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The MeerKAT Fornax Survey

II. The rapid removal of H I from dwarf galaxies in the Fornax cluster

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ABSTRACT

We present MeerKAT Fornax Survey atomic hydrogen (H I) observations of the dwarf galaxies located in the central $\sim 2.5 \times 4 \text{ deg}^2$ of the Fornax galaxy cluster ($R_{\text{vir}} \sim 2^\circ$). The H I images presented in this work have a 3σ column density sensitivity between 2.7 and $50 \times 10^{18} \text{ cm}^{-2}$ over 25 km s^{-1} for spatial resolution between 4 and 1 kpc. We are able to detect an impressive $M_{\text{HI}} = 5 \times 10^5 M_\odot$ 3σ point source with a line width of 50 km s^{-1} at a distance of 20 Mpc. We detected H I in 17 out of the 304 dwarfs in our field, with 14 out of the 36 late-type dwarfs (LTDs) and three out of the 268 early-type dwarfs (ETDs). The H I-detected LTDs have likely just joined the cluster and are on their first infall as they are located at large clustocentric radii, with comparable M_{HI} and mean stellar surface brightness at fixed luminosity as blue, star-forming LTDs in the field. By contrast, the H I-detected ETDs have likely been in the cluster longer than the LTDs and acquired their H I through a recent merger or accretion from nearby H I. Eight of the H I-detected LTDs host irregular or asymmetric H I emission and disturbed or lopsided stellar emission. There are two clear cases of ram pressure shaping the H I, with the LTDs displaying compressed H I on the side closest to the cluster centre and a one-sided, starless tail pointing away from the cluster centre. The H I-detected dwarfs avoid the most massive potentials (i.e. cluster centre and massive galaxies), consistent with massive galaxies playing an active role in the removal of H I. We created a simple toy model to quantify the timescale of H I stripping in the cluster by reproducing the observed $M_{\text{HI}}-M_{\text{HI}}$ relation. We find that a $M_{\text{HI}} = 10^8 M_\odot$ dwarf is stripped in $\sim 240 \text{ Myr}$. The model is consistent with our observations, where low-mass LTDs are directly stripped of their H I from a single encounter and more massive LTDs can harbour a disturbed H I morphology due to longer times or multiple encounters being required to fully strip their H I. This is the first time dwarf galaxies with $M_{\text{HI}} \lesssim 1 \times 10^6 M_\odot$ have been detected and resolved beyond the local group and in a galaxy cluster.

Key words. galaxies: general – galaxies: clusters: individual: Fornax – galaxies: dwarf – galaxies: evolution – galaxies: interactions – radio lines: ISM

1. Introduction

Dwarf galaxies are the most abundant type of galaxy throughout the Universe. In general, dwarf galaxies have weak gravity, a metal poor interstellar medium (ISM), low interstellar pressure, insufficient dust shielding, and the disks tend to be thick and diffuse (see [Henkel et al. 2022](#), for a recent review). Due to these properties, dwarf galaxies are the most sensitive to internal processes, such as star formation and feedback, and non-secular processes, such as tidal and hydrodynamical interactions that strip the ISM and distort the stellar body. This makes them the best candidates to test and understand the rapid and efficient processes that evolve star-forming late-type dwarfs (LTDs) into quiescent early-type dwarfs (ETDs). It is well established that the dense environment of galaxy clusters provide hostile conditions that efficiently produce quiescent galaxies of any mass via ISM

removal and the quenching of star formation (for a comprehensive review, see e.g. [Boselli & Gavazzi 2006, 2014](#); [Cortese et al. 2021](#); [Boselli et al. 2022](#)). Direct stripping of the cold ISM either through hydrodynamical, tidal, or a combination of both forces is required during cluster infall to explain the lack of ISM observed in cluster galaxies (e.g. [Valluri & Jog 1990, 1991](#); [Mayer et al. 2006](#); [De Rijcke et al. 2010](#); [Gavazzi et al. 2018](#); [Boselli et al. 2022](#)). Both observations and simulations show that low-mass galaxies falling into clusters have their ISM fully stripped close to or just past their first pericentric passage, resulting in a more rapid quenching (e.g. [Boselli et al. 2008, 2014a](#); [Pasquali et al. 2019](#); [Lotz et al. 2019](#)).

Atomic neutral hydrogen (H I) remains one of the best observational tracers for measuring environmentally driven galaxy transformations (e.g. [Cortese et al. 2010, 2011](#); [Catinella et al. 2018](#); [de Blok et al. 2018](#)). However, dwarf galaxies in all

environments are of a low mass, have small angular sizes, and are often diffuse in nature. Therefore, being able to simultaneously detect and resolve HI in these galaxies has remained laborious for many years. This challenge is even more acute inside galaxy clusters given the aforementioned processes of ISM removal. Despite these observational constraints, there are multiple studies in the literature that have detected HI in dwarf galaxies in a variety of environments. Surveys such as the Faint Irregular Galaxies GMRT Survey (FIGGS; Begum et al. 2008), the Local Irregulars That Trace Luminosity Extremes – The HI Nearby Galaxy Survey (LITTLE-THINGS; Hunter et al. 2012), and the Local Volume HI Survey (LVHIS; Koribalski et al. 2018) have detected HI in star-forming LTDs in the field and groups down to $M_{\text{HI}} \sim 10^7 M_{\odot}$ within 10 Mpc. Similarly, Kovač et al. (2009) conducted a blind HI survey of the Canes Venatici (CVn) volume (≤ 18 Mpc) that detected HI in 70 star-forming LTDs, with a HI mass range of $10^{6.5} - 10^{9.8} M_{\odot}$. The observations of nearby LTDs have shown that scaling relations such as the HI-to-optical-diameter ratio hold true in this mass range and that LTDs in low-density environments can host asymmetric and irregular HI morphologies. This is further supported by the irregular morphology and asymmetric velocity fields in the outer HI disk of 18 starbursting LTDs that was likely caused by recent interactions (Lelli et al. 2014). In the Virgo galaxy cluster, ~ 80 dwarfs between $M_{\text{HI}} = 10^{6.9} - 10^{7.7} M_{\odot}$ (with 25% completeness at the low-mass end) were detected in HI by the ALFALFA survey (Giovannelli et al. 2005; Huang et al. 2012; Hallenbeck et al. 2012). The majority were LTDs; although, a few were ETDs akin to transitional galaxies, which have little ongoing star formation despite the detectable HI reservoir (Hallenbeck et al. 2017). Recently, Poulain et al. (2022) used the ALFALFA and ATLAS^{3D} HI (Cappellari et al. 2011; Serra et al. 2012) surveys to detect HI in 145 dwarf satellite galaxies observed in the Mass Assembly of early-Type GaLaxies with their fine Structures (MATLAS) deep imaging survey and found that $\sim 80\%$ of dwarf satellites increase their HI mass as a function of projected distance from the central galaxy. While progress has been made in understanding the nature of HI in dwarf galaxies, the aforementioned work has been limited to the local volume with restrictions on the combined HI sensitivity and resolution.

If dwarf galaxies are observed in the dense cluster environment, we are able to constrain galaxy quenching mechanisms on galaxies that are transformed and quenched the quickest. The Fornax cluster is a small, low-mass ($M_{\text{vir}} = 5 \times 10^{13} M_{\odot}$; Drinkwater et al. 2001) cluster that is located 20 Mpc away (Blakeslee et al. 2009), making it the second closest galaxy cluster to Earth after Virgo. Fornax is more dynamically evolved than Virgo; although, it is not virialized as there are massive star-forming late types present as well as in-falling groups (e.g. NGC 1365 and NGC 1316 Drinkwater et al. 2001; Raj et al. 2019; Kleiner et al. 2021; Loni et al. 2021). The intracluster medium (ICM) of Fornax is half the density; although, the galaxy density is twice as high in comparison to Virgo (e.g. Jordán et al. 2007), implying that in-falling galaxies experience a lower ram-pressure force but are prone to tidal interactions with cluster galaxies. The galaxy population of Fornax is dominated by red, quiescent ETDs, implying that the cluster is driving the transformation of LTDs into ETDs (Drinkwater et al. 2001; Schröder et al. 2001; Waugh et al. 2002; Venhola et al. 2018; Loni et al. 2021).

In this paper we present a multi-wavelength analysis of the dwarf galaxies located in the central $\sim 2.5 \times 4 \text{ deg}^2$ of the Fornax galaxy cluster ($R_{\text{vir}} \sim 2^\circ$), focussing on the HI content that has been observed in the MeerKAT Fornax Survey (MFS;

Serra et al. 2023). The MeerKAT telescope (Camilo et al. 2018) has exquisite column density sensitivity and resolution. For this work, we detected and resolved dwarf galaxies for the first time with $M_{\text{HI}} \lesssim 10^6 M_{\odot}$ at a distance of 20 Mpc. Previously, that had only been possible for Milky-Way satellites and dwarf galaxies in the local group. We made use of the deep optical imaging from the Fornax Deep Survey (FDS; Iodice et al. 2016; Venhola et al. 2018; Peletier et al. 2020) to use the morphological and photometric properties of the dwarfs as well as their spatial distribution in the cluster. Our main goals are to determine which dwarfs have HI, which dwarfs are being transformed in the cluster, how they are being transformed, and on what timescale.

This paper is organised as follows: Section 2 describes the MFS observations and data reduction used in this work and the relevant optical images and catalogues. In Sect. 3, we define our dwarf galaxy sample. We present our main results in Sect. 4, where we display the HI images and show how the HI relates to the optical colour, luminosity, and mean surface brightness, along with the location of the HI detection relative to the cluster substructure and distance to the nearest massive galaxy. In Sect. 5 we discuss the mechanisms that remove HI from dwarf galaxies in the cluster, and in Sect. 6 we make a simple toy model to measure the timescale of HI removal. Finally, we summarise our results in Sect. 7. Throughout this paper, we assume a luminosity distance of 20 Mpc to the Fornax cluster and all associated galaxies (Blakeslee et al. 2009). At this distance, $1''$ corresponds to 97 pc.

2. Observations

2.1. MeerKAT radio data

With reference to Table 2 of Serra et al. (2023), in this work we measure HI parameters (e.g. flux, line width) from the MFS data products¹ at a resolution of $66''$, that has a 3σ column density sensitivity of $1.3 \times 10^{18} \text{ cm}^{-2}$ over 25 km s^{-1} , and is the most likely to recover all of the HI flux. Given the availability of multiple resolutions sensitive to different HI column densities, we present the HI images with a resolution of $11''$, $21''$ and $41''$, which cover a 3σ column density sensitivity range of $2.7 - 50 \times 10^{18} \text{ cm}^{-2}$ over 25 km s^{-1} for spatial resolutions between 4 and 1 kpc. The $41''$ cube has the lowest rms = $0.24 \text{ mJy beam}^{-1}$, that enables us to detect a 3σ point source with a line width of 50 km s^{-1} down to a $M_{\text{HI}} = 5 \times 10^5 M_{\odot}$ at the distance of the Fornax cluster.

We limit our analysis to the portion of the MFS that has been completed at the time of writing. To avoid unreliable HI parameterisation and poor image fidelity due to the rapidly changing rms at the mosaic edge, we only consider sources where the noise is at less than twice the minimum rms. The resulting field is $\sim 2.5 \times 4 \text{ deg}^2$ with our most sensitive 3σ HI mass limit of $M_{\text{HI}} = 5 \times 10^5 M_{\odot}$ and least sensitive 3σ HI mass limit of $M_{\text{HI}} = 1 \times 10^6 M_{\odot}$ for a 50 km s^{-1} point source.

2.2. Optical data

We use deep optical images from the FDS (Iodice et al. 2016; Venhola et al. 2018), which are excellent for detecting dwarfs and resolving low surface brightness galaxies. The FDS observed the Fornax cluster using the OmegaCAM wide-field imager on the VLT Survey Telescope (VST). Observations

¹ Available at the MFS website – <https://sites.google.com/inaf.it/meerkatfornaxsurvey>

were taken in the Sloan Digital Sky Survey (SDSS) u' , g' , r' , and i' -bands and are sensitive down to an impressive surface brightness limit of 28.4 mag arcsec² in the g' -band. We use the photometric catalogue from Su et al. (2021), who used a combination of aperture photometry, Sérsic + point spread function (PSF) and multi-component decompositions to quantify the light distribution of each massive cluster galaxy (Iodice et al. 2019; Raj et al. 2019, 2020) and the dwarf galaxies (Venholá et al. 2018).

One caveat of the FDS is that the VST is unable to image the Fornax cluster where extremely bright foreground stars are present. There are two dwarfs we detect in H I (ESO 358–G064 and FCC 323) located in a region of the cluster where FDS images are either unavailable or not observed in all bands. To rectify this, we use images and the photometric catalogue from the Dark Energy Camera Legacy Survey (DECaLS) DR9² (Dey et al. 2019), in the same SDSS filters. We adopt the same method as Su et al. (2021) and use the photometric-colour relation to estimate the stellar mass (Taylor et al. 2011), and aperture photometry to measure the projected surface brightness of ESO 358–G064 and FCC 323. The surface brightness limit of the DECaLS g' -band images in this area is ~ 27.9 mag arcsec², slightly shallower than the FDS does not present any issue for this work.

3. Dwarf galaxy sample definition

There are a variety of definitions in the literature of what constitutes a dwarf galaxy, and a common method is to use a simple magnitude cut (e.g. Hallenbeck et al. 2012) in a given volume. Venholá et al. (2018) created a catalogue of Fornax cluster dwarf galaxies in the FDS, using the definition of $M_r > -18.5$ with a minimum size of 2'', and separated the cluster dwarfs from background galaxies using the colour–magnitude, luminosity–radius and luminosity–concentration relations. We use this catalogue as the starting point in this work. However, the full catalogue contains some galaxies that are arguably not dwarfs. The best example is NGC 1427A, which may indeed appear as a dwarf given its irregular morphology, although it has a stellar and H I mass of $\sim 2 \times 10^9 M_\odot$ (Lee-Waddell et al. 2018; Loni et al. 2021), making it quite massive for a dwarf. The Venholá et al. (2018) catalogue used a combination of visual and parametric classifications to define the first order morphology of the Fornax cluster dwarfs.

It has been suggested that at $M_\star \approx 10^9 M_\odot$, galaxies transition from dwarfs to giants due to real structural differences (Watkins et al. 2023). In this work we are interested in the low-mass galaxy population that is the most efficiently transformed by the cluster environment. Therefore, using the stellar masses measured by Su et al. (2021), we define dwarf galaxies as $M_\star < 10^9 M_\odot$. We also define intermediate galaxies to be $10^9 \leq M_\star < 10^{10} M_\odot$ which are intended to represent the populations of galaxies that are neither obvious dwarfs nor obvious giant galaxies. Lastly, we define the massive galaxies those with $M_\star \geq 10^{10} M_\odot$. There are 304 dwarfs (36 LTDs and 268 ETDs), 13 intermediate mass and 15 massive galaxies within our field. We show the stellar mass distribution for each sample in Fig. 1 and their basic sample statistics in Table 1.

4. H I in cluster dwarfs

We detect H I in 17 out of the 304 dwarfs within our field and show their position in the cluster in Fig. 2. The detections include

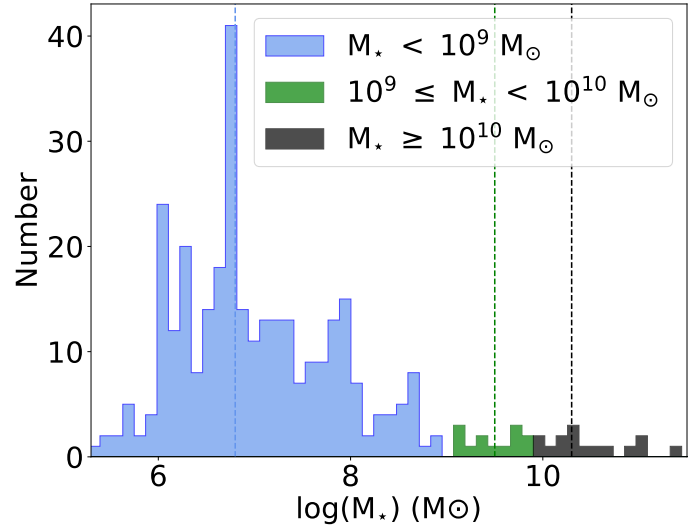


Fig. 1. Stellar mass distribution for our different galaxy samples. Each sample is defined by their stellar mass estimates (Su et al. 2021). The dwarfs are shown in light blue, the intermediate galaxies are in green and the massive galaxies are in black. The vertical dashed lines for each respective colour show the median of the samples.

Table 1. Number of galaxies, median, and standard deviation of the stellar mass for our galaxy samples.

Sample	Ngals	Median $M_\star$ ($M_\odot$)	Std. $M_\star$ ($M_\odot$)
Dwarf	304	6.3×10^6	1.1×10^8
Intermediate	13	3.2×10^9	1.9×10^9
Massive	15	2.0×10^{10}	6.1×10^{10}

Notes. Readers can refer to Fig. 1 for their stellar mass distribution.

14 out of the 36 LTDs, and 3 out of the 268 ETDs. We present a preview in Fig. 3 and all the dwarfs detected in H I in Fig. A.1 that shows the DECaLS 3-colour postage stamp, the FDS optical image overlaid with H I contours from the 11'', 21'' and 41'' MFS images and the 21'' velocity field. The basic H I properties for the H I detected dwarf galaxies are presented in Table 2 and the optical morphology, H I mass, stellar mass, photometry and projected distance to nearest massive galaxy in Table 3. Eight of the LTDs with H I show clear signs of irregular or asymmetric H I emission. We denote these galaxies as being (H I) disturbed as the irregularities can be seen in the form of H I tails, significant warps or non-rotating components (Fig. A.1) making them the most likely LTDs currently being shaped by the cluster environment. It is possible that more LTDs are disturbed; however, their signs are subtler. We refrain from labelling the ETDs as disturbed as H I-bearing ETDs are an unusual class of galaxy at a different evolutionary stage compared to LTDs.

4.1. New H I dwarfs and features

The following nine LTDs were detected in Loni et al. (2021): NGC 1437A, ESO 358–G015, ESO 358–G016, ESO 358–G060, FCC 090, FCC 102, FCC 120, FCC 306 and FCC 323. Therefore, this work has detected H I for the first time in an additional eight dwarfs; five LTDs – ESO 358–G064, FCC 128, FCC 299, FCC B905, and FDS6 100b; and three ETDs – FCC 134, FCC 207, and FCC 248.

² <https://www.legacysurvey.org/>

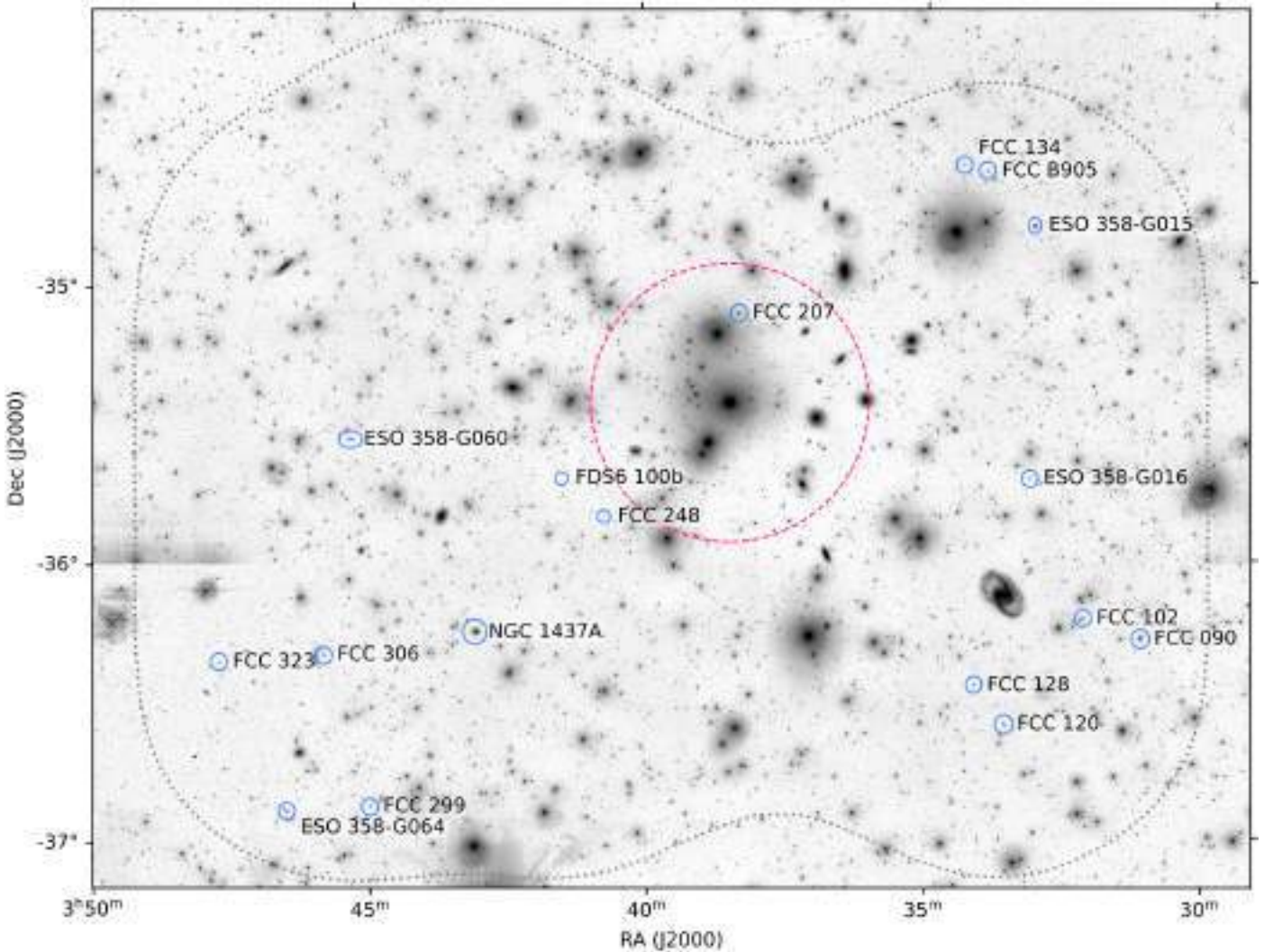


Fig. 2. HI-detected dwarfs in the Fornax cluster. The background is the FDS g' -band optical image, except in the region of FCC 323 that has been filled in with DECaLS images. The dwarfs detected in HI are shown in the blue ellipses with their respective labels. The dotted grey border denotes the $\sim 2.5 \times 4 \text{ deg}^2$ area analysed in this work. The pink dashed circle shows $0.25 R_{\text{vir}}$.

Due to the exquisite sensitivity and resolution of MeerKAT, we have detected and resolved new HI features in some previously HI-detected dwarfs. Here we describe the HI features of the disturbed LTDs. NGC 1437A and FCC 306 have starless, HI tails pointing away from the cluster centre, with velocity gradients aligned with the direction of the ram-pressure wind (Serra et al. 2023). In Fig. A.1, the 11'' image show a HI compression into the edge of the stellar body on the side of the galaxy pointing towards the cluster centre, exactly where the ram-pressure force is the strongest. ESO 358-G016 may also have a visible HI compression in the 11'' image and potentially a HI tail in the 21'' image on the opposite side of the compression. FCC 102 shows a compression on the west in the 11'' image as well as a potential tail in the form of extended HI emission to the east at all resolutions. FCC 102 has a prominent S-shaped iso-velocity contour which is the discernible sign of a changing position angle (PA) due to a warp. While FCCB 905 is not as well resolved as NGC 1437A and FCC 306, the 11'' image does suggest the HI is being compressed on the side that faces the cluster centre and there may be a small HI tail pointing in the opposite direction. ESO 358-G015 has asymmetric HI emission visible in the 21'' image and a U-shaped

iso-velocity contour. ESO 358-G060 has some clear asymmetries in the 11'' image with what appears to be an S-shaped iso-velocity contour containing kinks, also likely due to a warp. The well resolved velocity field of FCC 120 does not show any clear rotation, which is unusual as it appears to be a blue, edge-on LTD.

The stellar emission is also clearly irregular for some of the HI-disturbed LTDs (Fig. A.1). The following have been analysed by Raj et al. (2019): NGC 1437A, ESO 358-G015, ESO 358-G016, ESO 358-G060 and FCC 306. The stellar emission of NGC 1437A is asymmetric (Raj et al. 2019) and the DECaLS image (Fig. A.1) shows a diffuse stellar component in the north that may also be part of a halo. ESO 358-G015 has a lopsided stellar tail pointing towards the cluster centre, ESO 358-G016 appears to have a disturbed outer stellar disk, ESO 358-G060 has a warped stellar disk with irregular star-forming regions and FCC 306 is asymmetric with a disrupted outer stellar disk (Raj et al. 2019). The stellar emission for FCC B905 is clearly lopsided and this is also true, although to a lesser extent for FCC 102. This is confirmed by the non-parametric morphological indices (e.g. Asymmetry, Clumpiness and Gini) measured by Su et al. (2021) who show that FCC 102

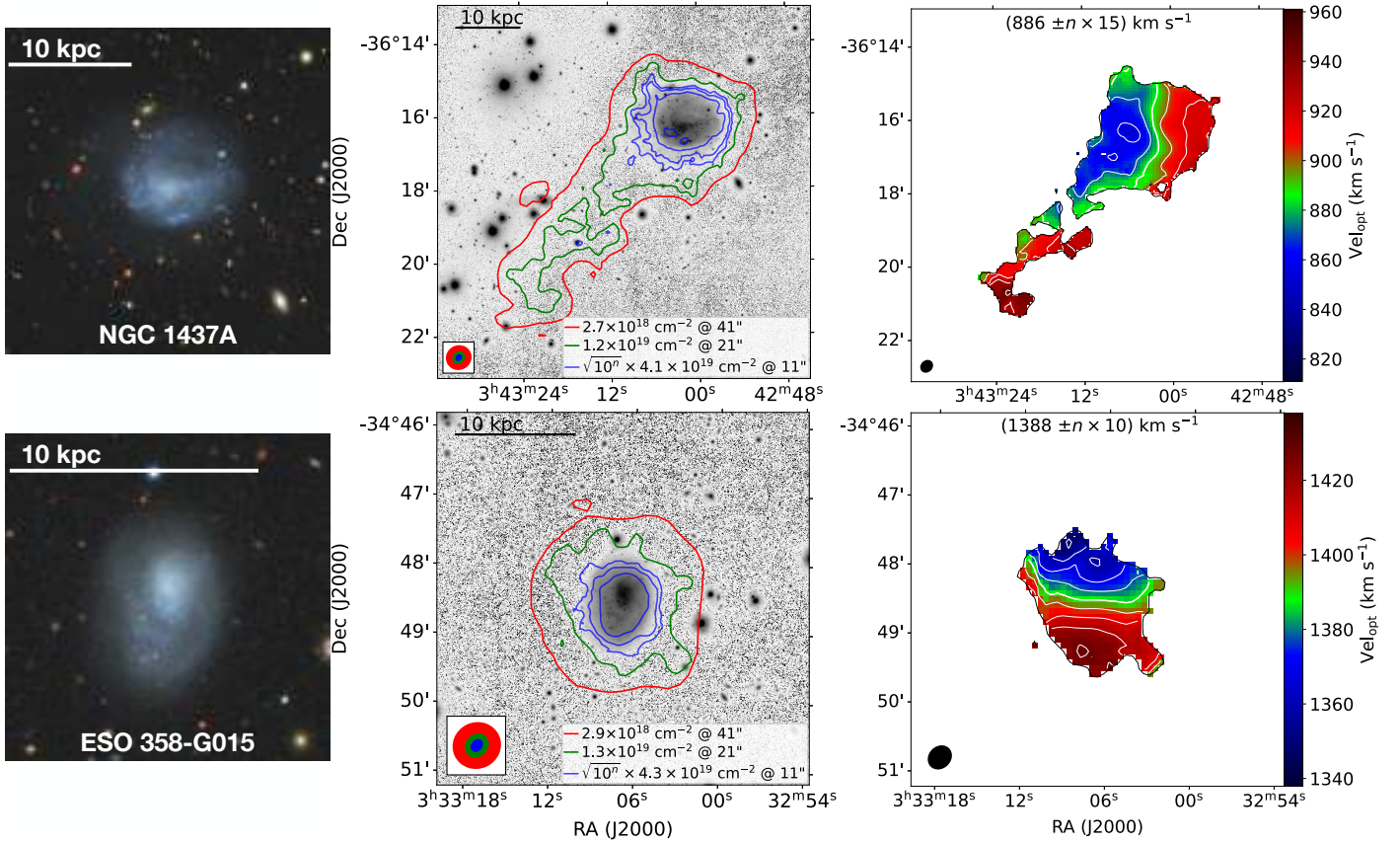


Fig. 3. Preview of the dwarf optical cutouts, multi-resolution MFS HI overlays and velocity fields. The full figure with all the HI detections are shown in Fig. A.1. The left panel shows the optical three-colour postage stamp of each dwarf, the middle panel shows the HI contours overlaid on the FDS (DECaLS for FCC 323) optical image where the blue contours show the 11'' HI emission above 3σ in steps of $\sqrt{10^n}$ and the green and red contours show the 3σ HI emission from the 21'' and 41'' images. The lowest column density sensitivities are labelled in accordance with the local rms over a 25 km s^{-1} line-width and the synthesised beam for each resolution are shown in the bottom left inset with their respective colours. The right panel shows the 21'' HI velocity field. The systemic velocity is the thick white contour with the value shown up the top and the synthesised beam is shown in black on the bottom left.

Table 2. HI properties of dwarfs measured from the 66'' MFS HI cube.

Name	RA (J2000) (hms)	Dec (J2000) (dms)	Integrated flux (Jy km s^{-1})	v_{sys} (km s^{-1})	w_{20} (km s^{-1})	Disturbed
NGC 1437A	03:43:02.42	-36:16:24.2	7.25 ± 0.7	886	128	Y
ESO 358-G015	03:33:07.06	-34:48:21.2	1.50 ± 0.2	1388	133	Y
ESO 358-G016	03:33:09.19	-35:43:05.2	1.26 ± 0.1	1700	131	Y
ESO 358-G060	03:45:12.29	-35:34:15.2	11.99 ± 1.2	803	129	Y
ESO 358-G064	03:46:28.51	-36:54:21.6	5.59 ± 0.7	1465	171	N
FCC 090	03:31:08.28	-36:17:24.4	0.59 ± 0.07	1825	129	N
FCC 102	03:32:10.56	-36:13:10.6	1.19 ± 0.1	1640	83	Y
FCC 120	03:33:34.18	-36:36:20.1	1.08 ± 0.1	846	60	Y
FCC 128	03:34:07.03	-36:27:57.6	0.14 ± 0.02	1543	68	N
FCC 134	03:34:21.70	-34:35:33.0	0.02 ± 0.01	1398	34	N
FCC 207	03:38:19.25	-35:07:44.8	0.0065 ± 0.006	1420	38	N
FCC 248	03:40:43.37	-35:51:38.9	0.010 ± 0.006	1859	15	N
FCC 299	03:44:58.61	-36:53:41.6	0.31 ± 0.04	2100	136	N
FCC 306	03:45:45.38	-36:20:46.0	2.03 ± 0.1	886	176	Y
FCC 323	03:47:37.46	-36:21:48.6	0.43 ± 0.05	1499	67	N
FCC B905	03:33:57.34	-34:36:43.2	0.52 ± 0.06	1248	99	Y
FDS6 100b	03:41:27.84	-35:43:13.1	0.14 ± 0.03	1983	50	N

Notes. The position, integrated flux, and w_{20} are from the SoFiA catalogue, while v_{sys} is estimated from where 11'' HI iso-velocity intercepts the optical centre. The uncertainty in the integrated flux is a combination of the statistical error and 10% uncertainty in the calibration of the flux scale (Serra et al. 2023).

Table 3. Optical morphology, stellar mass, H I mass, photometric properties, and projected distance to nearest massive galaxy for all the LTDs and the H I-detected ETDs.

Name	FCC number	Morphology	H I mass ($10^7 M_\odot$)	Stellar mass $\log(M_\odot)$	R' (mag)	$g' - r'$ (mag)	μ_e (mag arcsec $^{-2}$)	Dist. (kpc)
NGC 1437A	285	1	68.4 ± 6.9	8.8	-18.02	0.49	23.00	151
ESO 358-G015	113	1	14.2 ± 1.4	8.5	-17.03	0.5	22.87	181
ESO 358-G016	115	1	11.9 ± 1.2	7.7	-15.43	0.43	24.45	150
ESO 358-G060	302	1	113.2 ± 11.3	7.8	-16.02	0.15	24.92	148
ESO 358-G064	B1899	1	5.6 ± 0.7	8.5	-16.20	0.59	23.13	415
FCC 090	90	1	5.6 ± 0.6	8.9	-17.60	0.56	21.23	180
FCC 102	102	1	11.2 ± 1.2	7.8	-15.80	0.47	23.47	103
FCC 120	120	1	10.2 ± 1.0	7.9	-15.72	0.43	24.28	161
FCC 128	128	1	1.3 ± 0.2	8.0	-15.50	0.55	23.63	118
FCC 134	134	e	0.2 ± 0.1	7.7	-14.65	0.53	23.32	96
FCC 207	207	e	0.06 ± 0.03	8.5	-16.55	0.64	21.98	112
FCC 248	248	e	0.1 ± 0.06	7.2	-13.44	0.6	23.79	160
FCC 299	299	1*	2.9 ± 0.4	7.8	-15.18	0.5	23.83	372
FCC 306	306	1*	19.1 ± 2.0	8.1	-16.00	0.36	22.01	227
FCC 323	323	1	4.0 ± 0.5	7.4	-14.41	0.39	24.55	331
FCC B905	B905	1	4.9 ± 0.6	7.8	-15.12	0.49	23.69	125
FDS6 100b	–	1	1.3 ± 0.2	6.7	-13.10	0.43	24.72	128
FCC 247	247	1	≤ 0.06	7.2	-14.11	0.31	24.54	131
FDS2 393	–	1	≤ 0.07	6.3	-11.34	0.36	24.88	117
FDS5 010	–	1	≤ 0.06	5.7	-9.41	0.69	26.15	119
FDS5 137	–	1	≤ 0.06	6.6	-12.31	0.48	24.07	155
FDS5 159b	–	1	≤ 0.06	5.7	-9.64	0.56	26.17	223
FDS5 608b	–	1	≤ 0.10	5.7	-10.34	0.5	25.71	395
FDS7 021	–	1	≤ 0.08	6.3	-11.72	0.61	24.54	411
FDS7 085	–	1	≤ 0.06	6.0	-10.99	0.28	24.66	317
FDS7 275	–	1	≤ 0.06	6.1	-11.54	0.19	24.31	223
FDS7 461	–	1	≤ 0.06	6.3	-11.34	0.54	26.36	174
FDS10 130	–	1	≤ 0.08	6.1	-11.13	0.46	25.30	118
FDS10 167	–	1	≤ 0.09	5.5	-9.97	0.61	27.13	77
FDS11 403	–	1	≤ 0.06	6.4	-12.07	0.27	26.15	102
FDS11 411b	–	1	≤ 0.06	6.2	-11.42	0.56	24.38	106
FDS12 052	–	1	≤ 0.09	6.8	-12.98	0.4	23.84	287
FDS12 135	–	1	≤ 0.07	5.6	-10.12	0.33	25.64	274
FDS15 085	–	1	≤ 0.06	5.8	-10.03	0.25	26.54	58
FDS15 148	–	1	≤ 0.07	6.1	-11.02	0.38	24.76	122
FDS15 155	–	1	≤ 0.06	6.0	-11.25	0.26	24.92	184
FDS16 141	–	1	≤ 0.06	6.0	-11.21	0.4	25.04	105
FDS20 100	–	1	≤ 0.11	6.9	-13.24	0.29	23.98	285
FDS20 502	–	1	≤ 0.09	5.9	-10.25	0.5	24.71	162

Notes. The optical morphologies are taken from (Venhola et al. 2018) and those with an asterisk are nucleated. The uncertainties in the H I mass are propagated from the flux uncertainties in Table 2 and the upper limits are calculated as 3σ of the local rms for a 50 km s^{-1} wide point source. The stellar masses, M_r magnitudes, $g' - r'$ colours and r' -band mean surface brightness out to one effective radius (μ_e) are taken from the Su et al. (2021) catalogue. The photometry for ESO 358-G064 and FCC 323 is taken from the DECaLS DR9 catalogue (Dey et al. 2019).

is more asymmetric and has a higher Gini value compared to FCC 120.

We do not include FCC 090 and FCC 128 as disturbed LTDs, although they have some unusual features. The optical morphology of FCC 090 was originally classified as a peculiar elliptical (Ferguson 1989) although it is a LTD in Venhola et al. (2018). The FDS image shows some structure in the stellar body, such as a possible bar in the core. The H I emission of FCC 090 extends beyond the stellar body with hints of asymmetries and the velocity field does not appear to show rotation. There is a hint of a H I tail for FCC 128 shown in $21''$ and $41''$ images, and the velocity field suggests the presence of non-rotating components, although the H I is not resolved well enough to state this definitively.

4.2. The relation between H I and stellar emission

In the following section, we discuss the galaxy properties listed in Table 3 and how they relate to each other. In Fig. 4, we present the galaxy colour-magnitude diagram. We qualitatively separate the blue cloud from the red sequence by the relation $g' - r' = -0.03r' - 0.1$, that results in 10 out of the 36 LTDs in the red sequence (at the faint end) and 25 out of the 268 ETDs in the blue cloud. We detect all LTDs brighter than $M_r = -14.2$ in H I and as expected, M_{HI} increases for bluer, brighter LTDs. Two out of the three H I-detected ETDs reside on the red sequence, although the third (FCC 134) is significantly bluer and has the ~same colour as the LTDs FCC 090 and ESO 358-G064.

Figure 1 consists of two panels. The top panel is a histogram showing the number of galaxies (y-axis, 0 to 10) versus $\log(M_{\text{HI}})$ (x-axis, -10 to -20). The distribution peaks around $\log(M_{\text{HI}}) \approx -11.5$. The bottom panel is a scatter plot of $\log(M_{\text{HI}})$ (y-axis, 5 to 10) versus M_r (x-axis, -10 to -20). The plot shows various galaxy populations: CVn (grey diamonds), HI-detected LTDs (green circles), HI-detected ETDs (orange stars), Disturbed HI (open circles), and HI-undetected LTDs (blue circles). A color bar on the right indicates the evolution $g' - r'$ from 0.1 (blue) to 0.7 (red). Two dashed lines represent the evolution of the HI mass function, and a red dashed line is at $\log(M_{\text{HI}}) \approx 5.8$.

We show the luminosity-HI relation in Fig. 5 and compare our LTDs and HI-detected ETDs to the CVn dwarf galaxies (Kovac 2007; Kovač et al. 2009). Previous scaling relations have shown that M_{HI} is correlated with luminosity for late types brighter than $M_{r'} = -17$ (e.g. Dénes et al. 2014). The CVn dwarfs extends this trend down to fainter magnitudes of $M_{r'} = -12$. We quantify the luminosity-HI relation of the CVn dwarfs with a linear fit in the range of $-18 < M_{r'} \leq -12$ as our brightest HI-detected LTD has a $M_{r'} = -18$. The dashed and dotted lines shows our linear fit and the 3σ scatter, respectively (Fig. 5). In the luminosity-HI relation, we show the $g'-r'$ colours of our LTDs and HI-detected ETDs. The HI detections show that HI-deficiency is correlated with redder colours. All our HI-detected LTDs are within 3σ of the CVn fit and are therefore comparable to LTDs in the field. The HI-disturbed LTDs are all brighter than $M_{r'} = -15$, have blue colours and mostly reside $\sim$ on the CVn fit, implying that a significant amount of HI has not been lost yet. Irrespective of HI morphology, if a LTD has HI, it is within the scatter of the CVn fit, that is we do not detect any LTDs in HI below the $M_{r'}-M_{\text{HI}}$ relation. There are five HI-undetected, blue LTDs between $M_{r'} = -12$ and -14.2 that are $>3.5\sigma$ from the CVn fit. Given their morphology, luminosity and colour, we would expect to detect these LTDs in HI. The most extreme case is FCC 247, which is 5.8σ below the CVn fit. The upper limits for the HI-undetected LTDs were calculated by using 3σ of the local rms for a 50 km s^{-1} wide point source. The HI-undetected LTDs fainter than $M_{r'} = -12$ are within $\sim 3\sigma$ of the extrapolated CVn fit and are $\sim 50\%$ split between red and blue colours. The HI-detected ETDs are well below 3σ of the CVn fit, although this is unsurprising given that the gas-fraction of ETDs and LTDs are significantly different, independent of luminosity. The histogram in the upper panel of Fig. 5 shows all of the dwarfs (black) and the HI-detected dwarfs (blue) as a function of luminosity in our field and show that there are many dwarfs (mainly ETDs) that span the full luminosity range and are not detected in HI.

HI-detected ETDs in this comparison, even though we expect them to be a different, peculiar type of object. In Fig. 6, we can see that the LTDs and HI-detected ETDs lie within the scatter of the $M_{r'}-\mu_e$ relation of the CVn galaxies. The five HI-undetected, blue LTDs between $M_{r'} = -12$ and -14.2 show typical surface brightness values for their luminosity. They have low stellar masses that range $(0.3-1.6) \times 10^7 M_{\odot}$, making it likely that their HI has been very recently stripped such that neither the stellar structure nor blue stars have had a chance to fade yet. We show the DECaLS 3-colour optical images for these LTDs in Fig. 7.

The number density (i.e. local environment) of galaxies is crucial in shaping their evolution. The mass of the nearby galaxies (i.e. central vs satellite) also plays a vital role in H I removal and star formation quenching (e.g. [Spekkens et al. 2014](#); [Poulain et al. 2022](#); [Zhu & Putman 2023](#)). With this in mind, we create a projected density field to reveal the substructure of the cluster. At the central coordinates of each galaxy in our field, we assign its relative luminosity, normalised to the BCG (NGC 1399), and zero everywhere else. The image is then smoothed by a Gaussian kernel to produce the projected density field shown in Fig. 8. The normalised luminosity factors in brighter (i.e. massive) galaxies

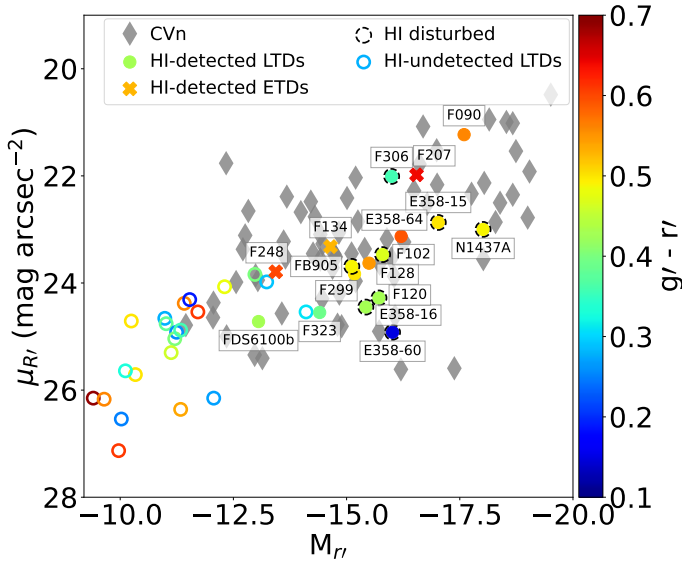


Fig. 6. Luminosity-projected surface brightness relation. The markers and colours are the same as Fig. 5. The surface brightness for all Fornax LTDs and HI-detected ETDs fall within the scatter of the CVn dwarfs. The five HI-undetected LTDs between $M_r = -12$ and -14.2 show typical surface brightness values for their luminosity.

having a stronger influence on nearby galaxies than the standard projected density (e.g. Σ_5). Therefore, this projected density field is a proxy for the cumulative, stellar potential throughout the cluster. In Fig. 8, the massive galaxies (white triangles) are clearly contributing the most to the density field. Easily identifiable features include the BCG (at the centre of the pink circle showing $0.25 R_{\text{vir}}$), the concentration of bright early-type galaxies west of it, and the bright spiral NGC 1365 towards the south-west.

Figure 8 shows that HI-detected dwarfs avoid the most massive potentials, whether that be the cumulative substructure in the cluster centre (and just west of it), or individual massive galaxies. This is consistent with the well-known morphology-density relation (Dressler 1980) and with the Next Generation Fornax Survey (Muñoz et al. 2015), who show a higher number of nucleated (i.e. more evolved) dwarfs in the cluster centre (Ordenes-Briceño et al. 2018a,b). The dwarfs with HI are typically situated at large clustocentric radii (i.e. the majority are beyond $0.5 R_{\text{vir}}$), which holds true for 13 out of the 14 HI-detected LTDs and is what we would expect for LTDs on their first infall. Given that the LTDs with disturbed HI are also far from the cluster core, their HI is being actively shaped well before they reach the cluster centre. This is consistent with the star-forming dwarfs being concentrated in the outer regions of the cluster, shown by Drinkwater et al. (2001). There is a single HI-detected ETD (FCC 207) near the cluster centre, $<0.25 R_{\text{vir}}$ in projection. This is a peculiar case and we discuss it in the next section.

We quantify the projected distance to the closest massive galaxy for all the LTDs in Fig. 9. The HI-detected LTDs are situated at least 100 kpc from their nearest massive galaxy in projection. There are two HI-undetected LTDs within 100 kpc of their nearest massive galaxy. The LTDs with disturbed HI tend to reside closer (e.g. between 100 and 190 kpc) to their nearest massive galaxy when compared to the LTDs with undisturbed HI. While this is not definite evidence, and projection effects introduce a lot of uncertainty, it does support the notion that LTDs

can be stripped of their HI by central galaxies and maintain their structure for a short time. Previous studies have found gas poor and quenched dwarf satellites in low-density environments (such as the local group) and have suggested that centrals have a universal and effective satellite quenching mechanism, such as ram-pressure stripping by the central halo (e.g. Poulain et al. 2022; Zhu & Putman 2023).

5. The mechanisms and timescale of HI removal

There are multiple pieces of evidence that suggest the dwarf galaxies rapidly lose their HI in the Fornax cluster. Firstly, the most abundant galaxies in the Fornax cluster are red, ETDs (Venhola et al. 2018, 2019). However, the cluster is still assembling, which is evident by the substructure in Fig. 8 and the un-settled giant early types, located in a region where many galaxies host asymmetric halos and low surface brightness features (Iodice et al. 2016, 2019; Spavone et al. 2020, 2022). Furthermore, there are different populations of recent infallers; massive late-types that are gas-rich and have disturbed stellar bodies, moderate atomic and molecular gas ratios and star formation rate deficiencies (Raj et al. 2019; Zabel et al. 2019; Loni et al. 2021; Morokuma-Matsui et al. 2022), and a range of star-forming dwarfs, located at the cluster outskirts (Drinkwater et al. 2001). The dwarfs with redshifts have a high velocity dispersion compared to both the cluster and the giants, implying that dwarfs are also currently infalling into the cluster with $\sim 70\%$ on radial orbits (Drinkwater et al. 2001; De Rijcke et al. 2010; Maddox et al. 2019).

As described in Sect. 4, we detect HI in 14 out of the 36 LTDs, including all LTDs brighter than $M_r = -14.2$. Interestingly, all HI-detected LTDs contain the expected amount of HI as LTDs at the same luminosity in the field (e.g. Fig. 5). Because the HI-detected LTDs have a normal M_{HI} and they are located at large clustocentric distances (Fig. 8), they must have recently joined the cluster and are on their first infall. Furthermore, eight of the LTDs show disturbed or asymmetric HI morphologies. Considering that small perturbations, even from secular processes such as stellar feedback are able to induce asymmetries in the stellar and HI disk in field LTDs (e.g. Begum et al. 2008; Hunter et al. 2012; Zhang et al. 2012; Poulain et al. 2022), it is unsurprising that the majority of HI-detected LTDs in Fornax also display stellar and HI asymmetries. The Fornax LTDs are subjected to the same secular processes as those in the field, as well as interacting with the ram-pressure wind of the ICM and massive halos, tidal interactions with nearby galaxies and the potential of the cluster. The Fornax HI-detected LTDs are analogous to field LTDs; although, a clear difference is the unambiguous evidence of the ICM ram-pressure shaping the outer HI disk of some LTDs.

Serra et al. (2023) showed that NGC 1437A and FCC 306 have starless, HI tails pointing away from the cluster centre, with velocity gradients aligned with the direction of the ram-pressure wind. Using higher resolution HI images (Fig. A.1), we show that the HI on the side facing the cluster centre is being compressed into the stellar disk, an essential ingredient required to disentangle ram-pressure and tidal tails. We further show another three disturbed HI-detected LTDs (ESO 358-G016, FCC 102 and FCCB 905) that may also show evidence of ram-pressure shaping as they have potential HI compressions and tails pointing away from the cluster centre. Serra et al. (2023) showed that a combination of ram pressure and tidal forces are required to operate together in the Fornax cluster. First, the HI is pulled out of the stellar body by tidal forces.

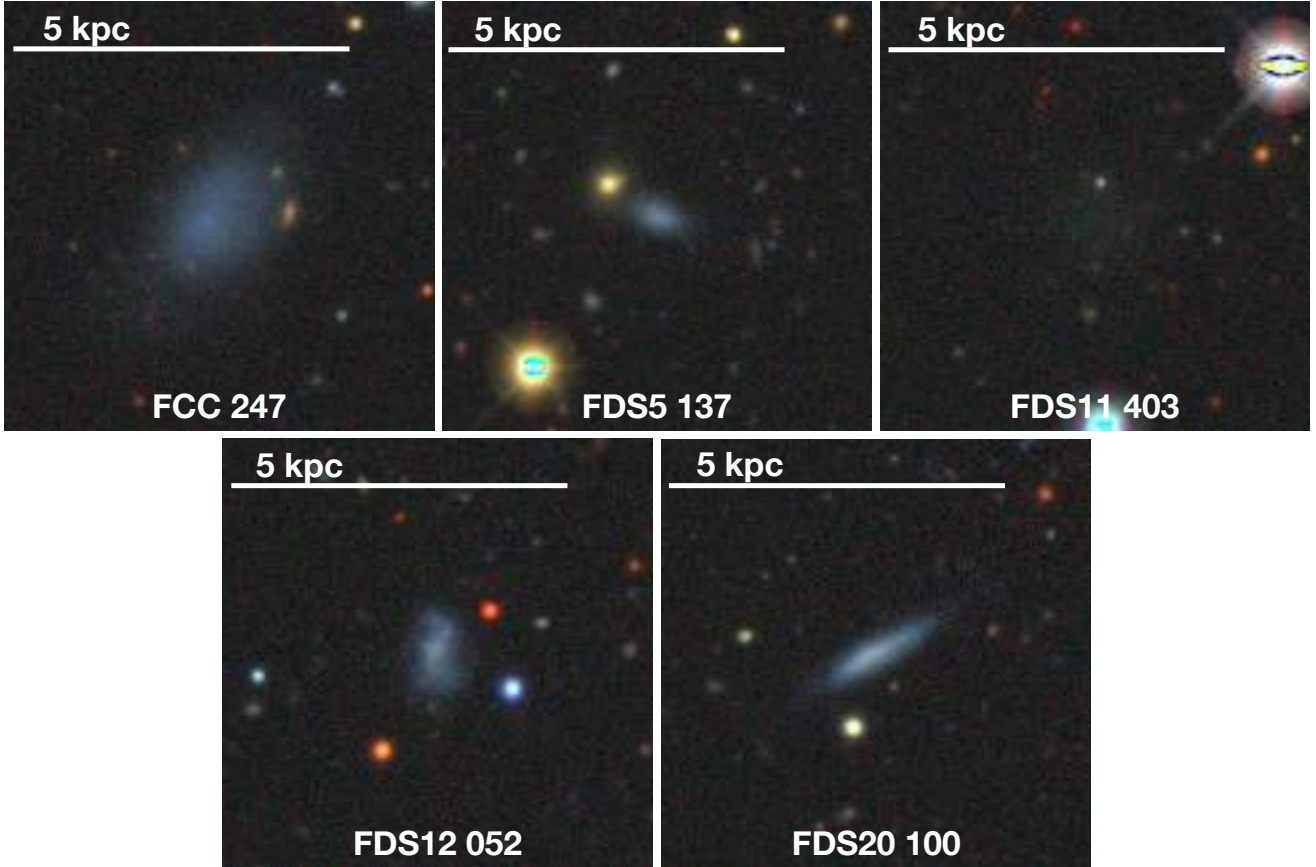


Fig. 7. DECaLS three-colour image of the H I-undetected LTDs. The optical three-colour image for the five blue, LTDs between $M_r = -12$ and -14.2 that we would expect to detect in H I. Their colour, luminosity, morphology and surface brightness is the same as H I-detected LTDs, suggesting that these five LTDs have very recently lost their H I.

Then, this weakly bound H I is further displaced and shaped by ram pressure as the galaxies travel through the ICM (Serra et al. 2023). This is consistent with the low ICM density and high number density of the Fornax cluster along with the asymmetric stellar bodies of the H I-disturbed LTDs. Raj et al. (2019) noted the irregular stellar body of NGC 1437A and suggested that the disturbed stellar emission in FCC 306, ESO 358–G015, ESO 358–G016 and ESO 358–G060 are due to interactions with the cluster potential. This is further supported by the simulations of Mastropietro et al. (2021) who showed that the tidal field of a Fornax-like cluster is able to cause damage to the stellar disks of low-mass galaxies and that perpendicular gaseous and stellar tails are explainable if they are subjected to different environmental effects. Due to the high galaxy number density in Fornax and influence of massive galaxies (Figs. 8 and 9), we cannot rule out tidal interactions between galaxies as a mechanism to distort the stellar bodies. For example, FCC 306 is located next to NGC 1437B in projection, an intermediate mass galaxy. Even though both have disturbed stellar disks, they are separated by 700 km s^{-1} in velocity and there is no H I connecting them down to a column density of 10^{18} cm^{-2} . Another example of tidal interactions between galaxies is that three (ESO 358–G016, FCC 102 and FCC 120) of the five H I-detected LTDs satellites of NGC 1365 have disturbed H I disks and stellar bodies. As previously suggested by Loni et al. (2021), these LTDs may be subjected to both pre-processing and early interactions with the cluster.

Regarding the H I-undetected LTDs, 17 out of 22 are fainter than $M_r = -12$ and they lie within the 3σ scatter of the CVn luminosity–H I relation (Fig. 5). They also have reasonable mean surface brightness values that follow the trend in Fig. 6. These LTDs are so faint that even if they do contain H I along the extrapolated CVn relation, it is below our detection limit. We would expect to detect H I in the five remaining LTDs (between $M_r = -12$ and -14.2) given their morphology, colour, luminosity and surface brightness. However, they all lie $>3.5\sigma$ below the CVn fit (Fig. 5) and would need between ~ 18 and 58 times more H I to lie on the median of the relation. Figure 7 shows their 3-colour image of these LTDs and they visually appear similar (e.g. blue, with structure and some asymmetric stellar emission) to some of the H I-detected LTDs. FDS11 403 has a significantly lower mean surface brightness ($\mu_e = 26.15$) than the H I-undetected LTDs in this luminosity range, and its diffuse nature is a plausible explanation for why no H I is detected. The other four H I-undetected LTDs may have very recently had their H I stripped and not enough time has passed for them to lose their structure and have the blue stars fade to red. Their stellar masses range between 0.3 and $1.6 \times 10^7 M_\odot$. A single shock to galaxies in this mass range can directly strip the H I and initiate the quenching process. Therefore, low-mass LTDs with H I may not have a chance to display a disturbed H I morphology given how rapidly it can be removed. Similarly, the H I-disturbed LTDs in Fornax are the more massive LTDs. Their irregular H I morphology will be visible for longer timescales as it will take a longer

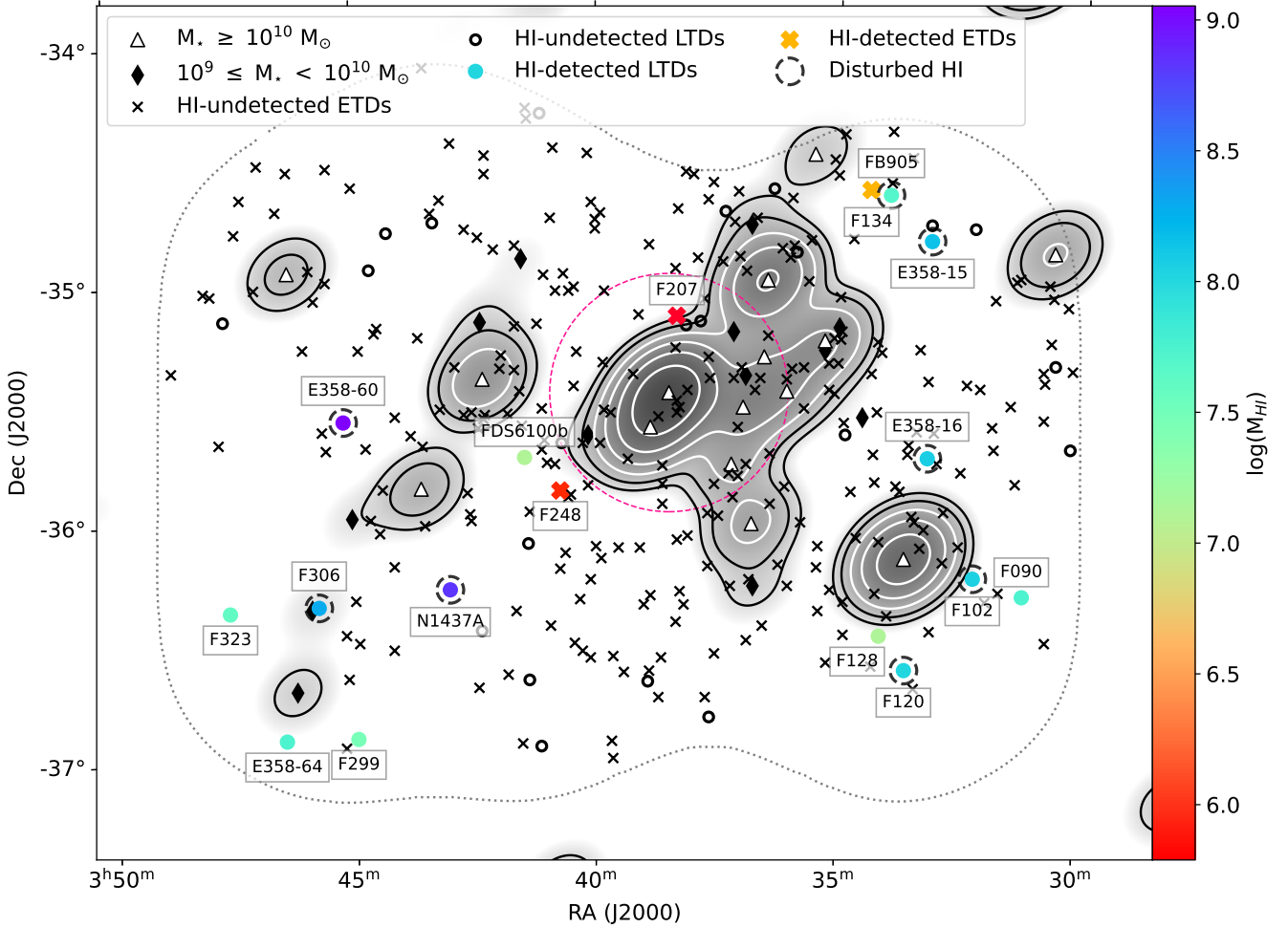


Fig. 8. Fornax cluster projected density field. The background image shows the projected density field on a log scale, produced by Gaussian smoothing the luminosity ratio relative to the BCG (NGC 1399), at the position of each galaxy. The (black and white) contours emphasise the density field, where the lowest shows the influence of $M_* \approx 5 \times 10^9 M_\odot$ and increases by $2''$. The grey dotted line shows the border of our observed area (same as Fig. 2), and the pink dashed circle is $0.25 R_{\text{vir}}$. Black crosses and open black circles are the HI-undetected ETDs and LTDs. The filled circles and thick crosses are the HI-detected ETDs and LTDs, which have been labelled and coloured by their M_{HI} . LTDs with disturbed HI morphologies are enclosed with dashed black circles. Black diamonds show the intermediate galaxies and the white triangles are the massive galaxies. The projected substructure is dominated by the brightest (hence most massive) galaxies, and the HI-detected dwarfs avoid the most massive potentials (i.e. cluster centre and massive galaxies) in the cluster.

(although still rapid) time or multiple events to strip all of the HI.

FCC 090 is an interesting case. While it is classified as a late type (Venhola et al. 2018), there are significant differences between it and the other HI-detected LTDs. It has the $\sim$ same brightness and stellar mass as NGC 1437A, although FCC 090 has a lower amount of HI by more than a factor of 10 (Table 3). FCC 090 is one of the least blue HI-detected LTDs, occupying the $\sim$ same colour region as the ETD FCC 134 (Figs. A.1 and 4). FCC 090 has a very high surface brightness (Fig. 6) and is not completely smooth, with a potential bar in the core. However, its stellar body is significantly smoother than the other HI-detected LTDs. It resides on the edge of the scatter in the CVn luminosity-HI fit (Fig. 5) and the velocity field (Fig. A.1) does not show any rotation, which is akin to the ETDs with HI. A tail of CO extending beyond the stellar body and pointing away from the cluster centre has been detected in FCC 090 (Zabel et al. 2019). FCC 090 has a low molecular gas fraction and ram-pressure was suggested to have created the unusual CO tail (Zabel et al. 2019), however, our HI image (Fig. A.1) does not show a ram-pressure tail. This does not exclude ram-pressure previously stripping

some HI gas and the mixed LTD and ETD characteristics suggest that FCC 090 is a transition type dwarf (TTD; Koleva et al. 2013). Even though the HI and CO have not been fully stripped, there is a small amount of HI in the core above a column density of 10^{21} cm^{-2} . FCC 090 has a blue core that coincides with the high-density HI, although there is no dense HI outside this region to collapse and form stars. This may explain why the stellar morphology is smoother than the other HI-detected LTDs. As FCC 090 is being stripped of gas while the blue stars and structure are fading, the morphological transformation may occur on a similar timescale to the stripping of cool gas.

ESO 358–G064 is similar to FCC 299 as they both are edge-on LTDs with HI that does not deviate much from symmetric emission. Both galaxies also display a regularly rotating disk in the velocity field (Fig. A.1). The clear difference is that FCC 299 is blue (as expected) and ESO 358–G064 is not. ESO 358–G064 is as red as FCC 090 and redder than the ETD FCC 134 (Figs. A.1 and 4). A red colour can indicate the quenching of star-formation, which may be a consequence of pre-processing. However, no other observations support this as ESO 358–G064 is beyond R_{vir} with typical M_{HI} and μ_e values for its luminosity

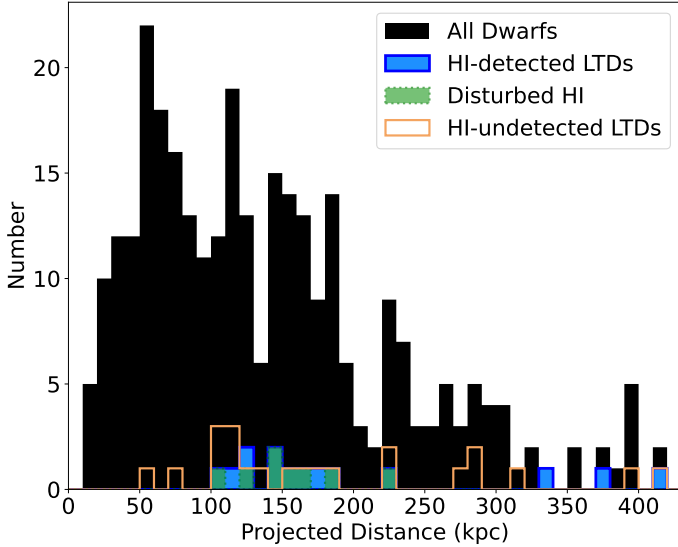


Fig. 9. Projected distance to nearest massive galaxy. All dwarfs in our observed area are shown as the filled-black histogram. HI-detected LTDs are shown in the blue-filled histogram, LTDs with disturbed HI are shown by the dashed green-filled histogram and HI-undetected LTDs are the open orange histogram. There are no HI-detected LTDs within a projected distance of 100 kpc to their nearest massive galaxy, implying that massive galaxies play an active role in the removal of HI from LTDs.

(Figs. 5, 6 and 8). The neighbouring galaxy is a confirmed background source (Fig. A.1), ruling out tidal interactions. Rather, in the case of ESO 358–G064 the red colour is likely a result of the presence of dust. Dust can obscure stellar emission and artificially redden the colour of a galaxy. Most dwarf galaxies are not dusty and the internal extinction is often insignificant (see Henkel et al. 2022, and references therein). FCC 090, FCC 306 and NGC 1437A have dust masses between 0.3 and $32 \times 10^5 M_{\odot}$ (Zabel et al. 2021). The optical image shows that ESO 358–G064 does indeed contain some dust. As it edge-on, the dust obscuration can have an even higher internal extinction than face-on LTDs such as NGC 1437A. A small (e.g. $\lesssim 0.07$) colour correction would mean that the true colour of ESO 358–G064 is the $\sim$ same as other LTDs with a comparable M_{HI} . Internal extinction from dust reddening is a more plausible explanation for the red colour of ESO 358–G064 than the star-formation being quenched.

The dense (e.g. $N_{\text{HI}} \geq 10^{21} \text{ cm}^{-2}$) HI is required to be stripped from a LTD so that star formation ceases. Only then can the blue stars and structure fade, thus transforming into an red, quiescent ETD. Therefore, a morphological transformation from a LTD to an ETD cannot be faster than the dense HI stripping time. However, as suggested above, it can be on a similar timescale, particularly if there are external tidal forces disrupting the stellar body. A possible explanation for the three HI-detected ETDs is that they are the end phase of quenched LTDs. If a LTD is on a tangential orbit, the HI will be slowly consumed to form stars. There may only be a small amount of low column density HI remaining that is not dense enough to form stars by the time they are ETDs. Another scenario is that these ETDs have recently acquired some HI through a ‘rejuvenation’ event. An example of this could be a minor merger between a quenched ETD with a LTD containing some HI. The ETD will appear blue from a merger-induced starburst and retain a small amount of HI. This has already been observed in the

Fornax A subgroup (De Rijcke et al. 2013; Kleiner et al. 2021) and is indeed a plausible scenario in the Fornax cluster. FCC 134 is the bluest (Fig. 4 and Table 3) HI-detected ETD and has been observed in the SAMI-Fornax Dwarfs survey (Scott et al. 2020). The average age of the stellar population of FCC 134 is 3 Gyr (Romero-Gómez et al. 2023). This is young for a ETD and is only possible if there has been recent star formation. FCC 248 is located at the end of the HI tail of NGC 1427A. It is plausible that it overlaps in projection and does not actually contain HI. However, at high resolutions the HI column density peaks exactly at the position of FCC 248 and the radial velocities of FCC 248 and the surrounding HI emission from the NGC 1427A tail are the same. Rather than being a projection effect, we believe that FCC 248 is a quenched, red ETD that accreted some HI from the tail of NGC 1427A, but is not dense enough to form new stars. FCC 207 is located near the cluster core, on the edge of the densest region of the cluster (Figs. 2 and 8). It is difficult to explain how this red ETD has any HI and off-centre molecular gas with disturbed kinematics (Zabel et al. 2019). The average age of the stellar population is 6 Gyr (Romero-Gómez et al. 2023), which does not support any recent star formation. There may be projection effects regarding its true location, although the velocity of FCC 207 is the same as the systemic cluster velocity. Perhaps FCC 207 favours the tangential orbit slow quenching scenario (similarly suggested by Conselice et al. 2003, in the Virgo cluster) or it is not actually that near to the cluster core.

6. HI removal toy model

The cluster crossing time is ~ 2 Gyr (Loni et al. 2021) and active gas removal that quench dwarf galaxies on quicker timescales than the crossing time has already been proposed (De Rijcke et al. 2010; Drinkwater et al. 2001; Loni et al. 2021). Similarly, previous works have suggested that dwarf galaxies will lose all of their HI soon after the crossing of the virial radius or at first pericentric passage in other (e.g. Virgo) clusters (Boselli et al. 2008, 2014b; Hester 2010; Cortese et al. 2021). This is very difficult to show observationally, as we do not have star formation histories (SFHs) for our sample of dwarfs, which can be used to constrain the gas removal and quenching timescales. Thanks to our very sensitive and deep MeerKAT observations, we are confident about which galaxies have HI down to $5 \times 10^5 M_{\odot}$. Therefore, we construct a simple toy model, that models the HI loss as dwarf galaxies fall into the Fornax cluster, to reproduce what we observe in Fig. 5.

In this toy model, we control only 3 parameters; (i) The accretion rate (N_{infall}) of dwarf galaxies falling onto the cluster i.e. number of accreted dwarfs per Gyr, (ii) $t_{1/2}$, which is the halving time of HI, assuming an exponential M_{HI} loss using $M_{\text{HI}}(t) = M_{\text{HI},0} e^{\lambda t/t_{1/2}}$ and $\lambda = \ln(0.5)$. Lastly (iii) the accretion clustering i.e. how grouped together the dwarfs being deposited onto the cluster are – no clustering would be individual galaxies deposited one-at-a-time and some clustering is when multiple dwarfs are deposited simultaneously.

Galaxies in the field that consume HI to form stars will have a $\sim$ constant M_{HI} over cosmic time, provided they have a supply of gas from the cosmic web. Strong environmental encounters are responsible for removing significant amounts of HI. Hydrodynamical simulations (e.g. Marasco et al. 2016; Rohr et al. 2023) have shown that ram pressure, tidal interactions and the combined effect of both, rapidly removes a significant amount of HI. After this initial phase of HI loss, the remaining HI will be removed at a lower rate as there is less HI to remove and it resides deeper in the potential well. We approximate this

process to first order in our toy model by using an exponential decrease of M_{HI} over time.

To reproduce Fig. 5, we impose the same observational constraints we observe in Fornax; we do not detect HI in dwarfs fainter than $M_r = -13$ nor do we detect any HI below $M_{\text{HI}} = 5 \times 10^5 M_\odot$ (Fig. 5). We are qualitatively looking to reproduce two fundamental things; the number of LTDs within the 3σ scatter of the CVn fit, which should be of the order of ~ 10 , and numerous objects below the detection limit, though in practice we accept having a few objects just above it to account for the low number statistics and error bars. If either of these criteria are not satisfied, then there will be too few or too many LTDs on the CVn fit or numerous detections between $M_{\text{HI}} = 5 \times 10^5 M_\odot$ and the M_r – M_{HI} relation, which are not observed in Fornax. We presume there is nothing special about observing the Fornax cluster at this time, and given that we are modelling the M_{HI} loss, we should be able to reproduce Fig. 5 in multiple snapshots.

We model infalling dwarf galaxies onto the cluster as blue, field LTDs that populate the M_r – M_{HI} plane according to the CVn distribution. To achieve this, a random FDS dwarf magnitude chosen and the corresponding M_{HI} is computed from a random normal distribution centred on the M_r – M_{HI} CVn fit with 3σ scatter (Fig. 5). Then, galaxies in the cluster exponentially lose their M_{HI} (as above) and additional dwarfs are continually deposited onto the cluster as time evolves. We explored a combination of parameters that spanned two orders of magnitude both in N_{infall} and $t_{1/2}$. In Fig. 10 we show the result with the best values to reproduce Fig. 5. A $N_{\text{infall}} = 100 \text{ Gyr}^{-1}$ is required to reproduce ~ 10 (i.e. not too few or too many) dwarfs on the CVn fit and is a reasonable value as we know that ~ 300 dwarfs are in the Fornax cluster within the observed MFS area at $z \sim 0$. The best value for the M_{HI} halving time is $t_{1/2} = 0.03 \text{ Gyr}$, as the dwarfs lose HI rapidly enough that numerous objects are below the detection limit with few objects just above it at any given snapshot. However, we are only able to simultaneously reproduce both these criteria by including some accretion clustering that simultaneously deposits 20 dwarfs onto the cluster. This is actually a more realistic scenario and akin to how the cosmic web funnels galaxy groups onto clusters. The majority of snapshots qualitatively reproduce what we observe in Fornax, although snapshots at $t = 0.4, 1.0, 1.8, 2.8, 4.0$ and 4.2 Gyr are the most alike Fig. 5.

In Table 4, we show the time it takes for galaxies in the M_{HI} range of 10^6 – 10^9 to be undetectable by the MFS in our toy model. This is a simplistic way of modelling the HI loss as it does not take into account how galaxies lose their HI, only the rate that HI is removed. Therefore, the timescales in Table 4 should not be overstated. Despite the simplicity, our timescales are consistent with low-mass dwarfs being directly stripped of their HI in a single shock (e.g. what we expect happened to the HI-undetected LTDs between $M_r = -12$ and -14.2), more massive LTDs harbouring disturbed HI morphologies due to longer times or multiple events being required to fully strip their HI, and not detecting any LTDs in HI below the M_r – M_{HI} relation. These timescales are significantly shorter than the 2 Gyr cluster crossing time and consistent with active and efficient HI removal from LTDs in the Fornax cluster. In the Virgo cluster it has also been estimated that LTDs with $M_{\text{HI}} \sim 4 \times 10^7 M_\odot$ will be stripped of HI in $\sim 150 \text{ Myr}$ (Boselli et al. 2008; Kenney et al. 2014). While we constrain the HI removal timescale for dwarf galaxies in the Fornax cluster, the morphological transformation from a LTD to an ETD in clusters and the field is not well constrained. As mentioned previously, it cannot be quicker than the HI removal timescale, although it may be able to occur on a sim-

ilar timescale, particularly with external tidal fields shaping the stellar body. FCC 090 may suggest that a morphological transformation can occur in parallel with the HI removal, although this is just one case and not direct evidence of both transformations occurring at the same rate.

7. Summary

We present results on the HI content of dwarf galaxies in the central $\sim 2.5 \times 4 \text{ deg}^2$ of the Fornax cluster located 20 Mpc away as part of the MeerKAT Fornax Survey. In this work we show three different HI images that correspond to a 3σ column density sensitivity range between 2.7 and $50 \times 10^{18} \text{ cm}^{-2}$ over 25 km s^{-1} for spatial resolutions between 4 and 1 kpc. Such impressive sensitivity allows us to detect a 3σ point source with a line width of 50 km s^{-1} down to a $M_{\text{HI}} = 5 \times 10^5 M_\odot$.

We detected HI in 17 out of the 304 dwarfs. Seven are new HI detections and we discovered new HI features in previously known detections. Our HI-detected dwarfs consist of 14 out of the 36 LTDs and three of the 268 ETDs. Eight of the HI-detected LTDs show irregular or asymmetric HI emission in the form of HI tails, significant warps, or non-rotating components. NGC 1437A and FCC 306 are clearly having their HI shaped by ram pressure as they have one-sided starless HI tails pointing away from the cluster centre and velocity gradients aligned with the ram-pressure wind direction (Serra et al. 2023). We show in the high-resolution images that their HI closest to the cluster centre is being compressed into the stellar disk along with an additional three LTDs that show evidence of HI compression and one-sided tails, albeit they are less obvious than NGC 1437A and FCC 306. Serra et al. (2023) suggest that tidal forces are required in addition to ram pressure to reproduce the observed properties of some HI-disturbed cluster galaxies and we observe the same as the eight HI-disturbed LTDs have asymmetric or lopsided stellar emission.

To reveal the substructure of the cluster, we created a projected density field and find that all the HI-detected dwarfs avoid the most massive potentials (i.e. cluster centre and massive galaxies). None of the HI-detected LTDs are within a projected distance of 100 kpc to their nearest massive galaxy, implying that massive galaxies play an active role in the removal of HI from LTDs.

The HI-detected LTDs have most likely recently joined the cluster and are on their first infall as they are located at large clustocentric radii, that is to say beyond $0.5 R_{\text{vir}}$, and have comparable M_{HI} and μ_e at fixed luminosity when compared to blue, star-forming LTDs in the field. We detected HI in all LTDs brighter than $M_r = -14.2$, which corresponds to a $M_\star$ between $\sim 1.6 \times 10^7$ and $10^9 M_\odot$, a range where LTDs can briefly display a disturbed HI morphology. At luminosities between $M_r = -12$ and -14.2 , $M_\star$ range $\sim (0.3\text{--}1.6) \times 10^7 M_\odot$, we detected one LTD with a normal amount of HI, which is again analogous to a field LTD, and five LTDs that may have very recently had their HI stripped. Given their colour, luminosity, morphology, and surface brightness, we would expect to detect HI, but do not down to $M_{\text{HI}} = 5 \times 10^5 M_\odot$. Direct stripping from a single shock is likely to occur in this mass range (and below) such that these LTDs transition between having a normal amount of HI to no HI and never show a disturbed HI morphology. No LTDs were detected in HI fainter than $M_r = -12$. Assuming that the M_r – M_{HI} relation can be extrapolated down to these low luminosities, our upper limits are consistent with those galaxies with a normal M_{HI} .

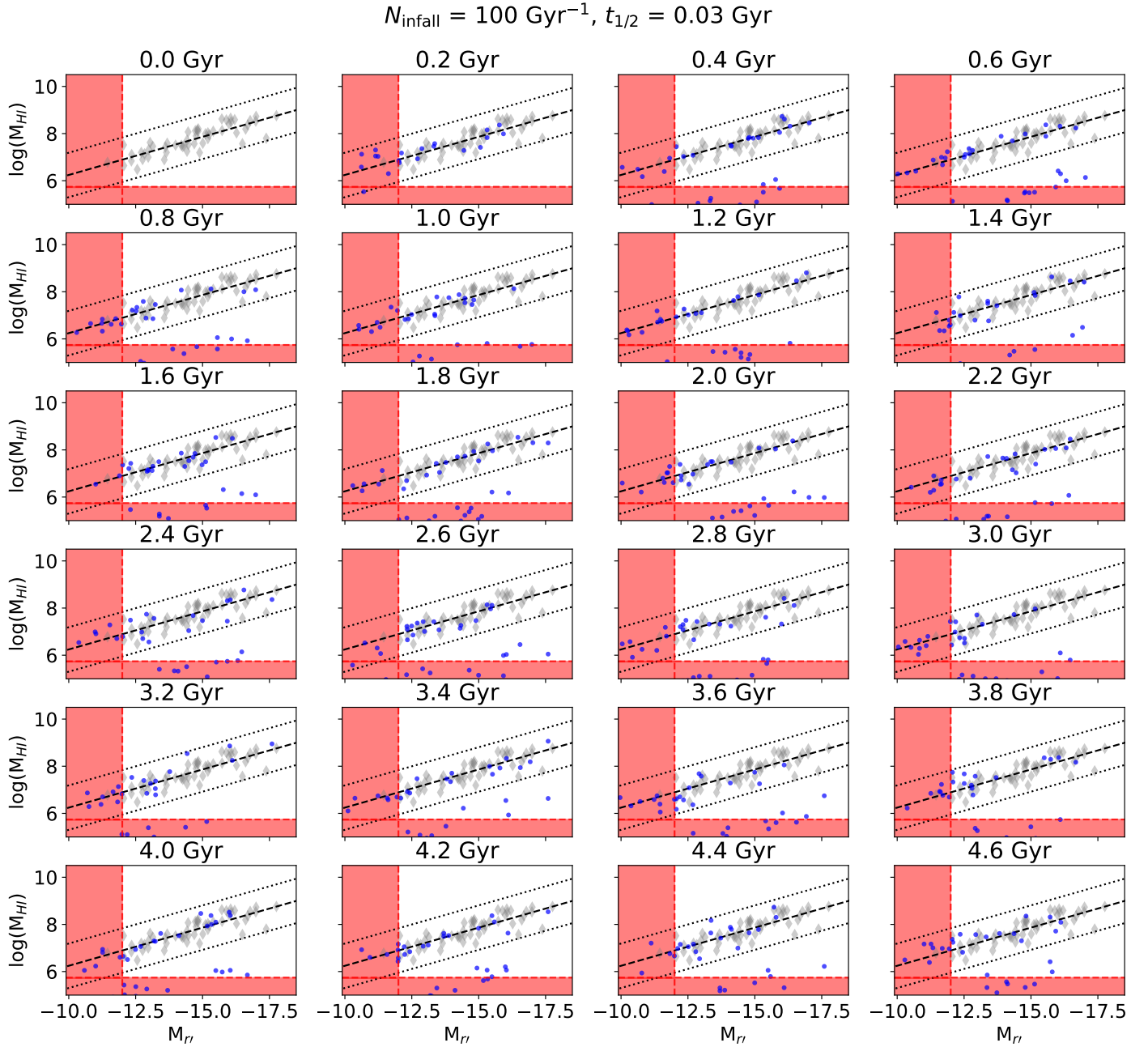


Fig. 10. Toy model of H I removal in the cluster. The best results to reproduce Fig. 5, showing what we would observe every 0.2 Gyr. The grey diamonds are the CVn galaxies, the dashed black line shows the linear fit used to populate infalling galaxies to the CVn galaxies and the dotted lines show the 3σ scatter. The blue filled circles are galaxies in the cluster and the red dashed lines are the observational constraints both in M_r and M_{HI} , such that we would not detect a galaxy in H I in the red shaded regions. With $N_{\text{infall}} = 100 \text{ Gyr}^{-1}$ and $t_{1/2} = 0.03 \text{ Gyr}$ we are able to qualitatively reproduce what we observe in Fig. 5.

We created a simple toy model that deposits field LTDs onto the Fornax cluster and removes their H I at an exponential rate. In order to reproduce what we observed in the Fornax cluster – that is to say enough LTDs on the M_r – M_{HI} plane and no LTDs detected in H I below the M_r – M_{HI} relation – we find that the H I of a $M_{\text{HI}} = 10^8 M_{\odot}$ dwarf is stripped in $\sim 240 \text{ Myr}$. As this is much shorter than the cluster crossing time, active H I removal must be occurring in the cluster. Our toy model is consistent with low-mass LTDs being directly stripped of their H I (e.g. through a single shock) and more massive LTDs harbouring disturbed H I morphologies due to longer times or multiple events being required to fully strip their H I.

FCC 090 is likely a TTD as it shares joint properties with LTDs and ETDs. It is one of the least blue with the smoothest stellar body of the H I-detected LTDs, does not show any rotation in the H I velocity field, and resides at the edge of the scatter below the median of the M_r – M_{HI} relation. This is tentative evidence that, in a cluster, a morphological transformation can occur in parallel on a similar timescale to the stripping of H I. It is difficult to discern the origin of the peculiar H I-detected ETDs. For two of them, we believe it is more likely that they have recently acquired some H I through a ‘rejuvenation’ event. FCC 134 is blue with a young average stellar population age for an ETD and may have acquired some H I through a minor merger

Table 4. Approximate time it takes for dwarfs with different M_{HI} to be undetectable in the MFS, as measured from our H I removal toy model.

$\log(M_{\text{HI}})$ ($M_{\odot}$)	$\sim t_{\text{HI loss}}$ (Myr)
6	30
7	150
8	240
9	330

that induced a starburst. FCC248 may have accreted some H I from the tail of NGC 1427A, but is not dense enough to form new stars. FCC 207 does not fit the scenario of recently acquiring H I. It is located near the cluster core with a typical ETD stellar population age. The systemic velocity of FCC 207 is the same as the systemic cluster velocity; although, there still may be projection effects. Otherwise, a slow quenching process from a tangential orbit could potentially produce an ETD with a small amount of H I near the cluster core.

In this work, we have demonstrated how active and efficient H I removal is required to reproduce the multi-wavelength properties of the dwarf galaxies in the Fornax cluster. The H I is rapidly stripped from infalling galaxies of the order of a few hundred million years, which is responsible for producing the hundreds of quiescent ETDs in the cluster. The H I is removed via both tidal and hydrodynamical forces operating in the cluster. While the MFS is ongoing, this work has covered the central and most efficient region of H I-removal in the cluster, out to $\sim R_{\text{vir}}$. This is the first time dwarf galaxies of this mass have been detected and resolved beyond the local group and in a galaxy cluster.

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Appendix A: Full Fig. 3

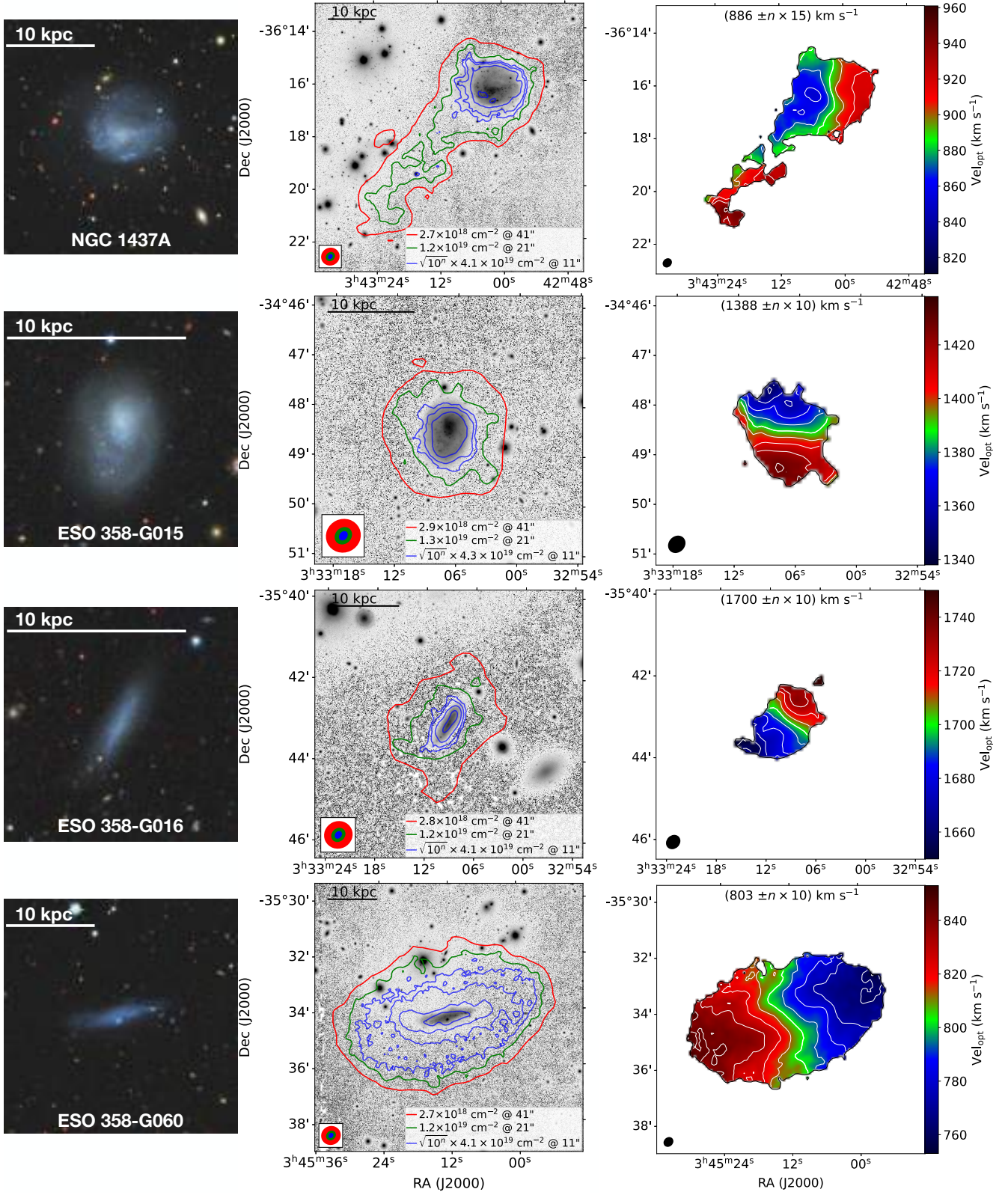


Fig. A.1. Full Fig. 3, but displaying all 17 H I detections.

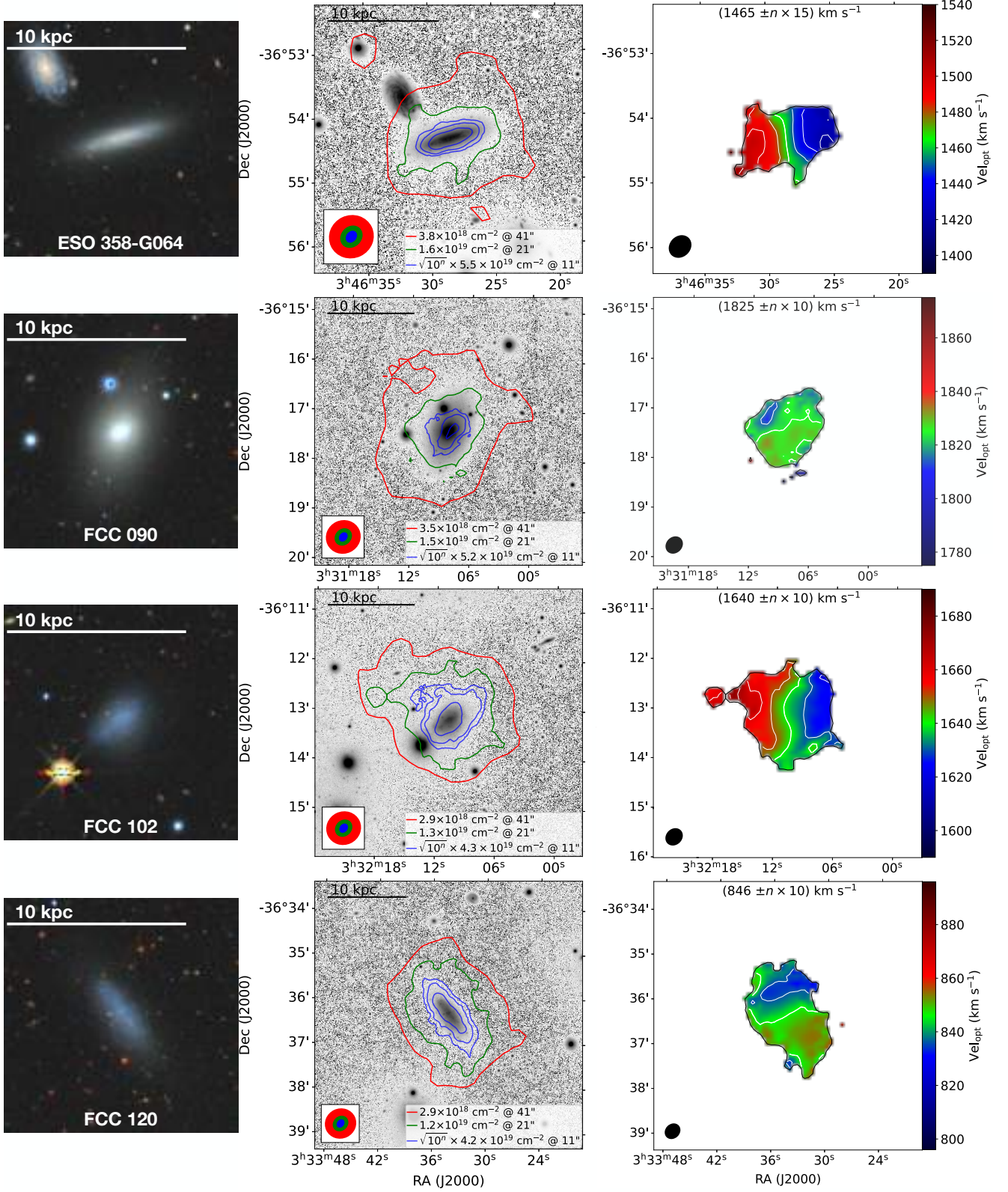


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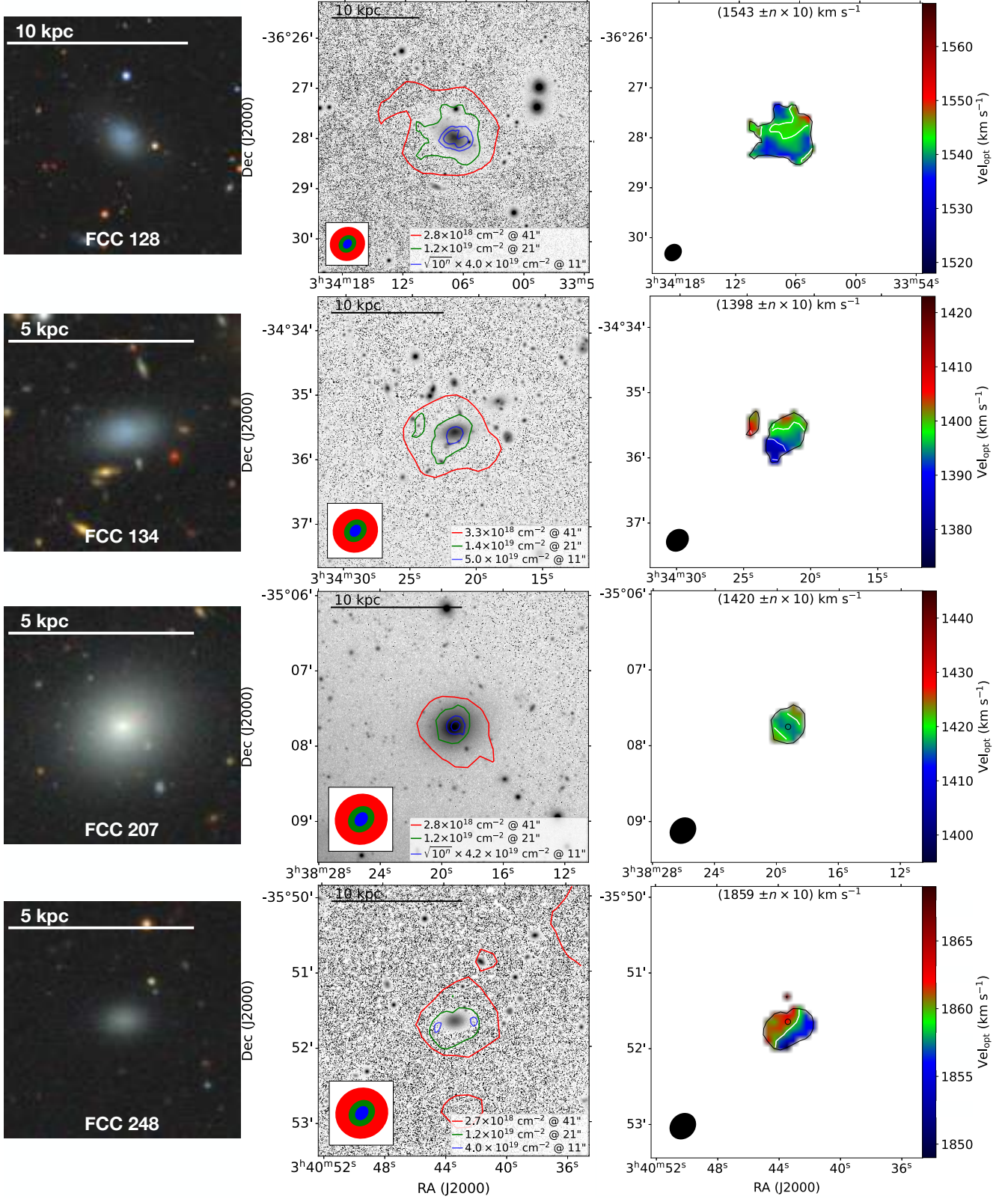


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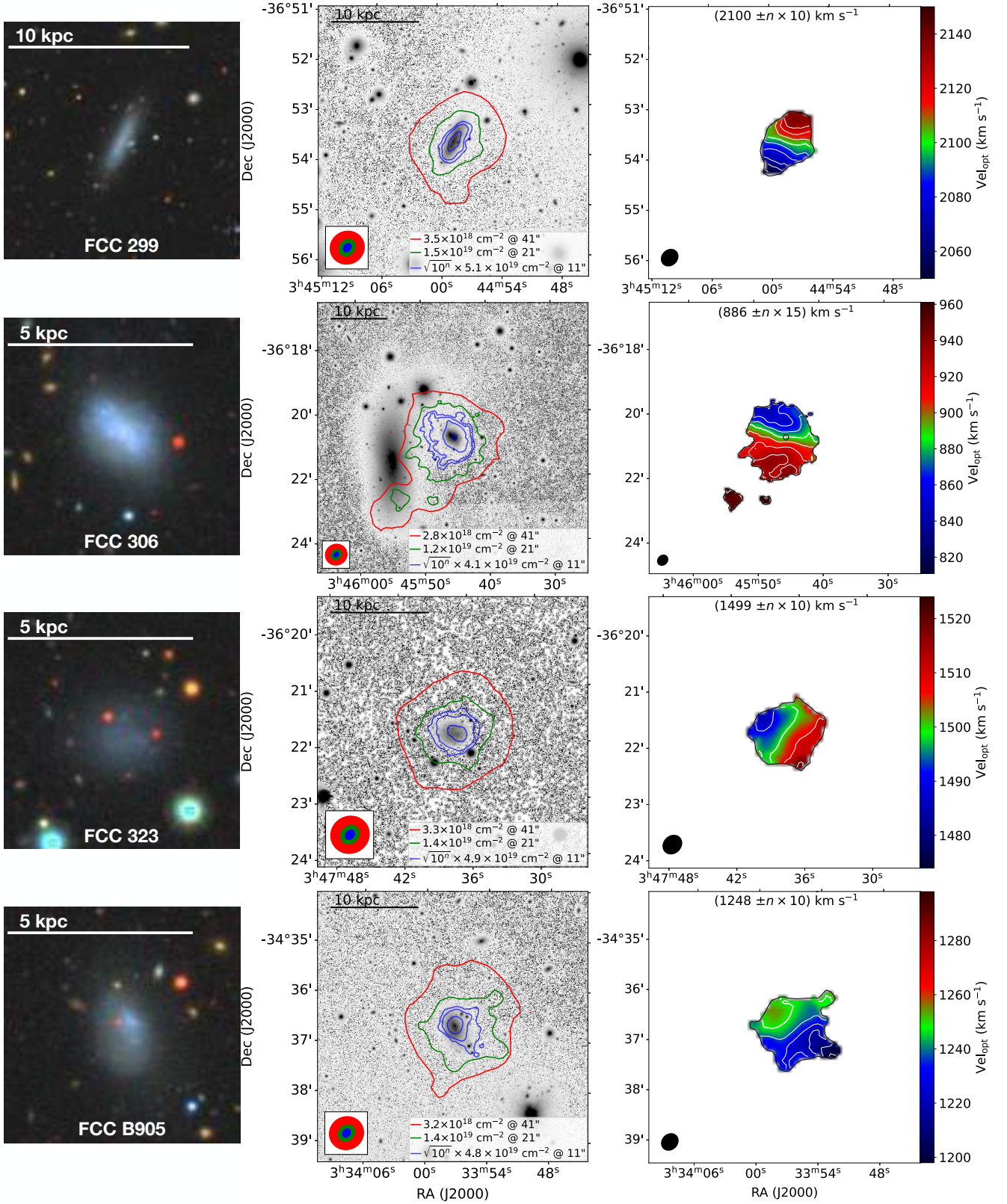


Fig. A.1. continued.

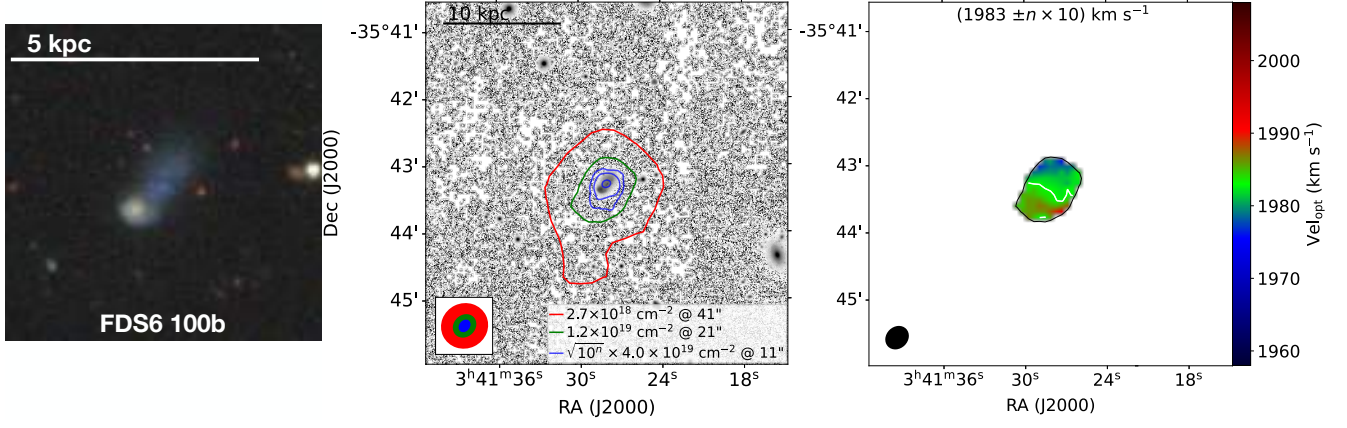


Fig. A.1. continued.