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Supersolar Ni/Fe production in the Type IIP SN 2012ec

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ABSTRACT

SN 2012ec is a Type IIP supernova (SN) with a progenitor detection and comprehensive photospheric phase observational coverage. Here, we present Very Large Telescope and Public ESO Spectroscopic Survey of Transient Objects observations of this SN in the nebular phase. We model the nebular [O I] $\lambda\lambda 6300, 6364$ lines and find their strength to suggest a progenitor main-sequence mass of 13–15 $M_{\odot}$. SN 2012ec is unique among hydrogen-rich SNe in showing a distinct line of stable nickel [Ni II] $\lambda 7378$. This line is produced by ^{58}Ni , a nuclear burning ash whose abundance is a sensitive tracer of explosive burning conditions. Using spectral synthesis modelling, we use the relative strengths of [Ni II] $\lambda 7378$ and [Fe II] $\lambda 7155$ (the progenitor of which is ^{56}Ni) to derive a Ni/Fe production ratio of 0.20 ± 0.07 (by mass), which is a factor 3.4 ± 1.2 times the solar value. High production of stable nickel is confirmed by a strong [Ni II] 1.939 μm line. This is the third reported case of a core-collapse SN producing a Ni/Fe ratio far above the solar value, which has implications for core-collapse explosion theory and galactic chemical evolution models.

Key words: line: formation – nuclear reactions, nucleosynthesis, abundances – stars: evolution – supernovae: general – supernovae: individual: SN 2012ec.

1 INTRODUCTION

Research on the elusive core-collapse supernova (SN) explosion mechanism is ongoing in multiple groups around the world (see Janka 2012, for a review). As the core of the star collapses to a neutron star, a shock wave is born that travels out through the infall of the mantle. The initial energy of the prompt shock is not enough to reverse the accretion, but a delayed explosion can be obtained if neutrino heating is efficient enough. Several successful explosions have been obtained in recent simulations with self-consistently cal-

culated neutrino luminosities, e.g. by Kitaura, Janka & Hillebrandt (2006) in spherically symmetric models of electron-capture SNe, and with additional support from multidimensional instabilities in more massive progenitors (Marek & Janka 2009; Müller, Janka & Heger 2012; Bruenn et al. 2013).

Important constraints on the explosive process can be obtained by observations and modelling of SNe in the nebular phase, when the ejecta become optically thin and the inner regions of nucleosynthesis products become visible. The deepest layers of iron-group nuclei, which are the ashes of explosive silicon burning, are directly associated with the critical gain layers where the neutrinos power the shock wave, and are therefore direct diagnostics of the explosion process.

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Table 1. Optical photometry. V , R , I magnitudes are template subtracted whereas B , r , z are not. B , V , R , I magnitudes are in the Landolt Vega system, whereas r , z are in the AB system. The phase is relative to the estimated explosion epoch of 2012 August 5 (MJD 56144).

Date	Phase	Telescope	B	V	R	r	I	z
2013 Feb 08	+187	SMARTS	19.29 ± 0.09	18.47 ± 0.08	17.29 ± 0.08	...	16.88 ± 0.08	...
2013 Feb 23	+202	LT	19.31 ± 0.10	18.62 ± 0.09	...	17.34 ± 0.15	17.19 ± 0.09	17.13 ± 0.10
2013 Mar 01	+208	SMARTS	...	18.62 ± 0.11	17.52 ± 0.09	...	17.25 ± 0.09	...
2013 Sep 11	+403	NTT	...	20.60 ± 0.11	19.19 ± 0.09	...	20.18 ± 0.09	...
2013 Oct 07	+429	NTT	...	20.78 ± 0.10	19.68 ± 0.06	...	20.38 ± 0.09	...
2013 Nov 22	+475	NTT	...	21.04 ± 0.08	20.87 ± 0.07	...	20.72 ± 0.09	...
2013 Dec 23	+506	NTT	...	22.38 ± 0.08	21.53 ± 0.11	...	21.13 ± 0.11	...
2014 Jan 31	+545	NTT	...	22.66 ± 0.15	22.16 ± 0.10	...	...	...

Type IIP (plateau) SNe are the most common core-collapse explosions (~ 50 per cent per unit volume; Li et al. 2011). Light-curve models (Chevalier 1976) and progenitor detections (Smartt et al. 2009) have shown that these are explosions of red supergiant (RSG) stars. Progenitor luminosities (Smartt et al. 2009) and nucleosynthesis yields (Jerkstrand et al. 2014, hereafter J14) suggest helium core masses $M_{\text{He}} \lesssim 5 M_{\odot}$ ($M_{\text{ZAMS}} \lesssim 18 M_{\odot}$) for the progenitor population, whereas hydrodynamical modelling favours a more extended mass range (e.g. Utrobin & Chugai 2009; Dall’Ora et al. 2014).¹

SN 2012ec is a Type IIP explosion that occurred in NGC 1084 in early 2012 August. A potential progenitor was reported by Maund et al. (2013), a high-luminosity star with $\log(L/L_{\odot}) = 5.0\text{--}5.4$. Stars with luminosity on the upper end of this range have not yet been seen to explode as RSGs (Smartt et al. 2009). The possibility that the progenitor was a luminous, high-mass star, motivated us to embark on a follow-up campaign to follow the SN through its photospheric and nebular phases. The photospheric phase Public ESO Spectroscopic Survey of Transient Objects (PESSTO) data are presented by Barbarino et al. (2014). Here we report on PESSTO and Very Large Telescope (VLT) observations in the nebular phase, and modelling of these data. For the analysis, we follow Maund et al. (2013) and assume an explosion epoch of 2012 August 5, a distance of 17.3 Mpc, an extinction $E(B - V) = 0.11$ mag, and a heliocentric recession velocity of 1407 km s^{-1} (Koribalski et al. 2004). All observed spectra displayed in the paper have been dereddened and redshift corrected.

2 OBSERVATIONS AND DATA REDUCTION

2.1 Photometry

Optical and near-infrared (NIR) imaging were obtained for the phase +176 to +551 d after explosion using the European Southern Observatory’s New Technology Telescope (NTT) with ESO Faint Object Spectrograph and Camera, v.2 (EFOSC2) and Son of ISAAC (SOFI), the Small and Moderate Aperture Research Telescope System (SMARTS) 1.3-m telescope (operated by the SMARTS consortium) using ANDYCAM, and the Liverpool 2-m Telescope (LT) using RATCam. Tables 1 and 2 present the photometry. The V , R , I magnitudes are template subtracted (see below) whereas B , r , z are not (as no templates were available in these filters). The NTT data were collected as part of the PESSTO program (Smartt et al. 2013), and are a continuation of the public monitoring campaign of

Table 2. NIR photometry. Magnitudes are in the 2MASS Vega system. The upper limits are 3σ values. The first two epochs are template subtracted.

Date	Phase	Telescope	J	H	K_s
2013 Jan 28	+176	NTT	16.59 ± 0.04	16.18 ± 0.04	16.29 ± 0.05
2013 Feb 21	+201	NTT	17.08 ± 0.04	16.55 ± 0.04	16.70 ± 0.05
2014 Feb 06	+551	NTT	> 20.3	> 19.6	> 19.8

SN 2012ec described in Maund et al. (2013) and Barbarino et al. (2014). The EFOSC2 images were reduced (trimmed, bias subtracted, and flat-fielded) using the PESSTO pipeline as described in Smartt et al. (2014, submitted), whereas the LT and SMARTS images were reduced automatically using their respective pipelines.

Photometric zero-points and colour terms were computed through observations of Landolt standard fields (Landolt 1992). Three of the seven nights in which optical imaging was carried out were photometric and we calibrated the magnitudes of a local stellar sequence using these data. We chose the reference stars 1, 11, and 12 presented in Barbarino et al. (2014), and found reasonable agreement in the computed magnitudes of the secondary standards (the differences were within 0.05 mag). The average magnitudes of the local-sequence stars were used to calibrate the photometric zero-points obtained in non-photometric nights, or when the colour terms were not retrieved.

The complex and high background galaxy flux at the position of SN 2012ec meant that image template subtraction was required on all the EFOSC2 VRI^2 data. Fortunately the large programme 184.D-1140 (PI: S. Benetti) had observed SN 2009H in the same galaxy and VRI images from 2009 October 10 with EFOSC2 were available to construct pre-discovery templates. Each filter had 5×120 s frames taken, which were co-added and then subtracted from the target frames using the HOTPANTS³ image subtraction software (which is based on the algorithm presented in Alard 2000).

For the NIR observations, we took multiple, dithered, on-source exposures; these images were then flat-fielded and median combined to create a sky frame. The sky frame was subtracted from each of the individual images, which were then aligned and co-added. The total exposure times in the two epochs which had detections were 400 s (J), 360 s (H), and 1080 s (K_s) on 2013 January 28 and 400 s (J), 360 s (H), and 720 s (K_s) on 2013 February 21. NIR photometry of the reference stars was calibrated using the Two Micron All Sky Survey (2MASS) catalogue magnitudes (Skrutskie et al. 2006). The full reduction procedure is described in Smartt et al.

¹ The light curve is most sensitive to the hydrogen envelope mass (see discussion in Dessart et al. (2013)), with best-fitting values being 20–30 $M_{\odot}$ for some SNe, favouring high-mass progenitors.

² EFOSC2 mounts the Gunn i #705 filter that we calibrated to Johnson–Cousin I .

³ <http://www.astro.washington.edu/users/becker/hotpants.html>

Table 3. Summary of spectroscopic observations.

Date	Phase	Telescope	Instrument	Wavelength	Slit width (arcsec)	Resolution (Å)
2013 Feb 07	+185	NTT	SOFI + Blue Grism	0.95–1.64 μm	1.0	23
2013 Aug 11	+371	VLT	X-shooter + UVB/VIS/NIR	0.32–2.45 μm	1.0/0.9/0.9	0.9/1.7/2.3
2013 Sep 11	+402	NTT	EFOSC2 + Grism #11	3380–7250 Å	1.0	13

(2014, submitted) and these reduced images have been released as part of the PESSTO Spectroscopic Survey Data Release 1 (SSDR1). Users should note that the on-source dithering pattern results in the oversubtraction of the sky background in some of the images, but the photometry of the SN is not affected. Further imaging was taken a year later on 2014 February 6 with exposure times of 200 s (*J*), 150 s (*H*), and 900 s (*K_s*). SN 2012ec was not detected on these images. We used these images as templates so that the values at 176 and 201 d in Table 2 are template subtracted.

Photometric flux measurements were performed using a point spread function (PSF) fitting technique. We simultaneously fitted the PSF of SN 2012ec and the sequence stars using the SNOOPY⁴ package within IRAF.⁵

2.2 Spectroscopy

We obtained three spectra in the nebular phase. The first was a NIR spectrum at +185 d using NTT and the SOFI NIR spectrograph, the second was a combined optical and NIR spectrum at +371 d using X-shooter on the ESO VLT (via programme 091.D-0608, PI: J. Sollerman). The third was an optical spectrum at +402 d using EFOSC2 on the NTT. The details of the wavelength coverage, grisms employed, and spectral resolutions are listed in Table 3.

The NTT SOFI spectrum was a combination of exposures taken with on-slit nodding to give a total exposure time of 3240 s with the blue grism (BG) and the grism blue filter (GBF) order blocking filter. The PESSTO observing sequence and reduction procedure are described in Smartt et al. (2014, submitted). After extraction and wavelength calibration, the spectrum was corrected for telluric absorption with the standard Hip010502 and flux calibrated with the spectrophotometric standard EG274.

The NTT EFOSC2 optical spectrum was reduced using the PESSTO pipeline. Wavelength calibration was performed using the spectra of comparison lamps and adjusted by checking the night sky line positions. The spectrum was corrected for telluric absorption by subtracting a model absorption spectrum and flux calibration was performed using the PESSTO selected spectrophotometric standard stars observed on the same night.

SN 2012ec was observed with the VLT+X-shooter at an airmass of 1.2 on 2013 August 9.3 UT, and again on August 11.3 UT at an airmass of 1.1. X-shooter consists of three echelle spectrographs, each covering a separate wavelength range [ultraviolet-blue (UVB): 300–560 nm; visible (VIS): 556–1024 nm; NIR: 1024–2480 nm], which are used simultaneously through a dichroic. A 1.0-arcsec slit was used in the UVB arm, and 0.9-arcsec slits were used in the VIS and NIR arms. Observations were taken while nodding the telescope, moving the SN between two positions in the slit

to facilitate sky subtraction in the NIR arm (using the template XSHOOTER_SLT_OBS_AUTO NOD ON SLIT). The spectrum from each arm was reduced using version 2.4.0 of the X-shooter pipeline⁶ running under the REFLEX environment. All data were pre-reduced (bias subtracted in the UVB and VIS arms, flat-fielded, and wavelength calibrated) within the pipeline, and the separate echelle orders were merged into a single spectrum. The bias level and sky background in the NIR arm were removed by subtracting pairs of consecutive spectra taken with the target at different locations in the slit. Finally, the spectra were optimally extracted, and flux calibrated using spectrophotometric standards.

3 SPECTRAL SYNTHESIS MODELS

As foundation for our analysis we use the spectral models presented by Jerkstrand et al. (2012, hereafter J12) and J14. These models compute the temperature and Non-Local-Thermodynamic-Equilibrium (NLTE) excitation/ionization solutions in each zone of the SN ejecta, taking all relevant thermal, non-thermal, and radiative rates into account. Specifically, the modelling consists of the following computational steps. (i) Transport and deposition of radioactive decay products (gamma-rays, X-rays, leptons). (ii) The distribution of non-thermal electrons created by the radioactivity. (iii) Thermal equilibrium in each compositional zone. (iv) NLTE ionization balance for the 20 most common elements. (v) NLTE excitation structure for about 50 atoms and ions. (vi) Radiative transfer through the ejecta. The solutions are generally coupled to each other and global convergence is achieved by iteration.

The calculations are applied to the SN ejecta models computed by Woosley & Heger (2007), with some modifications to mimic the effects of multidimensional mixing which is known to be important (e.g. Hammer, Janka & Müller 2010). Specifically, the metal zones (Fe/He, Si/S, O/Si/S, O/Ne/Mg, and O/C) are macroscopically mixed together with parts of the He/C, He/N, and H zones in a core region between 0 and 1800 km s^{−1}. Each zone has an individual filling factor in this core and is distributed over 10³ clumps. The model takes dust formation into account by applying a grey absorption coefficient over the core from 250 d, increasing with time such that the optical depth increases by 1.8E-3 per day (a calibration to the observed dust formation in SN 2004et; J14). Over the period covered here, this dust component has only a small effect on the optical/NIR spectrum as $\tau_{\text{dust}} < 0.25$ up to 400 d. Outside the core reside the remaining helium layers followed by the hydrogen envelope, whose density profile is determined by the one-dimensional hydrodynamic solutions.

In addition to these models, we here computed a $M_{\text{ZAMS}} = 15 M_{\odot}$ model at 370 d using a lower ⁵⁶Ni mass of 0.03 $M_{\odot}$, as appropriate for SN 2012ec (all models computed in J12 and J14 have ⁵⁶Ni masses of 0.062 $M_{\odot}$). The lower ⁵⁶Ni mass was achieved by reducing the mass of the Fe/He zone. The density of this zone was

⁴ Based on the IRAF DAOPHOT package.

⁵ IRAF is distributed by the National Optical Astronomy Observatory, which is operated by the Association of Universities for Research in Astronomy (AURA) under cooperative agreement with the National Science Foundation.

⁶ <http://www.eso.org/sci/software/pipelines/>

kept the same, and the filling factors of the other core zones were increased by a factor of 1.09 to fill the vacated volume.

4 ANALYSIS OF PHOTOMETRIC DATA

4.1 The ^{56}Ni mass

In the early tail phase of a Type II SN (~ 150 – 200 d), steady state sets in so the emergent luminosity follows the instantaneous energy input by radioactivity, dominated by the $^{56}\text{Ni} \rightarrow ^{56}\text{Co} \rightarrow ^{56}\text{Fe}$ decay chain. At the same time, the escape fraction of gamma-rays is negligible. This phase therefore offers an opportunity to determine the ^{56}Ni mass by estimating the bolometric luminosity.

We estimate the $BVRIJHK_s$ quasi-bolometric luminosity at 187 d by combining the observed optical (dereddened) photometry at 187 d with extrapolated (dereddened) NIR photometry from day 176, using an extrapolation factor $\exp(-11\text{ d}/111.4\text{ d})$. The magnitudes were converted to flux values at the effective wavelengths of the filters and linearly interpolated. Since no useful templates exist in B , the B magnitudes reported in Section 2.1 are not template subtracted. To obtain an estimate for the SN B magnitudes, we estimate the contribution from the galaxy by taking the $B - V$ colour from the galaxy model derived in Section 5.1 and apply this to the V -band template, obtaining an estimate of the galaxy B magnitude. Subtracting this artificial B template from the measured

B magnitudes gave estimates for the SN of $B = 19.36$ mag at 187 d and $B = 19.50$ mag at 202 d. The resulting $B - V$ colours of the SN are similar to other Type IIP SNe (Maguire et al. 2010; Inserra et al. 2012, 2013; Tomasella et al. 2013).

We obtain $L_{BVRIJHK_s}^{187\text{d}} = 7.0 \times 10^{40} \text{ erg s}^{-1}$. To estimate the full bolometric luminosity, we assume the fraction outside B to K_s to be the same as in SN 1987A at a similar epoch, which we compute as 19 per cent at 165 d using data from Hamuy et al. (1988) and Bouchet et al. (1989), and using $E(B - V) = 0.15$ mag for SN 1987A. We then obtain $L_{\text{bol}}^{187\text{d}} = 8.6 \times 10^{40} \text{ erg s}^{-1}$, which from using equation (6) in J12 corresponds to an initial ^{56}Ni mass of $0.033 M_{\odot}$. Doing the same evaluation using the optical data at 202 d combined with the NIR data at 200 d gives a ^{56}Ni mass of $0.030 M_{\odot}$. These values are in good agreement with derivations in Barbarino et al. (2014) based on earlier epochs ($0.04 \pm 0.01 M_{\odot}$). Doing the same analysis with V to K_s (to avoid the uncertain B band) gives very similar numbers. For the rest of the analysis in this paper we adopt a ^{56}Ni mass of $0.03 M_{\odot}$, and take the error in this to be $0.01 M_{\odot}$.

4.2 Photometric evolution

Fig. 1 shows the evolution of SN 2012ec (red circles) in B to K_s bands between 176 and 550 d post-explosion, compared to the models computed in J12 as well as the observed evolution of the well-studied SN 2004et (blue diamonds). All magnitudes have been

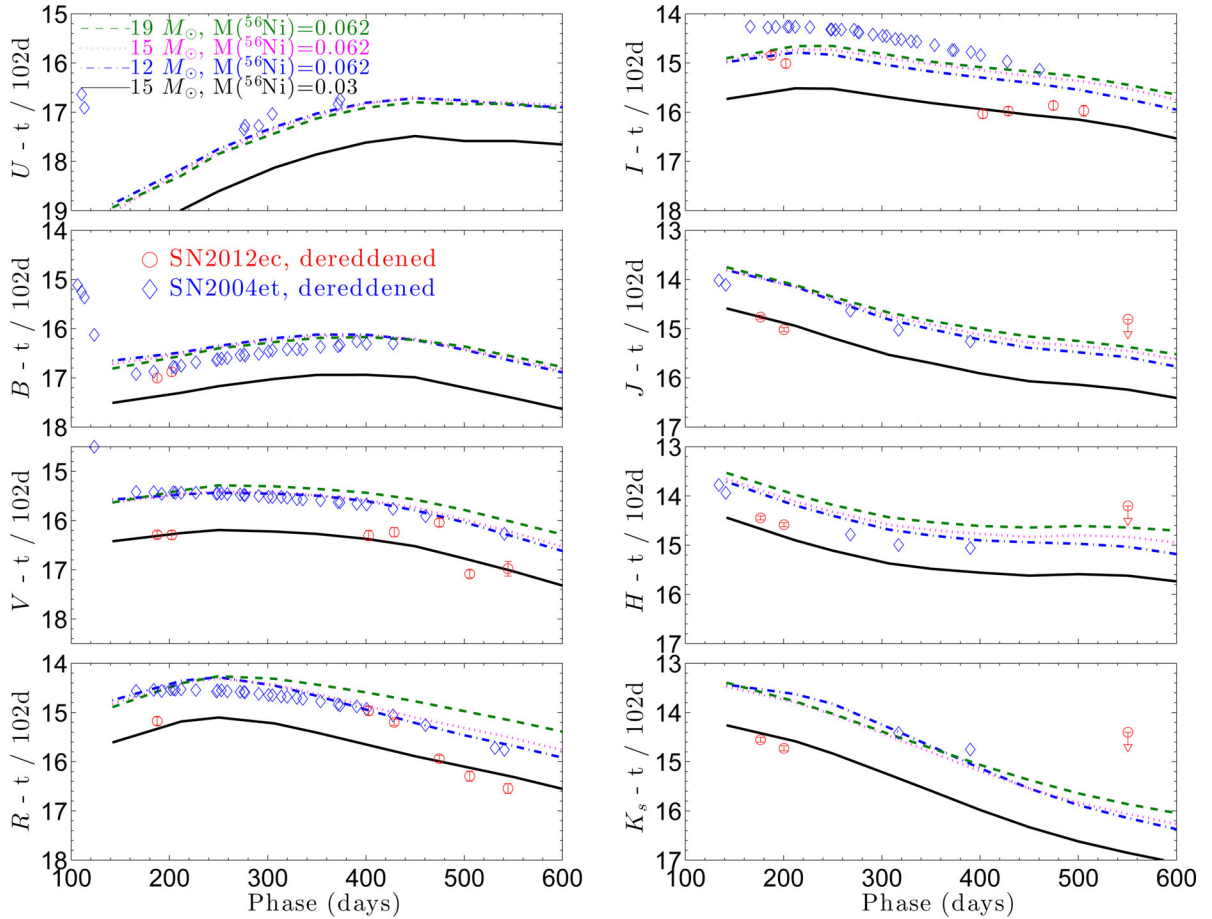


Figure 1. Evolution of B to K_s (subtracted by $t/102$ d to normalize to the ^{56}Co decay) of SN 2012ec (red points) and SN 2004et (blue points, data from Sahu et al. 2006; Maguire et al. 2010; Fabbri et al. 2011). The 12 , 15 , and $19 M_{\odot}$ models of J12 are plotted as green, magenta, and blue lines, and the $15 M_{\odot}$ model scaled with $0.03/0.062$ to adjust for a lower ^{56}Ni mass of $0.030 M_{\odot}$ is plotted as a black solid line. All observed magnitudes have been corrected for dust extinction and scaled to the distance of SN 2012ec.

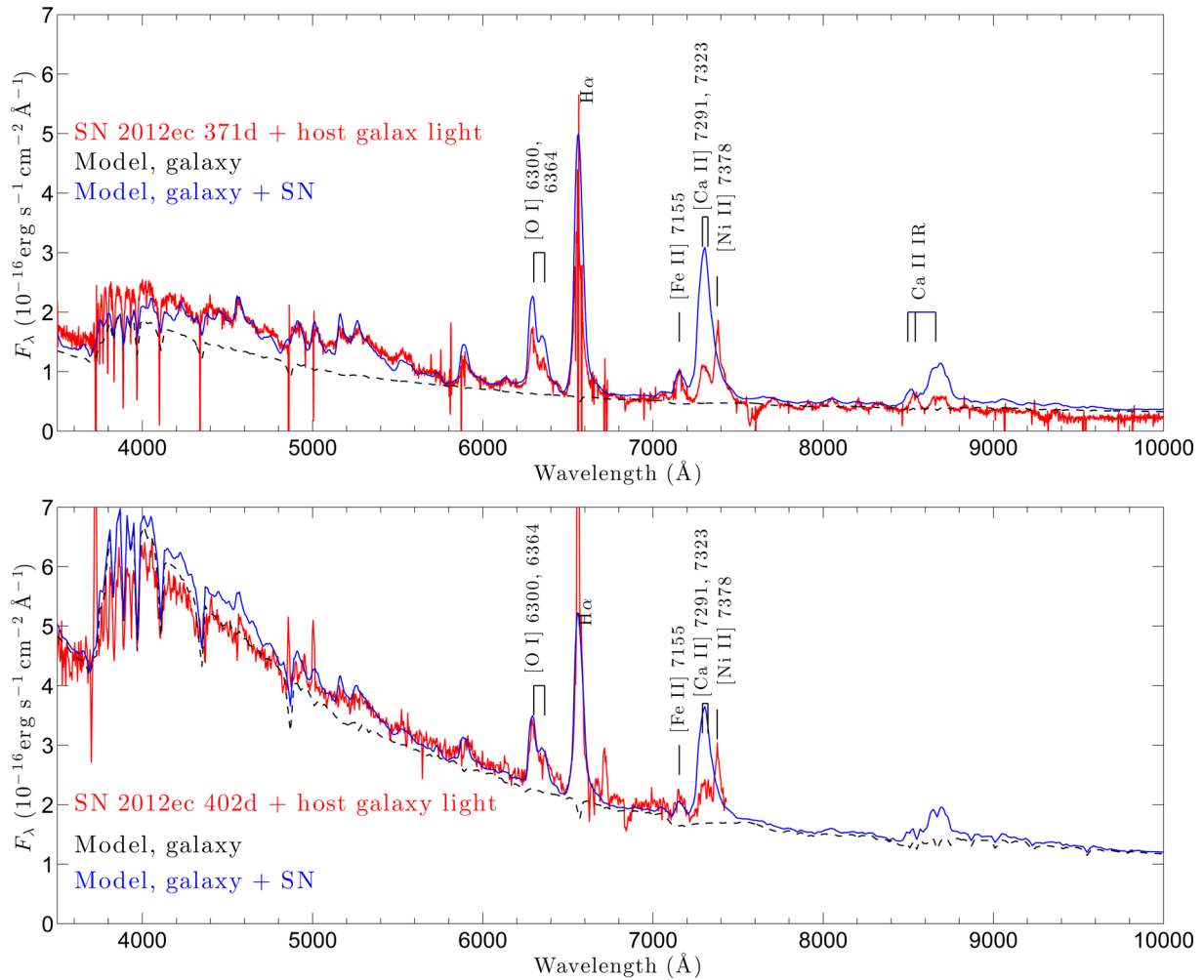


Figure 2. Top: smoothed X-shooter spectrum of the SN 2012ec location at 371 d (red), dereddened and redshift corrected. Also shown is the STARBURST99 galaxy model (black dashed), and this galaxy model plus the new $15 M_{\odot}$ SN model with a ^{56}Ni mass of $0.03 M_{\odot}$ (blue). Bottom: NTT spectrum of SN 2012ec at 402 d (red), dereddened and redshift corrected. Also shown is the same galaxy model as above, but scaled by a factor of 3.6 (black dashed), and the sum of this model and the same SN model as above, but scaled with a factor $\exp(-31 \text{ d}/111 \text{ d})$ to compensate for the later epoch (blue).

normalized to the ^{56}Co decay (1 mag in 102 d) by having a term $t/102 \text{ d}$ subtracted. The observed magnitudes have been corrected for dust extinction ($E(B - V) = 0.11$ mag for SN 2012ec and $E(B - V) = 0.41$ mag for SN 2004et) and SN 2004et has been scaled to the same distance as SN 2012ec.

The early phase photometry (176–202 d) shows reasonable agreement with a model scaled with a factor of 0.03/0.062 (to adjust for the ^{56}Ni mass of $0.03 M_{\odot}$ in SN 2012ec, see Section 4.1). The data between 400 and 550 d show quite a lot of scatter. This is likely due to difficulties in the template subtractions, as the background flux is significantly higher than the SN flux at these epochs (Section 5). The average flux levels are still in reasonable agreement with the model brightness for a ^{56}Ni mass of $0.03 M_{\odot}$.

5 ANALYSIS OF SPECTROSCOPIC DATA

5.1 Optical spectral modelling

The optical spectra from the two epochs at +371 d (VLT X-shooter) and +402 d (NTT EFOSC2) are shown in Fig. 2 (both dereddened and redshift corrected). To reduce noise in the X-shooter spectrum it was smoothed substantially using a Savitzky–Golay filter of second

order, 31 points. Given the position of SN 2012ec in a bright region of the host galaxy, the slits included significant light from nearby stellar populations. The slope of the spectra and observed narrow absorption due to the Balmer series indicate significant contamination from OB stars. To make a comparison of the observed spectrum with a SN model, we therefore need to add a model component for this background galaxy light.

To model the galactic background, we used STARBURST99 (Leitherer et al. 1999) to compute a model with a single starburst of standard parameters.⁷ The spectrum of this model at 30 Myr is plotted as a black dashed line in Fig. 3. Also plotted here (red line) is a spectrum of the star-forming region C9 in which SN 2012ec exploded, presented by Ramya, Sahu & Prabhu (2007). The Ramya spectrum has been dereddened with $E(B - V) = 0.45$ mag (as estimated by Ramya, Sahu & Prabhu 2007). The comparison shows that the model spectrum is a reasonable representation of the

⁷Two-component initial mass function (IMF) with $\alpha = 1.3$ for $M = 0.1\text{--}0.5 M_{\odot}$ and $\alpha = 2.3$ for $M = 0.5\text{--}100 M_{\odot}$, Geneva 2012 tracks with no rotation, wind model ‘evolution’, atmosphere model ‘Pauldrach/Hillier’, RSG microturbulence 3 km s^{-1} , and solar metallicity.

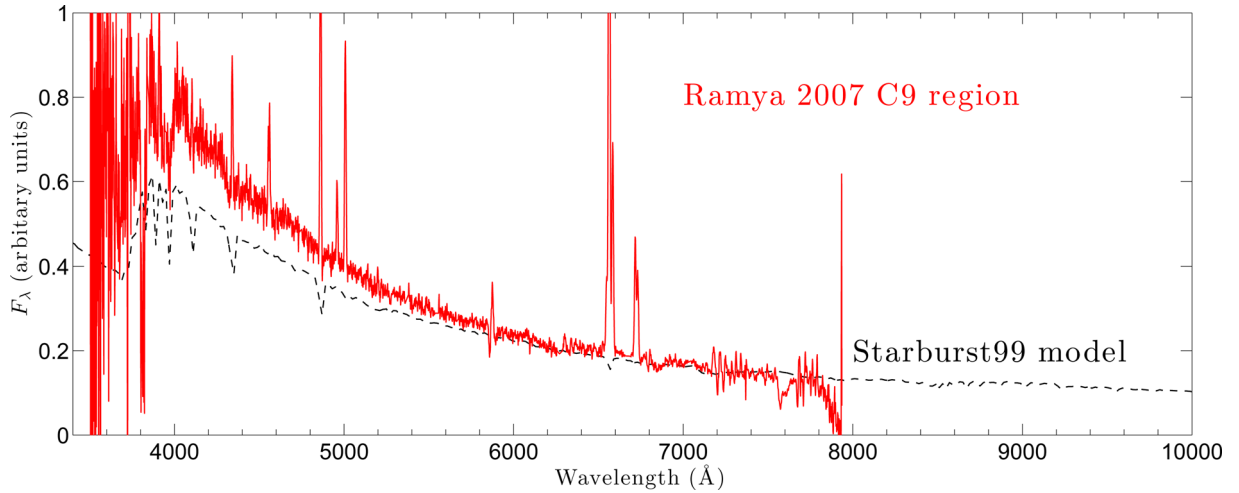


Figure 3. STARBURST99 galaxy model (black dashed line), and a spectrum of the C9 region (in which SN 2012ec exploded) of NGC 1084, from Ramya et al. (2007), dereddened with $E(B - V) = 0.45$ mag (red line).

galactic light around SN 2012ec, although the C9 region is a quite large region of the galaxy (~ 30 arcsec) and we lack knowledge of how much conditions change locally within this region (both intrinsic spectra and extinction).

Plotted in Fig. 2, as solid blue lines, are the sum of the galaxy model and the spectrum of the new $15 M_{\odot}$ model computed with a ^{56}Ni mass of $0.03 M_{\odot}$. The galaxy component is plotted as a black dashed line. The flux scale of the galaxy model is arbitrary and has been chosen to give overall good fits of the galaxy+SN model at each epoch. As the different epoch spectra have different slit widths, slit angles, and image qualities, the amount of galaxy light within the SN extraction window varies between them. In the +402 d spectrum the flux (and thus our galaxy model scaling factor) is a factor of 3.6 higher than for the +371 d spectrum.

The sum of the galaxy and SN models fits the observed spectrum reasonably well at both epochs. The continuum slope is reproduced across a fairly wide wavelength range and the Balmer absorption lines and Balmer jump are reasonably well matched. The classic nebular emission lines seen in Type IIP SNe at these stages (see J12 and J14) are reproduced satisfactorily, but there are two major and important discrepancies. The SN model overproduces the calcium lines ([Ca II] $\lambda\lambda 7291, 7323$ and the Ca II NIR triplet) and underproduces [Ni II] $\lambda 7378$. We investigate the formation of these lines further in Section 5.1.2.

5.1.1 Oxygen lines and progenitor mass

Using the method described in J12 to measure line luminosities, we measure the observed [O I] $\lambda\lambda 6300, 6364$ luminosity to $L = 2.8 \times 10^{38} \text{ erg s}^{-1}$ at 371 d and $L = 2.9 \times 10^{38} \text{ erg s}^{-1}$ at 402 d. The [O I] $\lambda 5577$ line is too weak and noisy for any meaningful luminosity measurements, which otherwise can provide a temperature constraint and a direct handle on the oxygen mass (see J14 for an example).

The J12/J14 model luminosities of [O I] $\lambda\lambda 6300, 6364$ at 371 d for $M_{\text{ZAMS}} = 12, 15, 19$, and $25 M_{\odot}$ are 4.0, 8.0, 20, and $21 \times 10^{38} \text{ erg s}^{-1}$ (for a ^{56}Ni mass of $0.062 M_{\odot}$). Assuming a

direct proportionality with the ^{56}Ni mass,⁸ these values for a ^{56}Ni mass of $0.030 M_{\odot}$ are 1.9, 3.9, 9.5, and $10 \times 10^{38} \text{ erg s}^{-1}$. The observed luminosity of $2.8 \times 10^{38} \text{ erg s}^{-1}$ thus corresponds to a progenitor zero-age main sequence (ZAMS) mass of $13\text{--}14 M_{\odot}$. At 402 d, the measured value of $2.9 \times 10^{38} \text{ erg s}^{-1}$ should be compared with the model values of 1.4, 2.9, 7.1, and $7.7 \times 10^{38} \text{ erg s}^{-1}$, giving $M_{\text{ZAMS}} = 15 M_{\odot}$ as the best fit. We conclude that the nebular oxygen lines suggest a progenitor mass $M_{\text{ZAMS}} = 13\text{--}15 M_{\odot}$.

5.1.2 Model lines in the 7100–7500 Å region

We now aim to analyse in more detail the spectral region between 7100 and 7500 Å, where the observed line strengths show the strongest discrepancies with the spectral model. Inspection of the models shows that lines from calcium, iron, and nickel are the only ones produced in any significant strength in this region.

Calcium lines. The [Ca II] $\lambda\lambda 7291, 7323$ and Ca II NIR lines are weak in the observed spectra of SN 2012ec, both in comparison to the models as well as to most observed Type IIP SNe at this epoch. In the Maguire et al. (2012) sample of nine Type IIP SNe, the [Ca II] $\lambda 7291, 7323$ /[Fe II] $\lambda 7155$ ratio at 370 d spanned an interval 4–20, whereas for SN 2012ec it is ~ 1 . Comparison with other emission lines gives a similar picture; the calcium lines are intrinsically weak in SN 2012ec.

As originally demonstrated by Li & McCray (1993) (see also Kozma & Fransson 1998), the calcium lines originate as cooling emission from the hydrogen zone. We find the same conclusion with our models here; over 2/3 of the [Ca II] $\lambda\lambda 7291, 7323$ luminosity, and over 9/10 of the Ca II NIR luminosity, originate from hydrogen gas between 200 and 500 d. The rest of the [Ca II] $\lambda\lambda 7291, 7323$ luminosity has contributions from the Fe/He, Si/S, and He/C zones, all giving about 10 per cent each. The typical fraction of the total SN luminosity emerging in these forbidden calcium lines are 5–10 per cent in the models.

Being formed primarily as cooling of hydrogen-rich gas, their strength depends on the mass of the hydrogen zone, the heating rate

⁸ An approximation we can validate by comparing the new $15 M_{\odot}$ model ($M(^{56}\text{Ni}) = 0.03 M_{\odot}$) with the old ($M(^{56}\text{Ni}) = 0.062 M_{\odot}$); the [O I] $\lambda\lambda 6300, 6364$ luminosity scales almost exactly with the ^{56}Ni mass.

per unit mass of this zone, and the fraction of cooling that is done by [Ca II]. Several different scenarios can thus lead to weak calcium lines, but some are more likely than others. A simple analysis is complicated by the fact that hydrogen zone material is present in many regions of the SN ejecta, both in the envelope and in the core (due to mixing).

If an unusually low hydrogen zone mass is responsible, we would expect to see unusually weak hydrogen lines as well. Both H α (Fig. 2) and Pa β –Pa δ (Fig. 7) are quite strong and reasonably well reproduced by the model. Pa α looks weak but is compromised by its location in the middle of the telluric band. An unusually low ejecta mass was also not indicated by hydrodynamical modelling by Barbarino et al. (2014).

The heating rate per unit mass depends on the ^{56}Ni mass and the mixing of ^{56}Ni and hydrogen material. As the line luminosities to first-order scale with the ^{56}Ni mass, the unusually low ratio of calcium lines to other lines in SN 2012ec cannot easily be understood as due to a low ^{56}Ni mass (which is also measured to be quite typical, 0.03 $M_{\odot}$). An unusually weak mixing of ^{56}Ni gas with hydrogen gas could possibly reduce the gamma-ray deposition in the H zone, but then again the hydrogen lines would be expected to decrease as well.

A low fraction of cooling done by Ca II could have three fundamental causes; a low calcium abundance compared to other metals, a low fraction of calcium in the Ca II state, or a temperature that does not favour Ca II emission. Starting with the last possibility, the models show that [Ca II] $\lambda\lambda 7291, 7323$ is a strong coolant over the whole nebular evolution through a broad range of temperatures (J14). This suggests that an unusually low temperature is not the main cause of the weak lines. The first possibility, an unusually low abundance of calcium relative to other metals, is also not a very plausible explanation as it would have to be lower by a factor of 3–4 which seems unlikely. A peculiarly low calcium abundance relative to the other elements is not a known feature of chemical abundances in star-forming disc galaxies, and as the star-forming clouds in spirals will have been enriched in metals by a great number of stars and SNe any strong deviations from IMF-weighted ratios is difficult to conceive of. We note that the metallicity itself of the progenitor will not affect the hydrogen zone cooling line strengths significantly. This is because cooling time-scales are short so thermal equilibrium is established and all heating is instantaneously reemitted as cooling. Line cooling is the only efficient channel so lower metallicity will just lead to somewhat higher temperature so that the collisional pumping rates increase to compensate for the lower amounts of metals.

The second possibility, an unusual ionization balance of calcium, looks like a more plausible solution, because the ionization equilibrium in the models is close to switching between Ca II and Ca III as the dominant ion. At 370 d the Ca II fraction is 25 per cent in the innermost H zone (75 per cent Ca III) and increases slowly outwards. Thus, the Ca II fraction is sensitive to small changes in physical conditions such as density and ionizing radiation field. It is quite plausible that these could vary with a factor of a few between different SNe. Dependency of the calcium ionization equilibrium on ejecta structure would be an interesting topic for further study. Other efficient coolants of the hydrogen gas around 400 d are Mg II $\lambda\lambda 2795, 2802$, Fe II (several UV, optical, and NIR lines), and [O I] $\lambda\lambda 6300, 6364$. Hydrogen and helium usually heat the gas rather than cool it as recombinations and non-thermal excitations populate metastable (intrinsically metastable as well as effectively metastable through optical depth) states that are then collisionally de-excited.

We finally mention the possibility that the line at 7378 Å (which the model identifies as [Ni II] $\lambda 7378$) is actually a redshifted (3000 km s $^{-1}$) [Ca II] $\lambda\lambda 7291, 7323$. Several aspects safely rule out this option; the [Ca II] $\lambda 7291, 7323$ doublet is clearly seen close to its rest wavelength (it is just weak); no other lines in the spectra (including the H lines, which come from the same zone) show any detectable redshifts; the redshift would have to give exact wavelength coincidence with [Ni II] $\lambda 7378$ (which is the second strongest line in the model in this region); the Ca II NIR lines are as much weakened as the [Ca II] $\lambda\lambda 7291, 7323$ lines and show no redshift; and finally a strong [Ni II] 1.939 μm line matches the theoretical prediction following a [Ni II] $\lambda 7378$ identification (Section 5.2).

Iron lines. In the J14 models at 370 d, there are two distinct iron lines in the 7100–7500 Å region; [Fe II] $\lambda 7155$ and [Fe II] $\lambda 7453$. These lines arise from the same upper level ($3d^7a^2G_{9/2}$) with Einstein A coefficients $A_{7155} = 0.146 \text{ s}^{-1}$ and $A_{7453} = 0.0477 \text{ s}^{-1}$, respectively. The lines are optically thin ($\tau_{7155} = 0.08$ and $\tau_{7453} = 0.02$), and so their luminosity ratio is $L_{7453}/L_{7155} = A_{7453}h\nu_{7453}/A_{7155}h\nu_{7155} = 0.31$.

There is also some emissivity in [Fe II] $\lambda 7172$ (24 per cent of [Fe II] $\lambda 7155$) and [Fe II] $\lambda 7388$ (19 per cent of [Fe II] $\lambda 7155$). These lines arise from $3d^7a^2G_{7/2}$, 0.07 eV above the $3d^7a^2G_{9/2}$ state, and have $A_{7172} = 0.0551 \text{ s}^{-1}$ and $A_{7388} = 0.0421 \text{ s}^{-1}$; for any temperature they will thus be significantly weaker than [Fe II] $\lambda 7155$. Both $3d^7a^2G_{9/2}$ and $3d^7a^2G_{7/2}$ are close to local thermodynamic equilibrium (LTE), with departure coefficients 0.83 and 0.85, respectively.

Nickel lines. The strongest nickel lines in this spectral region are [Ni II] $\lambda 7378$ and [Ni II] $\lambda 7412$. These lines come from $4s^2F_{7/2}$ and $4s^2F_{5/2}$ and have $A_{7378} = 0.23 \text{ s}^{-1}$ and $A_{7412} = 0.18 \text{ s}^{-1}$. The lines are optically thin ($\tau_{7378} = 0.04$ and $\tau_{7412} = 0.01$), and the parent states are close to LTE, with departure coefficients 0.94 and 0.92, respectively. The model ratio $L_{7412}/L_{7378} = 0.31$ at 370 d is therefore close to the theoretical value assuming LTE and optically thin conditions.

Other lines. Three additional lines are discernible in the models: He I $\lambda 7281$, Fe I $\lambda 7207$, and [Ni I] $\lambda 7393$. These lines are all quite weak, the [Ni I] $\lambda 7393$ line is for instance more than 10 times weaker than [Ni II] $\lambda 7378$. We can conclude that the spectral region 7100–7500 Å is dominated exclusively by the components [Ca II] $\lambda\lambda 7291, 7323$, [Fe II] $\lambda 7155$, [Fe II] $\lambda 7172$, [Fe II] $\lambda 7388$, [Fe II] $\lambda 7453$, [Ni II] $\lambda 7378$, and [Ni II] $\lambda 7412$.

5.1.3 Gaussian fits to the 7100–7500 Å region

As shown in Fig. 2, the SN nebular model does not reproduce well the relative strengths of calcium, iron, and nickel lines in this spectral region. The [Ni II] $\lambda 7378$ line is particularly interesting because it provides a rare opportunity to determine the mass of stable nickel produced in the explosion. Instead of a full forward modelling approach, we in this section measure the observed line luminosities and apply analytical line formation equations to determine the element abundance ratios.

The model analysis in Section 5.1.2 showed that there are seven dominant line transitions present in the 7100–7500 Å region, with constraints on some of the line ratios from the same species. We construct a fit to this spectral region using Gaussian components of these lines ([Ca II] $\lambda\lambda 7291, 7323$, [Fe II] $\lambda 7155$, [Fe II] $\lambda 7172$, [Fe II] $\lambda 7388$, [Fe II] $\lambda 7453$, [Ni II] $\lambda 7378$, and [Ni II] $\lambda 7412$). We force the relative luminosities of lines from a given element to have the same ratios as in the model, so the iron lines are constrained by $L_{7453} = 0.31L_{7155}$, $L_{7172} = 0.24L_{7155}$, $L_{7388} = 0.19L_{7155}$, and

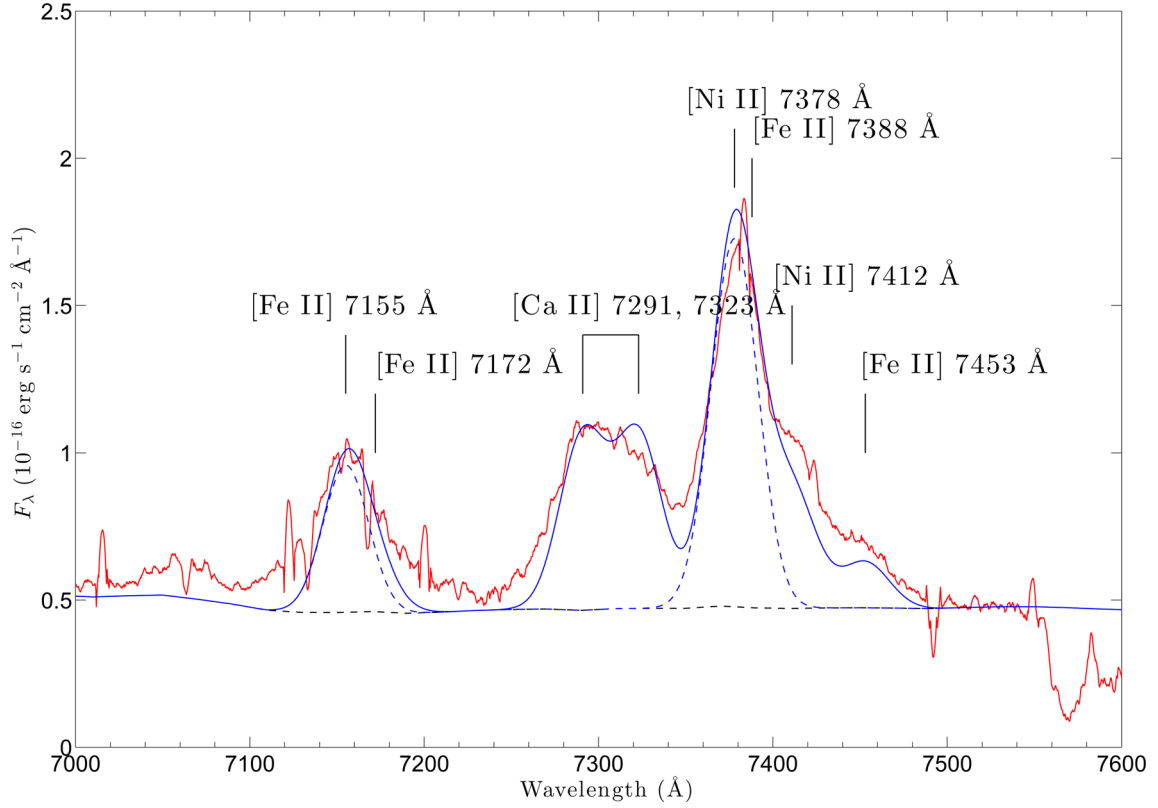


Figure 4. X-shooter spectrum between 7000 and 7600 Å (dereddened and redshift corrected, red), and the Gaussian fit described in Section 5.1.3 (blue). The components of the two lines that we use to derive the Ni/Fe ratio ([Fe II] λ 7155 and [Ni II] λ 7378) are plotted with dashed lines. Note that the [Fe II] λ 7172 and [Fe II] λ 7388 components are constrained to be weak from atomic structure arguments. Hence the emission lines at 7155 and 7378 Å are dominated by [Fe II] λ 7155 and [Ni II] λ 7378, respectively.

the nickel lines are constrained by $L_{7412} = 0.31L_{7378}$. We also use a single line width for all lines, the full width at half-maximum (FWHM) velocity ΔV . The free parameters are then $L_{7291, 7323}$, L_{7155} , L_{7378} , and ΔV . We also add the galactic model component described in Section 5.1 to the fit. As shown in Fig. 4, a good fit is obtained for $L_{7291, 7323} = 1.4 \times 10^{38}$ erg s $^{-1}$, $L_{7155} = 5.8 \times 10^{37}$ erg s $^{-1}$, $L_{7378} = 1.5 \times 10^{38}$ erg s $^{-1}$, and $\Delta V = 1300$ km s $^{-1}$. From this we determine a ratio $L_{7378}/L_{7155} = 2.6$.

5.1.4 The Ni II/Fe II ratio

The identification of distinct [Fe II] λ 7155 and [Ni II] λ 7378 lines, and their measured luminosities, can constrain the iron and nickel content in SN 2012ec with some analytic treatment. Assuming LTE and optically thin emission (which are good approximations according to the model, Section 5.1.2), the emissivity ratio of [Ni II] λ 7378 to [Fe II] λ 7155 is

$$\frac{L_{7378}}{L_{7155}} = \frac{n_{\text{Ni II}} g_{4s2SF7/2}^{\text{Ni II}} \exp\left(-\frac{1.68 \text{ eV}}{kT}\right) Z_{\text{Ni II}}(T)^{-1} A_{7378} h \nu_{7378}}{n_{\text{Fe II}} g_{3d7a2G9/2}^{\text{Fe II}} \exp\left(-\frac{1.96 \text{ eV}}{kT}\right) Z_{\text{Fe II}}(T)^{-1} A_{7155} h \nu_{7155}}, \quad (1)$$

where $n_{\text{Ni II}}$ and $n_{\text{Fe II}}$ are the number densities of Ni II and Fe II, g are statistical weights, and Z are partition functions. Computation of the partition functions $Z_{\text{Ni II}}(T)$ and $Z_{\text{Fe II}}(T)$ shows their ratio to vary little with temperature, staying between 0.24 and 0.26 over the temperature range 2000–6000 K, so we can with good accuracy

take $Z_{\text{Ni II}}/Z_{\text{Fe II}} = 0.25$. Using the atomic constants $g_{4s2SF7/2}^{\text{Ni II}} = 8$, $g_{3d7a2G9/2}^{\text{Fe II}} = 10$, $A_{7378} = 0.23$ s $^{-1}$, and $A_{7155} = 0.146$ s $^{-1}$, we then get

$$\frac{L_{7378}}{L_{7155}} = 4.9 \left(\frac{n_{\text{Ni II}}}{n_{\text{Fe II}}} \right) \exp\left(\frac{0.28 \text{ eV}}{kT}\right). \quad (2)$$

This line ratio is a powerful diagnostic of the Ni II to Fe II ratio and in turn the Ni to Fe ratio. It is relatively insensitive to temperature due to the similar excitation energies of the upper levels. It is also quite insensitive to density because the transitions have similar critical densities; deviation from LTE has therefore a similar impact on both lines. Furthermore, since iron and nickel have similar ionization potentials, they have similar ionization balances and the Ni II/Fe II ratio is likely to be close to the Ni/Fe ratio. Our models confirm this, showing that the Ni II/Fe II ratio deviates from the Ni/Fe ratio by less than 5 percent at 370 d.

Using the measured line ratio of $L_{7378}/L_{7155} = 2.6$, we can use equation (2) to compute the Ni II/Fe II ratio as function of temperature. This relationship is plotted in Fig. 5 together with curves for L_{7378}/L_{7155} ratios higher and lower by 30 percent than 2.6 (our estimated error). The temperature constraints from the [Fe II] λ 7155 luminosity ($2950 < T < 3350$ K, see below) give a (number) abundance ratio of Ni II/Fe II = 0.19 ± 0.07 , which we also then expect to be an accurate estimator of the Ni/Fe ratio. The solar abundance ratio is 0.06 (see Section 6), and in the $M_{\text{ZAMS}} = 15 M_{\odot}$ model of Woosley & Heger (2007), the Ni/Fe ratio is 0.04. Thus, the reason that the SN model significantly underproduces the [Ni II] λ 7378 line

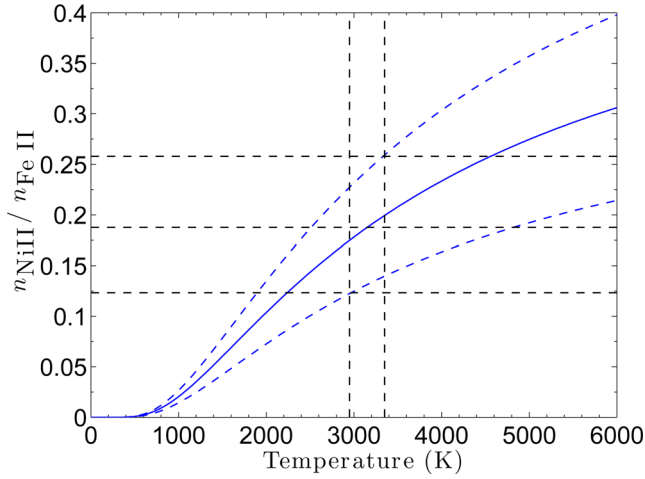


Figure 5. Derived Ni II/Fe II ratio as function of temperature (solid blue line shows the best estimate and the dashed blue lines give the error range). The temperature constraint from the [Fe II] $\lambda 7155$ luminosity ($2950 < T < 3350$ K) gives a ratio 0.19 ± 0.07 .

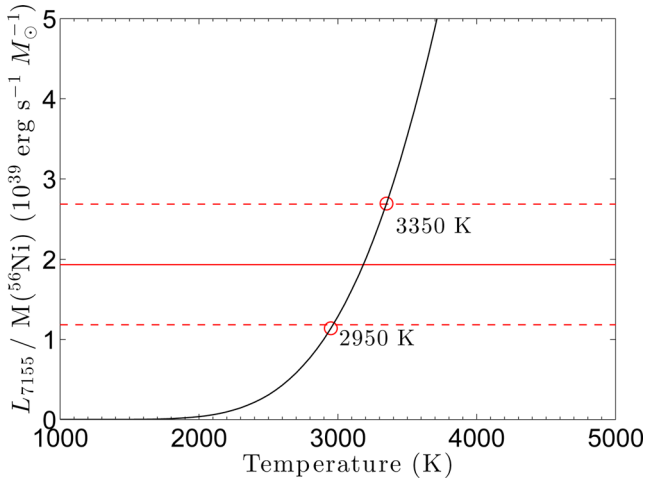


Figure 6. The measured value of $L_{7155}/M(^{56}\text{Ni})$ (red solid, errors as red dashed) compared to the theoretical value for LTE and optically thin emission as function of temperature (black line). To reproduce the observed value, the temperature is constrained to $2950 < T < 3350$ K.

in SN 2012ec is an abundance of stable nickel a factor of ~ 5 lower than needed.

Constraints on temperature. By 370 d all ^{56}Ni has decayed and over 95 per cent of the daughter isotope ^{56}Co has decayed to ^{56}Fe . As this channel dominates the iron production, the iron mass is constrained by the measurement of the initial ^{56}Ni mass from the early nebular phase (Section 4.1). Knowing the mass ($0.03 \pm 0.01 M_{\odot}$), we can use the LTE expression for the [Fe II] $\lambda 7155$ luminosity to constrain the temperature. This assumes that most of the iron is in the form of Fe II. The fraction is 0.9 in our models at 370 d so this is likely a good approximation. It also assumes LTE. The departure coefficient is close to unity in the model (0.83), also validating this approximation. Fig. 6 shows the measured value of $L_{7155}/M(^{56}\text{Ni})$ compared to the theoretical value as a function of temperature, which is given by

$$\frac{L_{7155}}{M(^{56}\text{Ni})} = \frac{A_{7155} h \nu_{7155} g_{\text{Fe II}}^{\text{Fe II}}}{56 m_u Z_{\text{Fe II}}(T)} \exp\left(\frac{-1.96 \text{ eV}}{kT}\right), \quad (3)$$

where m_u is the atomic mass unit. The partition function is well described by $Z_{\text{Fe II}}(T) = 15 + 0.006T$ between $T = 1000$ and 5000 K. The figure shows that the iron-zone temperature is constrained to $2950 < T < 3350$ K. This is in good agreement with the computed model temperature, which is $T = 3180$ K.

Contamination by primordial Fe and Ni. In J12 and Maguire et al. (2012), it was demonstrated that the total emission of primordial metals in the hydrogen zone is often comparable to the emission by newly synthesized metals in the core. We must therefore ask whether such emission may contribute to the [Fe II] $\lambda 7155$ and [Ni II] $\lambda 7378$ lines here.

At 370 d, the emission from primordial iron is about 40 per cent of the total [Fe II] $\lambda 7155$ emission. This contribution can be broken down into two parts: a contribution by hydrogen mixed into the core and having a similar velocity distribution as the synthesized iron, and the envelope part at higher velocities. These contributions give similar amounts of emission, about 20 per cent of the total each. The envelope component gives, however, a broad and flat-topped contribution to the line profile, which blends into the quasi-continuum as well as the [Ca II] $\lambda\lambda 7291, 7323$ doublet. Therefore, it does not become included in the luminosity we extract by fitting low-velocity Gaussians in Section 5.1.3. On the other hand, the fit will contain the contribution from the in-mixed hydrogen zone, which is then of order 20 per cent/60 per cent $\sim 1/3$ as strong as the synthesized iron component.

Also [Ni II] $\lambda 7378$ has a similar contribution by primordial nickel (~ 40 per cent). A solar abundance of Ni to Fe (as is presumably present in the hydrogen zone) gives however a primordial [Ni II] $\lambda 7378$ /[Fe II] $\lambda 7155$ ratio of only 0.7. The contribution by primordial nickel to the measured [Ni II] $\lambda 7378$ line luminosity (which is several times higher than the [Fe II] $\lambda 7155$ luminosity) must therefore be small (~ 10 per cent). The net effect of primordial contaminations is therefore that the measured Ni/Fe ratio may somewhat underestimate the Ni/Fe ratio in the iron zone, but not by more than about $\sim 1/3$.

5.2 Near-infrared

Fig. 7 shows the observed NIR spectra at 185 d (top) and 371 d (bottom), compared with the galaxy model plus the $15 M_{\odot}$ model of J14, scaled down with a factor of 0.03/0.062. The X-shooter spectrum has been smoothed to reduce noise, here with 61 points. The model provides identifications of the major emission lines, which are Pa β , Pa γ , Pa δ , He I 1.083 μm , O I 1.129 + 1.130 μm , O I 1.316 μm , and several blends of C I, Mg I, Si I, S II, Fe I, Fe II, and Co II lines.

Of particular interest is the emission line at 1.94 μm . In the model there are only two lines with significant emission at this wavelength: [Ni II] 1.939 μm and Br δ 1.944 μm . The line is observed to be much stronger than Br γ 2.166 μm , and is as narrow (FWHM $\approx 1300 \text{ km s}^{-1}$) as the iron-group lines between 7100 and 7500 $\text{\AA}$ (Section 5.1.3). This feature was broader and weaker in the NIR spectrum at +306 d of SN 2012aw (J14), and the model matched the line strength well. Given the quantitative difference, we suggest that the feature is dominated by [Ni II] 1.939 μm in SN 2012ec. With the caveat that the line is at the edge of the telluric band, we measure its luminosity to $L_{1.939 \mu\text{m}} = 2.5 \times 10^{37} \text{ erg s}^{-1}$. [Ni II] 1.939 μm arises from the same upper level as [Ni II] $\lambda 7378$. Their optically thin line ratio is

$$\frac{L_{1.939 \mu\text{m}}}{L_{7378}} = \frac{A_{1.939} h \nu_{1.939}}{A_{7378} h \nu_{7378}} = 0.14, \quad (4)$$

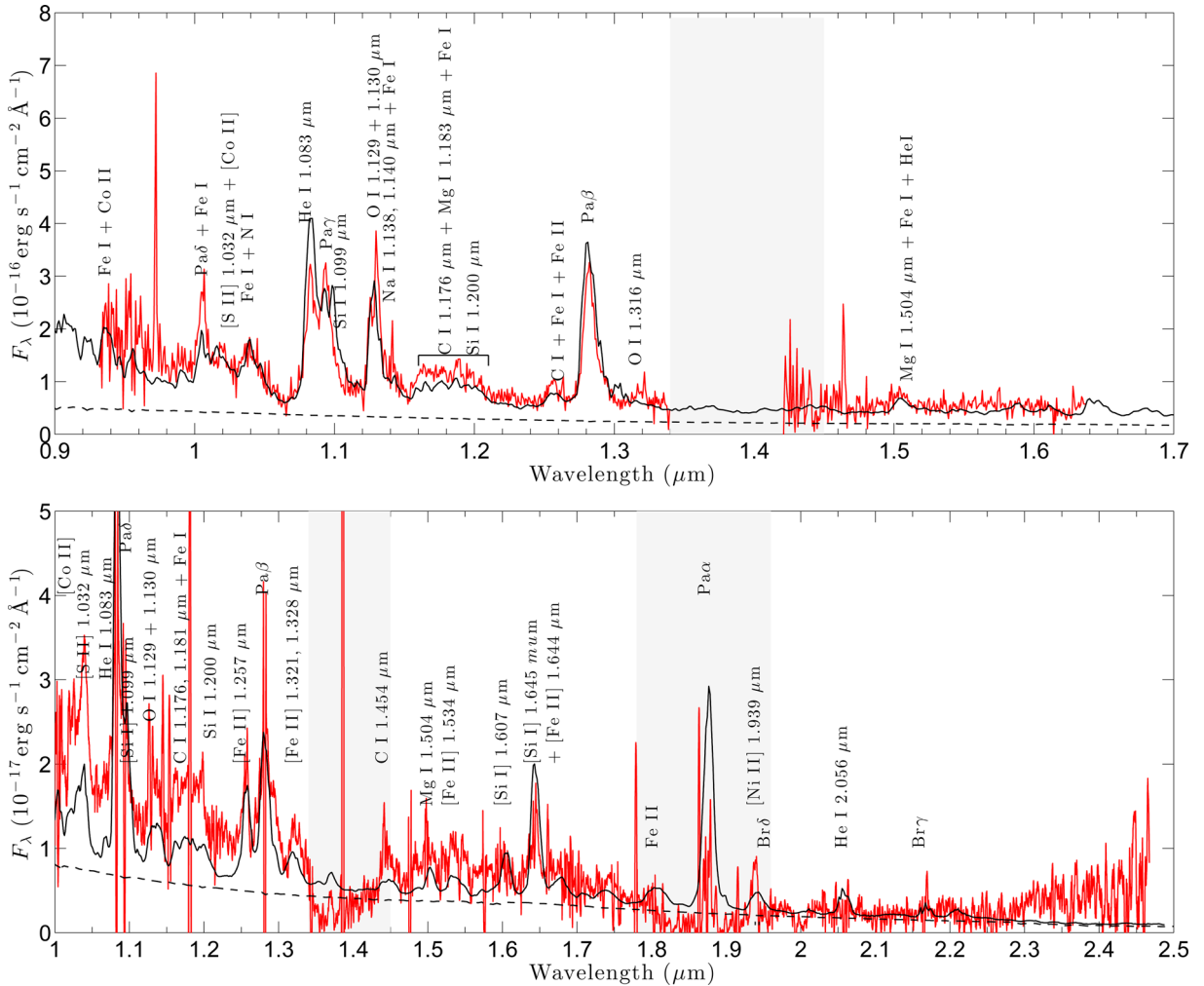


Figure 7. Top: SOFI NIR spectrum of SN 2012ec (dereddened and redshift corrected) at +185 d (red), and in black solid the sum of the galaxy model and the $M_{ZAMS} = 15 M_{\odot}$ SN model of J14 at 212 d scaled with 0.03/0.062 and a factor $\exp(-27/111.4)$. The galaxy model is shown as dashed black line. Line identifications are marked. Bottom: X-shooter NIR spectrum (dereddened and redshift corrected) at +371 d (red), and in black solid the sum of the galaxy model and the $M_{ZAMS} = 15 M_{\odot}$ model with a ^{56}Ni mass of $0.03 M_{\odot}$ computed here. The galaxy model is shown as a dashed black line. Telluric absorption bands are marked.

where we have used $A_{1.939} = 0.087 \text{ s}^{-1}$ and $A_{7378} = 0.23 \text{ s}^{-1}$. In SN 2012ec the measured ratio is $L_{1.939 \mu\text{m}}/L_{7378} = 0.17$, in excellent agreement. That the $[\text{Ni II}] 1.939 \mu\text{m}$ line is close to the expected strength relative to $[\text{Ni II}] \lambda 7378$, and is again much stronger than in the model with $\text{Ni/Fe} = 0.04$ (the model line feature in Fig. 7 is dominated by $\text{Br}\delta$), serves as an independent confirmation that the mass of stable nickel is high in SN 2012ec.

6 DISCUSSION AND SUMMARY

The solar abundances of iron and nickel are well known, with close agreement between photospheric and meteoritic measurements. Lodders (2003) reports $\log(N_{\text{Fe}}/N_{\text{H}}) + 12 = 7.48 \pm 0.03$ (meteoritic) and $\log(N_{\text{Fe}}/N_{\text{H}}) + 12 = 7.45 \pm 0.08$ (photospheric) for iron, and $\log(N_{\text{Ni}}/N_{\text{H}}) + 12 = 6.22 \pm 0.03$ (meteoritic) and $\log(N_{\text{Ni}}/N_{\text{H}}) + 12 = 6.22 \pm 0.13$ (photospheric) for nickel. The corresponding number ratio is $\text{Ni/Fe} = 0.056$ with an error of about 10 per cent, and the mass ratio is about 5 per cent higher (the mean atomic weights are 58.7 for nickel and 55.8 for iron), giving a mass abundance ratio of 0.059. Asplund et al. (2009) report similar values. Hinkel et al. (2014) analysed metal abundance variations in the

solar neighbourhood, reporting $[\text{Ni/Fe}] = -0.15^{+0.15}_{-0.07}$ dex at solar Fe/H metallicity, with similar values at other metallicities between $[\text{Fe/H}] = -0.6$ and $+0.6$ dex. This corresponds to Ni/Fe ratios of 0.6–1 times the solar value; the spread in the Ni/Fe ratio throughout the local Universe is thus small.

Analysis of strong emission lines of $[\text{Ni II}] \lambda 7378$ and $[\text{Ni II}] 1.939 \mu\text{m}$ in SN 2012ec indicates a strongly super-solar Ni/Fe production ratio of 0.20 ± 0.07 (by mass), or $(\text{Ni/Fe})/(\text{Ni/Fe})_{\odot} = 3.4 \pm 1.2$. Estimates of the Ni/Fe ratio have been presented (to our knowledge) for four previous core-collapse SNe. For SN 1987A, Rank et al. (1988) reported a measurement of the $[\text{Ni II}] 6.636 \mu\text{m}$ luminosity at 262 d. Assuming LTE, optically thin conditions, and that most nickel is singly ionized, this luminosity gave a Ni mass of $2 \times 10^{-3} M_{\odot}$, which with $M(\text{Fe}) = 0.075 M_{\odot}$ from the ^{56}Ni measurement gives $\text{Ni/Fe} = 0.027$, less than half the solar value. The analysis is insensitive to temperature due to the low excitation potential for the first excited state giving the $[\text{Ni II}] 6.636 \mu\text{m}$ line. Wooden et al. (1993) extended the analysis to later times, finding a similar value. In the ejecta models of J12/J14, LTE is valid for the $[\text{Ni II}] 6.636 \mu\text{m}$ line throughout this period. Regarding ionization, we find about 80 per cent of the nickel being

singly ionized between 250 and 450 d, with most of the remainder being Ni III. The most questionable assumption in the Rank et al. (1988) and Wooden et al. (1993) derivation is that of optically thin emission; in our models the [Ni II] 6.636 μm line has an optical depth of 3.2 at 250 d and 1.5 at 450 d. This optical depth leads to an underestimate of the Ni mass by a factor of 2–3 if one uses the optically thin formula. Taking the optical depth into account, the Ni/Fe ratio in SN 1987A is around solar or somewhat higher.

J12 calculated the model luminosity of [Ni II] 6.636 μm at 350 d using the Woosley & Heger (2007) nucleosynthesis models, finding good agreement with the observed line in SN 2004et (Kotak et al. 2009). As mentioned above, however, a caveat is that the nickel line has an optical depth of a few at these epochs, so the model luminosity depends not only on the nickel mass but also on the volume of the nickel-containing gas. This could fortunately be constrained from fits of various optically thick lines (see figs 11 and 12 in J12), so the fit of the [Ni II] 6.636 μm line is still constraining for the nickel mass in SN 2004et, giving a Ni/Fe ratio around solar or somewhat lower.

The [Ni II] $\lambda 7378$ line is one of the strongest lines emitted by the Crab nebula, with the Ni/Fe ratio estimated to be 60–75 times the solar value (MacAlpine et al. 1989, 2007). The variation with spatial position is relatively small, and imaging shows the nickel and iron lines to arise from the same filaments (MacAlpine et al. 1989; Hester et al. 1990). This extreme value is of interest in relation to the proposed scenario of an electron capture SN (Nomoto et al. 1982), which have been shown to produce very large Ni/Fe ratios (Wanajo et al. 2009).

A detection of [Ni II] $\lambda 7378$ was reported also for the broad-lined Type Ic SN 2006aj (Maeda et al. 2007; Mazzali et al. 2007) which was associated with an X-ray flash (Mazzali et al. 2006). The [Ni II] identification was somewhat hampered by high noise levels combined with very large widths of the emission lines. The nebular spectra were modelled to estimate a mass of stable nickel of 0.02–0.05 $M_{\odot}$, which with $M(^{56}\text{Ni}) = 0.2 M_{\odot}$ from the light curve gives Ni/Fe = 0.10–0.25, similar to the value derived here for SN 2012ec.

We repeated the analysis of SN 2012ec for two other recently observed Type IIP SN: SN 2012A and SN 2012aw. The 7000–7600 $\text{\AA}$ spectral region of these SNe around +400 d (data from Tomasella et al. 2013; J14), and Gaussian fits are presented in Fig. 8. The [Ca II] $\lambda\lambda 7291, 7323$ lines are strong in both of these SNe, and [Ni II] $\lambda 7378$ is not distinct. With our Gaussian fits, we can still obtain estimates

Table 4. Measurements of the Ni/Fe ratio in core-collapse SNe.

SN	Ni/Fe (times solar)	Reference
Crab	60–75	MacAlpine et al. (1989, 2007)
SN 1987A	0.5–1.5	Rank et al. (1988), Wooden et al. (1993), this work
SN 2004et	~ 1	J12
SN 2006aj	2–5	Maeda et al. (2007), Mazzali et al. (2007)
SN 2012A	~ 0.5	This work
SN 2012aw	~ 1.5	This work
SN 2012ec	2.2–4.6	This work

for the nickel line luminosities. We obtain $L_{7378}/L_{7155} \approx 0.4$ for SN 2012A and $L_{7378}/L_{7155} \approx 1.0$ for SN 2012aw. With the same analysis as for SN 2012ec (using the distances, extinction values, and ^{56}Ni masses from Tomasella et al. 2013; J14), this translates to Ni/Fe ratios of ~ 0.5 times solar in SN 2012A and ~ 1.5 times solar in SN 2012aw. We warn, however, that the ^{56}Ni mass in SN 2012A is only $\sim 0.01 M_{\odot}$ (Tomasella et al. 2013) so the physical approximations used (validated by models with ^{56}Ni masses of 0.062 and 0.03 $M_{\odot}$) may not be as good at this low ^{56}Ni mass. The ^{56}Ni mass in SN 2012aw is around 0.06 $M_{\odot}$, so that problem is not present for this SN. A roughly solar Ni/Fe ratio in SN 2012aw is supported by the good fit to the [Ni II] 1.939 μm + Br δ blend with the standard models in J14.

A summary of all reported measurements (that we are aware of) is presented in Table 4. Of the seven measurements, three (the Crab, SN 2006aj, and SN 2012ec) show significantly supersolar production of Ni/Fe, which should put strong constraints on the progenitor structure and explosion dynamics of these SNe. The Ni/Fe production is for many types of explosions dominated by the $^{58}\text{Ni}/^{56}\text{Ni}$ production. This ratio in turn depends on the neutron excess of the fuel, as well as the thermodynamic conditions for the burning (Woosley, Arnett & Clayton 1973). Higher neutron excess favours the production of neutron-rich isotopes such as ^{58}Ni , as does high entropy explosions which lead to large abundances of neutrons and α -particles. In Jerkstrand et al. (in preparation) we investigate the type of explosive silicon burning that can produce a Ni/Fe ratio as high as in SN 2012ec.

We have also analysed the nebular [O I] $\lambda\lambda 6300, 6364$ lines, which match models with $M_{\text{ZAMS}} = 13\text{--}15 M_{\odot}$. The progenitor analysis by Maund et al. (2013) found $M_{\text{ZAMS}} = 14\text{--}22 M_{\odot}$, and

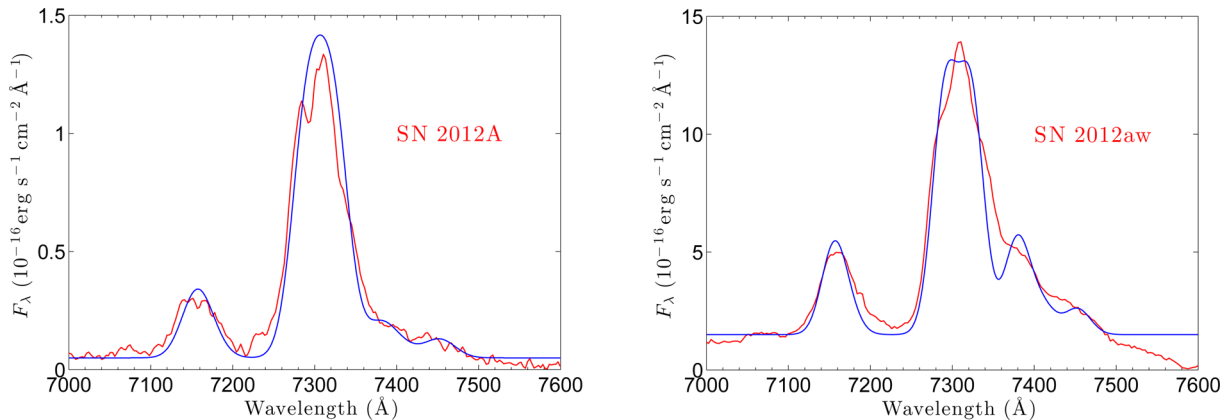


Figure 8. Spectra of SN 2012A at +393 d (left) and SN 2012aw at +369 d (right), and Gaussian fits following the same formalism as for the SN 2012ec fits. For SN 2012A the fit values are $\Delta V = 1700 \text{ km s}^{-1}$, $L_{7291, 7323} = 1.1 \times 10^{38} \text{ erg s}^{-1}$, $L_{7155} = 1.3 \times 10^{37} \text{ erg s}^{-1}$, and $L_{7378} = 5.3 \times 10^{36} \text{ erg s}^{-1}$. For SN 2012aw, the fit values are $\Delta V = 1500 \text{ km s}^{-1}$, $L_{7291, 7323} = 9.0 \times 10^{38} \text{ erg s}^{-1}$, $L_{7155} = 1.5 \times 10^{38} \text{ erg s}^{-1}$, and $L_{7378} = 1.6 \times 10^{38} \text{ erg s}^{-1}$.

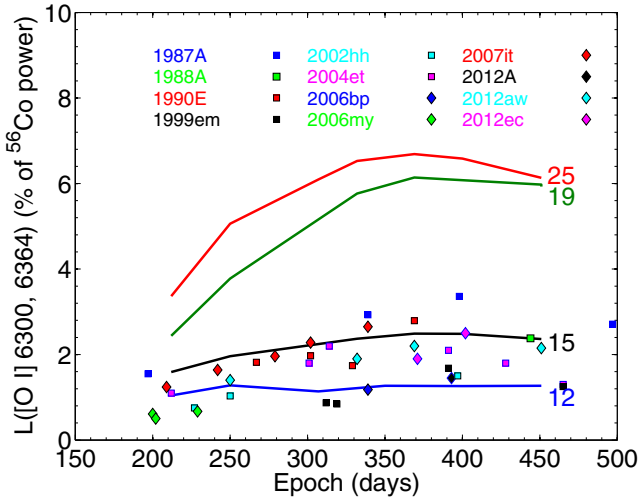


Figure 9. The [O I] $\lambda\lambda 6300, 6364$ luminosities (normalized to the ^{56}Co decay power) for 12 Type IIP SNe.

hydrodynamical modelling by Barbarino et al. (2014) favoured an ejecta mass of $\sim 13 M_{\odot}$ (to be compared with total ejecta masses for $M_{\text{ZAMS}} = 12, 15$, and $19 M_{\odot}$ stars which are 9, 11, and $14 M_{\odot}$ in the Woosley & Heger 2007 models). Within the errors, the methods seem to agree on an intermediate-mass progenitor. Fig. 9 shows the J14 model tracks for [O I] $\lambda\lambda 6300, 6364$ and the measured values for SN 2012ec and 11 other Type IIP SNe with high-quality nebular-phase spectra: SN 1987A⁹ (Phillips et al. 1990), SN 1988A (Turatto et al. 1993), SN 1990 (Benetti et al. 1994; Gomez & Lopez 2000), SN 1999em (Elmhamdi et al. 2003), SN 2002hh (Pozzo et al. 2006), SN 2004et (Sahu et al. 2006), SN 2006bp (Quimby et al. 2007), SN 2006my (Maguire et al. 2010), SN 2007it (Andrews et al. 2011), SN 2012A (Tomasella et al. 2013), and SN 2012aw (J14). Values for distances, extinctions, explosion epochs, and ^{56}Ni masses were taken as derived in each data paper, or in their references.¹⁰ All of these show O I lines suggesting $M_{\text{ZAMS}} \lesssim 17 M_{\odot}$ progenitors; a full statistical analysis with more data is underway.

SN 2012ec also shows an unusually weak [Ca II] $\lambda\lambda 7291, 7323$ feature (which indeed allowed the clear measurement of the [Ni II] $\lambda 7378$ line). According to the models, the [Ca II] $\lambda\lambda 7291, 7323$ doublet (as well as Ca II NIR triplet) originates mainly from the hydrogen zone in the SN. A plausible explanation for variation in this line between different SNe is varying ionization balance between Ca II and Ca III. However, other possibilities exist as well, and a more detailed study of Ca II line formation in Type IIP SNe would be desirable.

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⁹ SN 1987A is strictly speaking a Type Ipec rather than a Type IIP, but we include it here because its nebular properties are expected to be similar to those of Type IIP SNe.

¹⁰ For SN 1988A, SN 2006bp, and SN 2006my no ^{56}Ni masses have been presented, for these we used 0.030, 0.038 and $0.031 M_{\odot}$ based on calibration to SN 1987A flux levels (assuming distances of 17, 15, and 22 Mpc, respectively).

spectra on SN 2007it. We have made use of the SUSPECT and WISEREP (Yaron & Gal-Yam 2012) databases. This work is based in part on observations collected at the European Organisation for Astronomical Research in the Southern Hemisphere, Chile as part of PESSTO, (the Public ESO Spectroscopic Survey for Transient Objects) ESO program 188.D-3003. The research leading to these results has received funding from the European Research Council under the European Union's Seventh Framework Programme (FP7/2007-2013)/ERC Grant agreements no. [291222] (SJS) and no. [320360] and STFC grants ST/I001123/1 and ST/L000709/1 (SJS).

REFERENCES

- Alard C., 2000, A&AS, 144, 363
 Andrews J. E. et al., 2011, ApJ, 731, 47
 Asplund M., Grevesse N., Sauval A. J., Scott P., 2009, ARA&A, 47, 481
 Barbarino C. et al., 2014, preprint (arXiv:1410.8393)
 Benetti S., Cappellaro E., Turatto M., della Valle M., Mazzali P. A., Gouiffes C., 1994, A&A, 285, 147
 Bouchet P., Slezak E., Le Bertre T., Moneti A., Manfroid J., 1989, A&AS, 80, 379
 Bruenn S. W. et al., 2013, ApJ, 767, L6
 Chevalier R. A., 1976, ApJ, 207, 872
 Dall'Ora M. et al., 2014, ApJ, 787, 139
 Dessart L., Hillier D. J., Waldman R., Livne E., 2013, MNRAS, 433, 1745
 Elmhamdi A. et al., 2003, MNRAS, 338, 939
 Fabbri J. et al., 2011, MNRAS, 418, 1285
 Gomez G., Lopez R., 2000, AJ, 120, 367
 Hammer N. J., Janka H.-T., Müller E., 2010, ApJ, 714, 1371
 Hamuy M., Suntzeff N. B., Gonzalez R., Martin G., 1988, AJ, 95, 63
 Hester J. J., Graham J. R., Beichman C. A., Gautier T. N., III, 1990, ApJ, 357, 539
 Hinkel N. R., Timmes F. X., Young P. A., Pagano M. D., Turnbull M. C., 2014, AJ, 148, 54
 Inserra C. et al., 2012, MNRAS, 422, 1122
 Inserra C. et al., 2013, A&A, 555, A142
 Janka H.-T., 2012, Annu. Rev. Nucl. Part. Sci., 62, 407
 Jerkstrand A., Fransson C., Maguire K., Smartt S., Ergon M., Spyromilio J., 2012, A&A, 546, A28 (J12)
 Jerkstrand A., Smartt S. J., Fraser M., Fransson C., Sollerman J., Taddia F., Kotak R., 2014, MNRAS, 439, 3694 (J14)
 Kitaura F. S., Janka H.-T., Hillebrandt W., 2006, A&A, 450, 345
 Koribalski B. S. et al., 2004, AJ, 128, 16
 Kotak R. et al., 2009, ApJ, 704, 306
 Kozma C., Fransson C., 1998, ApJ, 497, 431
 Landolt A. U., 1992, AJ, 104, 340
 Leitherer C. et al., 1999, ApJS, 123, 3
 Li H., McCray R., 1993, ApJ, 405, 730
 Li W. et al., 2011, MNRAS, 412, 1441
 Lodders K., 2003, ApJ, 591, 1220
 MacAlpine G. M., McGaugh S. S., Mazzarella J. M., Uomoto A., 1989, ApJ, 342, 364
 MacAlpine G. M., Ecklund T. C., Lester W. R., Vanderveer S. J., Strolger L.-G., 2007, AJ, 133, 81
 Maeda K. et al., 2007, ApJ, 658, L5
 Maguire K. et al., 2010, MNRAS, 404, 981
 Maguire K. et al., 2012, MNRAS, 420, 3451
 Marek A., Janka H.-T., 2009, ApJ, 694, 664
 Maund J. R. et al., 2013, MNRAS, 431, L102
 Mazzali P. A. et al., 2006, Nature, 442, 1018
 Mazzali P. A. et al., 2007, ApJ, 661, 892
 Müller B., Janka H.-T., Heger A., 2012, ApJ, 761, 72
 Nomoto K., Sugimoto D., Sparks W. M., Fesen R. A., Gull T. R., Miyaji S., 1982, Nature, 299, 803

- Phillips M. M., Hamuy M., Heathcote S. R., Suntzeff N. B., Kirhakos S., 1990, *AJ*, 99, 1133
- Pozzo M. et al., 2006, *MNRAS*, 368, 1169
- Quimby R. M., Wheeler J. C., Höflich P., Akerlof C. W., Brown P. J., Rykoff E. S., 2007, *ApJ*, 666, 1093
- Ramya S., Sahu D. K., Prabhu T. P., 2007, *MNRAS*, 381, 511
- Rank D. M., Pinto P. A., Woosley S. E., Bregman J. D., Witteborn F. C., 1988, *Nature*, 331, 505
- Sahu D. K., Anupama G. C., Srividya S., Muneer S., 2006, *MNRAS*, 372, 1315
- Skrutskie M. F. et al., 2006, *AJ*, 131, 1163
- Smartt S. J., Eldridge J. J., Crockett R. M., Maund J. R., 2009, *MNRAS*, 395, 1409
- Smartt S. J. et al., 2013, *Messenger*, 154, 50
- Tomasella L. et al., 2013, *MNRAS*, 434, 1636
- Turatto M., Cappellaro E., Benetti S., Danziger I. J., 1993, *MNRAS*, 265, 471
- Utrobin V. P., Chugai N. N., 2009, *A&A*, 506, 829
- Wanajo S., Nomoto K., Janka H.-T., Kitaura F. S., Müller B., 2009, *ApJ*, 695, 208
- Wooden D. H., Rank D. M., Bregman J. D., Witteborn F. C., Tielens A. G. G. M., Cohen M., Pinto P. A., Axelrod T. S., 1993, *ApJS*, 88, 477
- Woosley S. E., Heger A., 2007, *Phys. Rep.*, 442, 269
- Woosley S. E., Arnett W. D., Clayton D. D., 1973, *ApJS*, 26, 231
- Yaron O., Gal-Yam A., 2012, *AJ*, 120, 367

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