



Publication Year	2019
Acceptance in OA	2020-12-29T16:21:10Z
Title	Interferometric observations of warm deuterated methanol in the inner regions of low-mass protostars
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Publisher's version (DOI)	10.1051/0004-6361/201936044
Handle	http://hdl.handle.net/20.500.12386/29304
Journal	ASTRONOMY & ASTROPHYSICS
Volume	632

Interferometric observations of warm deuterated methanol in the inner regions of low-mass protostars★

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Received 7 June 2019 / Accepted 19 September 2019

ABSTRACT

Methanol is a key species in astrochemistry because it is the most abundant organic molecule in the interstellar medium and is thought to be the mother molecule of many complex organic species. Estimating the deuteration of methanol around young protostars is of crucial importance because it highly depends on its formation mechanisms and the physical conditions during its moment of formation. We analyse several dozen transitions from deuterated methanol isotopologues coming from various existing observational datasets obtained with the IRAM-PdBI and ALMA sub-millimeter interferometers to estimate the methanol deuteration surrounding three low-mass protostars on Solar System scales. A population diagram analysis allows us to derive a $[\text{CH}_2\text{DOH}]/[\text{CH}_3\text{OH}]$ abundance ratio of 3–6% and a $[\text{CH}_3\text{OD}]/[\text{CH}_3\text{OH}]$ ratio of 0.4–1.6% in the warm inner (≤ 100 –200 AU) protostellar regions. These values are typically ten times lower than those derived with previous single-dish observations towards these sources, but they are one to two orders of magnitude higher than the methanol deuteration measured in massive hot cores. Dust temperature maps obtained from *Herschel* and *Planck* observations show that massive hot cores are located in warmer molecular clouds than low-mass sources, with temperature differences of ~ 10 K. The comparison of our measured values with the predictions of the gas-grain astrochemical model *GRAINOBLE* shows that such a temperature difference is sufficient to explain the different deuteration observed in low- to high-mass sources. This suggests that the physical conditions of the molecular cloud at the origin of the protostars mostly govern the present-day observed deuteration of methanol and therefore of more complex organic molecules. Finally, the methanol deuteration measured towards young solar-type protostars on Solar System scales seems to be higher by a factor of ~ 5 than the upper limit in methanol deuteration estimated in comet Hale-Bopp. If this result is confirmed by subsequent observations of other comets, it would imply that an important reprocessing of the organic material likely occurred in the solar nebula during the formation of the Solar System.

Key words. astrochemistry – molecular processes – ISM: abundances – ISM: molecules – submillimeter: ISM – stars: formation

1. Introduction

Low-mass protostars are known to be chemically rich with the detection of several dozen molecules (van Dishoeck et al. 1995; Cazaux et al. 2003; Herbst & van Dishoeck 2009; Caux et al. 2011; Caselli & Ceccarelli 2012; Jørgensen et al. 2016). A process of deuterium enrichment is also observed for several neutral species, in particular for molecules thought to be mostly formed in cold interstellar ices and then released in the warm gas surrounding protostars (Ceccarelli et al. 2014). So far, methanol and formaldehyde are the two molecules showing the highest deuterium fractionations, up to 50%, or 5 orders of magnitude higher than the cosmic D/H elemental abundance ratio of 1.6×10^{-5} (Linsky 2003; Ceccarelli et al. 1998; Parise et al. 2006). Other species that are thought to be also mostly formed in interstellar ices, such as water, show lower deuteration (Coutens et al. 2012;

Taquet et al. 2013a; Persson et al. 2014). These differences have been invoked to reflect different moments or different mechanisms of formation (Cazaux et al. 2011; Taquet et al. 2013a, 2014; Ceccarelli et al. 2014).

Gas-phase chemistry is known to be inefficient for the formation of gaseous methanol (Geppert et al. 2006; Garrod et al. 2006). Instead, methanol has been found to be efficiently formed at the surface of cold interstellar grains in dark cloud conditions through the hydrogenation of CO and H_2CO (Watanabe & Kouchi 2002; Rimola et al. 2014). During the collapse of the protostellar envelope, solid methanol is released into the gas phase when the temperatures exceed ~ 100 K. The deuteration of methanol is thought to remain constant after its evaporation in the warm-gas phase of the inner protostellar regions because gas-phase chemistry is likely too slow to significantly alter the ratios during the lifetime of the protostar (Charnley et al. 1997; Osamura et al. 2004). As a consequence, the deuteration observed around low-mass protostars would be a fossil of its formation in the precursor dense cloud where interstellar

* The reduced datacubes are only available at the CDS via anonymous ftp to cdsarc.u-strasbg.fr (130.79.128.5) or via <http://cdsarc.u-strasbg.fr/viz-bin/cat/J/A+A/632/A19>

ices formed. Taquet et al. (2012a, 2013a) showed that like other species, methanol deuteration is highly sensitive to the temperature and the density during ice formation. A deuteration higher than 1% is favoured by a methanol formation in cold ($T = 10$ K) and dense ($n_{\text{H}} \geq 10^5 \text{ cm}^{-3}$) molecular cloud conditions.

Most measurements of the methanol deuteration carried out so far have been performed using single-dish millimeter facilities with beams larger than $10''$ (or ~ 2000 AU at a typical distance of 200 pc). These large beams encompass the different components of the protostellar system: the small hot core where all the icy content is thermally evaporated, but also the cold external envelope and molecular outflows driven by the central source(s). The advent of modern millimeter interferometers, such as the IRAM Plateau de Bure Interferometer (IRAM-PdBI), now called the NOthern Extended Millimeter Array (NOEMA), and the Atacama Large Millimeter/submillimeter Array (ALMA) now allows astronomers to zoom towards the inner protostellar regions, in the so-called hot corinos, to measure the deuteration of the whole icy content released in the gas phase. New measurements towards low-mass protostars within different nearby molecular clouds suggest lower methanol deuterations than those originally measured a decade ago with single-dish telescopes. Bianchi et al. (2017) derived a $[\text{CH}_2\text{DOH}]/[\text{CH}_3\text{OH}]$ abundance ratio of 2% in HH212, a low-mass protostar located in the Orion molecular cloud, whilst Jørgensen et al. (2018) estimated a $[\text{CH}_2\text{DOH}]/[\text{CH}_3\text{OH}]$ ratio of 7% in Source B of IRAS 16293-2422, a low-mass protobinary system located in the Ophiuchus cloud. In addition to being lower than previous single-dish measurements of 20–60%, they also differ within each other by a factor of ~ 3 . This difference might be due to different properties of the precursor clouds. The Orion molecular cloud is the closest massive star-forming region and is known to be more active and slightly warmer than the more quiescent Ophiuchus cloud. However, this assumption remains to be confirmed for sources located in other clouds.

This article presents an analysis of different published observational datasets from the IRAM-PdBI towards NGC 1333-IRAS 2A and -IRAS 4A, two bright Class 0 low-mass protostars located in the nearby Perseus molecular cloud. We also re-analyse partially published data from the ALMA towards HH212 in order to measure the $[\text{CH}_2\text{DOH}]/[\text{CH}_3\text{OH}]$ and $[\text{CH}_3\text{OD}]/[\text{CH}_3\text{OH}]$ abundance ratios in their hot corinos. In this article, we aim to investigate whether the warm methanol deuteration observed in hot cores is regulated by the dust temperature of precursor dark clouds. To this aim, we compare our data with interferometric observations towards more massive protostellar systems and with the predictions of a modern astrochemical model.

2. Observations

2.1. Observational details

The two low-mass Class 0 protostars IRAS 2A and IRAS 4A located in the NGC 1333 cloud at 299 ± 14 pc (Zucker et al. 2018) were observed with the IRAM PdBI at 143, 165, and 225 GHz. Observations at 143 and 165 GHz were carried out on 2010 July 20, July 21, August 1, August 3, November 24, and 2011 March 10, in the C and D configurations of the array. Observations at 225 GHz were performed on 2011 November 27 and 28, on 2012 March 12, 15, 21, and 27, and on 2012 April, 2 in the B and C configurations of the array. A more detailed description of the observational setups is provided in Taquet et al. (2015) and Persson et al. (2014) for configurations B and C, respectively. The amplitude calibration uncertainty is estimated to be

$\sim 20\%$. The WIDEX correlator has been used at three frequency settings, providing a bandwidth of 3.6 GHz each with a spectral resolution of 1.95 MHz (4, 3.5, and 2.6 km s^{-1} at 145, 165, and 225 GHz, respectively). The data calibration and imaging were performed using the CLIC and MAPPING packages of the GILDAS software¹. Continuum images were produced by averaging line-free channels in the WIDEX correlator before the Fourier transformation of the data. The line spectral cubes were then obtained by subtracting the continuum visibilities from the whole (line+continuum) datacube, followed by natural weighted cleaning of individual channels. For IRAS 4A, the baseline has also been flattened by importing the data cubes into CLASS and subtracting a polynomial function of low order to each individual spectrum. Positions are given with respect to the continuum peaks of IRAS 2A and IRAS 4A1 located at $\alpha(\text{J2000}) = 03^{\text{h}} 28^{\text{m}} 55^{\text{s}}.57$, $\delta(\text{J2000}) = +31^\circ 14' 37''.22$ and $\alpha(\text{J2000}) = 03^{\text{h}} 29^{\text{m}} 10^{\text{s}}.52$, $\delta(\text{J2000}) = +31^\circ 13' 31''.06$, respectively.

The HH212-MM1 (hereafter HH212) protostar located in the Orion B cloud at about 400 pc (Kounkel et al. 2017) was observed in Band 7 with ALMA using 34 12 m antennas between June 15 and July 19, 2014, during the Cycle 1 phase and 44 12 m antennas between October 6 and November 26, 2016, during the Cycle 4 phase. The maximum baselines for the Cycle 1 and 4 observations are 650 m and 3 km, respectively. In Cycle 1, the spectral windows between 337.1–338.9 GHz and 348.4–350.7 GHz were observed using spectral channels of 977 kHz (or 0.87 km s^{-1}), subsequently smoothed to 1.0 km s^{-1} to increase sensitivity. Calibration was carried out following standard procedures, using quasars J0607-0834, J0541-0541, J0423-013, and Ganymede. For the Cycle 4 data, a single spectral window between 334.1–336.0 GHz with a 488 kHz (or 0.42 km s^{-1}) spectral resolution, successively smoothed to 1 km s^{-1} , was used. Calibration was carried out using quasars J0510+1800, J0552+0313, J0541-0211, and J0552-3627. Spectral line imaging was achieved using CASA², whilst data analysis was performed using the GILDAS package. Positions are given with respect to the MM1 protostar continuum peak located at $\alpha(\text{J2000}) = 05^{\text{h}} 43^{\text{m}} 51^{\text{s}}.41$, $\delta(\text{J2000}) = -01^\circ 02' 53''.17$. In addition to $^{13}\text{CH}_3\text{OH}$ and CH_2DOH images that have been presented by Bianchi et al. (2017), we also report the imaging of the $\text{CH}_3\text{OD}(6_2-6_{1+})$ transition at 335.1 GHz (see also Codella et al. 2019). The size of the synthesised beams and the rms noise are reported in Table 1.

2.2. Spectroscopic parameters

We here focus on the CH_2DOH , CH_3OD , and CHD_2OH deuterated isotopologues. We used the CH_2DOH JPL entry from Pearson et al. (2012). The new band strengths published by Pearson et al. (2012) can differ by a factor of a 1.5–2.0 from the values from Parise et al. (2002) that are used in most previous studies before 2012. With a vibrational correction factor of 1.15 at 160 K and 1.46 at 300 K based on the torsional data of Lauvergnat et al. (2009), the partition function is 1.3–2.0 higher than the partition function estimated by Parise et al. (2002). We re-analysed the CH_2DOH transitions observed towards the IRAS 16293-2422 low-mass protostellar system with the IRAM 30-m telescope by Parise et al. (2002) to investigate the effect of these new spectroscopic data on the estimation of the CH_2DOH column densities. We derived a CH_2DOH column density 2.2 lower than the one found by these authors. For CH_3OD , we prepared a catalogue entry based on the band

¹ <http://www.iram.fr/IRAMFR/GILDAS>

² <http://casa.nrao.edu>

Table 1. Parameters of the observational data.

NGC 1333-IRAS 2A (PdBI)			
Frequency	145 GHz	165 GHz	225 GHz
Beam size (")	2.1×1.7	2.3×1.7	1.2×1.0
Beam PA (°)	25	110	22
rms(spectrum) ^(a)	2.6	3.5	2.9
NGC 1333-IRAS 4A (PdBI)			
Frequency	145 GHz	165 GHz	225 GHz
Beam size (")	2.2×1.7	2.4×1.8	1.1×0.8
Beam PA (°)	25	114	14
rms(spectrum) ^(a)	3.3	4.0	4.8
HH212 (ALMA)			
	335 GHz	338 GHz	349 GHz
Beam size (")	0.15×0.12	0.52×0.34	0.41×0.33
Beam PA (°)	−88	−63	−63
rms(spectrum) ^(b)	1	5–6	5–6

Notes. ^(a)Units of mJy beam^{−1} channel^{−1} for a channel width of 1.95 MHz. ^(b)Units of mJy beam^{−1} channel^{−1} for a channel width of 0.87 km s^{−1} (335 GHz) and 1.0 km s^{−1} (at 338 and 349 GHz).

strengths published by Anderson et al. (1988) with frequencies updated by Duan et al. (2003). A self-consistent calculation of the CH₃OD partition function has not been performed so far because the CH₃OD spectrum has only been partially studied in the laboratory. Instead, we used two estimates for the CH₃OD partition function Z to derive the CH₃OD column densities. The first function follows the estimation by Parise (2004) within the rigid-rotor approximation, giving $Z = 1.42T^{3/2}$ (see also Ratajczak et al. 2011), as used in most previous studies. The second follows the estimation by Belloche et al. (2016), whose partition function values are scaled with those of the CH₃OH partition function (see also Jørgensen et al. 2018). At 150 K, the partition function estimated by Parise (2004) is four times lower than the one estimated by Jørgensen et al. (2018), resulting in large differences in the derived CH₃OD column densities. For CHD₂OH, we used the spectroscopic data and the partition function published by Parise et al. (2002). The CHD₂OH partition function is found to be similar to that of CH₂DOH published by the JPL database, with differences lower than 20% between 10 and 300 K.

We also targeted transitions from CH₂DCN. The CH₂DCN spectroscopic data were taken from the CDMS catalogue, based on Nguyen et al. (2013).

3. Results

3.1. Integrated maps

For IRAS 2A and IRAS 4A, the integrated emission maps of all methanol isotopologue transitions were obtained by integrating the flux over $V_{\text{sys}} \pm \Delta V$, where V_{sys} is the source systemic velocity and ΔV is the transition full width at half-maximum (FWHM) line width. In practice, we integrated the line emission over three channels because the spectral resolution of the WIDEX data is low. Figures 1 and 2 give an overview of the maps obtained for CH₂DOH, CH₃OD, and CHD₂OH in the three frequency settings. For all isotopologues, the emission is limited to the inner regions around the protostars within the

synthesised beams. Towards IRAS 4A, the molecular emission originates from the north-western source IRAS 4A2, although the continuum emission of the south-eastern source IRAS 4A1 is brighter. These maps are in good agreement with previous interferometric observations of these sources, which suggest that the emission of complex organic molecules originates from hot corinos whose sizes are smaller than 0''.5 (or 150 au; Jørgensen et al. 2005; Persson et al. 2012; Maury et al. 2014; Taquet et al. 2015; López-Sepulcre et al. 2017; De Simone et al. 2017).

For HH212, we produced the integrated emission maps of ¹³CH₃OH, CH₂DOH, and CH₃OD transitions that are not blended with other transitions by integrating the flux over the entire line width after an analysis of their spectra towards the continuum peak. The Cycle 1 emission is spatially unresolved with a beam of 0''.5 × 0''.3, corresponding to a size of 225 × 135 au. However, the analysis of the higher angular resolution of Cycle 4 dataset provides a beam-deconvolved FWHM size of 0''.18 × 0''.12 (81 × 54 au) for the ¹³CH₃OH emission and 0''.20 × 0''.10 (90 × 45 au) for CH₂DOH. As shown in Fig. 3, the emission originates from the hot-corino region and is confined to a rotating structure that extends at ±45 au from the equatorial plane and is elongated along and rotates around the jet axis. The actual interpretation of this emission is still debated and has been attributed to disc wind or a disc atmosphere from accretion shocks (Bianchi et al. 2017; Lee et al. 2017).

We measured the flux of all methanol isotopologue transitions integrated over an elliptical mask with similar size and position angle to the synthesised beam and with a centre at the molecular emission peak following the method described in Taquet et al. (2015). Tables A.1–A.6 summarise the properties and the measured flux of the targeted transitions towards the two sources.

3.2. Population diagrams

We used the so-called population diagram (PD) method described in Goldsmith et al. (1999) and Taquet et al. (2015) to investigate the effect of optical depth on the column densities of each level for the different methanol isotopologues. We performed a reduced chi-square (χ^2_{red}) minimization by running a grid of models covering a large parameter space in rotational temperature T_{rot} between 50 and 350 K, and total column density in the source N_{tot} between 10^{15} and 10^{20} cm^{−2}. We assumed the source sizes θ_s derived in Taquet et al. (2015) and Bianchi et al. (2017). A source size of 0''.36 and 0''.20 in IRAS 2A and IRAS 4A, respectively, was needed to reproduce the scatter of the population distribution of the low CH₃OH upper energy levels (see Taquet et al. 2015), and a source size of 0''.19 was derived by Bianchi et al. (2017) from the analysis of the spatially resolved ¹³CH₃OH Cycle 4 emission map. At local thermal equilibrium (LTE), the column density of every upper state N_{up} can be derived for each set of N_{tot} , T_{rot} , and source solid angle Ω_s . The best-fit model populations are plotted together with the observed populations of the levels in Figs. 4–6 and are marked by crosses. Table 2 summarises the parameters of the best-fit models and their associated uncertainties. The column densities of CH₃OH and CH₃CN have been derived in Taquet et al. (2015) by analysing the emission from 28 CH₃OH and 13 ¹³CH₃OH transitions assuming a ¹²C/¹³C elemental ratio of 70 and from six CH₃CN transitions detected in the 145 and 165 GHz frequency settings, respectively. Taquet et al. (2015) used the same PD method by considering the rotational temperature T_{rot} , the total column density in the source N_{tot} , and the source size θ_s as free parameters.

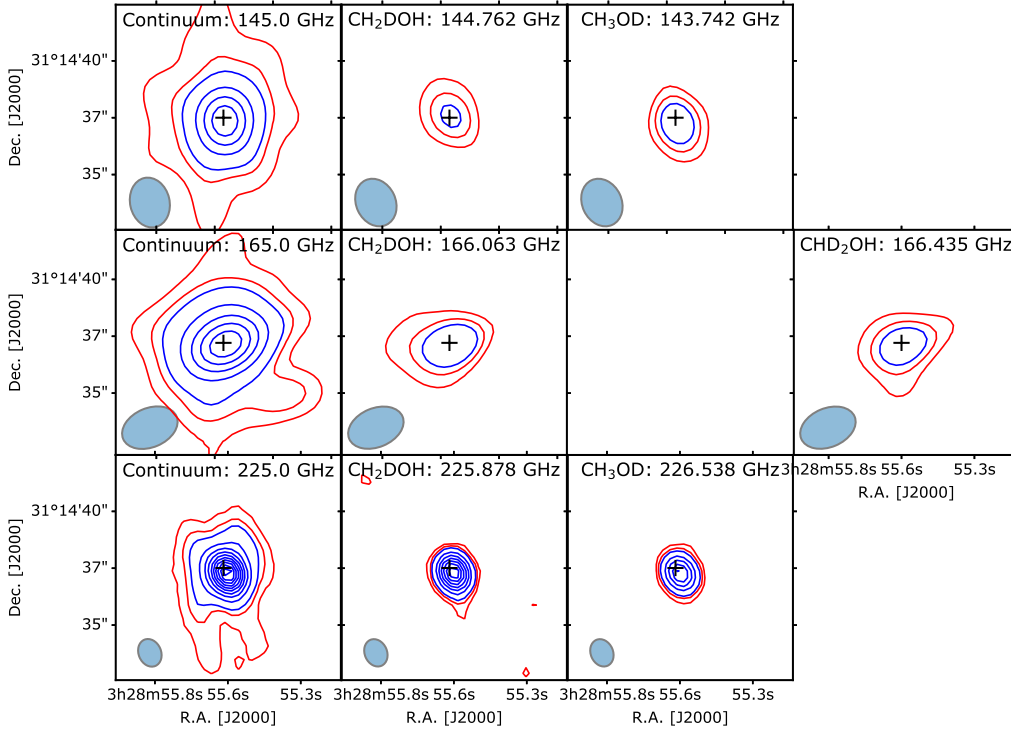


Fig. 1. Integrated continuum maps, CH_2DOH , CH_3OD , and CHD_2OH emission observed towards IRAS 2A for the three PdBI frequency settings (*top*: 145 GHz, *middle*: 165 GHz, and *bottom*: 225 GHz). Red contours show the 3σ and 6σ levels, whilst blue contours are in steps of 9σ . The synthesised beams are shown in the bottom left corner of each panel. The black cross depicts the position of the protostar.

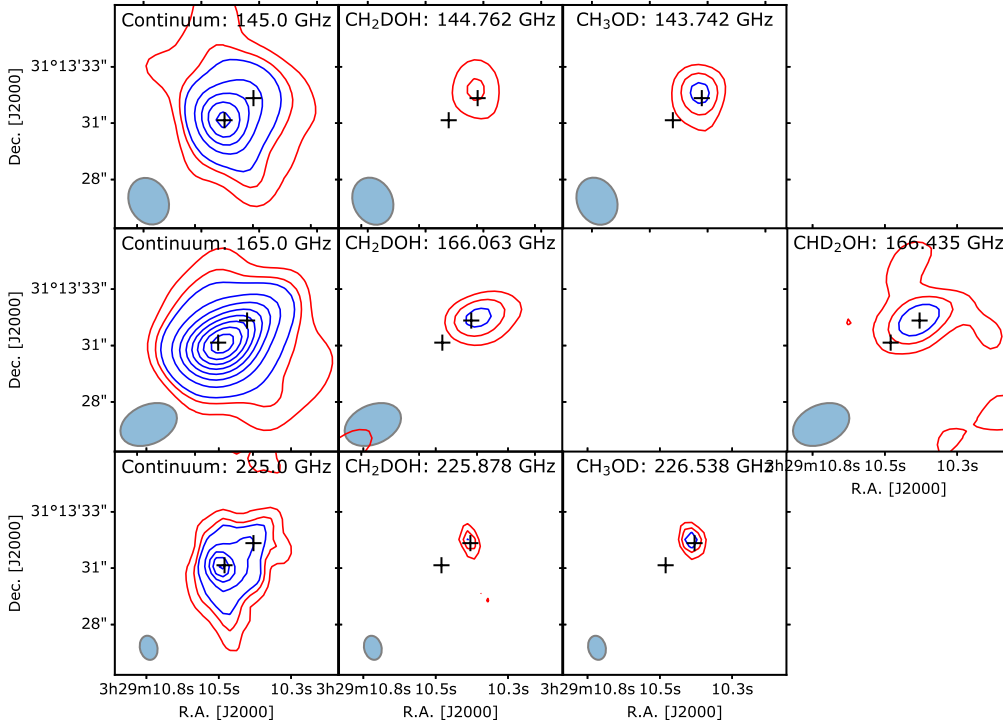


Fig. 2. Integrated continuum maps, CH_2DOH , CH_3OD , and CHD_2OH emission observed towards IRAS 4A for the three PdBI frequency settings (*top*: 145 GHz, *middle*: 165 GHz, and *bottom*: 225 GHz). Red contours show the 3σ and 6σ levels, whilst blue contours are in steps of 9σ . The synthesised beams are shown in the bottom left corner of each panel. The black crosses depict the position of the protostars.

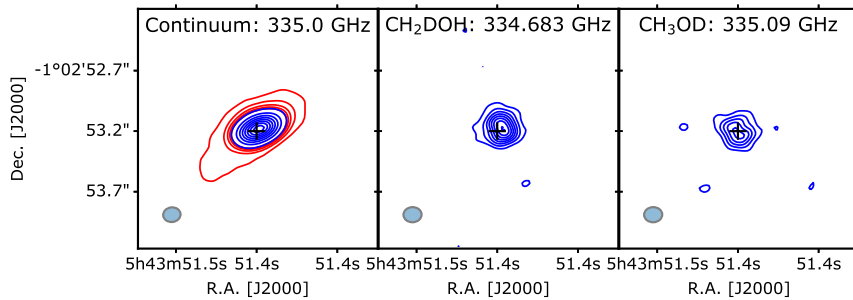


Fig. 3. Integrated continuum maps, CH_2DOH , and CH_3OD emission observed towards HH212 for the ALMA frequency setting at 335 GHz. For the continuum map, red contours show the 5 first 20σ levels and blue contours are in steps of 100σ . For the molecular maps, blue contours are in steps of 3σ . The synthesised beams are shown in the bottom left corner of each panel.

Table 2. Results from the PD analysis for the methanol and methyl cyanide deuterated isotopologues.

Molecule	N_{hc} (cm^{-2})	T_{rot} (K)	Source size ($''$)	$N(\text{XD})/N(\text{XH})$ ^(a) (%)	$N(\text{XD})/N(\text{XH})$ ^(b) (Single-dish, %)
IRAS 2A					
CH ₃ OH	$(5.0^{+2.9}_{-1.8})(+18)$	140^{+20}_{-20}	$0.36^{+0.04}_{-0.04}$ ^(c)	—	—
¹³ CH ₃ OH	$(7.1^{+4.2}_{-2.6})(+16)$	140^{+20}_{-20}	$0.36^{+0.04}_{-0.04}$ ^(c)	—	—
CH ₂ DOH	$(2.9^{+1.2}_{-0.6})(+17)$	166^{+58}_{-48}	0.36 ^(d)	$5.8^{+1.3}_{-1.5}$	62^{+71}_{-33}
CH ₃ OD (P04 ^(e))	$(7.9^{+2.1}_{-1.6})(+16)$	166 ^(f)	0.36 ^(d)	$1.6^{+0.6}_{-0.8}$	≤ 8
CH ₃ OD (J18 ^(e))	$(3.6^{+0.9}_{-0.4})(+17)$	166 ^(f)	0.36 ^(d)	$7.1^{+1.2}_{-1.6}$	≤ 8
CHD ₂ OH	$(2.2^{+0.6}_{-0.5})(+17)$	164^{+44}_{-32}	0.36 ^(d)	$4.4^{+0.9}_{-1.3}$	25^{+29}_{-14}
CH ₃ CN	$(2.0^{+1.2}_{-0.4})(+16)$	130^{+230}_{-40}	0.36 ^(c)	—	—
CH ₂ DCN	$(7.2^{+1.5}_{-0.9})(+14)$	130 ^(g)	0.36 ^(g)	$3.6^{+0.5}_{-1.2}$	—
IRAS 4A					
CH ₃ OH	$(1.6^{+0.6}_{-0.8})(+19)$	140^{+30}_{-30}	$0.20^{+0.08}_{-0.04}$ ^(c)	—	—
¹³ CH ₃ OH	$(2.3^{+1.3}_{-1.1})(+17)$	140^{+30}_{-30}	$0.20^{+0.08}_{-0.04}$ ^(c)	—	—
CH ₂ DOH	$(5.9^{+2.2}_{-1.8})(+17)$	152^{+62}_{-78}	0.20 ^(d)	$3.7^{+1.2}_{-0.9}$	65^{+30}_{-21}
CH ₃ OD (P04 ^(e))	$(1.1^{+0.5}_{-0.2})(+17)$	152 ^(f)	0.20 ^(d)	$0.7^{+0.5}_{-0.4}$	$4.7^{+2.9}_{-2.1}$
CH ₃ OD (J18 ^(e))	$(5.0^{+1.3}_{-1.0})(+17)$	152 ^(f)	0.20 ^(e)	$3.1^{+1.0}_{-0.8}$	$4.7^{+2.9}_{-2.1}$
CHD ₂ OH	$(3.3^{+0.9}_{-0.7})(+17)$	138^{+34}_{-38}	0.20 ^(d)	$2.1^{+0.8}_{-0.6}$	17^{+9}_{-6}
CH ₃ CN	$(6.3^{+3.6}_{-1.3})(+16)$	200^{+110}_{-40}	0.20 ^(c)	—	—
CH ₂ DCN	$(1.7^{+0.7}_{-0.7})(+15)$	200 ^(g)	0.20 ^(g)	$2.7^{+0.7}_{-1.1}$	—
HH212					
¹³ CH ₃ OH	$(3.2^{+1.4}_{-1.4})(+16)$	155^{+69}_{-65}	0.19 ^(h)	—	—
CH ₂ DOH	$(6.4^{+1.3}_{-1.3})(+16)$	155 ⁽ⁱ⁾	0.19 ^(h)	$2.9^{+0.8}_{-0.8}$	—
CH ₃ OD (P04 ^(e))	$(9.9^{+1.2}_{-1.2})(+15)$	155 ⁽ⁱ⁾	0.19 ^(h)	$0.4^{+0.3}_{-0.3}$	—
CH ₃ OD (J18 ^(e))	$(4.4^{+0.5}_{-0.5})(+16)$	155 ⁽ⁱ⁾	0.19 ^(h)	$2.0^{+0.6}_{-0.6}$	—

Notes. ^(a)Column density ratio between a deuterated species and its main isotopologue. ^(b)From [Parise et al. \(2006\)](#). ^(c)From [Taquet et al. \(2015\)](#). ^(d)The source size was assumed to be equal to that of CH₃OH. ^(e)P04 refers to the partition function computed by [Parise \(2004\)](#) within the rigid-rotor approximation, whilst J18 refers to the partition function used in [Jørgensen et al. \(2018\)](#) obtained from a scaling of that of CH₃¹⁸OH. ^(f)The rotational temperature was assumed to be equal to that of CH₂DOH. ^(g)The rotational temperature was assumed to be equal to that of CH₃CN. ^(h)From [Bianchi et al. \(2017\)](#). ⁽ⁱ⁾The rotational temperature was assumed to be equal to that of ¹³CH₃OH.

For IRAS 2A and IRAS 4A, we started the analysis with CH₂DOH because this is the isotopologue with the highest number of detected transitions and with the widest excitation range, with 25 detected transitions of upper level energies E_{up} between 33 and 364 K. The population distributions can be reproduced with rotational temperatures of 166^{+58}_{-48} and 152^{+62}_{-78} K in IRAS 2A and IRAS 4A, respectively. These temperatures are slightly higher, but within the uncertainties, than the temperature of 140 K needed to reproduce the CH₃OH and ¹³CH₃OH emissions ([Taquet et al. 2015](#)). Transitions with upper level energies lower than 100 K are optically thick with opacities of 0.2–0.4 towards IRAS 2A and 0.5–1.0 towards IRAS 4A, depending on their properties. For CH₃OD, we assumed the same rotational temperature as for CH₂DOH because of the lower range of excitation (E_{up} from 33 to 104 K) and the high scatter and uncertainties in the population distributions for both sources. The strong scatter of the CH₃OD population is not entirely reproduced, possibly because of non-LTE effects due to high

critical densities. CH₃OD transitions are also optically thick with opacities ranging from 0.3 to 1.3 towards the two sources. The CHD₂OH population distribution shows much less scatter than CH₃OD. The comparison of the column densities derived in this work with those found for CH₃OH in [Taquet et al. \(2015\)](#) allows us to derive the methanol deuterations in the two sources. We find a [CH₂DOH]/[CH₃OH] abundance ratio of 5.8 ± 0.8 and $3.2 \pm 0.7\%$ towards IRAS 2A and IRAS 4A, respectively. Using the CH₃OD partition function estimated by [Parise \(2004\)](#), we find [CH₂DOH]/[CH₃OD] abundance ratios of 3.6 ± 1.6 and 5.3 ± 3.4 in IRAS 2A and IRAS 4A, respectively. However, as expected, we find lower [CH₂DOH]/[CH₃OD] abundance ratios of 0.8 ± 0.3 and 1.2 ± 0.5 in IRAS 2A and IRAS 4A, respectively, when the partition function from [Jørgensen et al. \(2018\)](#) is used. As discussed in Sect. 4.3, a [CH₂DOH]/[CH₃OH] abundance ratio of ~ 3 is consistent with the statistical value, whereas a ratio of ~ 1 remains puzzling. In addition to methanol, we also detect three transitions from the deuterated methyl cyanide

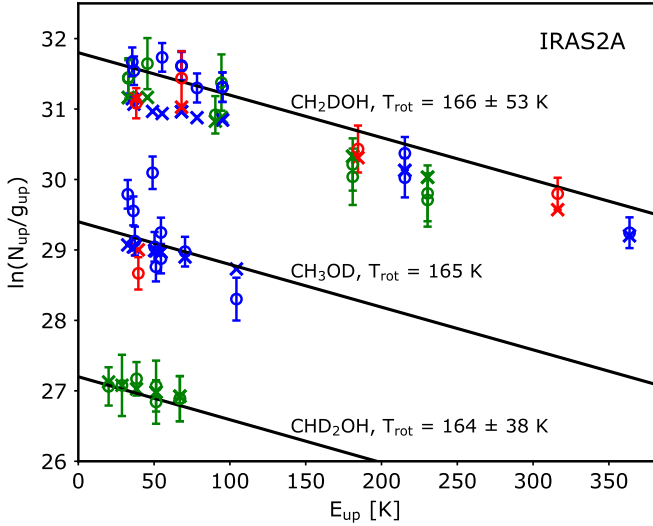


Fig. 4. Rotational and population diagrams of deuterated methanol isotopologues (CH_2DOH , CH_3OD , and CHD_2OH) derived towards NGC 1333-IRAS 2A for source sizes derived from the PD analysis of the methanol population distribution ($0''.36$; see Taquet et al. 2015). Red, green, and blue symbols represent transitions at 145, 165, and 225 GHz, respectively. The CH_3OD and CHD_2OH levels have been artificially shifted by -1.5 and -4 for clarity. Observational data are depicted by diamonds. Crosses show the best fit of the PD to the data. Blue and green symbols represent transitions at 2 and 1.3 mm, respectively. Error bars are derived assuming a calibration uncertainty of 20% in addition to the statistical error. Black straight lines represent the best fit of the rotational diagram analysis to the data, assuming optically thin emission.

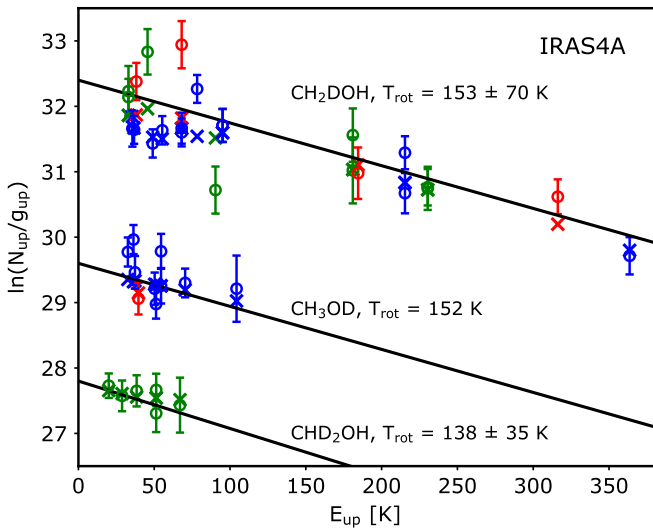


Fig. 5. Rotational and population diagrams of deuterated methanol isotopologues (CH_2DOH , CH_3OD , and CHD_2OH) derived towards NGC 1333-IRAS 2A for source sizes derived from the PD analysis of the methanol population distribution ($0''.20$; see Taquet et al. 2015). Red, green, and blue symbols represent transitions at 145, 165, and 225 GHz, respectively. The CH_3OD and CHD_2OH levels have been artificially shifted by -1.5 and -4 for clarity. Observational data are depicted by diamonds. Crosses show the best fit of the PD to the data. Blue and green symbols represent transitions at 2 and 1.3 mm, respectively. Error bars are derived assuming a calibration uncertainty of 20% in addition to the statistical error. Black straight lines represent the best fit of the rotational diagram analysis to the data, assuming optically thin emission.

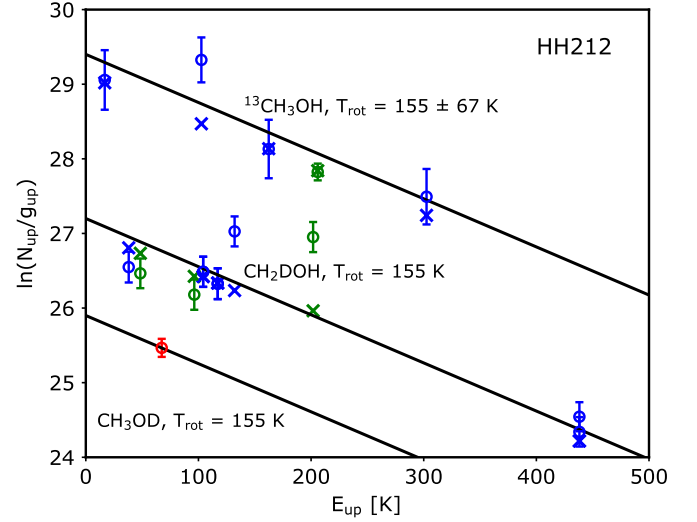


Fig. 6. Rotational and population diagrams of methanol isotopologues ($^{13}\text{CH}_3\text{OH}$, CH_2DOH , and CH_3OD) derived towards HH212 for the source sizes derived from the ALMA Cycle 4 data ($0''.19$; see Bianchi et al. 2017). The CH_2DOH and CH_3OD levels have been artificially shifted by -3 for clarity. Red, green, and blue symbols represent transitions at 335, 338, and 349 GHz, respectively. Observational data are depicted by diamonds. Crosses show the best fit of the PD to the data. Black straight lines represent the best fit of the rotational diagram analysis to the data, assuming optically thin emission.

isotopologue CH_2DCN at 225 GHz. We obtain the CH_2DCN column densities by assuming that the CH_2DCN rotational temperature is equal to 200 K, the rotational temperature of CH_3CN measured in Taquet et al. (2015). Comparing the CH_2DCN column densities derived in this work with those of CH_3CN from Taquet et al. (2015) allows us to obtain $[\text{CH}_2\text{DCN}]/[\text{CH}_3\text{CN}]$ ratios of 3.6 ± 0.8 and $2.7 \pm 0.9\%$ in IRAS 2A and IRAS 4A, respectively.

For HH212, we re-derived the column densities of $^{13}\text{CH}_3\text{OH}$ and CH_2DOH that have been estimated by Bianchi et al. (2017). Our method differs slightly from that of Bianchi et al. (2017) because the PD diagram is built upon the flux extracted from an elliptical mask instead of the continuum peak position in order to use consistent methods for the different sources. Assuming a source size of $0''.19$ following Bianchi et al. (2017), we started with $^{13}\text{CH}_3\text{OH}$ and then used the derived rotational temperature to estimate the CH_2DOH and CH_3OD column densities. We obtain a $[\text{CH}_2\text{DOH}]/[\text{CH}_3\text{OH}]$ abundance ratio of $2.9 \pm 0.8\%$. The $[\text{CH}_2\text{DOH}]/[\text{CH}_3\text{OD}]$ abundance ratio is found to be 7.2 ± 5.3 and 1.5 ± 0.6 with the CH_3OD partition function estimated by Parise (2004) and Jørgensen et al. (2018), respectively.

4. Discussion

4.1. Comparison with previous observations

Deuterated methanol has previously been detected with the IRAM 30-m single-dish telescope towards IRAS 2A and IRAS 4A. Parise et al. (2006) derived $[\text{CH}_2\text{DOH}]/[\text{CH}_3\text{OH}]$ abundance ratios of 62 and 65% towards IRAS 2A and IRAS 4A, respectively, which are 10 and 17 times higher than the deuteration values found in this study. The $[\text{CH}_3\text{OD}]/[\text{CH}_3\text{OH}]$ and $[\text{CH}_2\text{DOH}]/[\text{CH}_3\text{OD}]$ abundance ratios derived in this work highly depend on the partition function that is used (see

Sect. 2.2). With the same partition function as Parise (2004), the $[\text{CH}_3\text{OD}]/[\text{CH}_3\text{OH}]$ ratio of 4.7% found by Parise et al. (2006) towards IRAS 4A is 6.7 times higher than our derived value, whilst the $[\text{CH}_2\text{DOH}]/[\text{CH}_3\text{OD}]$ is found to decrease by a factor of 4 between Parise et al. (2006) and this work. For doubly deuterated methanol, the $[\text{CHD}_2\text{OH}]/[\text{CH}_3\text{OH}]$ ratios derived in this work are also much lower than the abundances derived by Parise et al. (2006): by a factor of 5.6 and 8.1 for IRAS 2A and IRAS 4A, respectively.

The large differences found between Parise et al. (2006) and the present work can be explained by several factors. First, the new CH_2DOH spectroscopic data by Pearson et al. (2012) used in this work decrease by a factor of two the derived CH_2DOH column density following a re-analysis of the single-dish data towards IRAS 2A by Parise et al. (2006) using the most recent spectroscopy data entry. This also naturally explains the increase of the $[\text{CHD}_2\text{OH}]/[\text{CH}_2\text{DOH}]$ ratio in this work by a factor of about two, from 41 to 75% in IRAS 2A and from 26 to 56% in IRAS 4A. Second, the methanol deuteration derived from single-dish observations by Parise et al. (2006) has been estimated from the detection of various transitions from the main CH_3OH isotopologue. However, the targeted transitions have low upper level energies ($E_{\text{up}} < 150$ K) and are optically thick in the inner protostellar regions surrounding these two low-mass protostars (see Taquet et al. 2015). Therefore, the rotational diagram analysis, assuming optically thin emission, carried out in Parise et al. (2006) could have likely led to an underestimation of the CH_3OH column density. Finally, the IRAM 30-m telescope has a large beam of 9–30'' depending on the frequency. As suggested by the low rotation temperatures of 55 and 27 K found by Parise et al. (2006) for CH_2DOH towards IRAS 2A and IRAS 4A, respectively, the methanol emission would mostly come from the large external envelope and/or from the molecular outflows driven by the targeted sources where methanol deuteration is higher. Methanol is mostly produced through surface chemistry through CO hydrogenation as gas phase chemistry is known to be inefficient (Watanabe & Kouchi 2002; Geppert et al. 2006). Non-thermal evaporation processes could release a fraction of solid methanol formed in ices into the gas phase. As shown by Taquet et al. (2014), the deuteration of methanol and other molecules mostly produced in ices like water is higher in external protostellar envelopes than in the inner hot cores because the cold gaseous deuteration reflects the deuteration at the highly deuterated ice surface.

Deuterated methanol has not previously been targeted with single-dish facilities towards HH212. Our PD analysis gives a $[\text{CH}_2\text{DOH}]/[\text{CH}_3\text{OH}]$ abundance ratio of $2.9 \pm 0.8\%$, which agrees well with the value of $2.4 \pm 0.4\%$ found by Bianchi et al. (2017), who used a slightly different method.

4.2. Deuteration from low- to high-mass protostars

Figure 7 and Table 3 compare the deuteration of CH_2DOH , CH_3OD , and CH_2DCN , including their statistical (i.e. taking into account the statistical ratios for the functional groups with multiple H-atoms) correction, between four low-mass, one intermediate-mass, and three massive hot cores obtained only through interferometric observations of methanol isotopologues. We also show the water deuteration for comparison when available. Because the main CH_3OH isotopologue is known to be optically thick in hot cores, we only chose observations whose deuteration values were derived from $^{13}\text{CH}_3\text{OH}$ or $\text{CH}_3^{18}\text{OH}$ column densities assuming $^{12}\text{C}/^{13}\text{C}$ and $^{16}\text{O}/^{18}\text{O}$ elemental ratios

of 70 and 560, respectively (Wilson & Rood 1994). IRAS 16293–2422, a low-mass protobinary system located in the Ophiuchus molecular cloud, has been observed with ALMA by Jørgensen et al. (2018) in the context of the PILS survey that covered the whole ALMA band 7 between 329.15 and 362.90 GHz (Jørgensen et al. 2016). NGC 7129 FIRS 2, an intermediate-mass protostar, has been observed with the IRAM-PdBI at 220 GHz by Fuente et al. (2014). The methanol deuterium fractionation towards the Orion KL massive star-forming region has been estimated by Peng et al. (2012) through interferometric observations using the IRAM-PdBI at 103 and 225 GHz. The Sgr B2(N2) massive hot core has been observed by Bellocche et al. (2016) in the context of the EMOCA ALMA survey of the whole 3 mm band. The NGC 6334 massive protocluster has recently been observed by Bøgelund et al. (2018) through ALMA observations at 300 GHz, and we selected the deuteration derived using the $\text{CH}_3^{18}\text{OH}$ column density. The error bars of the CH_3OD deuteration derived in this work include the different column density values derived with the two partition functions used for CH_3OD .

The dust temperature in the dark clouds surrounding protostars is mostly governed by the external interstellar radiation field with local variations due to bright stars and/or dense cores. All the hot cores considered in Fig. 7 and Table 3 are likely relatively young, that is, have a lifetime of 10^5 yr at most, the dust temperatures of their precursor dark clouds are therefore likely similar to the current observed values. The deuterium fractionation is plotted against the current dust temperature of their surrounding cloud as measured by the *Planck* observatory for NGC 7129 FIRS2 and by the PACS and SPIRE instruments on board the *Herschel* Space Observatory for other sources. The Ophiuchus, Perseus, and Orion molecular clouds have been observed with the Gould Belt Survey key program (André et al. 2010), the Sgr B2 region has been observed by Etzaluze et al. (2013) as part of the Hi-GAL key program (Molinari et al. 2010), whilst the NGC 6334 massive complex has been observed with the HOBYS key program (Motte et al. 2010; Russeil et al. 2013; Tige et al. 2017). To estimate the dust temperature, we extracted 10 arcmin maps surrounding the selected sources. We then generated temperature histograms of all pixels showing visual extinctions A_V higher than 3 mag when the H_2 column density maps are available (for the Ophiuchus, Perseus, and NGC 6334 clouds) or showing 353 GHz opacities higher than ten times the rms noise (for the Orion and NGC 7129 clouds). To derive a temperature representative of the entire region, we filtered out the hot regions that are locally heated by stars or cold regions of dense cores by fitting the temperature histogram around the main peak with a Gaussian function. Figure B.1 gives an example for the NGC1333 region containing IRAS 2A and IRAS 4A.

It is found that the statistical deuterium fractionations of CH_2DOH and CH_3OD tend to decrease from a few percent to less than $\sim 0.1\%$ with the dust temperature between ~ 13 and ~ 24 K. With the same partition function as estimated by Bellocche et al. (2016) and Jørgensen et al. (2018), the $[\text{CH}_2\text{DOH}]/[\text{CH}_3\text{OD}]$ abundance ratio seems to show a large dispersion because it typically varies from 1–4 in low-mass protostars to 0.1–2 in high-mass hot cores (0.8 ± 0.3 in IRAS 2A, 1.2 ± 0.5 in IRAS 4A, 3.9 ± 1.1 in IRAS 16293, 1.5 ± 0.6 in HH212, 1.7 in Sgr B2, 0.73 ± 0.27 in Orion KL, and 0.13 ± 0.03 in NGC6334-IMM1).

It has been shown that warm hot-core gas-phase chemistry is likely too slow to significantly alter the deuteration after the evaporation of ices in hot cores (Charnley et al. 1997; Osamura et al. 2004). Thus, the methanol deuterium

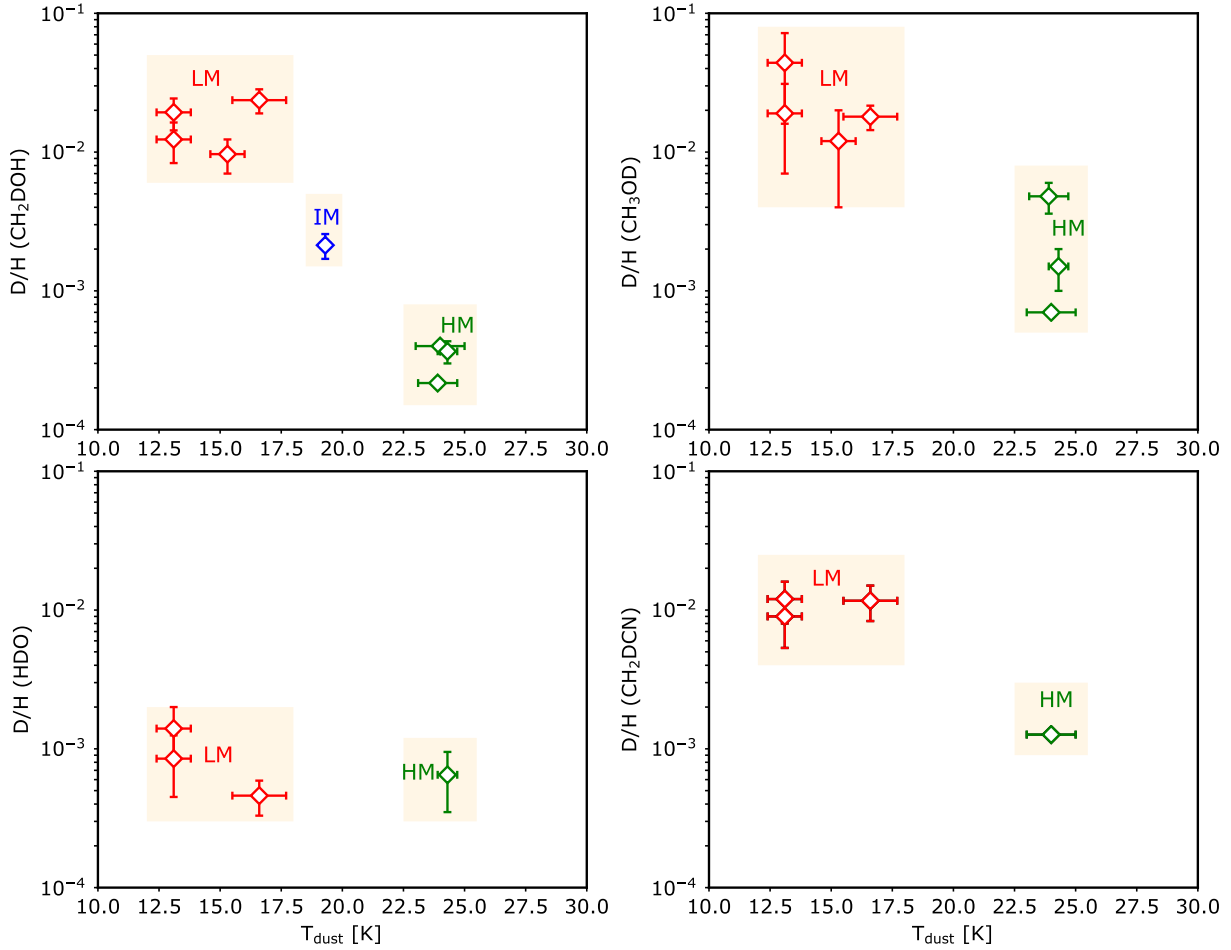


Fig. 7. Statistical (i.e. taking into account the statistical ratios for the functional groups with multiple H-atoms) CH_2DOH (top left), CH_3OD (top right), HDO (bottom left), and CH_2DCN (bottom right) deuterium fractionations observed towards the hot core of a sample of low-mass (red), intermediate-mass (blue), and high-mass (green) protostars with sub-millimeter interferometers as function of the dust temperature of the surrounding cloud measured with the *Herschel* or *Planck* telescopes (see values in Table 3).

Table 3. Surrounding cloud dust temperatures and statistical deuterium fractionations observed towards protostellar hot cores.

Source	T_{dust} (K)	D/H (CH_2DOH) (%)	D/H (CH_3OD) (%)	$[\text{CH}_2\text{DOH}]/$ $[\text{CH}_3\text{OD}]$	D/H (CH_2DCN) (%)	D/H (HDO) (%)	Ref.
IRAS 16293-B	16.6 ± 1.1	2.4 ± 0.5	1.8 ± 0.4	3.9 ± 1.1	1.2 ± 0.3	0.046 ± 0.013	[1, 2, 3]
IRAS 2A	13.1 ± 0.7	1.9 ± 0.5	$1.6 - 7.1$	$0.8 - 3.6$	1.2 ± 0.4	0.085 ± 0.040	[4, 5]
IRAS 4A	13.1 ± 0.7	1.2 ± 0.4	$0.7 - 3.1$	$1.2 - 5.3$	0.90 ± 0.37	0.17 ± 0.08	[4, 6]
HH212	15.3 ± 0.7	1.0 ± 0.3	$0.4 - 2.0$	$1.5 - 7.2$	–	–	[7, 4]
NGC 7129 FIRS1	19.3 ± 0.2	0.21 ± 0.04	–	–	–	–	[8]
Sgr B2	24 ± 1	0.04	0.07	1.7	0.13	–	[9]
Orion KL	24.3 ± 0.4	0.037 ± 0.007	0.15 ± 0.05	0.73 ± 0.27	–	0.13 ± 0.06	[10, 11]
NGC 6334 IMM1I	23.9 ± 0.8	0.020 ± 0.007	0.48 ± 0.12	0.13 ± 0.03	–	–	[12]

Notes. For the CH_3OD column densities in IRAS 2A, IRAS 4A, and HH212, two estimates are given using the partition functions estimated by Parise (2004) and Jørgensen et al. (2018). The highest CH_3OD column densities (and the lowest $[\text{CH}_2\text{DOH}]/[\text{CH}_3\text{OD}]$ ratios) are those with the partition function estimated by Jørgensen et al. (2018).

References. [1] Jørgensen et al. (2018); [2] Calcutt et al. (2018); [3] Persson et al. (2013); [4] this work; [5] Coutens et al. (2014); [6] average between $\text{HDO}/\text{H}_2\text{O}$ ratios of 5×10^{-3} (Taquet et al. 2013a) and 5.4×10^{-4} (Persson et al. 2014, M. Persson, priv. comm.); [7] Bianchi et al. (2017); [8] Fuente et al. (2014); [9] Belloche et al. (2016); [10] Peng et al. (2012); [11] Neill et al. (2013); [12] Bøgelund et al. (2018).

fractionation observed in the warm gas-phase of low-mass and massive hot cores would rather reflect the deuteration of methanol formed in interstellar ices within their progenitor dense clouds.

4.3. Comparison with theoretical predictions

We compared the observed methanol deuterations with the predictions of the GRAINOBLE astrochemical code by Taquet et al. (2012a, 2013a, 2014), which follows the formation and

the deuteration of ices in cold dense cores. The code has been extensively described in previous studies and has recently been used to interpret the deuteration towards NGC 6334 measured by Bøgelund et al. (2018). We investigated the effect of the temperature, total density, and time on the methanol deuteration using the same code and same chemical network as in Taquet et al. (2014). However, unlike in Taquet et al. (2014), we ran here a series of pseudo-time-dependent simulations in which the chemistry evolves over time for constant physical properties such as temperature or density. Figure C.1 shows the deuteration of methanol and water in ices as a function of temperature between 10 and 30 K for three different dense cloud densities, $n_{\text{H}} = 10^4$, 10^5 , and 10^6 cm^{-3} at three different times, $0.1 \times t_{\text{FF}}$, t_{FF} , and $10 \times t_{\text{FF}}$, where t_{FF} is the free-fall time. t_{FF} is equal to 4.4×10^5 , 1.4×10^5 , and $4.4 \times 10^4 \text{ yr}$ at densities of 10^4 , 10^5 , and 10^6 cm^{-3} , respectively.

The methanol deuteration strongly depends on the density, the considered temperature, and time. Methanol and its deuterated isotopologues are formed through addition reactions of atomic H and D on CO and H_2CO on ices possibly supplemented by abstraction and substitution reactions (Nagaoka et al. 2005; Hidaka et al. 2007, 2009). The methanol deuteration is therefore governed by the atomic $[\text{D}]/[\text{H}]$ abundance ratio in the gas phase during ice formation. Atomic D is mostly formed through electronic recombination of H_3^+ isotopologues, which in turn are formed through exothermic reactions between H_3^+ and HD. The efficiency of the backward reaction increases with increasing temperature and with the ortho/para ratio of H_2 . In addition, reactions between H_3^+ or with its isotopologues and HD are in competition with reactions involving CO. The production of atomic deuterium is therefore enhanced at low temperatures, close to 10 K, when the abundances of CO and ortho state of H_2 are low (see Fig. 1 of Taquet et al. 2012b, for a scheme detailing the gas-phase deuteration network). As a consequence, the strong decrease of up to two orders of magnitudes of the methanol deuteration predicted at $t \geq t_{\text{FF}}$ from 10 to 30 K is essentially due to the decreased efficiency of deuterium chemistry in the gas phase of cold dense cores.

The observed $[\text{CH}_3\text{OD}]/[\text{CH}_3\text{OH}]$ ratios observed both towards low-mass and massive protostars are relatively well reproduced by the model predictions at $t \geq t_{\text{FF}}$, despite the large dispersion of values among massive protostars. The $[\text{CH}_2\text{DOH}]/[\text{CH}_3\text{OH}]$ abundance ratios of low-mass protostars can also be explained by the models at $t \geq t_{\text{FF}}$, but the massive protostar values can only be explained by a shorter time of $0.1 \times t_{\text{FF}}$ or a low density of 10^4 cm^{-3} . The predicted $[\text{CH}_2\text{DOH}]/[\text{CH}_3\text{OD}]$ ratio remains close to the statistical value of 3 in most cases and can increase up to ~ 6 at 10^6 cm^{-3} , $T \sim 10\text{--}15 \text{ K}$, and $t = 10 \times t_{\text{FF}}$. Our model only focuses on cold surface chemistry and is not able to explain $[\text{CH}_2\text{DOH}]/[\text{CH}_3\text{OD}]$ ratios lower than 1 observed in massive hot cores. CH_3OD would then need to be produced through other processes such as warm gas-phase chemistry following methanol evaporation within massive hot cores. Osamura et al. (2004) showed that the reaction $\text{H}_2\text{DO}^+ + \text{CH}_3\text{OH} \rightarrow \text{CH}_3\text{OHD}^+ + \text{H}_2\text{O}$, followed by dissociative recombination, could increase the $[\text{CH}_3\text{OD}]/[\text{CH}_3\text{OH}]$ ratio by a factor of ~ 5 at 100 K. However, this increase is obtained at times longer than 10^5 yr and assuming $[\text{HDO}]/[\text{H}_2\text{O}]$ abundance ratios of ~ 0.1 that are about 100 times higher than the observed values. It remains to be tested whether this process could play a major role at higher temperatures and for realistic water deuteration values. On the other hand, hydrogen-deuterium exchanges in warm ices have been proven to be efficient, but only on the hydroxyl functional group

of methanol. CH_3OD would thus give its deuterium to water, inducing a decrease in CH_3OD abundance in ices before its evaporation (Ratajczak et al. 2009; Faure et al. 2015).

Unlike methanol, water does not show any significant decrease of its observed deuteration with the dust temperature because it remains around 0.1% between 13 and 24 K. The low water deuteration has been interpreted as due to an early formation of solid water in the translucent phase at low visual extinctions, when the dark cloud is still lukewarm, with temperatures of 15–20 K, and when the abundances of CO and ortho H_2 are still high (Taquet et al. 2013a, 2014). Methyl cyanide is thought to be mostly produced either on warm ices from recombination between the CN and CH_3 radicals or in the gas phase through the radiative association between HCN and CH_3^+ , after the prestellar stage. Comparing the methyl cyanide deuteration with the model predictions with constant physical conditions of cold cores is therefore not relevant.

4.4. Deuteration from low-mass protostars to comets

Solar System comets mostly contain pristine material that has poorly evolved since the end of the Solar System formation 4.5 billions years ago. Their chemical composition should reflect the chemical composition of the solar nebula (see Mumma & Charnley 2011, for instance). Comparing the deuteration of water and organics measured around young solar-type protostars on Solar System scales and in Solar System comets should allow us to follow the chemical evolution throughout the star formation process.

Figure 8 compares the (statistical) deuterium fractionation of methanol, methyl cyanide, and water observed in low-mass protostars with the deuteration of HCN in protoplanetary discs and of CH_3OH , HCN, and water in comets. Deuteration of water has been estimated in four nearby protostars on Solar System scales by Taquet et al. (2013a), Persson et al. (2014) and Coutens et al. (2014) through interferometric observations of HDO and H_2^{18}O transitions and about a dozen comets (see the list of references in the caption of Fig. 7). Water deuteration in protostars is only higher by about a factor of two than the deuteration measured in comets. Unlike water, cometary deuterated methanol has not been detected so far. The most stringent upper limit in the methanol deuteration has been given by Crovisier et al. (2004) towards comet Hale-Bopp, with a $[\text{CH}_2\text{DOH}]/[\text{CH}_3\text{OH}]$ abundance ratio lower than 0.8%. The $[\text{CH}_2\text{DOH}]/[\text{CH}_3\text{OH}]$ ratio of 3–6% estimated with interferometers towards young protostars is higher by a factor of 4–8 than the estimate made in comet Hale-Bopp. A low methanol deuteration in comets remains to be confirmed with a clear detection of deuterated methanol, either through the analysis of the high-resolution ROSINA mass spectrometer data at mass 33 on board the Rosetta space probe (see for instance Le Roy et al. 2015), or through future sensitive sub-millimeter or millimeter ground-based surveys of nearby and bright comets using interferometers like ALMA or NOEMA.

Assuming that the observed deuteration towards nearby low-mass protostars is representative for the deuteration of the young Sun, this comparison suggests that an important reprocessing of the organic material occurred in the solar nebula whilst little water reprocessing in the solar nebula is required. Three reasons can be suggested. First, methanol observed in the hot core of protostars on $\sim 100 \text{ AU}$ scales is reprocessed before entering in the disc mid-plane through high-temperature chemistry involving reactions with H or OH because they have moderate energy barriers of 2200–3000 K (Li & Williams 1996). However, the temperature of the hot corino is likely too low to trigger

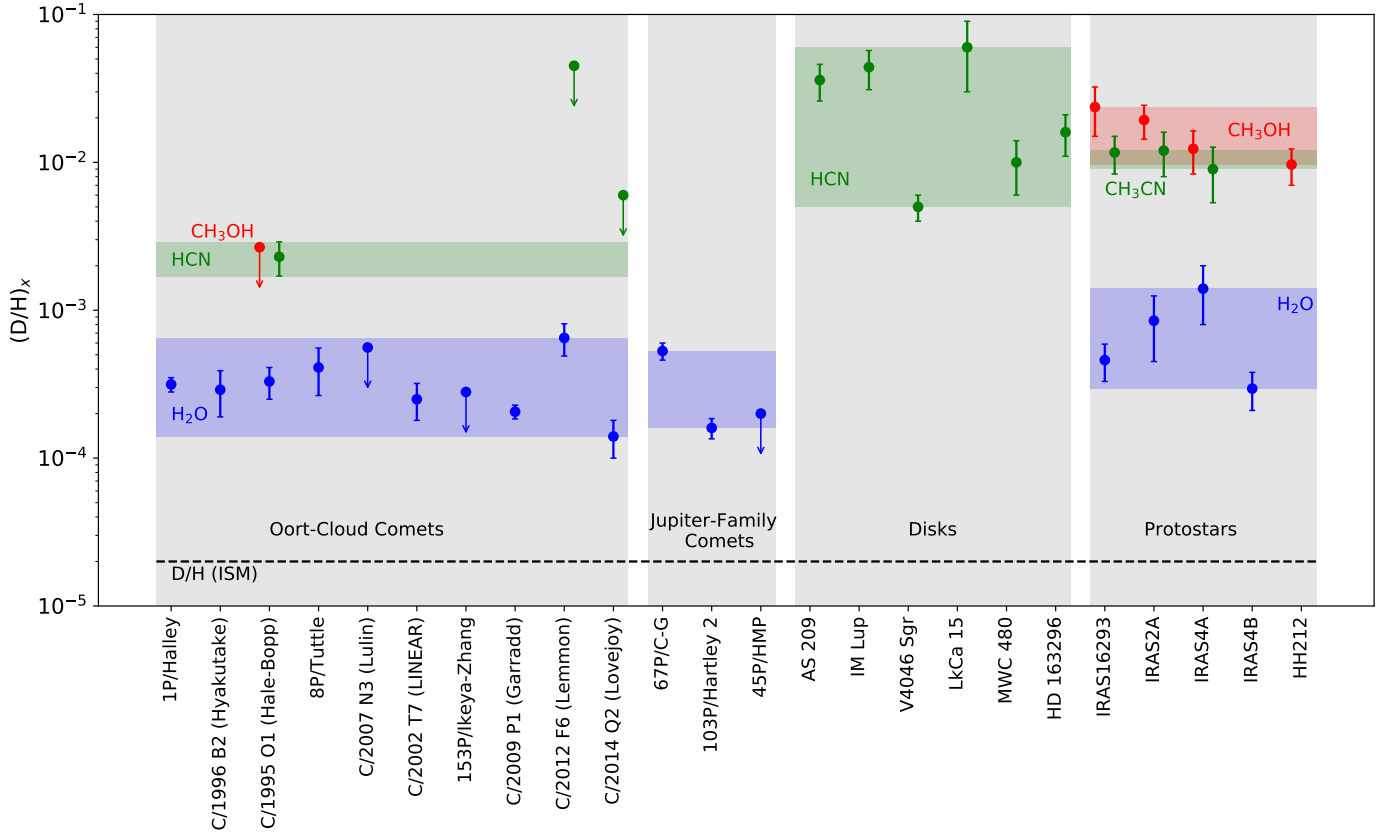


Fig. 8. Statistical D/H ratios in Solar System comets, protoplanetary discs, and low-mass Class 0 protostars for methanol (red), water (blue), and HCN or CH₃CN (green). References: 1P/Halley: Eberhardt et al. (1995); C/1996 B2 (Hyakutake): Bockelée-Morvan et al. (1998); C/1995 O1 (Hale-Bopp): Meier et al. (1998); Crovisier et al. (2004); 8P/Tuttle: Villanueva et al. (2009); C/2007 N3 (Lulin): Gibb et al. (2012); C/2002 T7 (LINEAR): Hutsemékers et al. (2008); 153P/Ikeya-Zhang: Biver et al. (2006); C/2009 P1 (Garradd): Bockelée-Morvan et al. (2012); C/2012 F6 (Lemmon): Biver et al. (2016); C/2014 Q2 (Lovejoy): Biver et al. (2016); 67P/C-G: Altwegg et al. (2015); 103P/Hartley 2: Hartogh et al. (2011); 45P/HMP: Lis et al. (2013); All discs: Huang et al. (2017); IRAS 16293: Persson et al. (2014); Jørgensen et al. (2018); Calcutt et al. (2018); IRAS 2A: Coutens et al. (2014), this work; IRAS 4A: average values derived by Taquet et al. (2013b) and Persson et al. (2014, M. Persson priv. comm.), this work; IRAS4B: Persson et al. (2014).

such reactions and other processes that increase the gas temperature to ~ 1000 K, such as accretion shocks that form at the interface of the envelope and the disc in formation (Aota et al. 2015), would be needed. Second, physical processes such as turbulence in the disc could transport the material from the disc mid-plane to the atmosphere where efficient photolytic processes can photo-dissociate organic molecules. Albertsson et al. (2014) and Furuya et al. (2014) demonstrated that radial and vertical turbulence can decrease the water deuteration by typically one order of magnitude in the cometary zone. However, it is yet to be demonstrated that this process could also affect the deuteration of more complex organic species with slightly different chemistries. A third possibility is that the cloud at the origin of the Solar System was initially slightly warmer than the temperatures of nearby dark clouds in which the observed protostars are located. This has recently been suggested by Taquet et al. (2016), who concluded that a dark cloud temperature of ~ 20 K, together with a dark cloud density higher than 10^5 cm $^{-3}$, was needed to explain the high abundance of O₂ and its strong correlation with water in comet 67P/CG. This modelling work confirmed the studies of meteoritic data, which suggest that the Solar System was born in a dense cluster of stars (see Adams 2010). According to Fig. 7, an increase in dark cloud temperature from 10 to 20 K would decrease the methanol deuteration from 2–6 to 0.2–0.6%, depending on the density, which seems to be in good agreement with the value found in comet 67P/CG.

5. Conclusions

We analysed several existing observational datasets obtained with the PdBI and ALMA sub-millimeter interferometers to estimate the methanol deuterations in the hot cores surrounding three protostars on Solar System scales. For this purpose, we analysed several dozen deuterated methanol transitions with a PD analysis in order to measure the [CH₂DOH]/[CH₃OH] and [CH₃OD]/[CH₃OH] abundance ratios. The obtained [CH₂DOH]/[CH₃OH] ratios of 3–6% and [CH₃OD]/[CH₃OH] ratios of 0.4–1.6% are typically one order of magnitude lower than previous estimates derived from single-dish observations towards the same sources, and are in good agreement with the recent ALMA measurements by Jørgensen et al. (2018) towards the low-mass protostar IRAS 16293-B. We then compared our methanol deuteration estimates with previous measurements of intermediate- and high-mass hot cores with similar observational properties and analysis methods. Unlike water, which does not show strong variation of its deuteration between low-mass and high-mass protostars, we find that the methanol deuteration around massive hot cores is much lower by one to two orders of magnitude. This strong difference could be attributed to a different physical or chemical history of the sources. Methanol observed around protostars is mostly formed in interstellar ices during the previous molecular cloud phase. Dust temperature maps derived with the *Herschel* or *Planck* space observatories

suggest that the observed massive protostars were born in warm molecular clouds of $T \sim 24$ K, whilst low-mass protostars are located in more quiescent and cold regions with temperatures lower than 15 K. Comparing the observed deuteration values with the predictions of the GRAINOBLE astrochemical model on ice formation and deuteration, an increase of 10 K in the dust temperature is enough to explain a decrease of the observed methanol deuteration, depending on the density and the chemical time. Finally, the methanol deuterations measured towards young solar-type protostars at high angular resolution on Solar System scales seem to be higher by a factor of ~ 5 than the upper limit in methanol deuteration estimated in comet Hale-Bopp by Crovisier et al. (2004). If this result is confirmed by subsequent observations of other comets, it would imply that an important reprocessing of the organic material likely occurred in the solar nebula during the formation of the Solar System.

Acknowledgements. This work is based on observations carried out with the ALMA Interferometer under project numbers ADS/JAO.ALMA#2012.1.00997.S and ADS/JAO.ALMA#2016.1.01475.S data (PI: C. Codella) and with the IRAM PdBI/NOEMA Interferometer under project numbers V05B and V010 (PI: M.V. Persson) and U003 (PI: V. Taquet). ALMA is a partnership of ESO (representing its member states), NSF (USA) and NINS (Japan), together with NRC (Canada) and NSC and ASIAA (Taiwan), in cooperation with the Republic of Chile. The Joint ALMA Observatory is operated by ESO, AUI/NRAO and NAOJ. IRAM is supported by INSU/CNRS (France), MPG (Germany) and IGN (Spain). V.T. acknowledges the financial support from the European Union's Horizon 2020 research and innovation programme under the Marie Skłodowska-Curie grant agreement n. 664931. E.B., Ce.Ce., C.K., A.L. and Cl.Co acknowledge the funding from the European Research Council (ERC) under the European Union's Horizon 2020 research and innovation programme, for the Project "The Dawn of Organic Chemistry" (DOC), grant agreement No 741002. Cl.Co acknowledges the funding from PRIN-INAF 2016 "The Cradle of Life – GENESIS-SKA (General Conditions in Early Planetary Systems for the rise of life with SKA)". Ce.Ce. and Cl.Co acknowledge the financial support from the European MARIE SKŁODOWSKA-CURIE ACTIONS under the European Union's Horizon 2020 research and innovation programme, for the Project "Astro-Chemistry Origins" (ACO), Grant No 811312.

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Appendix A: Line parameters of the targeted transitions

Table A.1. Line parameters of CH₂DOH lines observed towards IRAS 2A and IRAS 4A.

Number	Frequency (GHz)	Transition	E_{up} (K)	A_{ul} (s ⁻¹)	Flux (IRAS 2A) (Jy km s ⁻¹)	Flux (IRAS 4A) (Jy km s ⁻¹)
1	166.063164	2 _{2,0} e ₁ –1 _{1,0} o ₁	33.0	2.40(–5)	0.285 ± 0.081	0.199 ± 0.076
2	166.787448	2 _{2,1} e ₁ –1 _{1,1} o ₁	33.0	2.43(–5)	0.292 ± 0.071	0.184 ± 0.051
3	225.878232	3 _{1,3} o ₁ –2 _{0,2} o ₁	35.6	3.23(–5)	0.607 ± 0.123	0.204 ± 0.056
4	226.818248	5 _{1,4} e ₀ –4 _{1,3} e ₀	36.7	3.58(–5)	0.931 ± 0.187	0.349 ± 0.074
5	144.762399	4 _{1,3} e ₁ –3 _{0,3} o ₁	38.2	8.80(–6)	0.132 ± 0.028	0.153 ± 0.043
6	164.108467	5 _{0,5} e ₁ –4 _{1,4} e ₁	45.6	2.67(–6)	0.086 ± 0.031	0.090 ± 0.031
7	223.898819	3 _{2,2} o ₁ –4 _{2,2} e ₁	48.3	1.56(–6)	0.450 ± 0.095	0.219 ± 0.068
8	225.667709	5 _{1,4} e ₁ –4 _{1,3} e ₁	49.0	4.44(–5)	0.272 ± 0.062	0.351 ± 0.075
9	224.928016	5 _{1,4} e ₂ –4 _{1,3} e ₂	55.3	4.42(–5)	1.390 ± 0.280	0.428 ± 0.092
10	144.134719	5 _{3,2} e ₀ –6 _{2,5} e ₀	68.1	2.64(–6)	0.069 ± 0.026	0.098 ± 0.035
11	223.691457	5 _{3,3} e ₀ –4 _{3,2} e ₀	68.1	2.17(–5)	0.601 ± 0.122	0.203 ± 0.043
12	223.697110	5 _{3,2} e ₀ –4 _{3,1} e ₀	68.1	2.17(–5)	0.608 ± 0.124	0.217 ± 0.051
13	225.848108	7 _{0,7} o ₁ –6 _{1,6} o ₁	78.3	2.17(–5)	0.604 ± 0.123	0.540 ± 0.114
14	165.861895	8 _{0,8} e ₁ –8 _{1,7} e ₀	90.4	8.99(–6)	0.217 ± 0.057	0.056 ± 0.020
15	164.577869	8 _{1,7} e ₁ –7 _{2,6} e ₁	94.4	1.52(–6)	0.058 ± 0.023	0.071 ± 0.029
16	223.616196	5 _{4,2} e ₀ –4 _{4,1} e ₀	95.1	1.26(–5)	0.522 ± 0.107	0.264 ± 0.065
	223.616210	5 _{4,1} e ₀ –4 _{4,0} e ₀	95.1	1.26(–5)		
17	164.374775	10 _{4,7} e ₀ –10 _{3,7} o ₁	181.0	5.91(–6)	0.073 ± 0.029	0.062 ± 0.031
18	164.388327	10 _{4,6} e ₀ –10 _{3,8} o ₁	181.0	5.91(–6)	0.087 ± 0.032	0.106 ± 0.043
19	143.566661	12 _{2,10} e ₀ –11 _{3,9} e ₀	184.5	3.76(–6)	0.082 ± 0.027	0.041 ± 0.016
20	224.273919	10 _{5,6} e ₀ –11 _{4,7} e ₀	215.4	1.16(–5)	0.126 ± 0.035	0.082 ± 0.025
21	224.285457	10 _{5,5} e ₀ –11 _{4,8} e ₀	215.4	1.16(–5)	0.178 ± 0.041	0.152 ± 0.038
22	166.950015	12 _{4,9} e ₀ –12 _{3,9} o ₁	230.4	6.80(–6)	0.079 ± 0.031	0.065 ± 0.018
23	166.995876	12 _{4,8} e ₀ –12 _{3,10} o ₁	230.4	6.80(–6)	0.072 ± 0.027	0.064 ± 0.021
24	142.838827	16 _{1,15} o ₁ –16 _{0,16} o ₁	316.2	9.85(–6)	0.138 ± 0.031	0.099 ± 0.026
25	225.551640	17 _{2,15} o ₁ –17 _{1,16} o ₁	363.6	3.28(–5)	0.273 ± 0.059	0.149 ± 0.042

Table A.2. Line parameters of CH₃OD lines observed towards IRAS 2A and IRAS 4A.

Number	Frequency (GHz)	Transition	E_{up} (K)	A_{ul} (s ⁻¹)	Flux (IRAS 2A) (Jy km s ⁻¹)	Flux (IRAS 4A) (Jy km s ⁻¹)
1	226.53867	5 ₀₊ –4 ₀₊ A	32.7	4.31(–5)	0.869 ± 0.175	0.290 ± 0.064
2	226.350191	5 ₀ –4 ₀ E	36.4	4.66(–5)	0.744 ± 0.151	0.380 ± 0.084
3	226.185930	5 _{–1} –4 _{–1} E	37.3	4.16(–5)	0.433 ± 0.088	0.206 ± 0.051
4	143.741650	5 _{1–} –5 ₀₊ A	39.6	3.52(–5)	0.237 ± 0.054	0.121 ± 0.029
5	226.92258	5 _{–2} –4 _{–2} E	50.2	3.71(–5)	0.355 ± 0.074	0.143 ± 0.035
6	226.892864	5 ₂ –4 ₂ E	54.4	3.71(–5)	0.267 ± 0.055	0.113 ± 0.025
7	226.94283	5 ₂₊ –4 ₂₊ A	54.4	3.70(–5)	0.436 ± 0.091	0.253 ± 0.067
8	226.825536	5 ₃ –4 ₃ E	70.5	3.71(–5)	0.254 ± 0.053	0.119 ± 0.026
9	226.70660	5 _{2–} –4 _{2–} A	100	2.84(–5)	0.300 ± 0.062	0.149 ± 0.040
10	226.73886	5 _{–4} –4 _{–4} E	104.3	1.60(–5)	0.073 ± 0.022	0.062 ± 0.031

Table A.3. Line parameters of CHD₂OH lines observed towards IRAS 2A and IRAS 4A.

Number	Frequency (GHz)	Transition	E_{up} (K)	A_{ul} (s ⁻¹)	Flux (IRAS 2A) (Jy km s ⁻¹)	Flux (IRAS 4A) (Jy km s ⁻¹)
1	166.435	4 ₀ –3 ₀ e ₀	20.0	1.67(–5)	0.245 ± 0.066	0.151 ± 0.028
2	166.327	4 ₀ –3 ₀ o ₁	28.8	1.67(–5)	0.248 ± 0.107	0.129 ± 0.030
3	166.234	4 ₀ –3 ₀ e ₁	38.4	1.66(–5)	0.272 ± 0.064	0.139 ± 0.033
4	166.271	4 _{2–} –3 _{2–} e ₁	51.3	1.25(–5)	0.184 ± 0.066	0.106 ± 0.026
5	166.304	4 ₂₊ –3 ₂₊ e ₁	51.3	1.25(–5)	0.147 ± 0.045	0.074 ± 0.021
6	166.297	4 ₃₊ –3 _{3–} e ₁	67.0	7.14(–6)	0.176 ± 0.055	0.048 ± 0.020
	166.298	4 _{3–} –3 _{3–} e ₁	67.0	7.14(–6)		

Table A.4. Line parameters of CH₂DCN lines observed towards IRAS 2A and IRAS 4A.

Number	Frequency (GHz)	Transition	E_{up} (K)	A_{ul} (s ⁻¹)	Flux (IRAS 2A) (Jy km s ⁻¹)	Flux (IRAS 4A) (Jy km s ⁻¹)
1	224.754530	13 _{1,13} –12 _{1,12}	80.9	9.72(–4)	0.159 ± 0.032	0.081 ± 0.030
2	225.723769	13 _{3,11} –13 _{3,10}	124.4	9.38(–4)	0.130 ± 0.030	0.089 ± 0.018
3	225.724053	13 _{3,10} –13 _{3,9}	124.4	9.38(–4)	0.130 ± 0.030	0.089 ± 0.018
4	225.726540	13 _{3,12} –13 _{3,11}	97.4	9.68(–4)	0.162 ± 0.037	0.069 ± 0.018
5	225.781540	13 _{2,11} –13 _{2,10}	97.4	9.98(–4)	0.162 ± 0.037	0.069 ± 0.020

Table A.5. Line parameters of CH₂DOH lines observed towards HH212.

Number	Frequency (GHz)	Transition	E_{up} (K)	A_{ul} (s ⁻¹)	Flux (HH212) (Jy km s ⁻¹)
1	348.16076	4 _{1,3} e ₁ –4 _{0,4} e ₀	38.1	2.03(–4)	0.152 ± 0.031
2	338.95711	6 _{1,6} e ₀ –5 _{0,5} e ₀	48.4	1.59(–4)	0.157 ± 0.031
3	337.34866	9 _{0,9} e ₀ –8 _{1,8} o ₀	96.3	1.46(–4)	0.159 ± 0.032
4	350.09024	5 _{4,2} e ₁ –5 _{3,2} o ₁	104.2	7.58(–5)	0.065 ± 0.065
5	350.09038	5 _{4,1} e ₁ –5 _{3,3} o ₁	104.2	7.58(–5)	0.065 ± 0.065
6	350.02735	6 _{4,3} e ₁ –6 _{3,3} o ₁	117.1	9.06(–5)	0.078 ± 0.016
7	350.02777	6 _{4,2} e ₁ –6 _{3,4} o ₁	117.1	9.06(–5)	0.078 ± 0.016
8	349.95168	7 _{4,4} e ₁ –7 _{3,4} o ₁	132.1	1.00(–4)	0.201 ± 0.040
9	338.86898	13 _{1,12} e ₀ –12 _{0,12} o ₁	201.8	2.81(–5)	0.080 ± 0.016
10	347.76728	18 _{4,15} e ₁ –18 _{3,15} o ₁	438.2	1.31(–4)	0.046 ± 0.009
11	347.95281	18 _{4,14} e ₁ –18 _{3,16} o ₁	438.2	1.31(–4)	0.057 ± 0.011

Table A.6. Line parameters of the CH₃OD line observed towards HH212.

Number	Frequency (GHz)	Transition	E_{up} (K)	A_{ul} (s ⁻¹)	Flux (HH212) (Jy km s ⁻¹)
1	335.089661	6 ₂ –6 ₁ +	67.4	1.65(–4)	0.025 ± 0.003

Appendix B: Column density and dust temperature maps

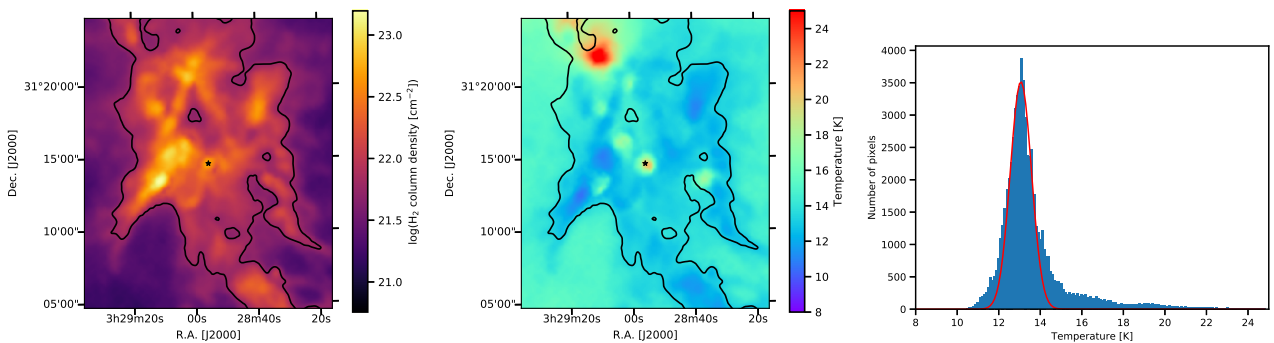


Fig. B.1. H₂ column density (*left*), dust temperature (*centre*) maps, and dust temperature histogram of the map pixels with $A_V > 3$ mag (*right*) of the NGC 1333 star-forming region surrounding IRAS 2A as observed with the *Herschel* Space Observatory by the Gould Belt survey (André et al. 2010). The black contour in the maps depicts the 3 mag level. The red curve in the histogram depicts the Gaussian fit of the histogram around the histogram peak.

Appendix C: Comparison with model predictions

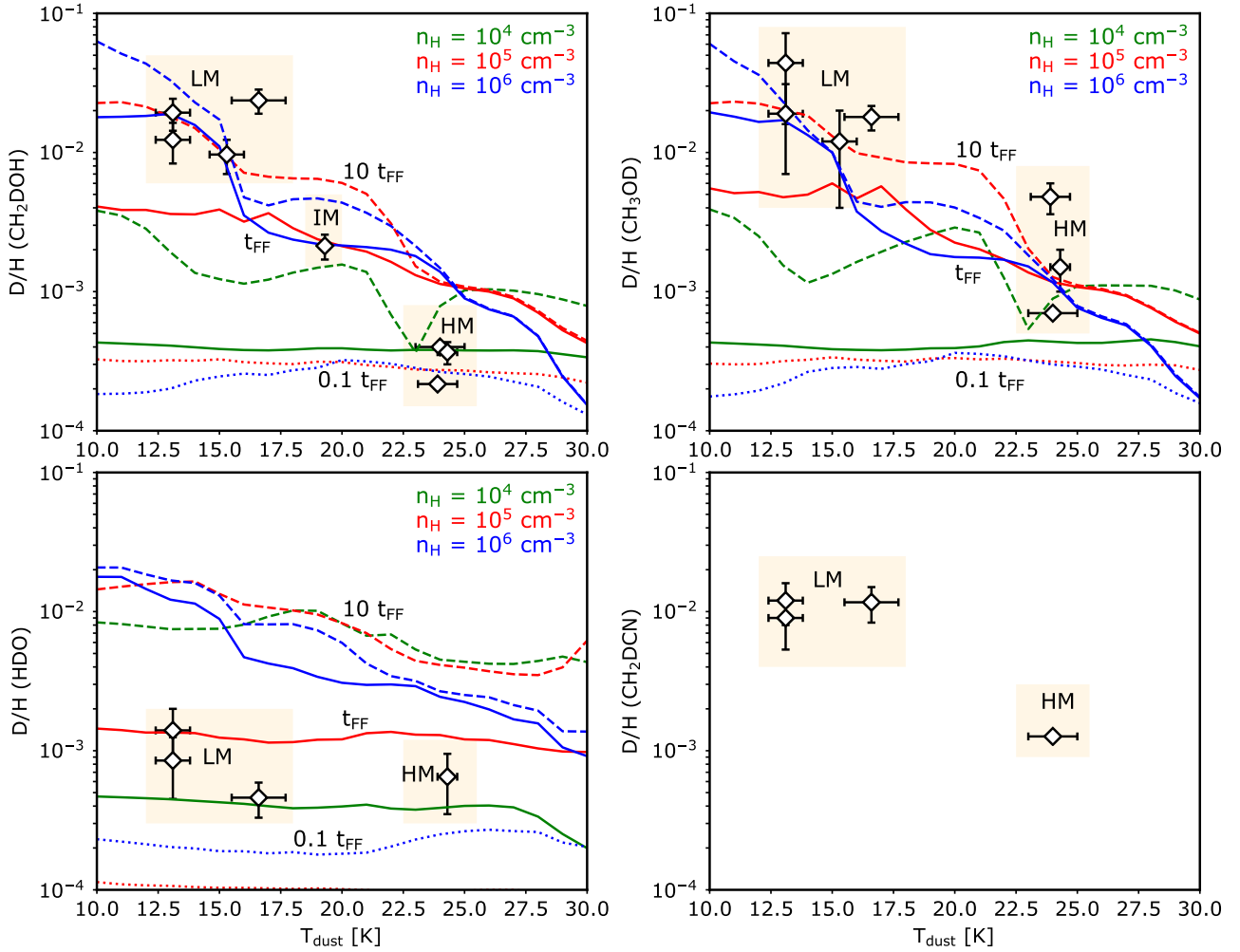


Fig. C.1. Statistical (i.e. taking into account the statistical ratios for the functional groups with multiple H-atoms) CH_2DOH (top left), CH_3OD (top right), HDO (bottom left), and CH_2DCN (bottom right) deuterium fractionations observed towards the hot core of a sample of low-mass, intermediate-mass, and high-mass protostars with sub-millimeter interferometers as a function of the dust temperature of the surrounding cloud measured with the *Herschel* or *Planck* telescopes (see values in Table 3). Green, red, and blue curves show the deuterations predicted by the GRAINOBLE astrochemical model in ices as a function of the dust temperature assumed in the dark cloud simulation at $n_{\text{H}} = 10^4$, $n_{\text{H}} = 10^5$ and 10^6 cm^{-3} , respectively, for at $0.1 \times t_{\text{FF}}$ (dotted lines), t_{FF} (solid), and $10 \times t_{\text{FF}}$ (dashed), where t_{FF} is the free-fall time for the corresponding density. No predictions are given for the CH_3CN deuteration because this molecule is thought to be mostly produced in the vicinity of protostars either on ices or in the gas phase through warmer chemistry. LM, IM, and HM stand for low-mass, intermediate-mass, and high-mass protostars, respectively. See text for more details.