



Publication Year	2023
Acceptance in OA	2024-12-09T15:40:33Z
Title	Spectroscopy of QUBRICS quasar candidates: 1672 new redshifts and a golden sample for the Sandage test of the redshift drift
Authors	CRISTIANI, Stefano, PORRU, Matteo, GUARNERI, FRANCESCO, CALDERONE, GIORGIO, Boutsia, Konstantina, GRAZIAN, Andrea, CUPANI, Guido, D'ODORICO, Valentina, FONTANOT Fabio, Martins, Carlos J. A. P., Marques, Catarina M. J., MAITRA, Soumak, Trost, Andrea
Publisher's version (DOI)	10.1093/mnras/stad1007
Handle	http://hdl.handle.net/20.500.12386/35429
Journal	MONTHLY NOTICES OF THE ROYAL ASTRONOMICAL SOCIETY
Volume	522

Spectroscopy of QUBRICS quasar candidates: 1672 new redshifts and a golden sample for the Sandage test of the redshift drift

Stefano Cristiani^{1,2,3★}, Matteo Porru,¹ Francesco Guarneri^{1,4,5★}, Giorgio Calderone^{1,6},
Konstantina Boutsia^{1,6}, Andrea Grazian,⁷ Guido Cupani¹, Valentina D’Odorico^{1,2,8},
Fabio Fontanot^{1,2}, Carlos J.A.P. Martins^{9,10}, Catarina M. J. Marques,^{9,10,11} Soumak Maitra^{1,4}
and Andrea Trost^{1,4}

¹INAF–Osservatorio Astronomico di Trieste, Via G.B. Tiepolo, 11, I-34143 Trieste, Italy

²IFPU–Institute for Fundamental Physics of the Universe, via Beirut 2, I-34151 Trieste, Italy

³INFN–National Institute for Nuclear Physics, via Valerio 2, I-34127 Trieste, Italy

⁴Dipartimento di Fisica, Sezione di Astronomia, Università di Trieste, via G.B. Tiepolo 11, I-34131 Trieste, Italy

⁵ESO–European Southern Observatory, Karl-Schwarzschild-Strasse 2, D-85748 Garching bei München, Germany

⁶Las Campanas Observatory, Carnegie Observatories, Colina El Pino, Casilla 601, La Serena, Chile

⁷INAF–Osservatorio Astronomico di Padova, Vicolo dell’Osservatorio 5, I-35122 Padova, Italy

⁸Scuola Normale Superiore, P.zza dei Cavalieri, I-56126 Pisa, Italy

⁹Centro de Astrofísica da Universidade do Porto, Rua das Estrelas, P-4150-762 Porto, Portugal

¹⁰Instituto de Astrofísica e Ciências do Espaço, CAUP, Rua das Estrelas, P-4150-762 Porto, Portugal

¹¹Faculdade de Ciências, Universidade do Porto, Rua do Campo Alegre, P-4150-007 Porto, Portugal

Accepted 2023 March 30. Received 2023 March 10; in original form 2023 January 27

ABSTRACT

The QUBRICS (QUasars as BRiGht beacons for Cosmology in the Southern hemisphere) survey aims at constructing a sample of the brightest quasars with $z \gtrsim 2.5$, observable with facilities in the Southern Hemisphere. QUBRICS makes use of the available optical and IR wide-field surveys in the South and of Machine Learning techniques to produce thousands of bright quasar candidates of which only a few hundred have been confirmed with follow-up spectroscopy. Taking advantage of the recent *Gaia* Data Release 3, which contains 220 million low-resolution spectra, and of a newly developed spectral energy distribution fitting technique, designed to combine the photometric information with the *Gaia* spectroscopy, it has been possible to measure 1672 new secure redshifts of QUBRICS candidates, with a typical uncertainty of $\sigma_z = 0.02$. This significant progress of QUBRICS brings it closer to (one of) its primary goals: providing a sample of bright quasars at redshift $2.5 < z < 5$ to perform the Sandage test of the cosmological redshift drift. A Golden Sample of seven quasars is presented that makes it possible to carry out this experiment in about 1500 h of observation in 25 yr, using the ANDES spectrograph at the 39m ELT, a significant improvement with respect to previous estimates.

Key words: methods: data analysis – methods: statistical – astronomical data bases: miscellaneous – surveys – quasars: general.

1 INTRODUCTION

Quasars (QSOs), as the brightest non-transient sources, can be observed at very high redshifts and shed light on fundamental topics such as the formation and evolution of galactic structures and massive black holes, the Big Bang nucleosynthesis, Cosmology, reionizations, and the variation of the fundamental constants. As cosmic lighthouses, they provide a unique view of the Universe through the observation of absorption features, and the brightest QSOs are coveted as precious tools of investigation.

The QUBRICS (QUasars as BRiGht beacons for Cosmology in the Southern hemisphere) survey (Calderone et al. 2019) has been conceived with the aim of making up for the scarcity of bright QSOs in the Southern Hemisphere, which is due to the historical paucity of

all-sky surveys in the South, and has produced several hundreds new spectroscopically confirmed bright QSOs. Various methods for the selection of QSOs have been used: in Calderone et al. (2019, hereafter Paper I) candidates have been selected using a canonical correlation analysis (CCA; Anderson 2003), in Guarneri et al. (2021, hereafter Paper III) the Probabilistic Random Forest (PRF; Reis, Baron & Shahaf 2019) has been adopted, with modifications introduced to properly treat upper limits and missing data. In Guarneri et al. (2022, Paper VI) the PRF selection has been further improved, in particular adding synthetic data to the training sets. Calderone et al. (in preparation, Paper VII) have developed a method that takes advantage of the extreme gradient boosting technique (Chen & Guestrin 2016; XGB) to significantly improve the recall¹ of the

* E-mail: stefano.cristiani@inaf.it (SC); francesco.guarneri@inaf.it (FG)

¹Recall is defined as the fraction of relevant instances that are retrieved by the selection algorithm.

selection algorithms even in the presence of severely imbalanced data sets, with the aim of completing the QUBRICS survey up to $z \sim 5$.

While refining the methods of selection, a continuous effort has been dedicated in QUBRICS to the follow-up spectroscopy (Boutsia et al. 2020, hereafter [Paper II](#)), testing the selection procedures, and leading to statistically well-defined subsamples that allowed us to address the topics of the QSO luminosity function (LF) and cosmic re-ionization(s) (Boutsia et al. (2021), hereafter [Paper IV](#); Grazian et al. (2022), Paper V; and Fontanot et al. (2023), Paper IX). Rare objects, such as extreme broad absorption line QSOs (BALs), discovered in the course of QUBRICS, have been described in Cupani et al. (2022, Paper VIII). One of the main goals of QUBRICS is to provide a sample of bright targets for the Sandage test of the cosmological redshift drift (Sandage 1962). The redshift drift ($\dot{z} = dz/dt_{\text{obs}}$) is a small, dynamical change in the redshift of objects following the Hubble flow. Measuring it provides a direct, real-time, and model-independent mapping of the expansion rate of the Universe. It is fundamentally different from other cosmological probes: instead of mapping our (present-day) past light cone, it directly compares different past light cones. Being independent of any assumptions on gravity, geometry, or clustering, it directly tests the pillars of the Λ CDM paradigm. Recent theoretical studies have uncovered unique synergies with other cosmological probes, in particular for the characterization of the physical properties of dark energy (Martins et al. 2016; Alves et al. 2019; Esteves et al. 2021). This measurement is a flagship objective of the Extremely Large Telescope (ELT) (Liske et al. 2008), specifically observing the Lyman forest of QSOs with its high-resolution spectrograph, ANDES (Marconi et al. 2022). The effect is tiny, expected to be of the order of $\text{cm s}^{-1}\text{yr}^{-1}$ at the redshifts of interest, and to carry out the measurement a high signal-to-noise ratio (SNR) is a necessary condition, implying, even with bright QSOs and large telescopes, a huge investment of observing time. Having the brightest possible cosmic beacons is a key ingredient to make it possible to perform a measurement which is presently at the edge of feasibility. Since the observability of the Lyman forest with ground-based spectrographs requires a substantial redshift, the QUBRICS has been focused on QSOs at $z \geq 2.5$.

A strategic feature of QUBRICS is the continuous updating, after each observation cycle, of the training set, paying attention to identify and correct the non-insignificant fraction of erroneous spectroscopic identifications found in the literature, which may affect the training of machine learning (ML) techniques. This allows us to improve the success rate and the completeness, while keeping the list of candidates manageable.

In 2022 June, the list of candidates derived from Papers III, VI, and VII still lacking a spectroscopic confirmation included 5469 targets. On 2022 June 13 the Gaia Data Release 3 (DR3) was published, providing, among a wealth of data, low-resolution spectra for about 220 million objects, selected to have a reasonable number of *Gaia* observations and to be sufficiently bright to ensure good SNR (De Angeli et al. 2022). In this paper, we make use of the DR3 spectra and of the QUBRICS photometric data base (Section 3.1), combined with a spectral energy distribution (SED) fitting technique, to obtain secure spectroscopic identifications and redshifts of a significant fraction of the 5469 candidates of the QUBRICS survey.

Unless stated otherwise, magnitudes are given in the AB magnitude system; uncertainties represent 68 per cent confidence intervals. We adopt a flat Λ CDM cosmology, with $\Omega_m = 0.31$, $\Omega_\Lambda = 0.69$, and $H_0 = 67.7 \text{ km s}^{-1} \text{ Mpc}^{-1}$, in agreement with the Planck Collaboration VI (2020).

Table 1. Breakdown of sources with secure spectroscopic redshifts in the QUBRICS data base. Spectroscopically confirmed stars are few compared to QSOs or galaxies: stars are identified mostly using *Gaia* proper motion and parallax criteria, and are not followed up spectroscopically. The 31 observed stars are misclassified as QSOs by our selection algorithms.

Object type	Number of redshifts	Observed by QUBRICS
QSO (Type 1)	781 298	856
QSO (Type 2)	6074	–
BLLac	809	–
Galaxy	33 614	23
Star	31	31

2 MATCHING THE QUBRICS DATA BASE TO THE GAIA DR3 LOW-RESOLUTION SPECTRA

The QUBRICS data base is made of a collection of spectroscopic and photometric data from the literature, and all spectroscopic follow-ups carried out in the framework of the QUBRICS survey. The photometric data base includes optical and infrared data from several public catalogues:

- (i) u, v, g, r, i, z magnitudes from the SkyMapper DR1 survey (Wolf et al. 2018)
- (ii) $G, G_{\text{BP}}, G_{\text{RP}}$ magnitudes from the Gaia eDR3 catalogue (Gaia Collaboration 2021);
- (iii) J, H, K from 2MASS (Skrutskie et al. 2006);
- (iv) W1, W2, W3, W4 from the AllWise survey (Wright et al. 2010);
- (v) g, r, i, z, Y magnitudes from the PanSTARRS1 DR2 survey (Chambers et al. 2016);
- (vi) g, r, i, z, Y magnitudes from the DES survey (Sevilla-Noarbe et al. 2021);

From these data sets we extract convenient subsets on which the selection is performed, and we refer to them as Main Samples (MS). MS are created based on magnitude and data-quality cuts, and positional match distance. For instance, the MS used in Calderone et al. (2019) includes sources with (i) a magnitude measurement in six bands (SkyMapper i, z , *WISE* W1, W2, W3, and *Gaia* G); (ii) a magnitude i limited between $14 < i < 18$; (iii) $\text{SNR} > 3$ in the AllWise bands; (iv) matching distance between *Gaia* and SkyMapper, and AllWise and SkyMapper lower than 0.5 arcmin.

The Gaia catalogue additionally provides parallax and proper-motion measurements. This information is used to identify stars, which are assumed to be objects with parallax and proper motion significantly different from zero ($> 3\sigma$).

Spectroscopic redshifts and classifications are collected from several catalogues and added to the data base:

- (i) the Sloan Digital Sky Survey (SDSS) DR16q (Lyke et al. 2020);
- (ii) the Veron-Cetty catalogue (Véron-Cetty & Véron 2010), including some of the correction by Flesch (2013);
- (iii) the 2dF (Colless et al. 2001);
- (iv) the 6dF (Jones et al. 2009);
- (v) Onken et al. (2022);
- (vi) Schindler et al. (2019a);
- (vii) Schindler et al. (2019b);
- (viii) Wolf et al. (2020);
- (ix) Yang et al. (2016).

Finally, QUBRICS identifications are added to the data base after every observing run. Combining the contribution of all these

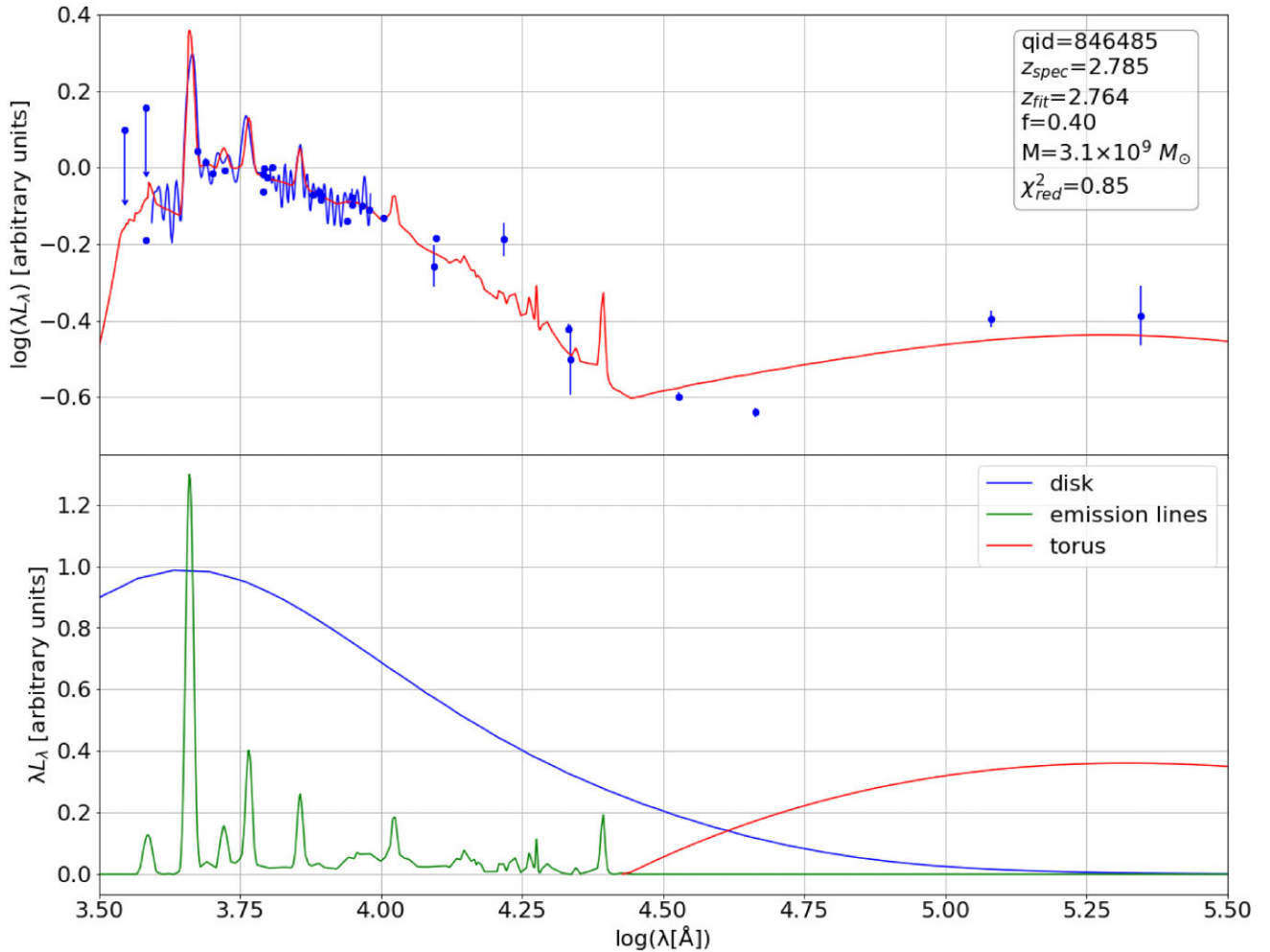


Figure 1. An example of SED fit (red line, upper panel) for a QSO with spectroscopic redshift $z_{\text{spec}} = 2.785$. In the upper panel the blue points show the photometric data, while the blue line is the *Gaia* low-resolution spectrum. The parametrization of the SED is described in the text and the resulting optimal parameters, f (adimensional) and M (in solar masses), are shown in upper right part of the figure. ‘qid’, here and in the following, is the unique identifier of each source in the QUBRICS data base.

catalogues provides us, as of the end of 2022 June, 821 992 objects with a secure redshift estimate and classification (Table 1).

These identifications are used to train the selection algorithms (CCA, PRF, XGB), which are then applied on unclassified sources to find QSO candidates. Combining the lists produced by the three algorithms leaves 5469 candidates still lacking a spectroscopic confirmation. These candidates are scattered over a large area of the sky and observing all of them is a daunting task.

The publication of *Gaia* low-resolution spectra gave us the chance to significantly speed up this process. We cross-matched the candidates list to the *Gaia* DR3 source table, using a 0.75 arcmin matching radius: 2635 of the 5469 candidates turned out to have a low-resolution spectrum available from the *Gaia* archive.

The *Gaia* spectra have been processed using GaiaXPY²: each spectrum was calibrated and sampled on a wavelength grid using the `calibrate` routine; we chose a fixed step of 20 Å and discarded the regions with wavelength < 3900 and > 9600 Å, as these are typically very noisy.

The goal of the next section is to use these data to obtain a secure redshift for a large fraction of the 2635 QUBRICS candidates with a *Gaia* spectrum.

3 OBTAINING SECURE SPECTROSCOPIC IDENTIFICATIONS

Having a set of 2635 objects confidently classified as $z \gtrsim 2.5$ QSOs (Guarneri et al. 2021, 2022), we derive, for each of them, the redshift z , comparing a robust model of the SED to their spectro + photometric data in the following way (described in detail in Sections 3.1 and 3.2.):

- (i) for each object an approximate redshift estimate is obtained by χ^2 -fitting a QSO SED to the photometric information available in QUBRICS complemented with the *Gaia* low-resolution spectroscopy;
- (ii) the redshift estimate is refined by processing the GAIA DR3 low-resolution spectrum with the MARZ package (Hinton et al. 2016), which uses a cross-correlation algorithm to match the spectrum to a set of templates (in our case a QSO template).

²<https://gaia-dpci.github.io/GaiaXPY-website/>, 10.5281/zenodo.6637762

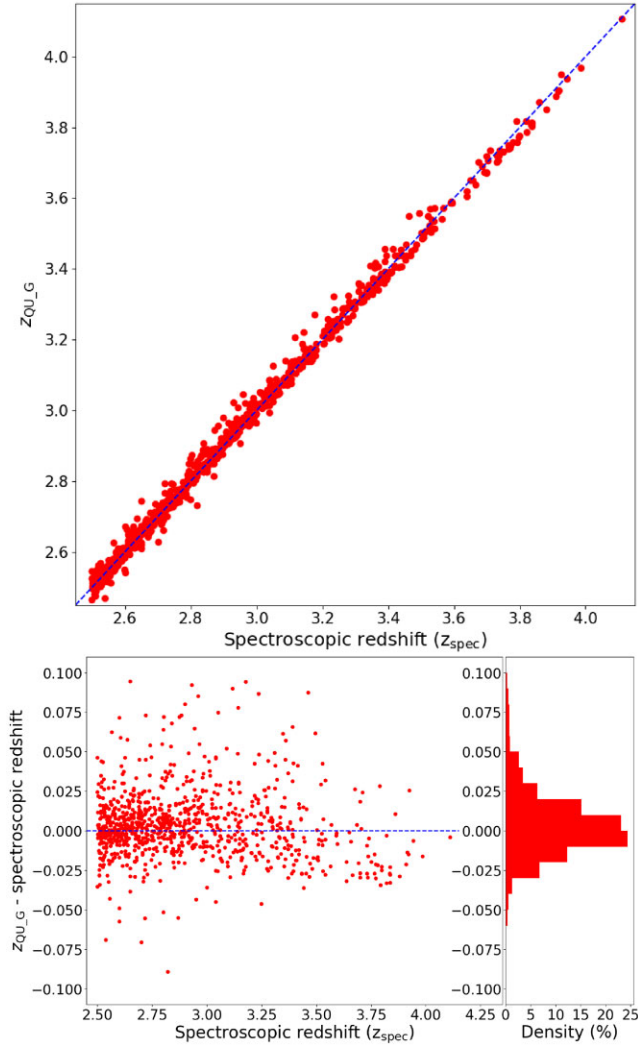


Figure 2. Top panel: the redshifts determined on the basis the *Gaia* low-resolution spectra with the procedure described in Section 3 versus the spectroscopic redshifts for 938 known QSOs with $z > 2.5$ in the QUBRICS data base. Bottom panel: the difference between the redshifts determined on the basis the *Gaia* low-resolution spectra and the spectroscopic redshifts, Δz , as a function of the spectroscopic redshift, z_{spec} .

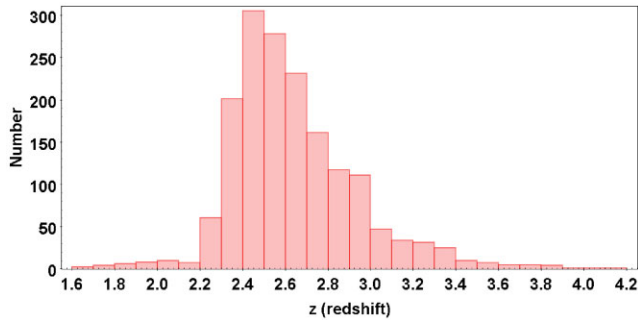


Figure 3. Histogram of the 1672 new redshifts determined with *Gaia* spectroscopy.

Consequently, three different types of redshift are used in this paper:

- (i) z_{spec} , the spectroscopic redshift obtained with follow-up spectroscopy, as, for example, described in Section 5;
- (ii) $z_{\text{QU,G}}$, the redshifts derived from low-resolution *Gaia* spectra (De Angeli et al. 2022) combined with photometric data, as described in this section;
- (iii) z_{Gaia} , redshifts estimated by the *Gaia* Collaboration (De Angeli et al. 2022), as reported in Section 6.

3.1 SED fitting to estimate the redshift of a QSO

For each object we compare the observed spectro + photometric data with synthetic SEDs of QSOs.

We use the same photometric catalogues listed in Section 2. In addition, when available, optical data from the SDSS DR16q survey (Lyke et al. 2020), J , H , K magnitudes from the VHS survey (McMahon et al. 2013) and NUV magnitude from the *GALEX* survey (Morrissey et al. 2007) are included. Finally, we add *Gaia* low-resolution spectra, rebinned with a step of 20 Å in wavelength.

The SED of quasars is parametrized in the following way:

(i) a ‘blue bump’ component, $F_{\text{BB}}(M, \dot{m})$, representing the accretion disc (Laor & Netzer 1989; Sun & Malkan 1989), modelled with the PYAGN package (Kubota & Done 2018). The shape of this spectral component depends on the mass M of the supermassive black hole (SMBH) and on its accretion rate $\dot{m}$: in particular, a larger M tends to move the blue bump to longer wavelengths, while a higher $\dot{m}$ shifts it to shorter wavelengths;

(ii) an ‘IR bump’ component, F_{IR} , representing a dusty torus (Pier & Krolik 1993; Mor, Netzer & Elitzur 2009), starting from a rest wavelength of 8000 Å, that we derived from the QSO SED by Richards et al. (2006);

(iii) a component representing the emission features, F_{em} , obtained with the QSFIT package (Calderone et al. 2017) using the line fluxes and equivalent widths from Vanden Berk et al. (2001) to normalize the intensity of each emission line with respect to the other lines and to the blue bump component. It should be noted that in this way the EW of the emission lines is fixed to the Vanden Berk et al. (2001) template. For QSOs with more exotic spectra (e.g. BALs) the procedure might result less reliable. This is further discussed in Section 3.2.

Different SEDs can be generated by choosing a different SMBH mass (M), accretion rate ($\dot{m}$) and fractional contribution (f) of the ‘IR bump’ with respect to the ‘blue bump’ and emission features components, according to the following formula:

$$\text{SED} = (1 - f)(F_{\text{BB}}(M, \dot{m}) + F_{\text{em}}) + f \cdot F_{\text{IR}} \quad (1)$$

Finally, according to the given redshift, the mean absorption of the intergalactic medium (IGM) is applied shortwards of 1210 Å (Inoue et al. 2014), and the SED is shifted to the observer’s frame.

F_{BB} , F_{em} , and F_{IR} , dimensionally, are all fluxes. The resulting SED is renormalized and compared χ^2 -wise with the available spectro + photometric data. In this way a χ^2 is obtained as a function of the parameters z , M , $\dot{m}$, and f . An example of a fit of a synthetic SED to the spectro + photometric data of a redshift $z_{\text{spec}} = 2.785$ QSO is shown in Fig. 1.

To increase the robustness of the fitting procedure, we explored whether the number of fitting parameters in equation (1) could be reduced. Since the QUBRICS survey is focused on finding QSOs at $z > 2.5$, we have extracted from the QUBRICS data base the 938

Table 2. QUBRICS candidates with a new reliable spectroscopic identification derived from *Gaia* low-resolution spectra. When available, the redshift obtained with follow-up spectroscopy is reported, together with the date and the used spectrograph. The i magnitude is in the AB photometric system. Full table available on the online.

qid	RA J2000	DEC J2000	i_{psf}	$z_{\text{QU.G}}$	Class	z_{spec}	Obs. date	Instrument
888971	00:01:23.91	−58:57:22.2	17.982	3.255	QSO	—	—	—
889687	00:02:09.76	−60:59:07.2	17.064	2.664	QSO	—	—	—
7896285	00:03:43.91	−18:54:25.0	18.510	2.603	QSO	—	—	—
1039775	00:04:35.62	−25:07:07.6	17.248	1.898	QSO	—	—	—
33517238	00:04:49.35	−36:29:39.6	18.973	2.864	QSO	—	—	—
1169052	00:04:57.90	−52:45:30.4	18.297	3.143	QSO	—	—	—
7930596	00:04:59.18	−16:46:42.1	18.509	2.562	QSO	—	—	—
863587	00:05:02.01	−59:25:13.1	17.858	2.421	QSO	—	—	—
1172305	00:05:29.01	−46:55:21.7	18.405	2.579	QSO	—	—	—
1163878	00:06:28.87	−48:31:32.0	18.369	3.602	QSO	—	—	—

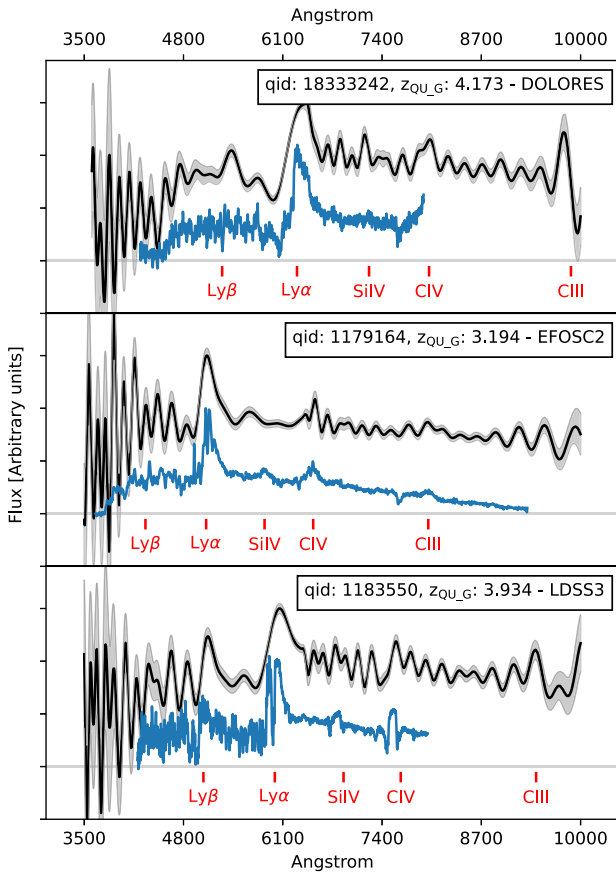


Figure 4. Example of discovery spectra for three of the QSO in the published catalogue. Spectra were taken with the DOLORES, EFOSC2 and LDSS3 spectrographs. In each panel, the black line shows the low-resolution *Gaia* spectrum, the blue line the spectrum taken at the telescope; prominent emission lines are marked in red.

spectroscopically confirmed QSOs with $z_{\text{spec}} > 2.5$ and a *Gaia* low-resolution spectrum. We fitted their spectro + photometric data fixing the redshift at the spectroscopic value and letting the parameters f , M , and $\dot{m}$ vary. Most QSO SEDs are fitted with an $\dot{m} \sim 1$ (and a non-negligible degeneracy exists between M and $\dot{m}$). Therefore, we decided to fix $\dot{m} = 1$ (i.e. Eddington accretion) in the following analysis and fit for each object three parameters: f , M , and the redshift z_{fit} .

It should be noted that in the present context the aim of the spectro + photometric data-fitting is to provide a robust estimation of the redshift for sources classified as QSOs by our ML techniques. The focus is not on measuring other properties as the SMBH mass or accretion rate. The present parametrization of the QSO SED is also different with respect to Guarneri et al. (2022) and has been chosen because it produces better and more physically motivated results.

3.2 Refining the redshift estimate by cross-correlation with the MARZ package

The redshift estimate derived from the SED fitting described in Section 3.1 is finally passed to the MARZ package (Hinton et al. 2016), that matches the *Gaia* low-resolution spectrum of the object with the MARZ QSO template (see Hinton et al. (2016) for the details of the MARZ matching technique) and produces our final estimate for the redshift, $z_{\text{QU.G}}$. A quality operator (QOP; Hinton et al. 2016) is also assigned to each spectrum in a human-supervised way. The QOP scale varies from a value of 1 for inconclusive spectra to 4 for ‘great’ spectra, with an absolutely certain redshift.

Fig. 2 shows the comparison between the spectroscopic redshifts, z_{spec} , present in the QUBRICS data base and the $z_{\text{QU.G}}$, as well as their difference, Δz , as a function of the spectroscopic redshift. The distribution is characterized by a median $\langle \Delta z \rangle = 0.001$ and a standard deviation $\sigma_{\Delta z} = 0.02$. The combination of the photometric and spectrophotometric data and of the χ^2 fitting and template matching techniques turns out to be effective in avoiding the misidentification of spectral features, while providing relatively precise redshifts. In particular, potential errors due to QSOs with exotic spectra are checked and corrected/eliminated with the cross-correlation MARZ procedure and the reliability, tested on the sample of 938 QSOs with known z_{spec} (including a significant fraction of BALs), turns out to be reassuring (Fig. 2).

4 APPLYING THE REDSHIFT MEASUREMENT TO QUBRICS CANDIDATES

The procedure described in Section 3 has been applied to 2635 quasar candidates produced in Papers III, VI, and VII of the QUBRICS survey with a *Gaia* low-resolution spectrum. The properties of these 2635 QSO candidates are not significantly different, either in photometric coverage or in magnitude/photometric uncertainty, from the 938 spectroscopically confirmed QSOs in the QUBRICS data base with $z_{\text{spec}} > 2.5$, described in Section 3. Therefore, we

Table 3. Setup for each observing run.

# of objects	Instrument	Telescope	Grism	Slit	Resolution
13	DOLORES	TNG	LR-B	1 arcsec	600
9	LDSS-3	Clay	VPH-all	1 arcsec	800
3	EFOSC-2	NTT	Grism#13	1.5 arcsec	1000

expect to obtain the same reliability (a test on a limited number of 25 candidates is carried out in Section 5).

For 1672 objects the procedure produces a redshift estimate of sufficient quality ($\text{QOP} \geq 2$) to be considered secure. They are listed in Table For 963 objects the absence of spectral features with sufficient SNR prevented the determination of a redshift with the requested confidence.

Fig. 3 shows the redshift distribution of the 1672 new identifications. It corresponds to the expected distribution for the QUBRICS candidates (see for example fig. 7 in Guarneri et al. 2021).

5 TESTING THE NEW REDSHIFT MEASUREMENTS WITH FOLLOW-UP SPECTROSCOPY

After 2022 June, 25 candidates contained in Table 2 (typically chosen on the basis of sky accessibility) have been observed with follow-up spectroscopy at the Telescopio Nazionale Galileo (TNG, La Palma), Las Campanas Observatory, and ESO-NTT using the DOLORES, LDSS-3 (Clay Telescope), and EFOSC-2 spectrographs, respectively.

Table 3 summarizes the observing setups and significant information about each observing run. Fig. 4 shows three spectra representative of the three telescopes.

For the 13 candidates observed with the DOLORES instrument, mounted on the Telescopio Nazionale Galileo, exposures have been taken during the AOT45 period, in 2022 August, under a proposal with PI F. Guarneri. The LR-B grism (covering a wavelength range between 3600 and 8000 Å at a resolution ~ 600) with a 1 arcmin slit aperture was used with an exposure time between 300 and 600 s.

For the nine targets observed with LDSS-3 at the Clay Telescope, in 2022 July and August, observations were obtained on several nights with varied conditions (e.g. bright time, variable weather conditions). The VPH-all grism with the 1 arcmin central slit and no blocking filter was employed, covering a wavelength range between 4000 and 10000 Å with a low resolution of $R \sim 800$. Exposure times ranging between 800 and 1800 s were used, depending on the candidate magnitude and seeing conditions.

Finally, three candidates were observed in 2022 November, as a part of an observing program at the ESO NTT (PI. F. Guarneri, proposal 110.23WP.001), employing the EFOSC-2 instrument and Grism #13 (wavelength range $\lambda \sim 3700\text{--}9300$ Å and resolution ~ 1000), with typical exposure times ranging between 500 and 1200 s.

Data obtained with the DOLORES, LDSS-3, and EFOSC-2 instruments were reduced with a custom pipeline based on MIDAS scripts (Banse et al. 1988). Each spectrum has been processed to subtract the bias and normalized by the flat; wavelength calibration is achieved using helium, neon, and argon lamps, finding a rms of ~ 0.5 Å. Observing conditions have not always been photometric.

All the 25 objects turned out to be QSOs and Fig. 5 shows the good agreement of the spectroscopic redshifts with the redshifts obtained with the procedure described in Section 3: the distribution is characterized by a median $\langle \Delta z \rangle = -0.0004$ and a standard deviation $\sigma_{\Delta z} = 0.02$, thus confirming the Section 3.2 estimates.

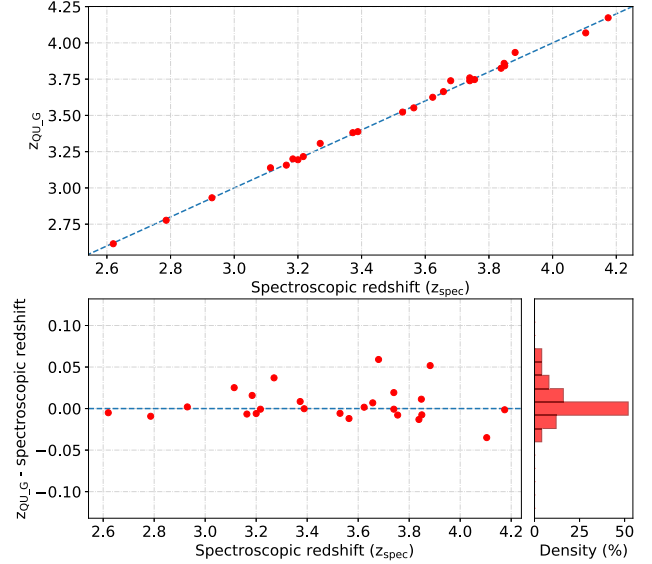


Figure 5. Top panel: the redshifts determined on the basis the *Gaia* low-resolution spectra with the procedure described in Section 3 versus the spectroscopic redshifts for 25 objects observed spectroscopically. Bottom panel: the difference between the redshifts determined on the basis the *Gaia* low-resolution spectra and the spectroscopic redshifts as a function of the spectroscopic redshift.

6 DISCUSSION

The tests described in Section 3.2 and the spectroscopic follow-up observations described in Section 5 demonstrate that the procedure described in this paper produces secure redshifts from the *Gaia* low-resolution spectra, with an uncertainty of $\sigma_z \sim 0.02$.

Thus, we can confidently add 1672 new spectroscopic redshifts to the QUBRICS data base, reducing to 3797 the number of remaining candidates still to be observed. The large majority (~ 80 per cent) of the 3797 remaining candidates have magnitudes fainter than $i = 18$ and expected redshift below $z = 3$. We will continue to follow them up with spectroscopic observations, prioritizing the construction of complete samples at relatively bright magnitudes ($z > 2.5$ and $i < 18.0$, ~ 500 candidates) and at higher redshift ($z > 3$ and $Y < 18.5$, ~ 180 candidates).

A comparison of the redshifts determined in this work with the redshifts estimated from *Gaia* (Gaia Collaboration 2021) shows a generally good correspondence: out of the 1672 redshift determinations of this work, 1668 also have a redshift estimated by the Gaia collaboration (Gaia Collaboration 2021) and in 1663 cases the agreement is within a $|\Delta z| \leq 0.06$ with a $\sigma_{\Delta z} = 0.009$. However, in five cases a catastrophic disagreement occurs. They are shown in Fig. 6. We do not have follow-up spectroscopy for these objects, so only the *Gaia* low-resolution spectra are shown. It is remarkable that the five objects are relatively bright, *Gaia* $G \leq 17.63$, and according to our analysis, based on the typical EW of the emission lines and the continuum slope, in these five cases the *Gaia* pipeline misclassifies some emission line, most notably the Lyman- α or C IV lines as a C III₁₉₀₉, as shown in Fig. 6.

Fig. 7 shows the new diagram with the known quasars at $z > 2.5$ in the Southern hemisphere versus their i magnitude. The contribution of the QUBRICS survey, especially at the brighter magnitudes, is apparent.

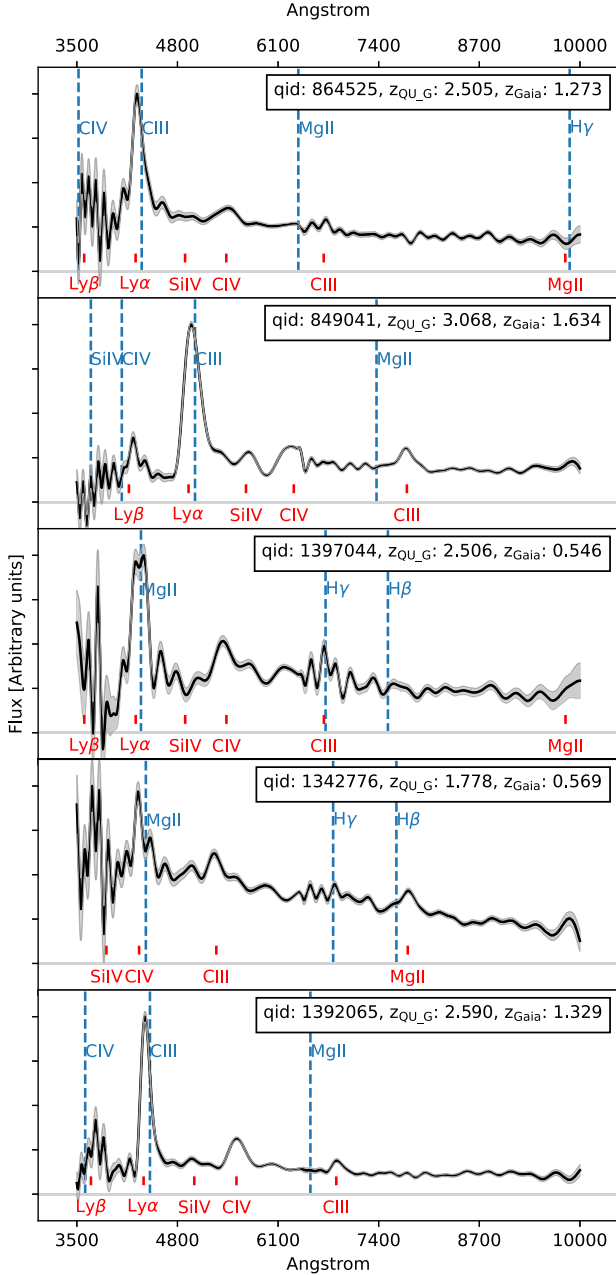


Figure 6. The five catastrophic disagreements between QUBRICS’ and *Gaia*’s redshifts. The black line shows the *Gaia* low-resolution spectra with associated error (grey shaded area). The red marks indicate the position of prominent emission lines at the QUBRICS’ redshift $z_{\text{QU,G}}$, blue dashed lines at the redshift estimated by the *Gaia* pipeline (z_{Gaia}).

6.1 The Sandage test of the redshift drift

In [Paper II](#), we showed that, thanks to the new bright QSO identifications at high redshift, the total time required to carry out the Sandage test of the cosmological redshift drift (Sandage 1962) could be reduced to less than 2500 h of observations with an ANDES-like (Marconi et al. 2022) spectrograph at the ELT. In particular we envisaged the use of the fibre-fed VIS-BLUE (UBV) and VISRED (RIZ) modules of ANDES to obtain in natural seeing conditions high-resolution ($R \sim 100\,000$), high-fidelity spectra of 30 targets, with a temporal separation of 25 yr. We are now in the position to

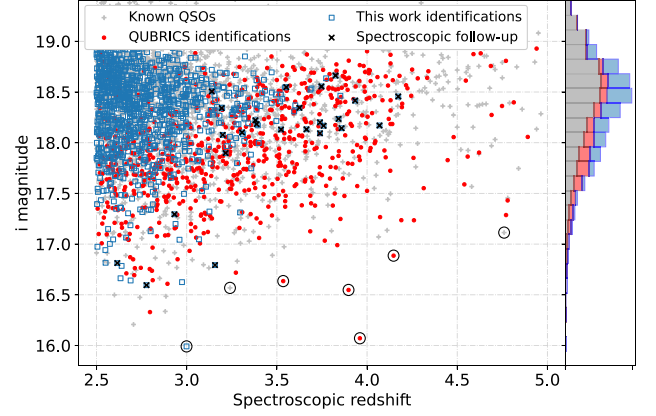


Figure 7. Known QSO in the Southern Hemisphere from the literature (grey ‘+’), identifications from QUBRICS before this work (red dots) and identifications from this work (blue, empty squares). Black crosses (‘x’) mark objects with a redshift determination from *Gaia* confirmed through spectroscopic follow-up (see last three columns in Table 2). The histogram on the right highlights the significant contribution of QUBRICS in the search for bright beacons in the south. Circled QSOs represent the new Golden Sample for the Sandage Test, updated with respect to [Paper II](#) (see Section 6.1).

update these estimates by defining a new ‘Golden Sample’: two main refinements are the inclusion of new bright QSOs discovered with QUBRICS and a realistic observing strategy (Dong et al. 2022) with a reduced number of targets, but still sufficient to have at least one suitable QSO available for observations from the ELT site at any time of the year.

In the following, we adopt the same assumptions as in [Paper II](#) (updated from Liske et al. 2008), namely the uncertainty in the radial velocity measurement, σ_v :

$$\sigma_v = g \times 1.35 \left(\frac{S/N}{3350} \right)^{-1} \left(\frac{1+z_{\text{QSO}}}{5} \right)^{-\gamma} \left(\frac{N_{\text{QSO}}}{30} \right)^{-0.5} \text{ cm s}^{-1}, \quad (2)$$

where the symbol ‘S/N’ refers to the total S/N per 0.0125 Å pixel per object accumulated over all observations, N_{QSO} is the number of QSOs in the sample, z_{QSO} is the redshift of the QSO, the γ exponent is 1.7 for $z_{\text{QSO}} \leq 4$ and 0.9 above. The form factor g is equal to 1 if all the targets are observed twice, at the beginning and at the end of the experiment, and becomes larger if the measurements are distributed in time, reaching 1.7 for a uniform distribution. The S/N per pixel for photon-noise limited observations can be written as:

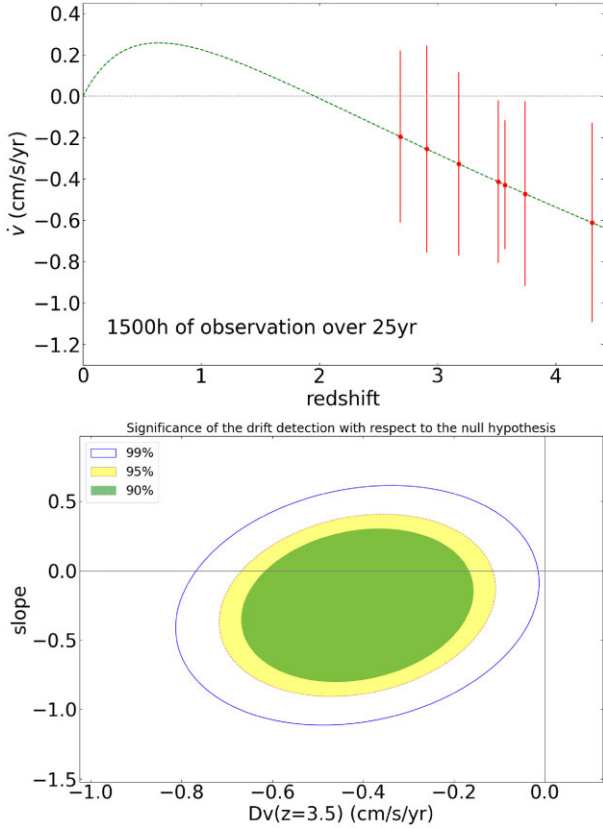
$$S/N = 650 \left[\frac{Z_X}{Z_r} 10^{0.4(16-m_X)} \left(\frac{D}{39\text{m}} \right)^2 \frac{t_{\text{int}}}{10\text{h}} \frac{\epsilon}{0.25} \right]^{\frac{1}{2}}, \quad (3)$$

where D , t_{int} , and ϵ are the telescope diameter, total integration time, and total efficiency, Z_X and m_X are the zeropoint and apparent magnitude of the source in the X -band, respectively, and $Z_r = (8.88 \times 10^{10}) \text{ s}^{-1} \text{ m}^{-2} \mu\text{m}^{-1}$ is the AB zeropoint for an effective wavelength of 6170 Å (corresponding to the SDSS r -band). The normalization of the above equation assumes a pixel size of 0.0125 Å and a central obscuration of the telescope’s primary collecting area of 10 per cent.

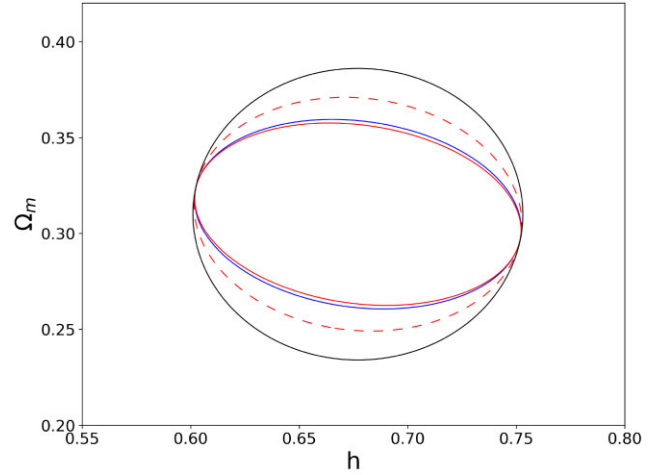
With respect to [Paper II](#) new brighter quasars (e.g. qid 1128023) have been found and we have reduced the Golden Sample to the seven brightest QSOs in the range $2.9 < z < 4.8$. The seven QSOs cover a wide range of right ascension and this favours observability/scheduling. We are also assuming a fixed allocation of $1500/7 = 214\text{h}$ of observation for each QSO (and not a variable

Table 4. The golden sample of Southern QSOs for the Sandage test.

qid	RA J2000	DEC J2000	r_{psf} (mag)	i_{psf} (mag)	G_{GAIA} (mag)	Redshift	σ_v in 214h (cm s ⁻¹)
956 603	00:41:31.44	-49:36:11.7	16.675	16.568	16.6981	3.240	12.5
1128023	05:29:15.81	-43:51:52.1	16.264	16.071	16.3452	3.960	7.8
1074326	05:48:03.20	-48:48:13.1	17.215	16.886	17.2414	4.147	11.2
1368250	14:51:47.05	-15:12:20.2	18.884	17.113	18.0754	4.760	12.1
860666	21:25:40.97	-17:19:51.4	16.644	16.548	16.8730	3.897	9.8
552772	22:47:08.93	-60:15:45.3	16.135	15.991	16.0895	2.999	10.4
27375882	22:54:51.33	-05:29:24.0	16.641	16.635	16.7257	3.535	11.1

**Figure 8.** Simulated measurement of the redshift drift with the seven QSOs of the Golden Sample of Table 4. 3000 random realizations have been fitted each with a linear trend, $\dot{v} = \text{slope} \times (z - 3.5) + \Delta v(z = 3.5)$. In the upper panel, the red bars represent the expected standard deviation and the green dotted line shows the expected signal of the redshift drift in a PLANCK18 cosmology. The green, yellow, and blue-line-encircled regions in the lower panel show the 90 percent and 95 percent and 99 percent confidence regions, respectively, for the measured $\dot{v}$.

amount of time required to reach a $\sigma_v = 22.8 \text{ cm s}^{-1}$ as in Paper II). In this way, assuming the most optimistic form factor $g = 1$ (which holds if all the targets are observed twice, at the beginning and at the end of the experiment), for a total of 1500 hours of observations in 25 years, the redshift drift is expected to be detected with a confidence level of the order of 99 per cent (as shown in Fig. 8). Besides, the combination of the measurement of the redshift drift with the QSOs Lyman Forest, at $z \gtrsim 2.5$, will be plausibly complemented at $z \lesssim 2.5$ with radio observations by the SKAO, CHIME, and FAST (Moresco et al. 2022), providing an extended redshift leverage of the drift and significantly shortening the observing time required. The goal of the Sandage

**Figure 9.** One-sigma confidence ellipses from our Fisher Matrix analysis of the expected data of the Golden Sample obtained with 1500 h of observation at the ELT over 25 yr. In solid red we have the Golden Sample observed with a form factor $g = 1$ (all the targets are observed twice, at the beginning and at the end of the experiment) and dash red for $g = 1.7$ (uniform distribution of the observations). In blue we show the previous result reported in Paper II for a Sample of 30 QSO, with 2500 h of observation and a form factor $g = 1$. The black line is the one-sigma constraint from the priors only.

test is to provide a direct, real-time mapping of the expansion rate of the Universe, independent of assumptions on gravity, geometry, or clustering, so the precision of the determination of cosmological parameters is not its main focus. Nevertheless, recent theoretical studies (Martins et al. 2016; Alves et al. 2019; Esteves et al. 2021) have uncovered important synergies with other cosmological probes, including the characterization of the physical properties of dark energy. In this respect it is of interest to mention that, assuming a flat Λ CDM cosmological model with the Planck Collaboration VI (2020) best-fitting cosmological parameters, $\Omega_m = 0.31$ and $H_0 = h 100 = 67.7 \text{ kms}^{-1} \text{ Mpc}^{-1}$, and external priors for both of these parameters of $\sigma_h = 0.05$ and $\sigma_{\Omega_m} = 0.05$, a Fisher Matrix analysis leads to a one-sigma uncertainty for Ω_m of 0.03 when the form factor is $g = 1$ and of 0.04 with $g = 1.7$, while the h constraint recovers the prior. Fig. 9 compares these one-sigma constraints with the ones from the previous version of the Golden Sample in Paper II. A small gain in the matter density constraint, in spite of the 1000 h reduction of the observing time, is clearly visible (comparing for consistency the two $g = 1$ cases).

The spectra of the seven QSOs in the Golden Sample have also been checked for the suitability of their absorption spectra for the redshift drift measurement (indeed, an eighth QSO has been

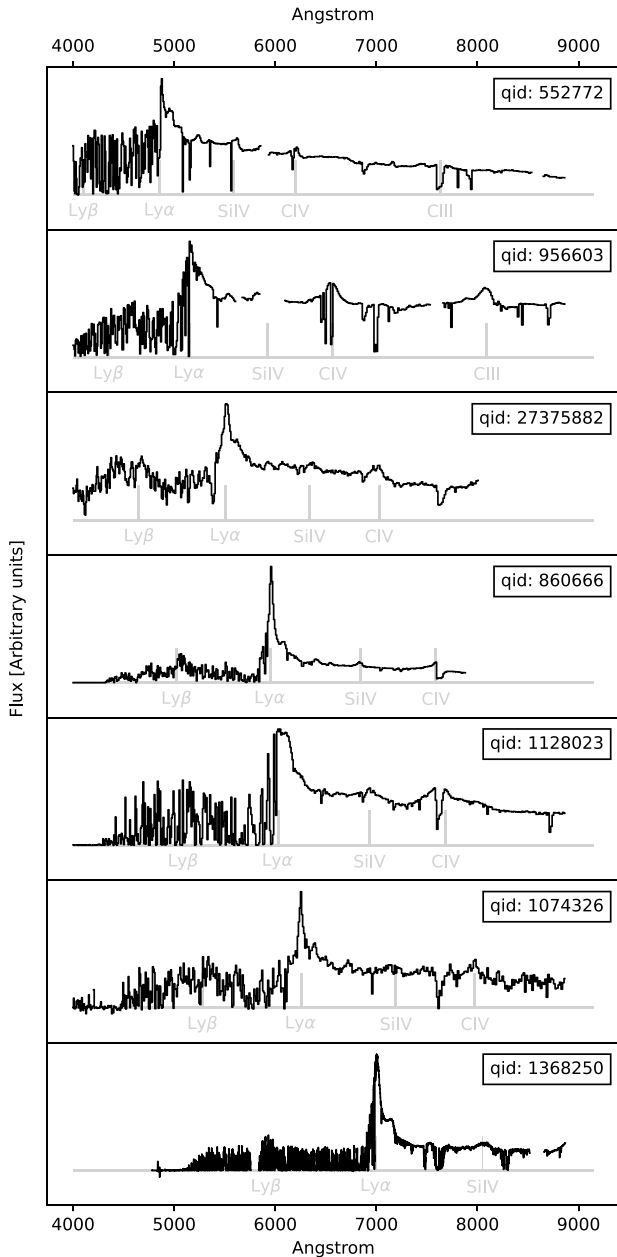


Figure 10. Spectra of the seven QSOs presented in Table 4. Spectra have been taken from the QUBRICS data base (qid 860666, 1128023, 1074326, and 27375882), the UVES SQUAD data set (qid 956603; Murphy et al. 2019), the SpecDB data base (qid 552772 and 1368250; Prochaska 2017). For visualization purposes, all spectra have been re-sampled to 500 km s^{-1} , with the exception of 1368250, sampled at 50 km s^{-1} .

excluded, due to the presence of BAL features) and are shown in Fig. 10.

The brightness of the QSOs of the Golden Sample makes the beginning of observations of this type conceivable in advance of the realization of the ELT, for example with the super-stable ESPRESSO spectrograph (Pepe et al. 2021) at the ESO VLT and a time investment of 3–4 nights per month. Indeed, a pilot program, 110.247Q.001 ‘An ESPRESSO Redshift Drift Experiment’, started at ESO in 2022.

ACKNOWLEDGEMENTS

We thank an anonymous referee for comments that prompted a significant improvement of this paper. This work was financed by Portuguese funds through FCT – Fundação para a Ciência e a Tecnologia in the framework of the project 2022.04048.PTDC. CJM also acknowledges FCT and POCH/FSE (EC) support through Investigador FCT Contract 2021.01214.CEECIND/CP1658/CT0001. We thank Società Astronomica Italiana (SAIt), Ennio Poretti, Gloria Andreuzzi and Marco Pedani for the observation support at TNG. Part of the observations discussed in this work are based on observations made with the Italian Telescopio Nazionale Galileo (TNG) operated on the island of La Palma by the Fundacion Galileo Galilei of the INAF (Istituto Nazionale di Astrofisica) at the Spanish Observatorio del Roque de los Muchachos of the Instituto de Astrofisica de Canarias.

The national facility capability for SkyMapper has been funded through ARC LIEF grant LE130100104 from the Australian Research Council, awarded to the University of Sydney, the Australian National University, Swinburne University of Technology, the University of Queensland, the University of Western Australia, the University of Melbourne, Curtin University of Technology, Monash University and the Australian Astronomical Observatory. SkyMapper is owned and operated by the Australian National University’s Research School of Astronomy and Astrophysics. The survey data have been processed and provided by the SkyMapper Team at ANU. The SkyMapper node of the All-Sky Virtual Observatory (ASVO) is hosted at the National Computational Infrastructure (NCI). Development and support the SkyMapper node of the ASVO has been funded in part by Astronomy Australia Limited (AAL) and the Australian Government through the Commonwealth’s Education Investment Fund (EIF) and National Collaborative Research Infrastructure Strategy (NCRIS), particularly the National eResearch Collaboration Tools and Resources (NeCTAR) and the Australian National Data Service Projects (ANDS).

This work has made use of data from the European Space Agency (ESA) mission *Gaia* (<https://www.cosmos.esa.int/gaia>), processed by the Gaia Data Processing and Analysis Consortium (DPAC, <https://www.cosmos.esa.int/web/gaia/dpac/consortium>). Funding for the DPAC has been provided by national institutions, in particular the institutions participating in the Gaia Multilateral Agreement. This job has made use of the Python package GaiaXPpy, developed and maintained by members of the Gaia Data Processing and Analysis Consortium (DPAC), and in particular, Coordination Unit 5 (CU5), and the Data Processing Centre located at the Institute of Astronomy, Cambridge, UK (DPCI).

This publication makes use of data products from the Two Micron All Sky Survey, which is a joint project of the University of Massachusetts and the Infrared Processing and Analysis Center/California Institute of Technology, funded by the National Aeronautics and Space Administration and the National Science Foundation.

This publication makes use of data products from the Wide-field Infrared Survey Explorer, which is a joint project of the University of California, Los Angeles, and the Jet Propulsion Laboratory/California Institute of Technology, funded by the National Aeronautics and Space Administration.

This paper includes data gathered with the 6.5 m Magellan Telescopes located at Las Campanas Observatory, Chile.

DATA AVAILABILITY

The data underlying this article will be shared on reasonable request to the corresponding author.

REFERENCES

- Alves C. S., Leite A. C. O., Martins C. J. A. P., Matos J. G. B., Silva T. A., 2019, *MNRAS*, 488, 3607
- Anderson T. W., 2003, *An Introduction to Multivariate Statistical Analysis* (Wiley Series in Probability and Mathematical Statistics), 3 edn. Wiley, Hoboken, NJ. Available at: <https://www.wiley.com/en-us/An+Introduction+to+Multivariate+Statistical+Analysis>
- Banase K., Ponz D., Ounnas C., Grosbol P., Warmels R., 1988, in Robinson L. B., ed., *Instrumentation for Ground-Based Optical Astronomy*. Springer, Berlin, p. 431
- Boutsia K. et al., 2020, *ApJS*, 250, 26 (Paper II)
- Boutsia K. et al., 2021, *ApJ*, 912, 111 (Paper IV)
- Calderone G., Nicastro L., Ghisellini G., Dotti M., Sbarrato T., Shankar F., Colpi M., 2017, *MNRAS*, 472, 4051
- Calderone G. et al., 2019, *ApJ*, 887, 268 (Paper I)
- Chambers K. C. et al., 2016, preprint ([arXiv:1612.05560](https://arxiv.org/abs/1612.05560))
- Chen T., Guestrin C., 2016, in *Proceedings of the 22nd ACM SIGKDD International Conference on Knowledge Discovery and Data Mining. KDD'16*, ACM, New York, NY, USA, p. 785. Available at: <http://doi.acm.org/10.1145/2939672.2939785>
- Colless M. et al., 2001, *MNRAS*, 328, 1039
- Cupani G. et al., 2022, *MNRAS*, 510, 2509
- De Angeli F. et al., 2022, *A&A*, preprint ([arXiv:2206.06143](https://arxiv.org/abs/2206.06143))
- Dong C., Gonzalez A., Eikenberry S., Jeram S., Likamonsavadi M., Liske J., Stelter D., Townsend A., 2022, *MNRAS*, 514, 5493
- Estevés J., Martins C. J. A. P., Pereira B. G., Alves C. S., 2021, *MNRAS*, 508, L53
- Flesch E., 2013, *Publ. Astron. Soc. Austr.*, 30, e004
- Fontanot F. et al., 2023, *MNRAS*, 520, 740
- Gaia Collaboration, 2021, *A&A*, 649, A1
- Grazian A. et al., 2022, *ApJ*, 924, 62
- Guarneri F., Calderone G., Cristiani S., Fontanot F., Boutsia K., Cupani G., Grazian A., D'Odorico V., 2021, *MNRAS*, 506, 2471 (Paper III)
- Guarneri F. et al., 2022, *MNRAS*, 517, 2436
- Hinton S. R., Davis T. M., Lidman C., Glazebrook K., Lewis G. F., 2016, *Astron. Comput.*, 15, 61
- Inoue A. K., Shimizu I., Iwata I., Tanaka M., 2014, *MNRAS*, 442, 1805
- Jones D. H. et al., 2009, *MNRAS*, 399, 683
- Kubota A., Done C., 2018, *MNRAS*, 480, 1247
- Laor A., Netzer H., 1989, *MNRAS*, 238, 897
- Liske J. et al., 2008, *MNRAS*, 386, 1192
- Lyke B. W. et al., 2020, *ApJS*, 250, 8
- Marconi A. et al., 2022, in Evans C. J., Bryant J. J., Motohara K. eds, *Proc. SPIE Conf. Ser. Vol. 12184, Ground-based and Airborne Instrumentation for Astronomy IX*. SPIE, Bellingham, p. 1218424
- Martins C. J. A. P., Martinelli M., Calabrese E., Ramos M. P. L. P., 2016, *Phys. Rev. D*, 94, 043001
- McMahon R. G. et al., 2013, *Msngr*, 154, 35
- Mor R., Netzer H., Elitzur M., 2009, *ApJ*, 705, 298
- Moresco M. et al., 2022, *Living Rev. Relat.*, 25, 6
- Morrissey P. et al., 2007, *ApJS*, 173, 682
- Murphy M. T., Kacprzak G. G., Savorgnan G. A. D., Carswell R. F., 2019, *MNRAS*, 482, 3458
- Onken C. A. et al., 2022, *MNRAS*, 511, 572
- Pepe F. et al., 2021, *A&A*, 645, A96
- Pier E. A., Krolik J. H., 1993, *ApJ*, 418, 673
- Planck Collaboration VI, 2020, *A&A*, 641, A6
- Prochaska J. X., 2017, *Astron. Comput.*, 19, 27
- Reis I., Baron D., Shahaf S., 2019, *AJ*, 157, 16
- Richards G. T. et al., 2006, *ApJS*, 166, 470
- Sandage A., 1962, *ApJ*, 136, 319
- Schindler J.-T. et al., 2019a, *ApJS*, 243, 5
- Schindler J.-T. et al., 2019b, *ApJ*, 871, 258
- Sevilla-Noarbe I. et al., 2021, *ApJS*, 254, 24
- Skrutskie M. F. et al., 2006, *AJ*, 131, 1163
- Sun W.-H., Malkan M. A., 1989, *ApJ*, 346, 68
- Vanden Berk D. E. et al., 2001, *AJ*, 122, 549
- Véron-Cetty M. P., Véron P., 2010, *A&A*, 518, A10
- Wolf C. et al., 2018, *Publ. Astron. Soc. Austr.*, 35, e010
- Wolf C. et al., 2020, *MNRAS*, 491, 1970
- Wright E. L. et al., 2010, *AJ*, 140, 1868
- Yang J. et al., 2016, *ApJ*, 829, 33

SUPPORTING INFORMATION

Supplementary data are available at *MNRAS* online.

NewSpec.pdf

Please note: Oxford University Press is not responsible for the content or functionality of any supporting materials supplied by the authors. Any queries (other than missing material) should be directed to the corresponding author for the article.

This paper has been typeset from a $\text{\LaTeX}$ file prepared by the author.