 12–36 μm , using three spectroscopic channels: low-resolution ($R = 50\text{--}120$, 17–36 μm); mid-resolution ($R = 1300\text{--}2300$; 18–36 μm); and high-resolution ($R = 28000$, 12–18 μm). SMI can also obtain large field-of-view ($10' \times 12'$) images at 34 μm . A full description of the SMI instrument can be found in Kaneda et al. (2016).

This article describes how SPICA will be able to address the study of galaxy formation and evolution, while a companion article focuses on the studies of the ISM in nearby galaxies (van der Tak et al. 2017). It is organised as follows. Section 2 introduces the current knowledge that has been accumulated so far in studying galaxy evolution and identifies a number of open questions that can be addressed and answered with high sensitivity IR spectroscopic and photometric observations.

We show in Sect. 3 how IR spectroscopy is able to separate the two main energy production mechanisms of star formation and black hole accretion driving galaxy evolution up to redshift of $z=3\text{--}4$. Finally, Sect. 4 illustrates the synergies of the SPICA observations with current and future facilities at other frequencies and sect. 5 gives the conclusions.

2 The rise and fall of galaxy formation

The bulk of star formation and supermassive black hole accretion in galaxies appears to have taken place more than six billion years ago, dropping sharply towards the present epoch (e.g., Madau & Dickinson 2014, and ref-

ferences therein, see Fig. 1). The Madau & Dickinson compilation indicates that the most of the UV/optical photons from hot young stars characteristic of active star-forming areas that combine to make star-forming galaxies have been absorbed by dust and re-radiated. Thus as a general condition, rest-frame ultraviolet and optical does not access these crucial regions where gas forms massive stars and vice versa. Moreover, since around half of the energy emitted by stars and accreting super-massive black holes is absorbed and re-emitted by dust (e.g. Hauser & Dwek 2001), understanding the physics of galaxy evolution requires IR observations of large, unbiased samples of galaxies spanning a wide range in luminosity, redshift, environment and nuclear activity. From *Spitzer* and *Herschel* photometric surveys the star-formation rate (SFR) and black-hole accretion rate (BHAR) density functions have been *estimated* through the bolometric luminosities of galaxies (Le Floc'h et al. 2005; Gruppioni et al. 2013; Delvecchio et al. 2014). However, these estimates should be treated with caution, because they are typically based on observations of only a few, broad IR bands, making the IR luminosities and the relative contribution of star formation and black hole accretion, highly uncertain. This crucial separation has been attempted so far through modelling of the spectral energy distributions (SEDs) and relies on model-dependent assumptions and local templates, with large uncertainties and degeneracies. Indeed *Spitzer*- and *Herschel*-based studies have been successful in estimating counts and evolving galaxy luminosity functions (Le Floc'h et al. 2005; Pérez-González et al. 2005; Oliver et al. 2012; Magnelli et al. 2013; Lapi et al. 2011), but only through statistical techniques applied to bulk galaxy populations as a function of redshift.

Determinations of the SFR from ultraviolet (UV) (e.g. Bouwens et al. 2007) and optical spectroscopy (e.g., from the *Sloan* Digital Sky Survey, Eisenstein et al. 2011) are based on measurements of only around 10% of the total integrated light, which escapes the dust absorption. These must therefore be corrected upwards by large, uncertain extinction factors (Fig. 1). X-ray analyses of the BHAR, similarly, are prone to large uncertainties, because deriving the bolometric luminosity from the X-ray luminosity depends on uncertain bolometric corrections (e.g., Vasudevan & Fabian 2009) and on estimates of the Compton-thick population contribution.

IR emission line spectra can be used to *physically* separate these contributions to the total integrated light on a galaxy-by-galaxy basis, directly measure redshifts, SFRs, BHARs, metallicities and physical properties of gas and dust in galaxies. With SPICA, we will be able to do this for lookback times up to about 12 Gyrs.

With its large, cold telescope and powerful instruments, SPICA will peer into the dust-enshrouded phases of galaxy formation and evolution, revealing the physical, dynamical and chemical states of the gas and dust. With such a unique power, SPICA will probe the histories of star-formation and black-hole accretion rates through cosmic time and how these processes drive galaxy evolution. It will provide detailed answers to the following questions:

1. How does accretion and feedback from star formation and AGN shape galaxy evolution? (see section 2.1)
2. How are metals and dust produced and destroyed in galaxies? What is the metallicity evolution in galaxies as a function of redshift? (see section 2.2)
3. How and when do early black-holes and starburst appear close to the epoch of re-ionization? (see section 2.3)
4. How did primordial gas clouds collapse into the first galaxies and black holes? (see section 2.4)

In the following four sections we briefly address each of the first four questions, reserving detailed analysis of the role of SPICA in answering these to four companion articles: i) the role of feeding and feedback in galaxy evolution (González-Alfonso et al. 2017b); ii) the chemical evolution of galaxies and the rise of metals and dust (Fernández-Ontiveros et al. 2017); iii) dust obscured star-formation and accretion histories from re-ionization using SPICA unbiased photometric (Gruppioni et al. 2017) and low spectral resolution spectroscopic surveys (Kaneda et al. 2017); iv) the first stars and galaxies (Egami et al. 2017). We will detail in Sect. 3 the power of using the many IR fine structure lines which fall in the wavelength range of the SPICA instrument for distant AGN and starbursts up to a redshift of $z=4$.

2.1 AGN Feeding and Feedback in the context of galaxy evolution

The correlations between SMBH masses and their host galaxy properties (Magorrian et al. 1998; Ferrarese & Merritt 2000), and the bimodality of the colour distribution of local galaxies (e.g., Strateva et al. 2001; Baldry et al. 2004), suggest that the growth of BH and stellar mass are related throughout the lifetime of a galaxy. In other words, feedback between SMBH and galactic star formation may be in part responsible for the $M_{BH}-\sigma$ relationship seen in the local Universe (e.g., Silk & Rees 1998; Di Matteo et al. 2005; Springel et al. 2005). It is precisely this feedback on the dense, circum-nuclear ISM that we can study with SPICA in the far-IR.

Herschel spectroscopic observations of far-IR OH lines (Fischer et al. 2010; Sturm et al. 2011; González-

Alfonso et al. 2014, 2017a) have shown that fast, molecular outflows are common among AGN-powered local ULIRGs. SPICA spectroscopy will allow us to search for these outflows well beyond the nearest galaxies, reaching up to $z=1-2$, and therefore providing an estimate of the impact of AGN-driven feedback at the peak epoch of star formation. With the SAFARI instrument on SPICA we will be able to detect the spectral signatures of outflowing dense molecular gas (through P-Cygni profiles on the OH 79 μm and 119 μm lines and blue-shifted high velocity wings in the OH 65 μm line) in ULIRGs like Mrk 231 out to $z=1.5$ in a few hours of integration. This will enable surveys of hundreds of galaxies at these redshifts, providing a measure of the demographics of molecular feedback in IR-luminous galaxies at the peak of the SFR density.

Far-IR spectroscopy can also provide direct evidence for the feeding of the SMBHs and central starbursts. Inflowing gas can be identified through inverse P-Cygni profiles or redshifted absorption wings of OH and [OI]63 μm , as shown by *Herschel* (González-Alfonso et al. 2012; Falstad et al. 2015). SPICA will therefore be able to measure the dynamics of the molecular gas in and around the nuclei of rapidly evolving, dusty galaxies and study both AGN accretion and energetic feedback in significant samples of dusty galaxies over the past 10 Gyr. For a more detailed discussion of the capabilities of SPICA to detect galaxy feedback, see the companion paper by González-Alfonso et al. (2017b).

2.2 The Rise of Metals and Dust

Galactic evolution is intimately tied to the production of metals and dust. Elements heavier than He (i.e. “metals”) play a major role in gas cooling and, as a consequence, they are critical in determining the conditions of cloud collapse and ultimately of star and planet formation. The metallicity in galaxies is determined by the cumulative effects of star formation, outflows, accretion and the radial redistribution of matter. Traditional metallicity diagnostics, based on optical lines, are biased towards dust-free regions, yielding metallicities significantly lower than those inferred from the dust mass or IR lines (e.g., Pereira-Santaella et al. 2017; Santini et al. 2010; Croxall et al. 2013). Furthermore, temperature variations within galaxies, which can strongly influence optically derived abundances, have little effect on IR-derived values, which are determined from lines lying close to the ground state (Bernard Salas et al. 2001). IR fine-structure lines observed with SPICA provide extinction-free measurements of the metallicity of galaxies, nearly independent of ionization, density and temperature out to redshift $z\sim 3$. Moreover, SPICA can detect mid-IR Hydrogen recombination lines in galaxies at intermediate redshifts ($z\sim 1.5-2$), which, together

with the forbidden lines, allow for direct determination of the Ne, S, N, O and Fe abundances.

Dust also plays a critical role in heating and cooling the ISM in galaxies. It forms in the dense, enriched atmospheres of evolved stars, novae, supernovae and dense molecular clouds (Valiante et al. 2009; Zhukovska & Henning 2013; Marassi et al. 2015; Bocchio et al. 2016) and is destroyed by shocks, sputtering and intense radiation fields. However, its dominant formation and destruction channels in different environments are poorly understood (Kemper 2015; Matsuura et al. 2011; Micelotta et al. 2016). IR spectroscopy can uniquely measure the dust mass produced by evolved stars and supernovae in nearby galaxies, allowing a detailed study of the dust mineralogy and composition, via the mid-IR SiO₂, FeO, FeS and crystalline silicate features. SPICA will trace the abundance and evolution of the dust components within galaxies, constraining the local conditions, such as ionization, radiation field, dust structures and overall dust-to-gas ratios, giving us clues to the chemical evolution of galaxies (Sandstrom et al. 2012; Rémy-Ruyer et al. 2014). For the details of these investigations and the assessment of the SPICA observations, we refer to the companion papers by van der Tak et al. (2017); Fernández-Ontiveros et al. (2017), for local and distant galaxies, respectively.

2.3 Towards the Epoch of Re-ionization: early Black Holes and Starbursts

The deepest cosmological surveys with *Herschel* (e.g., Oliver et al. 2012; Gruppioni et al. 2013; Magnelli et al. 2013; Lutz 2014) mapped out the star-formation rate density to $z\sim 3$ for the first time, but detected only small numbers of the most luminous ($L \gtrsim 10^{12} L_{\odot}$) star-forming galaxies at $z>3$. High redshift ($z>6$) quasars have been detected, containing black holes as massive as $10^{10} M_{\odot}$ (Wu et al. 2015; Jun et al. 2015), but their origin, demographics and role in re-ionization are still unclear. Deep SPICA/SMI photometric surveys will extend the study of the black-hole accretion rate and the star-formation rate density well beyond $z\sim 3$, detecting, in the 34 μm observed wavelength, the hot dust around high-redshift QSOs, as well as starburst-dominated galaxies at redshifts out to $z\sim 6$. Due to its large field of view of $10' \times 12'$, SMI will map large sky areas to the confusion limit (around 5 μJy) in relatively short times, with an effective surveying speed hundred times faster than JWST.¹ Deep and wide photometric surveys with SPICA will allow us to study the build-up of the progenitors of the elliptical galaxies dominating

¹The JWST Design Reference Mission includes a MIRI (16 μJy , 5 σ) $10' \times 9'$ survey at 21 μm . With the MIRI field of view of $1.25' \times 1.88'$, this survey will need 160hrs. A similar SMI survey at 34 μm , would take 1 hr (12 μJy , 5 σ). For detecting mid-IR sources SPICA will be over 100 times faster than JWST.

local galaxy clusters, and thus to probe environmental-dependent evolution (e.g., Dannerbauer et al. 2014; Clements et al. 2016). For a more detailed description of potential photometric surveys with SPICA, and their role in extending our knowledge of the population of IR-bright galaxies at $z>3$, we refer to the companion paper by Gruppioni et al. (2017).

2.4 The First Stars and Galaxies

Through chemical enrichment and the production of dust, the earliest stars imprint their signature on the ISM of high-redshift galaxies. The first generation of quiescent, self-gravitating primordial or metal-poor gas clouds may be too faint to be detected directly with SPICA (e.g., Mizusawa et al. 2005; Santoro & Shull 2006; Gong et al. 2013). However, during the hierarchical structure formation process, both the merging of primordial clouds and feedback due to supernovae and stellar winds can produce pockets of shock-heated gas. It is postulated that molecular clouds can be efficiently formed in the cooling gas behind these shocks (Ferrara 1998; Ciardi & Ferrara 2001). If these H₂ clouds are warm and massive enough, (e.g., 10¹⁰M_⊙ of T=200 K gas and 10⁸M_⊙ of T=1000 K gas), they are detectable in their shock-excited H₂ line emission with SAFARI. Such systems of warm massive H₂ gas reservoirs are already known at lower redshift (Egami et al. 2006; Ogle et al. 2012), but can be detected by SPICA out to $z\sim 10$.

The ultimate challenge for SPICA will be to catch a glimpse of the first (i.e., Pop III) galaxies. Theoretical models predict that Pop III star clusters should produce a large amount of dust quickly, as massive pair instability supernovae (PISNe) explode (e.g., Schneider et al. 2004). When the dust is released and heated by hot, main sequence Pop III stars, the resultant mid-IR spectra exhibit strong quartz (SiO₂) emission lines, which SPICA may detect. This will give a rare glimpse into the dust-production mechanism of Pop III stars and provide strong constraints for their evolutionary models. For a detailed discussion of the potential for SPICA to detect the cooling signatures of young, luminous galaxies at high redshift, see the companion paper by Egami et al. (2017).

3 Infrared Spectroscopic Probes of Star Formation and Black Hole Accretion

To understand galaxy evolution, we need to measure the rate at which stars form and black holes accrete matter as a function of time. The apparently similar shapes of the histories of star formation and black hole accretion with redshift (see Fig. 1), the fact that most AGN also display enhanced star formation, along with the black hole-stellar mass correlation seen in the local Universe

Table 1 For the top nine most abundant metals in the universe, this table gives the fine-structure lines that will be observed by SPICA in the range $0 < z < 4$. Solar abundances (written as $\log X/H + 12$, from Grevesse et al. 2010) and critical densities for collisional de-excitation have been taken from Tielens & Hollenbach (1985); Genzel (1992); Greenhouse et al. (1993).

Atom	Solar abundance	Ion/Line	λ (μm)	I.P. (eV)	n_{crit} (cm^{-3})
O	8.69	[OI]	63.18	–	4.7×10^5
		[OI]	145.5	–	9.5×10^4
		[OIII]	51.81	35.12	3.6×10^3
		[OIII]	88.36	35.12	5.1×10^2
		[OIV]	25.89	54.93	1.0×10^4
C	8.55	[CII]	157.7	11.26	2.8×10^3
Ne	7.93	[NeII]	12.81	21.56	5.4×10^5
		[NeIII]	15.56	40.96	2.9×10^5
		[NeIII]	36.01	40.96	4.2×10^4
		[NeV]	14.32	97.12	3.8×10^5
		[NeV]	24.32	97.12	5.4×10^4
		[NeVI]	7.65	126.21	2.5×10^5
N	7.83	[NII]	121.9	14.53	3.1×10^2
		[NII]	205.2	14.53	4.8×10^1
		[NIII]	57.32	29.60	3.0×10^3
Mg	7.60	[MgIV]	4.49	80.14	6.3×10^6
		[MgV]	5.61	109.24	2.0×10^6
Si	7.51	[SiII]	34.82	8.15	3.4×10^5
		[SiVII]	6.49	205.05	5.0×10^6
Fe	7.50	[FeI]	24.04	–	3.1×10^6
		[FeI]	34.71	–	3.0×10^6
		[FeII]	25.99	7.90	2.2×10^6
		[FeII]	35.35	7.90	3.3×10^6
S	7.12	[SIII]	18.71	23.34	1.7×10^4
		[SIII]	33.48	23.34	2.0×10^3
		[SIV]	10.51	34.79	5.6×10^4
Ar	6.40	[ArII]	6.99	15.76	1.9×10^5
		[ArIII]	8.99	27.63	3.1×10^5
		[ArIII]	21.83	27.63	3.5×10^4

(Magorrian et al. 1998; Ferrarese & Merritt 2000), all suggest that the two processes are physically linked.

The mid- to far-IR spectral range includes a suite of atomic and molecular lines and features, covering a wide range of excitation, and tracing the physical conditions (excitation, density, ionization, radiation field, metallicity and dust composition) in galaxies (Fig. 2, and Spinoglio & Malkan 1992). Table 1 lists the set of atomic and ionic fine structure lines that will be covered by the SMI and SAFARI spectrometers onboard SPICA, for the top nine most abundant elements in the universe after H and He (the so-called “metals”). These lines reveal the detailed physics in the various phases of the interstellar medium (ISM), from HII and photo-dissociation regions (PDR), to the Narrow Line Regions excited by AGN. In the highly opaque, dust obscured ISM of actively star-forming galaxies and AGN, the IR lines are among the few probes of the physical

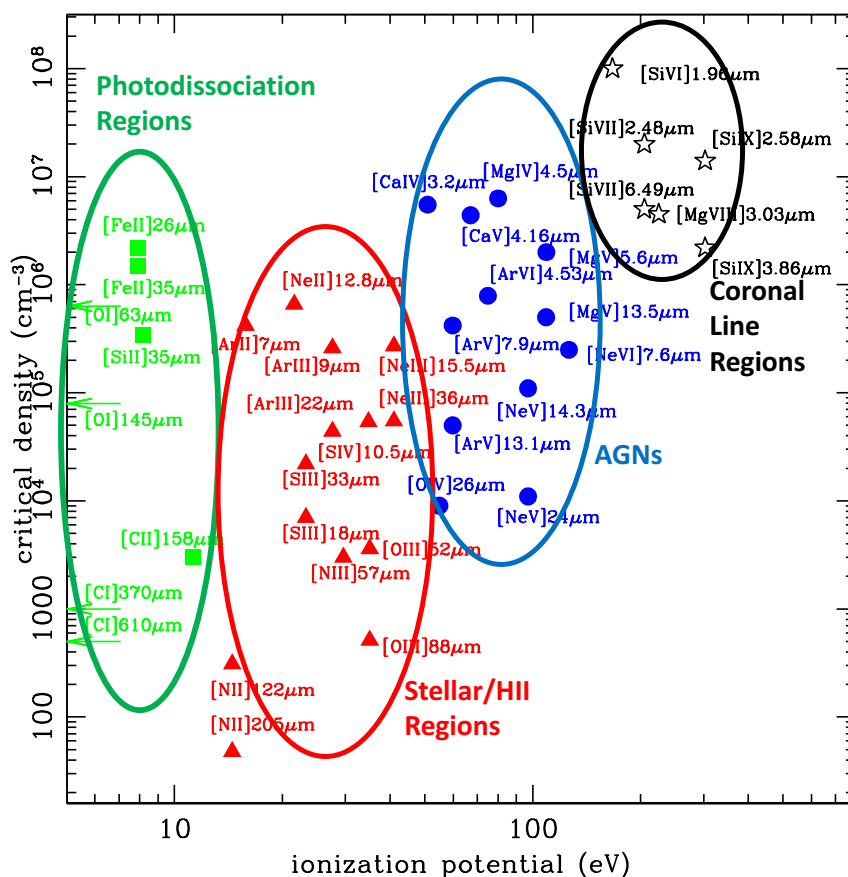


Figure 2. Critical densities for collisional de-excitation versus the ionization potential of the IR fine-structure lines (adapted from Spinoglio & Malkan 1992).

conditions in the gas and dust clouds surrounding the supermassive black holes (SMBH) or young, hot stars.

The neutral gas surrounding star-forming regions can be traced using the temperature sensitive [OI] lines at 63 μm and 145 μm , while the ionised gas can be studied using many different tracers to measure temperature, density and abundances. The [NII]122 μm to [NII]205 μm , [OIII]52 μm to [OIII]88 μm , [SiII]18.7 μm to [SiII]33.5 μm , and [NeIII]15.6 μm to [NeIII]36.0 μm line ratios are individually sensitive to the density (and not to the temperature, since the electron temperatures are generally higher than the excitation states of the lines above).

The strength and hardness of the radiation field can be derived from pairs of the same element in different ionization states and here again the SPICA wavelength range is ideal. Examples are the [NIII]57 μm to [NII]122 or 205 μm ratio, the [NeIII]15.6 μm to [NeII]12.8 μm ratio and the [OIV]26 μm to [OIII]52 or 88 μm ratio. The ratio of the [OIII] to [NIII] lines can be used to derive relative abundances of O and N, two key elements in the gas chemistry (see Fernández-Ontiveros et al. 2017).

Other strong lines, such as [SiII]34.8 μm or [CII]158 μm , probe the interface between the neutral and ionized gas.

The SFR can be obtained via low-ionization lines (e.g. [NeII]12.8 μm , [NeIII]15.5 μm , [SiII]18.7 μm), while the AGN accretion rate can be measured via high-ionization lines (e.g., [OIV]25.9 μm , [NeV]14.3 μm and 24.3 μm). Furthermore, fine-structure lines and measures of small grain dust, through the polycyclic aromatic hydrocarbons (PAH) bands, and the mid-IR thermal dust continuum can be used to disentangle AGN emission from star formation (Genzel et al. 1998; Armus et al. 2006, 2007; Farrah et al. 2007; Spoon et al. 2007; Ho & Keto 2007; Veilleux et al. 2009; Petric et al. 2011; Stierwalt et al. 2014) and AGN (Spinoglio et al. 2005; Lutz et al. 2008; Schweitzer et al. 2008; Tommasin et al. 2008, 2010; Alonso-Herrero et al. 2012; Spinoglio et al. 2015; Fernández-Ontiveros et al. 2016). As an example, we show in Fig. 3 for a sample of local active galaxies, the line ratios of [OIV]25.9 μm /[NeII]12.8 μm versus [NeV]14.3 μm /[NeII]12.8 μm , which both measure the strength of the AGN, while the equivalent width of the PAH emission feature at 11.25 μm , as well as the equivalent width of [NeII]12.8 μm , measure the strength

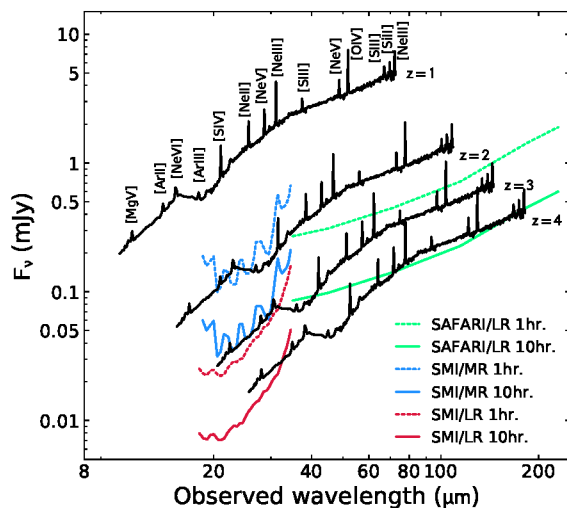


Figure 5. IR spectrum of MCG-3-34-64, a nearby active galaxy, rescaled to a luminosity of $L=10^{12}L_{\odot}$ at redshifts from 1 to 4. At $z=3$, the knee luminosity is $L^*=10^{12}L_{\odot}$, implying that SPICA will be able to map the bulk of the galaxy population up to this redshift. The 5σ sensitivities of SAFARI (in low resolution mode, LR) and SMI (in medium and low resolution, MR and LR) are shown for integration times of 1 hr and 10 hrs.

of the starburst component in galaxies (Tommasin et al. 2010).

Where *Spitzer* and *Herschel* detected only the brightest galaxies at redshifts $z>0.5$ (Sturm et al. 2010; Pope et al. 2008; Menéndez-Delmestre et al. 2009; Riechers et al. 2014; Zhao et al. 2016), SPICA will reveal the properties of the bulk of the IR galaxy population up to redshifts of $z\sim 4$ (Spinoglio et al. 2012). The detected lines and dust features will allow us to unambiguously quantify the contribution of star formation and AGN to the bolometric luminosity in each galaxy. In Fig. 4 we illustrate how specific IR line ratios can discriminate between nuclear activity from AGN, normal star formation and also low metallicity star formation. We notice that orders of magnitude in the line ratios separate the different objects, making this diagram very powerful, i.e. it is the analog of the classical optical BPT diagram (Baldwin et al. 1981), but using extinction-free IR lines.

The SAFARI grating spectrometer at low resolution ($R\sim 300$) will detect [OIV] $26\mu\text{m}$ at $z=1$ in galaxies with $L\sim 10^{11}L_{\odot}$, at $z=2$ in galaxies with $L\sim 3\times 10^{11}L_{\odot}$, and at $z=3$ in galaxies with $L\sim 10^{12}L_{\odot}$ in a few hours. These luminosities correspond to the knee of the luminosity function, L^* (or characteristic luminosity, Schechter 1976), at each redshift. Simultaneously, SAFARI will detect many other diagnostic lines of the ionized and neutral ISM, such as the star formation tracers of [NeII] $12.8\mu\text{m}$, [NeIII] $15.5\mu\text{m}$, [SiII] $18.7\mu\text{m}$

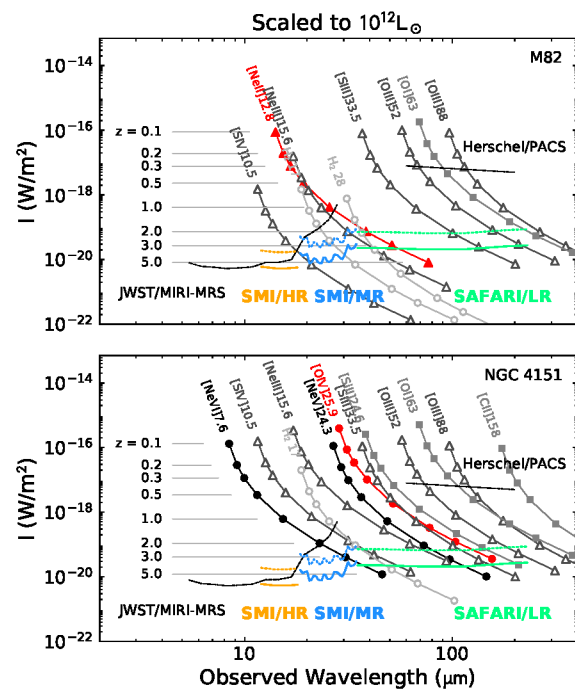


Figure 6. Predicted intensities for the mid- and far-IR fine-structure lines covered by SPICA for the starburst galaxy M82 (upper panel) and the AGN/Seyfert type 1 galaxy NGC 4151 (lower panel), scaled to a luminosity of $10^{12}L_{\odot}$. Filled circles correspond to tracers of AGN activity, triangles to lines typically dominated by star formation, squares to typical PDR lines, and open circles to transitions of warm molecular gas. The predicted line intensities are compared to the 5σ , 1 hour (dotted lines) and 10 hours (solid lines) sensitivities for SMI/HR (in yellow), SMI/MR (in blue), and SAFARI/LR (in green). Additionally, black-dashed lines indicate the sensitivities for JWST/MIRI-MRS (10 σ , 2.8 hours, from Glasse et al. 2015) and *Herschel*/PACS (5 σ , 1 hour). In less than few hours, SPICA will be able to detect the main star-formation and AGN tracers (e.g. [NeII] $12.8\mu\text{m}$ and [OIV] $25.9\mu\text{m}$) at the peak of star formation and SMBH accretion activity ($1 < z < 3$), for starburst galaxies and AGN at the knee of the luminosity function.

and $33.4\mu\text{m}$, and at longer wavelengths the photodissociation region tracers of [OI] $63\mu\text{m}$ and $145\mu\text{m}$ and [CII] $158\mu\text{m}$. Fig. 5 shows the IR spectrum of the local active galaxy MCG-3-34-64, rescaled to a luminosity of $L=10^{12}L_{\odot}$ at redshifts z from 1 to 4, illustrating that the SPICA spectrometers will be able to detect both the continuum and the brightest lines in such a galaxy up to $z\sim 4$ in a few hours. Fig. 6 shows the predicted intensities for the strongest mid- and far-IR lines in the nearby starburst galaxy M82 and in the AGN NGC 4151. For all galaxies, the line intensities have been scaled to a luminosity of $L=10^{12}L_{\odot}$. The lines of [NeII] $12.8\mu\text{m}$, [OIII] $52, 88\mu\text{m}$, [OIV] $25.9\mu\text{m}$, [NeV] $14.3\mu\text{m}$, $24.3\mu\text{m}$

and [NeVI]7.6 μm will be detected in less than a few hours for the bulk of both starburst galaxies and AGN ($L \sim 10^{12} L_{\odot}$) at the peak of star formation and SMBH accretion activity ($1 < z < 3$).

SPICA will also obtain deep 34–230 μm low-resolution spectra of around 1,000 galaxies, equally spaced in luminosity and redshift bins, up to a redshift of $z=3.5$ in a total integration time of around 2,000 hrs. Unbiased, deep low-resolution ($R \sim 50$ –120) SMI spectrophotometric surveys will be able to cover 10 deg^2 in 600 hrs. and detect 30,000 galaxies (Kaneda et al. 2017).

JWST will not cover wavelengths above $\lambda = 28.8 \mu\text{m}$, limiting its capability to measuring spectra of galaxies and AGN above $z \sim 2$ in the rest-frame mid-IR. This can be done with JWST with the [ArII]7 μm and [NeVI]7.6 μm lines, but these will be a factor of 3–4 fainter than the MIR lines accessible to SPICA. As a consequence of this, and incorporating the degradation of JWST MIRI sensitivity toward its long-wavelength band edge, we expect that SPICA will be 25 to 2000 times faster than JWST at diagnosing the power sources in dusty galaxies at $z=2$ and 3, respectively.

ALMA cannot trace far-IR cooling lines (e.g., the [OIII]88 μm) at redshifts lower than 3, and cannot probe the dust and gas spectral features shortwards of 300 μm , leaving for SPICA the unexplored territory which covers the peak of the SFR and BHAR density functions ($1 < z < 3$), at mid- to far-IR wavelengths.

3.1 Measuring the primary ionizing spectrum of AGN and starburst galaxies

The intrinsic primary unattenuated ionizing spectrum of AGN or starburst galaxies cannot be observed from the Lyman limit up to several hundred eV, because of Galactic and intrinsic absorption. Fortunately, for any region for which the ionization structure is dominated by photons, the primary ionizing spectrum can be inferred from the suite of the mid- and far-IR emission lines from the various ions, these traces are largely unaffected by dust. In the nearby Universe, SPICA will detect the full range of mid- to far-IR fine-structure lines, providing a powerful probe of the intrinsic, ionizing spectrum, while at higher redshift, (e.g. $z \sim 1$ –2, depending on the luminosity of the galaxy) SPICA will still be able to detect the brightest lines of, e.g., [NeII]12.8 μm , [NeIII]15.5 μm , [OIII]52 and 88 μm , [OIV]26 μm and [NeV] 14.3 and 24.3 μm , whose ionization potentials are positioned to well sample the ionizing spectrum.

Fig. 7 (upper panel) shows how the observed IR emission lines trace the primary spectrum of the Seyfert galaxy NGC4151, which is not directly observable from 10 to 200 eV, where the signature of the BH accretion disk (big blue bump) should become prominent (Malkan & Sargent 1982). This method has been applied to ISO spectra of bright Seyfert galaxies, e.g. NGC4151 and

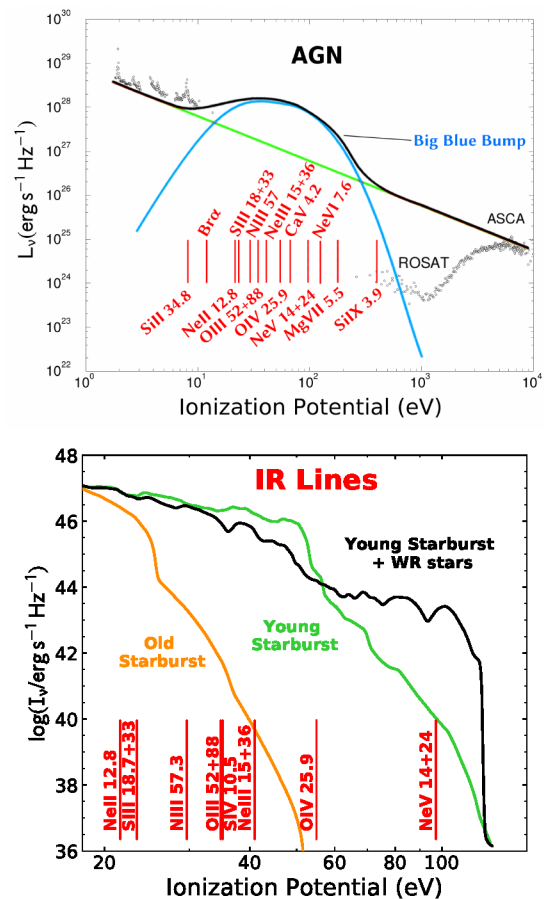


Figure 7. Diagrams showing how IR lines, which can be measured by SPICA up to $z \sim 1$ –2, can map the intrinsic ionizing spectra of both AGN and starburst galaxies. *Upper panel:* Overlay of the NGC4151 primary ionising spectrum (black points) with a schematic “blue bump” and power law (adapted from Alexander et al. 1999). The IR lines will trace the unobserved primary spectrum at the energies of their ionization potentials. *Lower panel:* IR lines probe the starburst’s primary ionizing spectrum, which is sensitive to the presence of Wolf-Rayet stellar winds or shocks and to the starburst age (models from Leitherer et al. 1999).

NGC1068 (Alexander et al. 1999, 2000), and *Spitzer* spectra of local AGN (Meléndez et al. 2011). Fig. 7 (lower panel) shows a similar plot for starburst galaxies, where the stellar spectrum can be traced by the IR lines, free from dust extinction. This often leads to unexpected results. For example, observations of the [OIV]26 μm line have revealed that starburst and dwarf galaxies can have, in particular conditions, a significant contribution of ionizing photons with energies above 50 eV (Lutz et al. 1998; Schaerer & Stasińska 1999; Fernández-Ontiveros et al. 2016), which is difficult to reconcile with current stellar population models (Stasińska et al. 2015). Observations of low-redshift

starburst and dwarf galaxies with SPICA will have the potential to reveal the physical mechanisms (e.g. strong winds or shocks, associated with very high mass stars) responsible for the high ionization tail of the underlying ionizing spectrum, constraining the upper mass cutoff and the stellar initial mass function.

3.2 Finding AGN in dwarf galaxies

It has been recently recognised that finding and characterising AGN in dwarf galaxies is a unique way to infer the properties of high-redshift BH seeds ($z > 5$), i.e. accreting BHs with masses of 10^5 – $10^6 M_\odot$, which are probably associated with low-metallicity environments (Baldassare et al. 2016a). Furthermore, dwarf galaxies are thought to be the local analogues of the high-redshift galaxies responsible for cosmic reionization, due to escaping photons produced by strong star formation (Ly α emitters; Sharma et al. 2016; Izotov et al. 2016; Schaerer et al. 2016) and/or by AGN activity from accreting BH of 10^5 – $10^6 M_\odot$ (Madau & Haardt 2015). The IR line ratio diagram shown in Fig. 3, which will be populated with SPICA spectroscopic observations, is ideal for measuring the AGN contribution in dwarf galaxies; this is because dwarf AGN are elusive in X-ray surveys and many of them are contaminated by strong star formation in classical optical “BPT” (Baldwin et al. 1981) diagnostics (Baldassare et al. 2016b; Simmonds et al. 2016). The [NeII] and [NeIII] lines will be detected at low redshift with SMI at high spectral resolution and at increasing redshift with SMI at medium resolution as well as with SAFARI.

4 Synergies with Future Facilities

SPICA will study the physical processes driving galaxy evolution, through sensitive mid- to far-IR observations of deeply embedded regions that characterize galaxy formation and evolution at the peak of the SFR and BHAR ($1 < z < 4$). JWST, due to its shorter wavelength range, will cover the same redshifts at shorter rest-frame wavelengths (e.g. at $z=3$ for $\lambda < 7 \mu\text{m}$ and at $z=2$ for $\lambda < 9 \mu\text{m}$), missing most of the fine-structure diagnostic lines of Table 1 and having a greatly reduced ability to detect most dust-enshrouded and obscured galaxies and AGN, due to the increasing extinction at shorter wavelengths. SPICA will be able to survey large areas of the sky, and find new samples of IR-bright, high-redshift galaxies. JWST will be able to study in greater spatial detail the low-redshift Universe ($z < 1$) and perform cosmological studies of the high-redshift Universe in the UV/optical rest-frame. In the submillimeter, ALMA and the Northern Extended Millimeter Array (NOEMA)², will map lower excitation emission-

line processes and colder dust continuum, with respect to SPICA. Ground-based interferometry with ALMA and the NOEMA can complement far-IR spectroscopy in tracing molecular outflows with CO (Feruglio et al. 2010, 2015; Ciccone et al. 2012, 2014; García-Burillo et al. 2015) and HCN-HCO⁺ (e.g., Sakamoto et al. 2009; Aalto et al. 2012; Imanishi et al. 2016). SPICA will, however, work in synergy with these facilities and the new generation ELTs, as well as with the X-ray telescope Athena, the *Advanced Telescope for High ENergy Astrophysics* (Nandra et al. 2013). This latter observatory will measure, in synergy with SPICA, the black hole accretion rate (BHAR) history as a function of cosmic time. However, it will be mostly limited to Compton thin sources, needing SPICA observations to access the obscured BHAR density. On AGN feedback, Athena will detect the high energy ionic winds in synergy with SPICA observations of massive molecular outflows.

If approved, SPICA will be operational in the late 2020s in the era of the ELTs and at the dawn of the Square Kilometer Array (SKA; Dewdney et al. 2009). SPICA will operate at the same time as Athena in the X-ray range. By then, both *Euclid* (Maciaszek et al. 2016) and the *Wide Field IR Survey Telescope* (WFIRST; Spergel et al. 2013) will have brought a vast legacy of deep near-IR images and spectra over large fields, reaching the peak of star formation and SMBH accretion activity for $1 < z < 3$, but limited to the unobscured populations of galaxies and AGN. From these large surveys, follow-up spectroscopic observations will be collected by SPICA mid and far-IR low-resolution ($R \sim 50$ – 300) spectroscopy, enabling a full characterization of the physical state of huge samples of galaxies.

5 CONCLUSIONS

The SPICA mission, with its 2.5-m actively-cooled telescope and powerful mid- and far-IR instrumentation, will make a huge step forward in understanding galaxy formation and evolution, through IR spectroscopy of individual galaxies and deep photometric and spectrophotometric surveys for large statistically significant samples. The mid- and far-IR spectral range is host to a powerful suite of diagnostic tools, able to penetrate even the dustiest regions. With SPICA astronomers will use this toolset to probe galaxy evolution in several unique ways, including:

- (i) obtaining the first physical determination through IR spectroscopy of the star-formation rate and of the black hole accretion rate histories across cosmic times up to a redshift of $z \sim 4$;
- (ii) studying AGN accretion and feedback and their impact in the evolution of star formation, through detection and characterisation of far-IR molecular and atomic line profiles;

²<http://www.iram-institute.org/EN/noema-project.php>

- (iii) measuring the rise of metals from the distant Universe to the present day through spectroscopy of IR lines, which minimises the effects of dust extinction and eliminates those due to temperature uncertainties and fluctuations in the ionized gas;
- (iv) studying early AGN and starburst-dominated galaxies in significant samples of dusty galaxies up to $z \sim 6$, through deep (spectro)-photometric surveys;
- (v) detecting and characterizing some of the youngest and most luminous galaxies in the early Universe when it was only half a billion years old.

With a large 2000-hr. spectroscopic survey, SPICA/SAFARI will obtain individual high S/N spectra of over 1000 galaxies to $z \sim 4$. These observations will simultaneously detect diagnostic lines of the ionized and neutral ISM, characterising the local environment and quantifying the contribution from young stars and AGN to the bolometric luminosity. SPICA/SMI will perform similar spectroscopic studies at lower redshifts ($z < 2$) with low-resolution, wide field surveys, to detect and characterize thousands of dusty galaxies via their PAH features and hot dust emission. No other currently planned telescope will be able to perform this type of detailed spectroscopic and wide-area spectro-photometric investigation. SPICA will fill the wide spectral gap left between ALMA (observing above $300 \mu\text{m}$) and JWST (below $28 \mu\text{m}$). In this spectral range, SPICA will observe the ionic fine structure lines, the H_2 as well as many other molecular lines (e.g. from H_2O , OH, and CO) and PAH features in the rest-frame mid-IR in galaxies at the peak epoch of star formation, as well as those from the molecular gas and photodissociation regions in the far-infrared.

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REFERENCES

- Aalto S., Garcia-Burillo S., Muller S., Winters J. M., van der Werf P., Henkel C., Costagliola F., Neri R., 2012, *A&A*, **537**, A44
- Alexander T., Sturm E., Lutz D., Sternberg A., Netzer H., Genzel R., 1999, *ApJ*, **512**, 204
- Alexander T., Lutz D., Sturm E., Genzel R., Sternberg A., Netzer H., 2000, *ApJ*, **536**, 710
- Alonso-Herrero A., Pereira-Santaella M., Rieke G. H., Rigopoulou D., 2012, *ApJ*, **744**, 2
- Armus L., et al., 2006, *ApJ*, **640**, 204
- Armus L., et al., 2007, *ApJ*, **656**, 148
- Baldassare V. F., Reines A. E., Gallo E., Greene J. E., 2016a, preprint, ([arXiv:1609.07148](https://arxiv.org/abs/1609.07148))
- Baldassare V. F., et al., 2016b, *ApJ*, **829**, 57
- Baldry I. K., Glazebrook K., Brinkmann J., Ivezić Ž., Lupton R. H., Nichol R. C., Szalay A. S., 2004, *ApJ*, **600**, 681
- Baldwin J. A., Phillips M. M., Terlevich R., 1981, *PASP*, **93**, 5
- Bernard Salas J., Pottasch S. R., Beintema D. A., Wesselius P. R., 2001, *A&A*, **367**, 949
- Bocchio M., Marassi S., Schneider R., Bianchi S., Limongi M., Chieffi A., 2016, *A&A*, **587**, A157
- Bouwens R. J., Illingworth G. D., Franx M., Ford H., 2007, *ApJ*, **670**, 928
- Ciardi B., Ferrara A., 2001, *MNRAS*, **324**, 648
- Cicone C., Feruglio C., Maiolino R., Fiore F., Piconcelli E., Menci N., Aussel H., Sturm E., 2012, *A&A*, **543**, A99
- Cicone C., et al., 2014, *A&A*, **562**, A21
- Clements D. L., et al., 2016, *MNRAS*, **461**, 1719
- Croxall K. V., et al., 2013, *ApJ*, **777**, 96
- Dannerbauer H., et al., 2014, *A&A*, **570**, A55
- Delvecchio I., et al., 2014, *MNRAS*, **439**, 2736
- Dewdney P. E., Hall P. J., Schilizzi R. T., Lazio T. J. L. W., 2009, *IEEE Proceedings*, **97**, 1482
- Di Matteo T., Springel V., Hernquist L., 2005, *Nature*, **433**, 604
- Egami E., Rieke G. H., Fadda D., Hines D. C., 2006, *ApJ*, **652**, L21
- Egami E., et al., 2017, in prep.
- Eisenstein D. J., et al., 2011, *AJ*, **142**, 72
- Falstad N., et al., 2015, *A&A*, **580**, A52
- Farrah D., et al., 2007, *ApJ*, **667**, 149
- Fernández-Ontiveros J. A., Spinoglio L., Pereira-Santaella M., Malkan M. A., Andreani P., Dasyra K. M., 2016, *ApJS*, **226**, 19
- Fernández-Ontiveros J., et al., 2017, PASA, this issue
- Ferrara A., 1998, *ApJ*, **499**, L17
- Ferrarese L., Merritt D., 2000, *ApJ*, **539**, L9
- Feruglio C., Maiolino R., Piconcelli E., Menci N., Aussel H., Lamastra A., Fiore F., 2010, *A&A*, **518**, L155
- Feruglio C., et al., 2015, *A&A*, **583**, A99
- Fischer J., et al., 2010, *A&A*, **518**, L41
- Franceschini A., Hasinger G., Miyaji T., Malquori D., 1999, *MNRAS*, **310**, L5
- Franceschini A., Rodighiero G., Vaccari M., 2008, *A&A*, **487**, 837
- García-Burillo S., et al., 2015, *A&A*, **580**, A35
- Gardner J. P., et al., 2006, *Space Sci. Rev.*, **123**, 485
- Genzel R., 1992, in Burton W. B., Elmegreen B. G., Genzel R., eds, *Saas-Fee Advanced Course 21: The Galactic Interstellar Medium*. pp 275–391
- Genzel R., et al., 1998, *ApJ*, **498**, 579
- Glasse A., et al., 2015, *PASP*, **127**, 686
- Gong Y., Cooray A., Santos M. G., 2013, *ApJ*, **768**, 130
- González-Alfonso E., et al., 2012, *A&A*, **541**, A4
- González-Alfonso E., et al., 2014, *A&A*, **561**, A27
- González-Alfonso E., et al., 2017a, *ApJ*, **836**, 11
- González-Alfonso E., et al., 2017b, PASA, this issue
- Greenhouse M. A., Feldman U., Smith H. A., Klapisch M., Bhatia A. K., Bar-Shalom A., 1993, *ApJS*, **88**, 23
- Grevesse N., Asplund M., Sauval A. J., Scott P., 2010, *Ap&SS*, **328**, 179
- Gruppioni C., et al., 2013, *MNRAS*, **432**, 23
- Gruppioni C., et al., 2017, PASA, this issue
- Hauser M. G., Dwek E., 2001, *ARA&A*, **39**, 249
- Ho L. C., Keto E., 2007, *ApJ*, **658**, 314
- Imanishi M., Nakanishi K., Izumi T., 2016, *AJ*, **152**, 218
- Izotov Y. I., Schaerer D., Thuan T. X., Worsack G., Gu-

PASA (2017)
doi:10.1017/pas.2017.xxx

- seva N. G., Orlitová I., Verhamme A., 2016, *MNRAS*, **461**, 3683
- Jun H. D., et al., 2015, *ApJ*, **806**, 109
- Kaneda H., et al., 2016, in Society of Photo-Optical Instrumentation Engineers (SPIE) Conference Series. p. 99042I, doi:10.1117/12.2232442
- Kaneda H., et al., 2017, PASA, this issue
- Kemper F., 2015, *Publication of Korean Astronomical Society*, **30**, 283
- Kessler M. F., et al., 1996, *A&A*, **315**, L27
- Lagache G., Puget J.-L., Dole H., 2005, *ARA&A*, **43**, 727
- Lapi A., et al., 2011, *ApJ*, **742**, 24
- Le Floch E., et al., 2005, *ApJ*, **632**, 169
- Leitherer C., et al., 1999, *ApJS*, **123**, 3
- Lilly S. J., Le Fevre O., Hammer F., Crampton D., 1996, *ApJ*, **460**, L1
- Lutz D., 2014, *ARA&A*, **52**, 373
- Lutz D., Kunze D., Spoon H. W. W., Thornley M. D., 1998, *A&A*, **333**, L75
- Lutz D., et al., 2008, *ApJ*, **684**, 853
- Maciaszek T., et al., 2016, in Society of Photo-Optical Instrumentation Engineers (SPIE) Conference Series. p. 99040T, doi:10.1117/12.2232941
- Madau P., Dickinson M., 2014, *ARA&A*, **52**, 415
- Madau P., Haardt F., 2015, *ApJ*, **813**, L8
- Magnelli B., et al., 2013, *A&A*, **553**, A132
- Magorrian J., et al., 1998, *AJ*, **115**, 2285
- Malkan M. A., Sargent W. L. W., 1982, *ApJ*, **254**, 22
- Marassi S., Schneider R., Limongi M., Chieffi A., Bochchio M., Bianchi S., 2015, *MNRAS*, **454**, 4250
- Matsuura M., et al., 2011, *Science*, **333**, 1258
- Meléndez M., Kraemer S. B., Weaver K. A., Mushotzky R. F., 2011, *ApJ*, **738**, 6
- Menéndez-Delmestre K., et al., 2009, *ApJ*, **699**, 667
- Micelotta E. R., Dwek E., Slavin J. D., 2016, *A&A*, **590**, A65
- Mizusawa H., Omukai K., Nishi R., 2005, *PASJ*, **57**, 951
- Murakami H., et al., 2007, *PASJ*, **59**, S369
- Nakagawa T., Shibai H., Onaka T., Matsuhara H., Kaneda H., Kawakatsu Y., Roelfsema P., 2014, in Space Telescopes and Instrumentation 2014: Optical, Infrared, and Millimeter Wave. p. 91431I, doi:10.1117/12.2055947
- Nandra K., et al., 2013, preprint, (arXiv:1306.2307)
- Neugebauer G., et al., 1984, *ApJ*, **278**, L1
- Ogle P., Davies J. E., Appleton P. N., Bertincourt B., Seymour N., Helou G., 2012, *ApJ*, **751**, 13
- Oliver S. J., et al., 2012, *MNRAS*, **424**, 1614
- Pastor C., et al., 2016, in Society of Photo-Optical Instrumentation Engineers (SPIE) Conference Series. p. 99043U, doi:10.1117/12.2232786
- Pereira-Santaella M., Rigopoulou D., Farrah D., Leboutteiller V., Li J., 2017, *MNRAS*, **470**, 1218
- Pérez-González P. G., et al., 2005, *ApJ*, **630**, 82
- Petric A. O., et al., 2011, *ApJ*, **730**, 28
- Pilbratt G. L., et al., 2010, *A&A*, **518**, L1
- Pope A., et al., 2008, *ApJ*, **675**, 1171
- Rémy-Ruyer A., et al., 2014, *A&A*, **563**, A31
- Riechers D. A., et al., 2014, *ApJ*, **786**, 31
- Roelfsema P., et al., 2017, *A&A*, submitted
- Rowan-Robinson M., et al., 2016, *MNRAS*, **461**, 1100
- Sakamoto K., et al., 2009, *ApJ*, **700**, L104
- Sandstrom K. M., et al., 2012, *ApJ*, **744**, 20
- Santini P., et al., 2010, *A&A*, **518**, L154
- Santoro F., Shull J. M., 2006, *ApJ*, **643**, 26
- Schaerer D., Stasińska G., 1999, *A&A*, **345**, L17
- Schaerer D., Izotov Y. I., Verhamme A., Orlitová I., Thuan T. X., Worseck G., Guseva N. G., 2016, *A&A*, **591**, L8
- Schechter P., 1976, *ApJ*, **203**, 297
- Schneider R., Ferrara A., Salvaterra R., 2004, *MNRAS*, **351**, 1379
- Schweitzer M., et al., 2008, *ApJ*, **679**, 101
- Sharma M., Theuns T., Frenk C., Bower R., Crain R., Schaller M., Schaye J., 2016, *MNRAS*, **458**, L94
- Sibthorpe B., Helmich F., Roelfsema P., Kaneda H., Shibai H., 2016, in EAS Publications Series. pp 411–417, doi:10.1051/eas/1575083
- Silk J., Rees M. J., 1998, *A&A*, **331**, L1
- Simmonds C., Bauer F. E., Thuan T. X., Izotov Y. I., Stern D., Harrison F. A., 2016, preprint, (arXiv:1609.07619)
- Spergel D., et al., 2013, preprint, (arXiv:1305.5425)
- Spinoglio L., Malkan M. A., 1992, *ApJ*, **399**, 504
- Spinoglio L., Malkan M. A., Smith H. A., González-Alfonso E., Fischer J., 2005, *ApJ*, **623**, 123
- Spinoglio L., Dasyra K. M., Franceschini A., Gruppioni C., Valiante E., Isaak K., 2012, *ApJ*, **745**, 171
- Spinoglio L., Pereira-Santaella M., Dasyra K. M., Calzoletti L., Malkan M. A., Tommasin S., Busquet G., 2015, *ApJ*, **799**, 21
- Spoon H. W. W., Marshall J. A., Houck J. R., Elitzur M., Hao L., Armus L., Brandl B. R., Charmandaris V., 2007, *ApJ*, **654**, L49
- Springel V., Di Matteo T., Hernquist L., 2005, *MNRAS*, **361**, 776
- Stasińska G., Izotov Y., Morisset C., Guseva N., 2015, *A&A*, **576**, A83
- Stierwalt S., et al., 2014, *ApJ*, **790**, 124
- Strateva I., et al., 2001, *AJ*, **122**, 1861
- Sturm E., et al., 2010, *A&A*, **518**, L36
- Sturm E., et al., 2011, *ApJ*, **733**, L16
- Swinyard B., et al., 2009, *Experimental Astronomy*, **23**, 193
- Tielens A. G. G. M., Hollenbach D., 1985, *ApJ*, **291**,

722

- Tommasin S., Spinoglio L., Malkan M. A., Smith H.,
González-Alfonso E., Charmandaris V., 2008, *ApJ*,
676, 836
- Tommasin S., Spinoglio L., Malkan M. A., Fazio G.,
2010, *ApJ*, 709, 1257
- Valiante R., Schneider R., Bianchi S., Andersen A. C.,
2009, *MNRAS*, 397, 1661
- Vasudevan R. V., Fabian A. C., 2009, *MNRAS*, 392,
1124
- Veilleux S., Rupke D. S. N., Swaters R., 2009, *ApJ*, 700,
L149
- Werner M. W., et al., 2004, *ApJS*, 154, 1
- Wootten A., Thompson A. R., 2009, *IEEE Proceedings*,
97, 1463
- Wright E. L., et al., 2010, *AJ*, 140, 1868
- Wu X.-B., et al., 2015, *Nature*, 518, 512
- Zhao Y., Yan L., Tsai C.-W., 2016, *ApJ*, 824, 146
- Zhukovska S., Henning T., 2013, *A&A*, 555, A99
- van der Tak F., et al., 2017, *PASA*, this issue

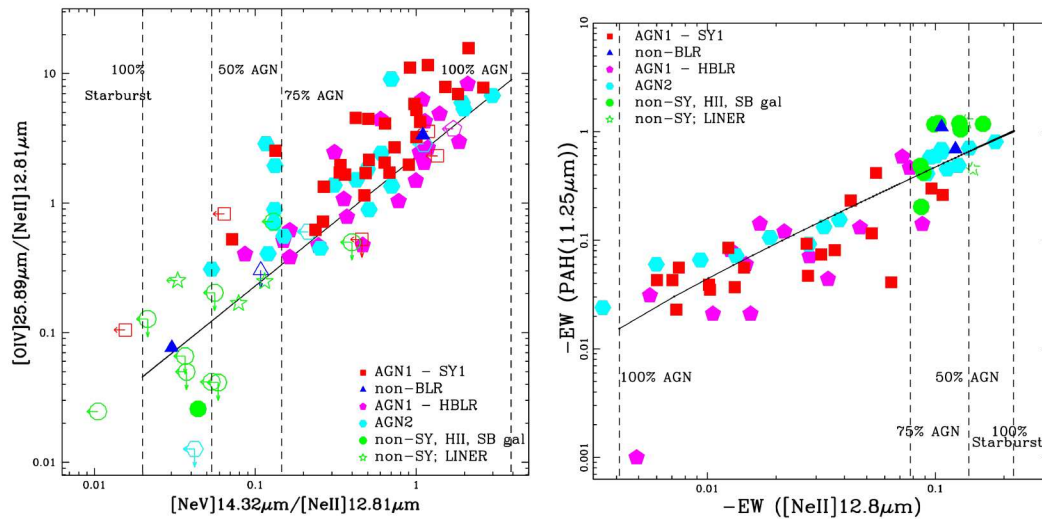


Figure 3. *Left panel:* $[\text{NeV}]14.3\,\mu\text{m}/[\text{NeII}]12.8\,\mu\text{m}$ line ratio versus $[\text{OIV}]25.9\,\mu\text{m}/[\text{NeII}]12.8\,\mu\text{m}$ line ratio. Both axes correlate with the strength of the AGN. The black line shows the behaviour of the analytical model (Tommasin et al. 2010). *Right panel:* $[\text{NeII}]12.8\,\mu\text{m}$ equivalent width versus PAH $11.25\,\mu\text{m}$ equivalent width. Both quantities in this plot correlate with the strength of the star-formation component in each galaxy. The black line shows the behaviour of the analytical model (Tommasin et al. 2010).

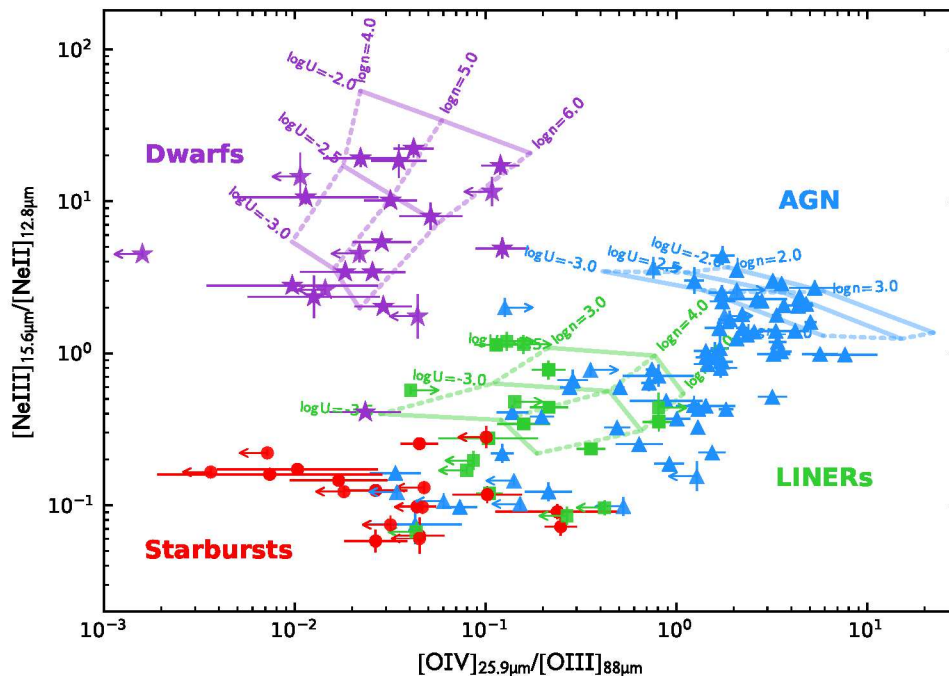


Figure 4. Line ratio diagram of $[\text{NeIII}]15.6\,\mu\text{m}/[\text{NeII}]12.8\,\mu\text{m}$ versus $[\text{OIV}]26\,\mu\text{m}/[\text{OIII}]88\,\mu\text{m}$, showing the comparison of the data for local Universe AGN, LINERs, starburst galaxies, and dwarf galaxies with some models. This diagram is able to disentangle nuclear activity from AGN, from normal star formation, as well as from low metallicity star formation, in e.g. dwarf galaxies (Fernández-Ontiveros et al. 2016). SPICA will be able to measure these line ratios and separate the power sources in active galaxies out to $z\sim 4$.