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A spectroscopic survey of Abell 1703: is it a rare relaxed cluster hosting a radio halo or a usual merging system?

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ABSTRACT

We present the study of the internal dynamics of the intriguing galaxy cluster Abell 1703, a system hosting a probable giant radio halo whose dynamical status is still controversial. Our analysis is based on unpublished spectroscopic data acquired at the Italian Telescopio Nazionale *Galileo* and data publicly available in the literature. We also use photometric data from the Sloan Digital Sky Survey. We select 147 cluster members and compute the cluster redshift $\langle z \rangle \sim 0.277$ and the global line-of-sight velocity dispersion $\sigma_v \sim 1300 \text{ km s}^{-1}$. We infer that **Abell 1703** is a massive cluster: $M_{200} \sim 1\text{--}2 \times 10^{15} M_\odot$. The results of our study disagree with the picture of an unimodal, relaxed cluster as suggested by previous studies based on the gravitational lensing analysis and support the view of a perturbed dynamics proposed by recent works based on *Chandra* X-ray data. The first strong evidence of a dynamically disturbed cluster comes from the peculiarity of the BCG velocity with respect to the first moment of the velocity distribution of member galaxies. Moreover, several statistical tests employed to study the cluster galaxies kinematics find significant evidence of substructure, being **Abell 1703** composed by at least two or three subclumps probably caught after the core–core passage. In this observational scenario, the suspected existence of a radio halo in the centre of this cluster is not surprising and well agrees with the theoretical models describing diffuse radio sources in clusters.

Key words: galaxies: clusters: general – galaxies: clusters: individual: Abell 1703 – galaxies: kinematics and dynamics.

1 INTRODUCTION

Radio haloes (also *giant radio haloes*, or GRHs) are diffuse sources found in the central regions of massive ($M_{200} \gtrsim 10^{15} M_\odot$) galaxy clusters. Extended over volumes of $\sim 1 \text{ Mpc}^3$, these low surface brightness features ($\sim 1 \mu\text{Jy arcsec}^{-2}$ at 1.4 GHz) have no obvious optical counterparts and roughly follow the intracluster medium (ICM) mass distribution. Their synchrotron steep-spectrum ($S(\nu) \sim \nu^{-\alpha}$; $\alpha > 1$) reveals the existence of a population of relativistic electrons and large-scale magnetic fields spread throughout the ICM (see e.g. Feretti et al. 2012 for a review).

Until recently, GRHs were always discovered in merging clusters (e.g. van Weeren et al. 2019). Indeed, in the hierarchical scenario of cosmic structure formation, it is a common fact that galaxy

groups and subclusters merge together into massive clusters. These processes release enormous amounts of gravitational energy (as large as 10^{64} erg ; Sarazin 2002) and induce turbulence in the ICM, which is considered the key mechanism able to accelerate particles to relativistic energies (e.g. Brunetti & Jones 2015).

However, in the last years this picture got more complicated following the discovery of diffuse radio emission in several dynamically relaxed clusters. The first, impacting, case was the extended radio source found in the cluster CL1821 + 643 ($z \sim 0.296$; Bonafede et al. 2014, Boschin et al. 2018), whose size, location, and power resemble that of typical GRHs despite the absence of any merging process responsible for its formation. Indeed, this case is not unique, since diffuse radio sources not powered by major mergers have been recently discovered in more clusters in the redshift range 0.1–0.3. For instance, Abell 2390 and Abell 2261 (Sommer et al. 2017) are two more examples of cool-core clusters with Mpc-scale radio sources. Intriguing are also the cases

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of PSZ1G139.61 + 24 and Abell 2142. The first one hosts both a mini-halo (a feature typically found in cool-core relaxed systems, e.g. Gitti et al. 2018) and an underluminous and ultrasteepest spectrum radio halo (Savini et al. 2018a). Abell 2142, on its hand, is affected by minor merging activity, as suggested by the cold fronts detected in its ICM. However, as in the case of PSZ1G139.61 + 24, it also exhibits a double-component extended radio emission (Venturi et al. 2017), with a spectral steepness increasing in the outer regions. These two last examples could be hybrid sources, with mini-halos evolving into GRHs or vice versa (van Weeren et al. 2019). It is expected that the number of these ‘intermediate’ cases will increase considerably in the near future thanks to observational facilities like LOFAR (van Haarlem et al. 2013).

Taking into account this new framework, in this paper we focus on Abell 1703 (hereafter **A1703**; Abell 1958, Abell, Corwin & Olowin 1989), a system at $z \sim 0.28$ (Allen et al. 1992) whose dynamical status and radio properties might look like the ones of CL1821 + 643. In fact, several authors claim the possible existence of a radio halo in this cluster. The first hint comes from Owen, Morrison & Voges (1999; see their table 2), who include **A1703** in a list of nine clusters with diffuse radio emission. Indeed, **A1703** is the brightest source in Owen et al. sample. More recently, the analysis of archival VLA data at 1.4 GHz showed some evidence of diffuse emission in the central region of the cluster (Govoni, private communication; see also our Fig. 1) despite contamination by radio point-like sources (see Rizza et al. 2003 and our Section 2.1). Finally, Wilber (2018) highlights the presence in **A1703** of possible radio halo emission at lower frequencies (120–168 MHz) in the LOFAR Two-Metre Sky Survey (LoTSS). This evidence is also confirmed by Savini et al. (2018b) in their fig. 8, where the contour levels at 144 MHz from LOFAR suggest the existence of a diffuse source in the location of the cluster.

On the other hand, the picture of **A1703** that emerges at other wavelengths is controversial and suggest conflicting interpretations about the real dynamical status of this cluster. From the optical point of view **A1703** is one of the richest clusters in the photometric catalogue of the SDSS (e.g. Koester et al. 2007) and presents a dominant giant elliptical cD galaxy in its centre (the brightest cluster galaxy; hereafter BCG). Moreover, this cluster shows impressive strong gravitational lensing features (Hennawi et al. 2008) which allowed to reconstruct its internal mass profile. For instance, Limousin et al. (2008), based on imaging data both from space (*HST*) and ground (Subaru), identified 13 systems forming highly magnified images. Thanks to spectroscopic (from Keck) and photometric redshifts of these images they found that all the lensing systems can be reproduced by a single NFW (Navarro, Frenk & White 1997) profile for the dark matter, thus suggesting that **A1703** is a regular, unimodal cluster. A follow-up study by Richard et al. (2009) based on further Keck spectroscopic measurements for seven multiple sources reinforced the evidence that **A1703** is a relaxed cluster. More studies based on strong and/or weak gravitational lensing analysis also provided estimates of the cluster mass and the NFW concentration parameter. Even though the results are quite discordant, they all coincide that **A1703** is a massive system: $M_{\text{vir}} \simeq 1.1\text{--}2 \times 10^{15} M_{\odot}$ and $c_{\text{vir}} \simeq 3.3\text{--}7.1$ (Oguri et al. 2009, Zitrin et al. 2010, Oguri et al. 2012).

Such a large mass is also consistent with estimates derived from *Planck* observations of the thermal Sunyaev–Zeldovich effect. In fact, based on measurements of the Compton parameter Y and the $Y - M_{500}$ scaling relation of Arnaud et al. (2010), the *Planck* Collaboration XIII (2016) reports an hydrostatic mass $M_{500}^{\text{SZ}} = 6.76_{-0.37}^{+0.35} \times$

$10^{14} M_{\odot}$, which corresponds to $M_{\text{vir}} \gtrsim 1 \times 10^{15} M_{\odot}$ assuming an NFW profile for the mass distribution.

As for the X-ray band, Piffaretti et al. (2011) report a measurement of the X-ray luminosity of **A1703** from *ROSAT* data: $L_X(R < R_{500} = 1.07 \text{ Mpc}) = 5.42 \times 10^{44} \text{ erg s}^{-1}$ in the 0.1–2.4 keV band. Using equation (10) of Böhringer, Chon & Collins (2014), this luminosity translates to a mass estimate $M_{200} \gtrsim 1 \times 10^{15} M_{\odot}$ (in our cosmology, see below), again indicative of a massive cluster. Very recently, Ge et al. (2019) used unpublished *Chandra* archival data to compute the temperature of the ICM within $0.15\text{--}0.75 R_{500}$: $kT_X = 9.63 \pm 0.75 \text{ keV}$. These data show that the ICM is elongated in the SSE–NNW direction (see contours in Fig. 1). However, despite the absence of obvious bimodality or more complex substructure, Ge et al. (2019) also report hints of a disturbed dynamics according to the measurements of several X-ray morphological parameters (see their table 5) and their locations in the morphological planes of Cassano et al. (2010) and Mantz et al. (2015).

Taking into account the scenarios described by studies based on gravitational lensing and X-ray data, it is important to definitively establish whether **A1703** is an anomalous relaxed cluster hosting diffuse radio emission or a common merging system. An exhaustive analysis of the kinematics of the cluster galaxy population could have the last word on the true dynamical status of this cluster. Moreover, this analysis would allow to study the possible presence of a merger along the line of sight, which would be difficult if not impossible to detect in the X-ray band. In fact, the spatial and kinematical study of member galaxies represent an effective tool to reveal substructures in clusters and put in evidence pre-merging subgroups or merger remnants (e.g. Boschin et al. 2004; Boschin, Girardi & Barrena 2013).

With this context in mind, we used archival spectroscopic data obtained at the Italian Telescopio Nazionale *Galileo* (TNG) in order to perform the first dynamical analysis of this cluster based on member galaxies. These data sample the central $\sim 1 \text{ Mpc}$ size region of the cluster characterized by the diffuse X-ray and radio emissions. More data found in the NED Data base, most of which obtained by Bayliss et al. (2014; hereafter B14) through spectroscopic measurements with MMT/Hectospec, allowed us to extend our spectroscopic sample to cover a wider area around the cluster (see Fig. 2).

This paper is organized as follows. Section 2 describes the TNG observations and data reduction and presents the velocity catalogue. In Section 3 we describe our member selection procedure. Sections 4 and 5 explain the results of the analysis of the cluster structure. Finally, in Section 6, we discuss our results and present a portrait of the dynamical status of **A1703**.

Unless otherwise stated, we indicate errors at the 68 per cent confidence level (hereafter c.l.). Throughout this paper, we use $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$ in a flat cosmology with $\Omega_m = 0.3$ and $\Omega_{\Lambda} = 0.7$. In the adopted cosmology, 1 arcmin corresponds to $\sim 253 \text{ kpc}$ at the cluster redshift.

2 GALAXY DATA AND VELOCITY CATALOGUE

We used unpublished spectroscopic data stored in the TNG archive (<http://archives.ia2.inaf.it/tng>) taken in 2010 May (program A21TAC_50; PI: F. Gastaldello). These data consist of five MOS masks mainly sampling the central region of the cluster. Another mask was taken in 2016 June during a technical night. In particular,

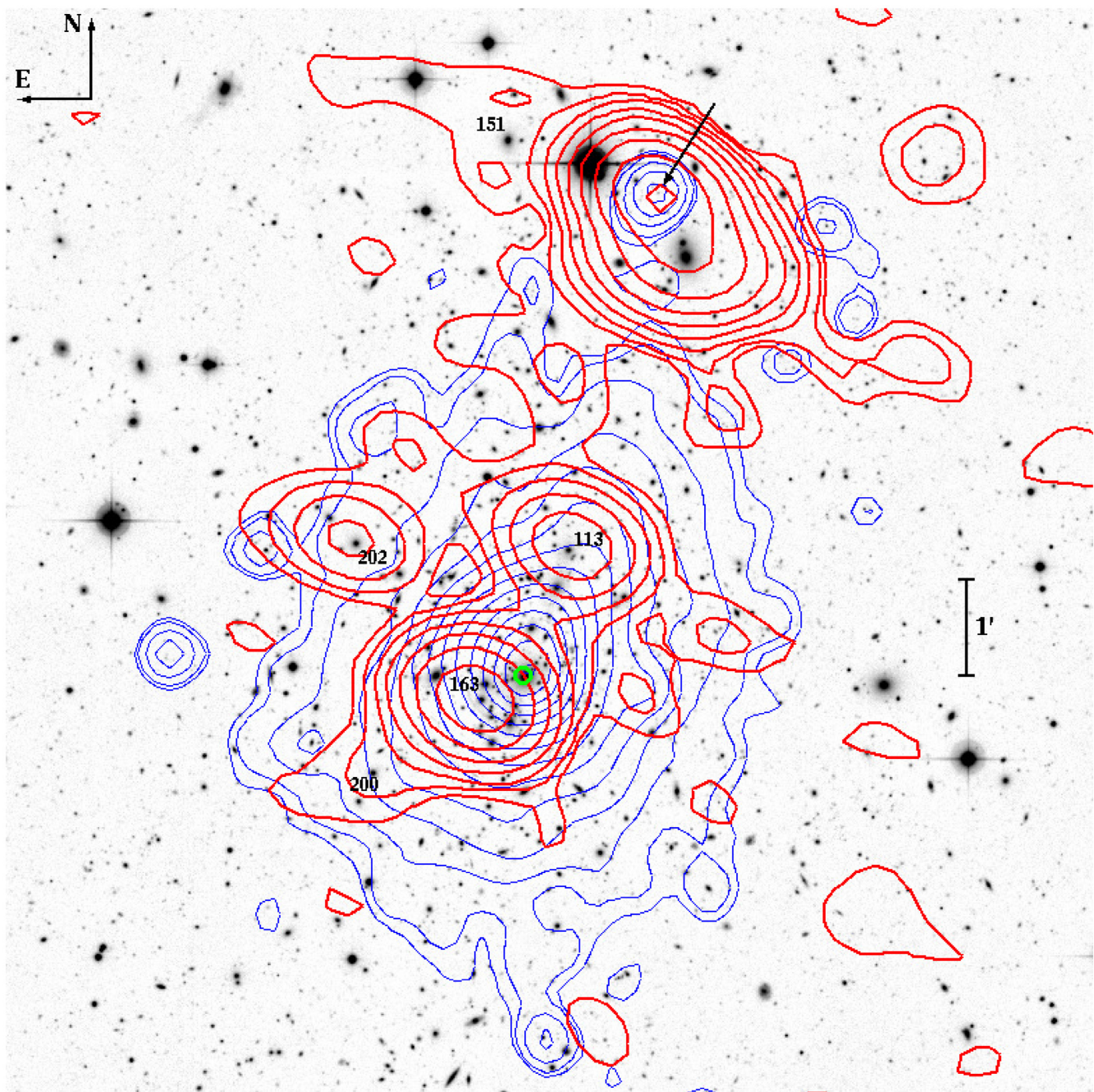


Figure 1. A multiwavelength picture of the cluster [A1703](#). The grey-scale image in the background corresponds to the optical r -band (CFHT/Megaprime archival data). The blue thin contours show the cluster X-ray emission in the 0.5–2 keV band (from *Chandra* archival image ID 16126; Texp: 48 ks). The thick red contours are the contour levels of a VLA 1.4 GHz low-resolution image (courtesy F. Govoni; from archival VLA observation program AM 469). Numbers highlight member galaxies mentioned in the text. The green circle is the BCG (ID 141; see Table 1).

this last mask allowed us to obtain the spectrum of the BCG. For all the six masks we used the LR-B Grism of the instrument DOLoRes¹ and obtained spectra for 131 objects. The total exposure times varied from 3600 to 7200 s.

We used standard IRAF tasks to reduce the spectra and adopted the cross-correlation technique (Tonry & Davis 1979) to compute redshifts for 104 targets. For 18 galaxies we obtained multiple

redshift determinations. They allowed us to obtain a better estimate for the redshift errors. In particular, we found that the nominal cross-correlation errors are underestimated and multiplied them by a factor 2.5 (see Girardi et al. 2011 for details of the redshift computation and their errors). For another five galaxies (IDs. 117, 152, 154, 162, and 180; see Table 1), their redshifts were computed by measuring the wavelengths of the emission lines in their spectra.

In order to extend our spectroscopic sample to the outskirts of the cluster, we searched the NED data base for galaxies with known

¹<http://www.tng.iac.es/instruments/lrs>

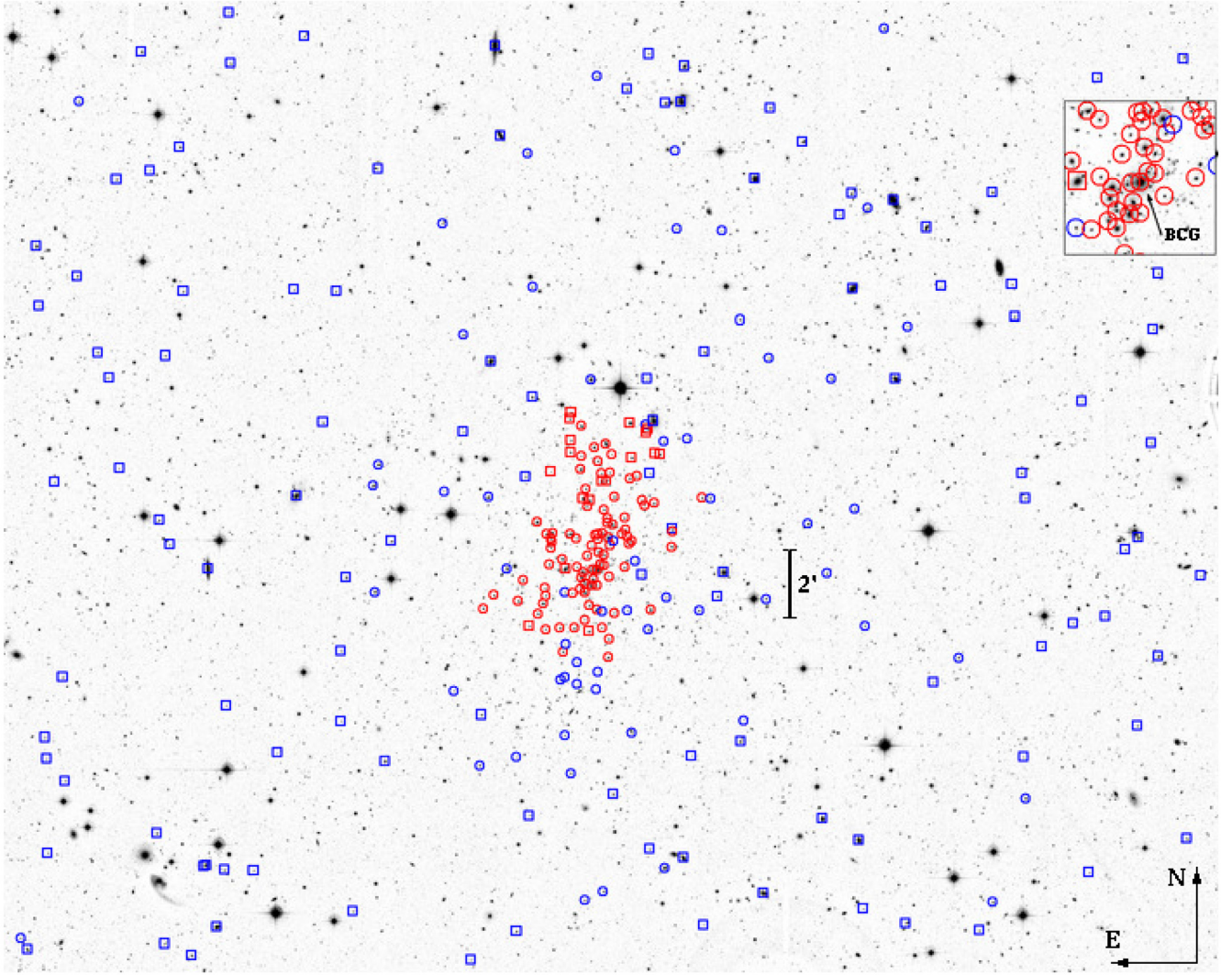


Figure 2. *r*-band wide field image (CFHT/Megaprime archival data) of the cluster A1703 showing the positions of all the galaxies of our spectroscopic catalogue. The circles and squares indicate cluster members and non-member galaxies, respectively. The red symbols refer to galaxies observed at the TNG (see the text and Table 1). Top right inset is a zoom on the central region of the cluster.

redshift in the field of A1703. We found 184 objects within a radius of ~ 20 arcmin from the cluster centre. Most of these objects (177) come from B14, ten of which are in common with our TNG data. To check for eventual systematic deviations, we performed a straight-line fit to the TNG and B14 redshift measurements taking into account the errors on both data sets (see chapter 15.3 of Press et al. 2007). We find an intercept $= (-2 \pm 8) \times 10^{-3}$ and a slope $= 1.01 \pm 0.03$ with a χ^2 probability $= 0.52$. Thus, we added the remaining 167 B14 galaxies and seven NED galaxies to our sample.

Our final spectroscopic catalogue includes 278 galaxies. The field of A1703 is covered by the SDSS. Its galaxy catalogue also provides us complete photometric information for all the galaxies of the spectroscopic sample in the magnitude bands *g*, *r*, and *i*.

Table 1 lists the velocity catalogue (see also Fig. 2): identification number of each galaxy, ID (Col. 1); redshift source (Col. 2; T:TNG, L:NED, B:B14); right ascension and declination, α and δ (J2000, Col. 3); (dereddened) SDSS *r* magnitude (Col. 4); heliocentric radial velocities, $V = cz_{\odot}$ (Col. 5) with errors, ΔV (Col. 6).

2.1 Notable galaxies

A1703 is dominated by the galaxy ID 141 (the BCG). It is by far the brightest cluster member, being the magnitude difference with the second brightest member galaxy ~ 1.3 . Its colours ($g - r = 1.58$ and $r - i = 0.59$) well match the red sequences of cluster early-type galaxies (see our analysis of Section 5.2.1) and, consistently, its optical spectrum is that of an elliptical galaxy free of emission lines.

The field of A1703 is populated by several point-like radio sources. Rizza et al. (2003; see their table 3) report six sources from high spatial resolution (~ 1.5 arcsec) VLA images at 20 cm. The counterparts of five of them are the member galaxies IDs. 113, 151, 163, 200, and 202 (see Fig. 1). The sixth one (n. 5 in table 3 of Rizza et al. 2003), at ~ 5.7 arcsec SW of the BCG, is not listed in our spectroscopic catalogue. Since its photometric redshift is 0.24 ± 0.04 (from the SDSS), it is unclear whether it is a cluster member or a foreground object. Finally, at ~ 5 arcmin NNW of the BCG there is another strong radio source whose possible optical

Table 1. Velocity catalogue of 278 spectroscopically measured galaxies in the field of **A1703**. IDs in italics refer to non-member galaxies. Galaxy ID 141 (in boldface) is the BCG.

ID	Source	α, δ (J2000)	r	V (km s ⁻¹)	ΔV
001	B	13 13 13.66, + 51 48 48.1	21.09	142 509	69
002	B	13 13 16.19, + 52 03 34.4	21.46	166 775	96
003	B	13 13 16.50, + 51 41 17.8	20.03	116 430	27
004	B	13 13 21.13, + 51 57 26.5	20.64	126 545	36
005	B	13 13 21.59, + 51 46 30.6	18.98	70 250	30
006	B	13 13 22.20, + 51 55 51.6	20.70	90 154	54
007	B	13 13 22.63, + 51 52 35.9	20.70	107 578	75
008	B	13 13 25.19, + 51 49 53.9	18.84	115 294	18
009	B	13 13 25.56, + 51 44 31.7	20.73	114 710	39
010	B	13 13 27.59, + 51 49 33.9	20.61	115 480	111
011	B	13 13 31.36, + 51 47 39.4	20.99	86 817	66
012	B	13 13 32.20, + 52 03 01.7	20.67	147 285	39
013	B	13 13 34.59, + 51 40 21.1	21.18	155 562	54
014	B	13 13 35.50, + 51 53 48.3	21.72	116 685	33
015	B	13 13 37.34, + 51 47 28.3	21.10	189 304	57
016	B	13 13 43.03, + 51 46 48.6	20.64	124 402	54
017	B	13 13 45.96, + 51 51 02.8	20.27	114 620	63
018	B	13 13 46.21, + 51 42 28.5	19.13	83 693	33
019	B	13 13 46.53, + 51 43 39.1	20.76	125 199	33
020	B	13 13 46.65, + 51 51 46.1	19.90	154 438	27
021	B	13 13 47.76, + 51 56 13.8	20.12	101 459	51
022	B	13 13 48.24, + 51 57 09.2	20.97	116 841	84
023	B	13 13 51.79, + 51 59 46.9	19.49	102 676	33
024	B	13 13 52.32, + 51 39 31.1	19.85	83 369	39
025	B	13 13 54.87, + 51 38 42.0	20.46	87 818	45
026	B	13 13 58.50, + 51 46 30.0	19.59	82 845	30
027	B	13 14 01.35, + 51 57 06.0	23.47	88 999	84
028	B	13 14 03.13, + 51 45 47.5	20.70	143 175	69
029	B	13 14 04.11, + 51 58 47.1	17.61	17 850	150
030	B	13 14 07.66, + 51 55 57.6	20.65	82 893	60
031	B	13 14 08.32, + 51 38 55.1	20.28	143 490	42
032	B	13 14 09.91, + 51 54 27.6	17.03	17 982	150
033	B	13 14 10.21, + 51 59 34.0	15.34	9875	150
034	B	13 14 11.77, + 52 04 28.2	19.49	83 207	36
035	B	13 14 14.92, + 51 59 20.6	20.54	81 654	48
036	B	13 14 15.76, + 51 47 25.5	20.39	84 610	42
037	B	13 14 16.22, + 51 39 20.6	21.10	151 044	45
038	B	13 14 16.99, + 51 41 18.2	17.37	30 336	150
039	B	13 14 17.60, + 51 50 45.7	20.69	84 365	42
040	B	13 14 17.76, + 51 57 03.2	15.98	17 586	150
041	B	13 14 17.93, + 51 59 45.5	18.29	230 894	150
042	B	13 14 20.11, + 51 59 09.1	20.87	103 548	63
043	B	13 14 21.80, + 51 54 28.3	19.60	82 689	27
044	B	13 14 22.78, + 51 48 55.1	20.03	83 906	42
045	B	13 14 23.69, + 51 41 55.0	17.35	30 183	150
046	B	13 14 26.19, + 51 50 20.3	19.25	82 893	24
047	B	13 14 27.08, + 52 01 14.4	19.79	106 429	24
048	B	13 14 32.96, + 52 02 12.2	18.98	759 027	150
049	B	13 14 33.36, + 51 55 04.2	20.39	82 611	51
050	B	13 14 34.01, + 51 48 11.9	19.85	82 707	30
051	B	13 14 34.53, + 51 39 47.5	17.56	53 363	150
052	B	13 14 35.86, + 52 00 11.3	16.88	25 995	150
053	B	13 14 38.16, + 51 44 43.7	21.72	84 428	69
054	B	13 14 38.65, + 51 56 10.3	20.17	82 281	54
055	B	13 14 38.71, + 51 44 07.6	19.56	92 369	27
056	L	13 14 41.80, + 51 48 57.5	17.12	56 244	10
057	B	13 14 42.01, + 51 58 42.5	21.28	83 045	72
058	L	13 14 42.92, + 51 48 16.3	19.95	143 644	34
059	B	13 14 44.28, + 51 51 03.7	20.34	84 679	48
060	B	13 14 45.30, + 51 55 15.2	19.64	80 920	30
061	B	13 14 45.60, + 51 38 52.8	21.01	102 172	108

Table 1 – *continued*

ID	Source	α, δ (J2000)	r	V (km s ⁻¹)	ΔV
062	T	13 14 45.78, + 51 51 05.3	19.96	85 369	130
063	B	13 14 46.37, + 51 47 52.0	20.15	84 134	39
064	B	13 14 47.76, + 51 43 43.1	21.59	143 346	54
065	B	13 14 48.48, + 51 52 46.5	20.90	82 956	69
066	B	13 14 48.94, + 52 03 24.4	17.77	90 741	150
067	B	13 14 49.29, + 51 40 48.0	17.63	29 904	150
068	B	13 14 49.50, + 52 02 23.0	16.78	52 874	150
069	B	13 14 50.39, + 51 58 46.2	20.58	82 485	42
070	B	13 14 50.66, + 52 00 58.7	21.08	82 488	72
071	T	13 14 51.28, + 51 50 07.2	20.29	83 357	77
072	L	13 14 51.28, + 51 50 11.3	20.41	168 269	38
073	T	13 14 51.41, + 51 49 41.3	19.28	81 636	50
074	B	13 14 52.38, + 51 48 14.6	20.60	85 846	66
075	B	13 14 52.39, + 52 02 20.8	19.17	102 700	18
076	B	13 14 52.73, + 51 40 31.2	18.23	85 060	21
077	B	13 14 52.84, + 51 52 41.4	20.34	83 657	51
078	T	13 14 53.56, + 51 52 19.4	19.79	52 842	62
079	T	13 14 54.53, + 51 52 21.5	20.72	17 931	145
080	T	13 14 54.68, + 51 50 55.8	19.58	81 561	58
081	L	13 14 54.73, + 51 53 16.9	16.85	53 548	10
082	T	13 14 55.33, + 51 47 53.0	18.79	84 662	75
083	L	13 14 55.38, + 51 51 46.9	21.25	184 791	104
084	B	13 14 55.51, + 51 41 03.0	19.65	576 687	150
085	B	13 14 55.52, + 52 03 45.2	20.44	90 810	51
086	B	13 14 55.69, + 51 47 19.0	19.91	81 514	30
087	T	13 14 55.79, + 51 53 02.6	20.32	53 080	100
088	B	13 14 55.85, + 51 54 28.4	20.72	80 851	66
089	T	13 14 56.12, + 51 52 57.2	19.38	53 618	90
090	B	13 14 56.25, + 51 53 09.9	20.99	81 151	60
091	T	13 14 56.40, + 51 50 52.4	20.87	80 925	122
092	L	13 14 56.83, + 51 48 53.0	20.79	733 766	263
093	T	13 14 56.90, + 51 51 00.5	19.84	82 293	77
094	T	13 14 57.88, + 51 51 41.8	20.46	81 642	115
095	B	13 14 58.09, + 51 49 16.2	21.19	86 520	180
096	T	13 14 58.76, + 51 52 14.0	19.39	53 872	41
097	B	13 14 58.77, + 51 44 22.6	20.42	83 594	33
098	T	13 14 58.77, + 51 49 50.9	19.30	84 116	70
099	T	13 14 59.06, + 51 53 13.5	18.20	52 547	39
100	T	13 14 59.15, + 51 51 38.4	20.15	81 889	72
101	T	13 14 59.32, + 51 49 46.6	20.01	80 833	115
102	T	13 14 59.54, + 51 49 57.8	21.11	83 505	137
103	B	13 14 59.56, + 51 47 52.3	21.06	82 725	63
104	B	13 14 59.56, + 52 02 44.4	19.70	102 763	27
105	T	13 14 59.91, + 51 50 10.7	20.59	85 323	125
106	T	13 15 00.12, + 51 49 05.7	20.00	83 368	90
107	T	13 15 00.12, + 51 50 30.5	19.89	81 193	97
108	T	13 15 00.56, + 51 50 03.0	20.84	86 270	190
109	T	13 15 02.01, + 51 51 06.5	19.72	81 248	53
110	T	13 15 02.02, + 51 47 47.7	20.40	84 399	102
111	B	13 15 02.22, + 51 42 37.4	19.99	90 645	45
112	B	13 15 02.25, + 51 49 51.3	21.27	81 154	180
113	T	13 15 02.41, + 51 50 19.3	19.09	81 157	65
114	T	13 15 02.51, + 51 52 19.6	21.10	84 816	135
115	T	13 15 02.87, + 51 51 48.3	20.50	81 548	145
116	T	13 15 02.89, + 51 49 43.8	20.41	84 247	105
117	T	13 15 03.02, + 51 47 02.8	20.07	84 860	100e
118	T	13 15 03.03, + 51 48 50.5	20.76	86 609	117
119	T	13 15 03.14, + 51 46 31.5	19.55	84 175	55
120	T	13 15 03.19, + 51 49 56.7	18.76	81 117	67
121	T	13 15 03.32, + 51 50 22.2	21.26	82 004	117
122	T	13 15 03.33, + 51 50 29.7	20.58	81 617	56
123	T	13 15 03.39, + 51 51 33.6	20.20	80 986	97
124	T	13 15 03.57, + 51 52 38.1	18.07	82 492	36

Table 1 – *continued*

ID	Source	α, δ (J2000)	r	V (km s^{-1})	ΔV
125	T	13 15 03.90, + 51 49 26.8	19.61	84 224	57
126	T	13 15 03.92, + 51 49 09.9	20.51	80 792	120
127	T	13 15 03.92, + 51 50 44.8	19.92	81 334	82
128	B	13 15 04.05, + 51 39 51.0	20.35	84 110	48
129	T	13 15 04.20, + 51 50 04.4	20.26	80 779	100
130	B	13 15 04.26, + 51 47 50.8	23.42	83 942	300
131	T	13 15 04.31, + 51 47 22.6	20.55	83 922	63
132	T	13 15 04.33, + 51 51 33.7	18.48	53 157	75e
133	T	13 15 04.35, + 51 51 47.1	20.66	82 945	92
134	T	13 15 04.50, + 51 49 11.1	21.57	84 697	145
135	T	13 15 04.84, + 51 49 32.0	18.74	81 410	65
136	T	13 15 05.00, + 51 50 02.7	19.33	83 533	72
137	B	13 15 05.04, + 51 46 06.3	22.58	87 210	150
138	T	13 15 05.04, + 51 52 08.2	20.27	81 517	105
139	T	13 15 05.10, + 51 49 54.0	20.50	81 889	102
140	B	13 15 05.18, + 52 03 06.9	20.05	81 331	30
141	T	13 15 05.24, + 51 49 02.6	16.76	84 697	75
142	T	13 15 05.28, + 51 48 35.3	20.92	83 155	80
143	T	13 15 05.29, + 51 47 53.1	18.37	79 194	32
144	B	13 15 05.32, + 51 45 36.5	22.04	82 923	150
145	T	13 15 05.50, + 51 50 01.7	20.46	81 449	72
146	T	13 15 05.75, + 51 52 31.7	19.15	83 298	54
147	T	13 15 05.92, + 51 48 45.1	19.16	82 249	65
148	T	13 15 06.02, + 51 49 01.4	21.22	79 868	97
149	T	13 15 06.17, + 51 49 43.0	20.60	83 816	75
150	T	13 15 06.24, + 51 48 34.9	18.34	82 684	117
151	L	13 15 06.36, + 51 54 27.9	18.13	82 178	12
152	T	13 15 06.52, + 51 51 00.5	19.72	26 235	100e
153	T	13 15 06.68, + 51 48 00.3	20.08	85 795	75
154	T	13 15 06.69, + 51 47 16.3	19.69	116 336	288e
155	T	13 15 06.96, + 51 49 25.8	19.69	85 901	72
156	T	13 15 07.05, + 51 50 49.9	20.46	83 602	100
157	T	13 15 07.27, + 51 51 20.6	20.07	84 067	80
158	T	13 15 07.41, + 51 48 22.6	19.05	85 381	37
159	T	13 15 07.43, + 51 48 37.9	21.34	83 013	102
160	B	13 15 07.45, + 51 39 35.6	20.00	84 730	51
161	T	13 15 07.47, + 51 47 36.2	20.81	83 184	105
162	T	13 15 07.71, + 51 51 03.4	17.80	25 921	100e
163	T	13 15 07.83, + 51 48 58.0	18.55	84 172	62
164	T	13 15 08.08, + 51 48 48.6	19.19	84 643	82
165	T	13 15 08.08, + 51 53 08.0	18.93	85 814	49
166	T	13 15 08.11, + 51 52 17.0	20.64	81 998	97
167	T	13 15 08.19, + 51 51 53.9	18.88	82 376	42
168	T	13 15 08.21, + 51 48 28.9	20.78	87 478	106
169	B	13 15 08.83, + 51 45 45.5	23.01	83 372	150
170	B	13 15 08.91, + 51 46 22.4	22.31	81 094	150
171	T	13 15 08.91, + 51 49 06.7	20.99	82 229	180
172	T	13 15 09.01, + 51 49 55.8	20.43	87 915	112
173	T	13 15 09.40, + 51 47 21.7	19.78	84 361	55
174	T	13 15 09.76, + 51 48 21.5	20.74	81 446	107e
175	T	13 15 09.97, + 51 53 31.5	21.04	184 767	60
176	T	13 15 10.01, + 51 52 43.6	20.31	124 393	165
177	B	13 15 10.05, + 51 43 13.3	20.20	83 252	45
178	T	13 15 10.11, + 51 52 23.2	19.87	93 275	107
179	T	13 15 10.15, + 51 50 03.2	19.31	82 684	117
180	T	13 15 10.18, + 51 53 20.0	19.93	52 960	100e
181	T	13 15 11.06, + 51 49 03.2	17.82	52 937	62
182	B	13 15 11.06, + 51 46 53.8	21.97	80 644	180
183	B	13 15 11.07, + 51 45 57.7	21.36	82 923	150
184	B	13 15 11.19, + 51 48 23.3	20.73	85 069	63
185	B	13 15 11.20, + 51 44 18.2	20.53	84 958	69
186	T	13 15 11.50, + 51 46 41.5	19.77	80 732	54
187	T	13 15 11.59, + 51 49 20.4	19.04	80 514	55

Table 1 – *continued*

ID	Source	α, δ (J2000)	r	V (km s^{-1})	ΔV
188	B	13 15 11.96, + 51 45 53.0	20.55	81 891	54
189	T	13 15 12.03, + 51 47 22.4	19.55	81 821	72
190	T	13 15 13.47, + 51 50 04.9	19.04	86 158	57
191	T	13 15 13.56, + 51 49 49.2	19.34	82 611	65
192	T	13 15 13.73, + 51 49 39.5	20.39	85 308	80
193	T	13 15 13.73, + 51 49 55.1	20.18	83 418	125
194	T	13 15 13.73, + 51 51 50.1	20.91	795 600	335e
195	T	13 15 13.98, + 51 49 10.8	19.75	80 841	90
196	T	13 15 14.53, + 51 50 03.3	19.63	81 590	50
197	T	13 15 14.69, + 51 48 17.7	20.18	82 642	100
198	T	13 15 14.72, + 51 47 20.1	20.43	82 698	102
199	T	13 15 14.81, + 51 48 30.1	20.36	79 793	72
200	T	13 15 15.18, + 51 48 04.1	18.90	82 526	62
201	T	13 15 16.11, + 51 47 46.8	19.66	85 026	65
202	T	13 15 16.29, + 51 50 23.4	18.86	82 002	52
203	B	13 15 17.14, + 51 57 05.6	19.90	84 035	33
204	B	13 15 17.17, + 51 53 57.6	19.44	80 425	24
205	B	13 15 17.62, + 51 42 00.2	22.14	207 031	45
206	T	13 15 17.69, + 51 47 25.6	20.32	116 791	88
207	B	13 15 18.03, + 52 00 55.0	19.80	82 938	30
208	B	13 15 18.38, + 51 51 41.3	20.21	102 454	30
209	T	13 15 18.91, + 51 48 43.7	19.48	85 388	60
210	T	13 15 19.72, + 51 48 08.1	19.66	82 082	70
211	B	13 15 19.98, + 51 38 42.0	19.81	91 425	42
212	B	13 15 20.01, + 51 43 40.5	21.30	84 344	102
213	B	13 15 21.87, + 51 49 03.7	19.06	82 020	33
214	B	13 15 23.11, + 52 01 25.3	16.85	17 784	150
215	B	13 15 24.09, + 52 03 59.1	15.75	18 020	150
216	T	13 15 24.22, + 51 48 18.1	19.67	82 459	77
217	B	13 15 24.90, + 51 54 58.9	16.80	34 545	150
218	B	13 15 25.33, + 51 51 06.1	19.12	86 208	27
219	T	13 15 26.22, + 51 47 55.6	19.80	85 366	82
220	B	13 15 26.50, + 51 44 53.0	20.06	79 502	30
221	B	13 15 26.86, + 51 43 26.6	19.66	84 155	30
222	B	13 15 28.43, + 51 37 52.7	21.40	137 335	111
223	B	13 15 29.91, + 51 52 57.7	21.08	135 929	48
224	B	13 15 29.93, + 51 55 44.0	19.58	84 053	30
225	B	13 15 31.56, + 51 45 34.0	21.26	82 521	54
226	B	13 15 33.48, + 51 51 16.1	21.35	84 607	63
227	B	13 15 33.75, + 51 58 54.8	20.39	81 948	39
228	B	13 15 43.19, + 51 49 50.2	20.19	116 071	54
229	B	13 15 44.27, + 51 43 33.4	19.18	86 715	30
230	B	13 15 45.61, + 51 52 01.9	19.19	82 593	27
231	B	13 15 45.69, + 52 00 28.3	22.39	304 781	93
232	B	13 15 46.20, + 51 48 23.5	19.14	84 032	24
233	B	13 15 46.60, + 51 51 25.2	19.36	82 686	24
234	B	13 15 50.10, + 51 39 16.0	20.85	129 262	57
235	B	13 15 51.52, + 51 48 47.4	19.80	78 333	33
236	B	13 15 52.41, + 51 44 41.9	21.29	159 684	45
237	B	13 15 52.41, + 51 46 41.6	20.63	102 907	51
238	B	13 15 53.47, + 51 56 58.1	20.38	132 484	18
239	B	13 15 55.87, + 51 53 13.3	20.10	56 832	36
240	B	13 15 59.50, + 52 04 14.3	20.44	95 229	45
241	B	13 16 00.79, + 51 51 07.8	17.12	56 613	150
242	B	13 16 01.28, + 51 57 00.4	20.38	132 517	45
243	B	13 16 04.09, + 51 43 47.0	21.54	117 054	63
244	B	13 16 08.41, + 51 40 25.6	20.93	89 371	75
245	B	13 16 13.15, + 52 03 27.8	19.93	116 766	30
246	B	13 16 13.47, + 52 04 54.6	20.95	100 044	63
247	B	13 16 13.62, + 51 45 08.0	21.08	102 475	57
248	B	13 16 13.79, + 51 40 25.9	20.12	89 287	48
249	B	13 16 15.13, + 51 38 48.8	17.37	7414	150
250	B	13 16 17.09, + 51 49 02.3	15.92	30 006	150

Table 1 – continued

ID	Source	α, δ (J2000)	r	V (km s ⁻¹)	ΔV
251	B	13 16 17.12, + 51 40 34.4	18.07	89 866	150
252	B	13 16 17.61, + 51 40 32.8	21.62	89 671	48
253	B	13 16 19.78, + 51 37 59.5	19.42	96 146	30
254	B	13 16 21.77, + 51 56 56.8	20.48	97 516	54
255	B	13 16 22.55, + 52 01 04.1	19.61	109 346	30
256	B	13 16 24.13, + 51 49 42.9	20.99	142 701	48
257	B	13 16 24.73, + 51 38 19.0	20.60	122 696	18
258	B	13 16 25.12, + 51 55 06.6	20.19	102 853	39
259	B	13 16 26.10, + 51 50 24.7	20.34	100 245	33
260	B	13 16 26.30, + 51 41 28.2	19.80	108 648	36
261	B	13 16 28.11, + 52 00 23.2	21.30	109 409	63
262	B	13 16 29.73, + 52 03 47.4	19.42	117 642	27
263	B	13 16 33.53, + 51 51 53.8	21.20	100 386	78
264	B	13 16 34.26, + 52 00 07.7	20.16	118 847	15
265	B	13 16 35.47, + 51 54 28.7	23.11	185 215	89
266	B	13 16 37.66, + 51 55 11.8	19.93	90 768	33
267	B	13 16 41.40, + 52 02 21.2	19.89	83 510	39
268	B	13 16 41.45, + 51 57 22.0	20.07	118 202	27
269	B	13 16 43.20, + 51 42 56.8	20.89	88 894	39
270	B	13 16 43.76, + 51 45 54.5	20.52	129 894	30
271	B	13 16 45.52, + 51 51 29.0	20.95	103 099	102
272	B	13 16 46.34, + 51 40 53.8	20.81	101 519	51
273	B	13 16 46.68, + 51 43 34.8	21.09	130 098	48
274	B	13 16 47.00, + 51 44 11.6	20.59	123 820	54
275	B	13 16 48.49, + 51 56 30.5	20.73	117 273	33
276	B	13 16 49.03, + 51 58 13.2	18.85	99 969	33
277	B	13 16 49.96, + 51 38 07.5	19.17	112 059	18
278	B	13 16 51.24, + 51 38 27.4	20.59	83 435	42

counterpart (highlighted by a black arrow on the top of Fig. 1) is also visible in the X-ray *Chandra* Image. Due to its faint magnitude ($r \sim 23$ from CFHT/Megaprime archival data) we argue that this object is a background AGN.

3 REMOVAL OF NON-MEMBERS

The removal of non-member galaxies was performed by using the two-step method called ‘P + G’ (see e.g. Biviano et al. 2013), which combines the 1D adaptive-kernel method DEDICA (1D-DEDICA; Pisani 1993) and the ‘shifting gapper’ method (Fadda 1996). For the centre of A1703 we adopted the position of the BCG (RA = 13^h15^m05^s.24, Dec. = +51°49′02.6″, see Table 1). The 1D-DEDICA method detected A1703 as a peak in the velocity space populated by 170 galaxies (see Fig. 3). Then, we rejected 23 galaxies from this provisional list of cluster members by using the ‘shifting gapper’, which combines the spatial and velocity information. The final sample contains 147 member galaxies (87 of which observed at the TNG), whose projected phase space is shown in Fig. 4.

4 GLOBAL PROPERTIES OF THE VELOCITY DISTRIBUTION

The biweight routines by Beers, Flynn & Gebhardt (1990) provide robust estimates of the first moments of the velocity distribution. Our measurement of the mean velocity is $\langle V \rangle = 83\,119 \pm 47$ km s⁻¹ (or $\langle z \rangle = 0.2773 \pm 0.0002$). The global line-of-sight (LOS) velocity dispersion is $\sigma_v = 1324^{+88}_{-71}$ km s⁻¹. Based on the σ_v – M_{200} relation of Munari et al. (2013), inferred from Λ -cold dark matter cosmological N -body and hydrodynamical simulations, we estimate a

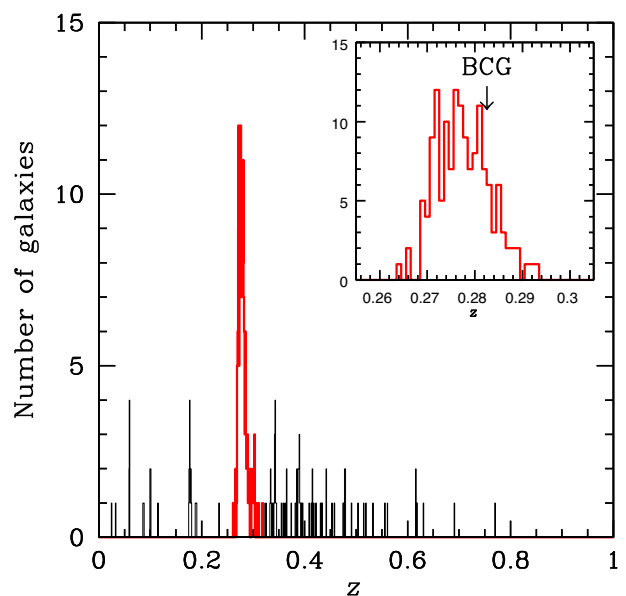


Figure 3. Redshift histogram of the galaxies of our spectroscopic sample. The solid red line histogram highlights the 170 galaxies assigned to A1703 by the 1D-DEDICA method. The distribution of the final 147 member galaxies with the indication of the BCG redshift is shown in the inset plot.

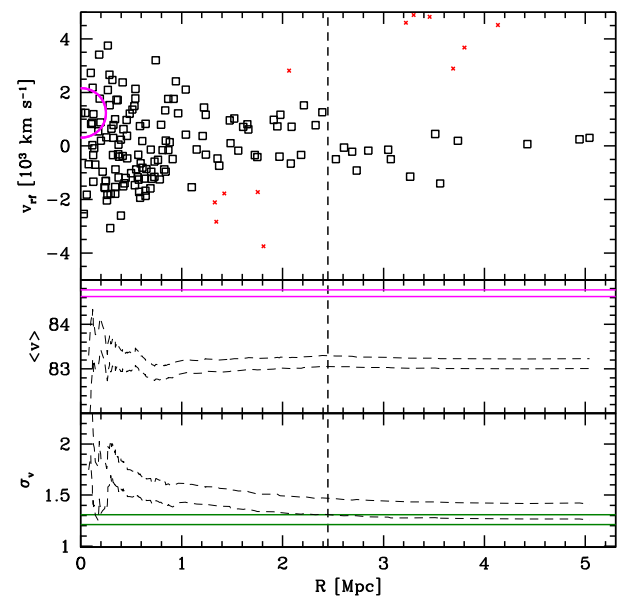


Figure 4. *Top panel:* Projected phase space diagram for the 147 member galaxies (squares) and non-members rejected with the shifting gapper procedure (small red crosses). Rest-frame LOS velocities are shown versus projected cluster-centric distance. The big magenta circle highlights the BCG. *Middle and Bottom panels:* Integral profiles of mean velocity and LOS velocity dispersion, respectively (only the one- σ error bands are shown). The mean (and dispersion) at a given (projected) radius R from the cluster-centre is estimated by considering all galaxies within that radius (the first value computed on the five galaxies closest to the centre). For all the panels, the vertical black dashed line indicates R_{200} . In the central panel, the horizontal magenta lines show the one- σ error band of the BCG velocity. In the bottom panel, the horizontal green lines show the one- σ error band of the *Chandra* X-ray temperature (from Ge et al. 2019, see Section 1) converted to σ_v (see Section 6 for details).

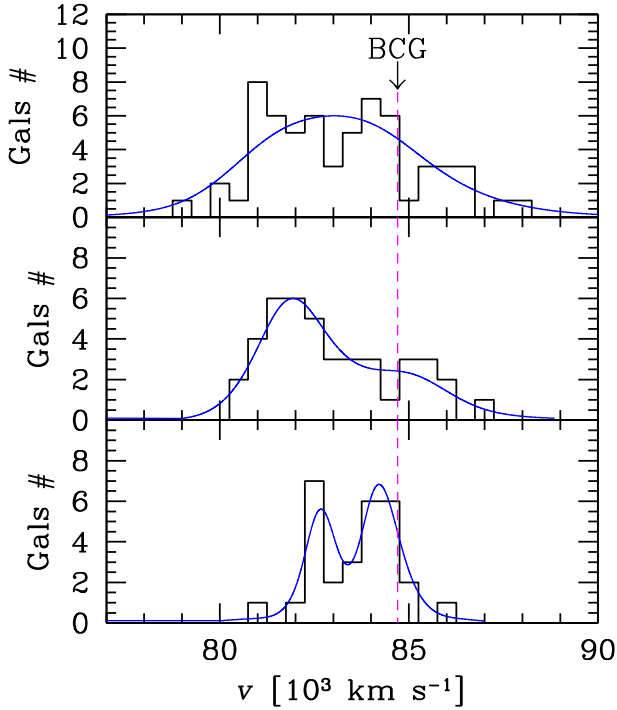


Figure 5. Velocity histogram of member galaxies in three projected cluster-centric distance intervals: 0–0.5 Mpc (*top panel*), 0.5–1 Mpc (*middle panel*), and 1–2.5 Mpc (*bottom panel*). For all panels, the blue curves show the number-galaxy density in the velocity space, as provided by the 1D-DEDICA method. The dashed magenta vertical line highlights the velocity of the BCG.

total cluster mass of $M_{200} = (2.2 \pm 0.6) \times 10^{15} M_{\odot}$ within $R_{200} = 2.45 \pm 0.15$ Mpc.

5 CLUSTER SUBSTRUCTURE

5.1 1D analysis of the velocity distribution

The 1D analysis refers to the study of the higher moments of the velocity distribution of member galaxies (Fig. 3). In particular, we find some evidence of deviation from Gaussianity according to the moments skewness and kurtosis (~ 90 –95 per cent and ~ 90 –99 per cent c.l., respectively). Moreover, there is also marginal evidence of asymmetry according to the asymmetry index (~ 90 –95 per cent c.l.; see Bird & Beers 1993 for details).

Very interestingly, the BCG has a significant (> 99 per cent c.l.) peculiar velocity ΔV_{BCG} relative to the cluster mean velocity according to the Indicator test by Gebhardt & Beers (1991). In particular, the absolute σ_v -normalized BCG peculiar velocity is $|\Delta V_{\text{BCG}}|/\sigma_v \sim 1.2$. This value puts the BCG of A1703 in the far tail of the $|\Delta V_{\text{BCG}}|/\sigma_v$ exponential distribution found by Lauer et al. (2014) studying a sample of 433 BCGs.

We then used the 1D-Kaye Mixture Model method (1D-KMM; Ashman, Bird & Zepf 1994). This test quantifies the statistical significance of bimodality (or more complex structure) in the velocity distribution with respect to a single Gaussian fit. The results are negative, i.e. there is no significant evidence of a two- or three-Gaussian partition. However, if we consider the velocity histograms obtained selecting member galaxies at different (projected) cluster-centric distances (see Fig. 5) we unveil a more complex reality. In fact, while in the centre of the cluster ($R \leq 0.5$ Mpc) the velocity

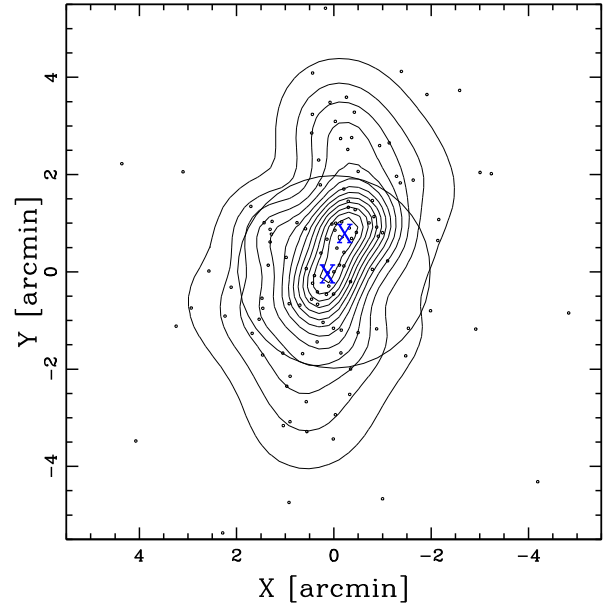


Figure 6. Spatial distribution of the spectroscopic cluster members in the central region of A1703 with, superimposed, the isodensity contour levels obtained with the 2D-DEDICA method. The blue crosses indicate the locations of the two galaxy peaks ‘S’ and ‘N’ (see also Table 2). The plot is centred on the cluster centre (the BCG) and circle contains the cluster within a radius equal to 0.5 Mpc.

Table 2. Substructures detected by the analysis of the 2D distribution of the spectroscopic members of A1703. For each subclump, the 2D-DEDICA method provides the number of assigned member galaxies N_S , right ascension, and declination of the density peak, the relative density with respect to the densest subclump ρ_S , and the χ^2 value of the galaxy peak.

Subclump	N_S	$\alpha(\text{J2000}), \delta(\text{J2000})$ h : m : s, ° : ' : ''	ρ_S	χ^2_S
N	60	13 15 03.8, + 51 49 50	1.00	33.9
S	59	13 15 06.1, + 51 49 00	0.95	27.4

distribution is fitted by 1D-DEDICA with a single peak curve, at larger cluster-centric distances signs of bimodality appear in the form of a fitted asymmetric curve. Then, for $R \gtrsim 1$ Mpc ($\sim 0.4 R_{200}$), the velocity distribution is clearly described by a two-peak curve. The significance of the two peaks, at $\sim 82\,700$ and $\sim 84\,200$ km s $^{-1}$, is > 99.4 per cent c.l. according to 1D-DEDICA.

5.2 2D analysis of the galaxy distribution

About the analysis of the 2D spatial distribution of the spectroscopic member galaxies, we employed the 2D adaptive-kernel method of Pisani et al. (1996, hereafter 2D-DEDICA). The results are shown in Fig. 6. The cluster is elongated along the SSE–NNW direction, in a similar way to the X-ray isophotes (see Fig. 1). Moreover, this test detects two dense galaxy peaks separated by only $\lesssim 1$ arcmin. The peak ‘S’ (see Table 2) is located close to the BCG (at ~ 8 arcsec ESE), the peak ‘N’ is found at ~ 9 arcsec SE of the galaxy ID 120. The eventual presence of luminosity segregation in a galaxy cluster can be a sign of disturbed dynamics (e.g. Maurogordato et al. 2011; see the discussion in Section 6). This motivates the exploration of the 2D galaxy distribution in different magnitude ranges. Very

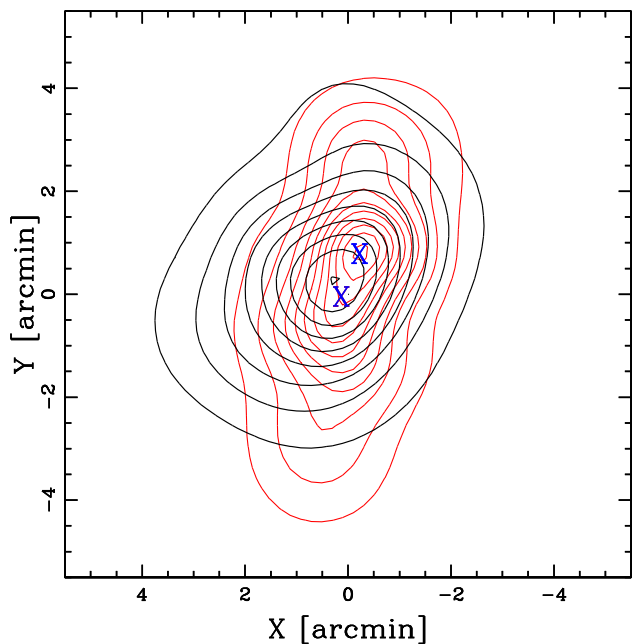


Figure 7. Isodensity contour map of the spectroscopic cluster members according to the 2D-DEDICA method. The black thick and red thin contours refer to bright ($r \leq 20$) and faint members, respectively (see the text). The blue crosses indicate the locations of the galaxy peaks ‘S’ and ‘N’ (see Table 2). The plot is centred on the cluster centre (the BCG).

interestingly, running 2D-DEDICA only on the bright spectroscopic members ($r \leq 20$) the result is completely different with respect to the analysis of the whole spectroscopic sample: the two peaks and the elongation disappear. The distribution of bright members is now much more circular (see Fig. 7) with a centre close to the BCG and to the peak ‘S’. Moreover, the distribution of faint spectroscopic members shows a peak in correspondence of the peak ‘N’.

5.2.1 Analysis of the photometric sample

We are aware that the spectroscopic sample suffers from magnitude incompleteness. This is caused by constraints in the production process of the TNG MOS masks and the positioning of Hectospec fibres for the galaxies observed by B14.

The SDSS photometry of the cluster field is deep enough to help us alleviate our incompleteness problems. In particular, we select likely members on the basis of both ($r - i$ versus r) and ($g - r$ versus r) colour-magnitude relations (hereafter CMRs). The CMRs allow us to identify the cluster ‘red’ early-type galaxies (i.e. the dominant cluster population; Dressler 1980) and to reduce the contamination by non-member galaxies. We determine the CMRs by applying the 2σ -clipping fitting procedure to the cluster members and obtain $r - i = 1.255 - 0.037 \times r$ and $g - r = 2.365 - 0.045 \times r$ (see Fig. 8). Then, within the photometric catalogue we consider as likely ‘red’ cluster members the galaxies with colour indexes $r - i$ and $g - r$ within 0.1 and 0.15 mag (i.e. the 1σ -error associated to the fitted intercept) of the respective CMRs.

Fig. 9 shows the contour map of the likely cluster members according to 2D-DEDICA. Again, we find that the distribution of bright galaxies ($r \leq 20$) shows just one peak and is only mildly elongated. Only if we consider fainter members (galaxies with 20

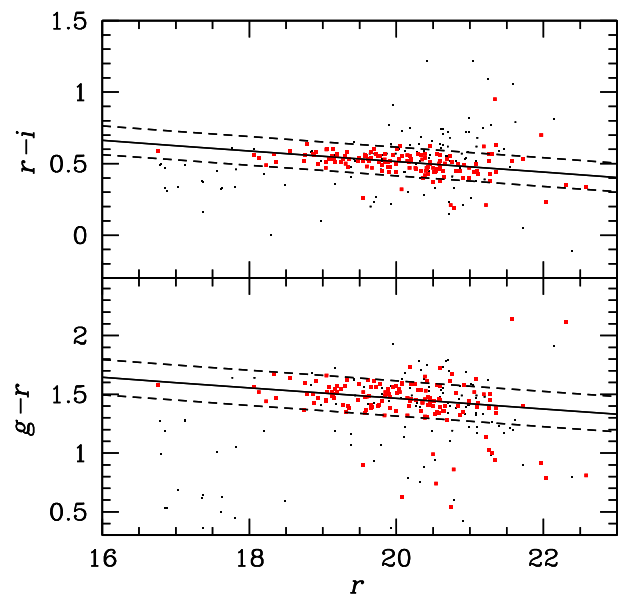


Figure 8. Upper panel: $r - i$ versus r diagram for galaxies with available spectroscopic data. The red squares are cluster members, while black points represent field galaxies. The solid line shows the best-fitting CMR as computed from cluster members; the dashed lines are drawn at $r - i \pm 0.1$ mag from the CMR (see the text). Lower panel: $g - r$ versus r diagram for galaxies with available spectroscopic data. As above, the solid line shows the best-fitting CMR and the dashed lines are drawn at $g - r \pm 0.15$ mag from the CMR.

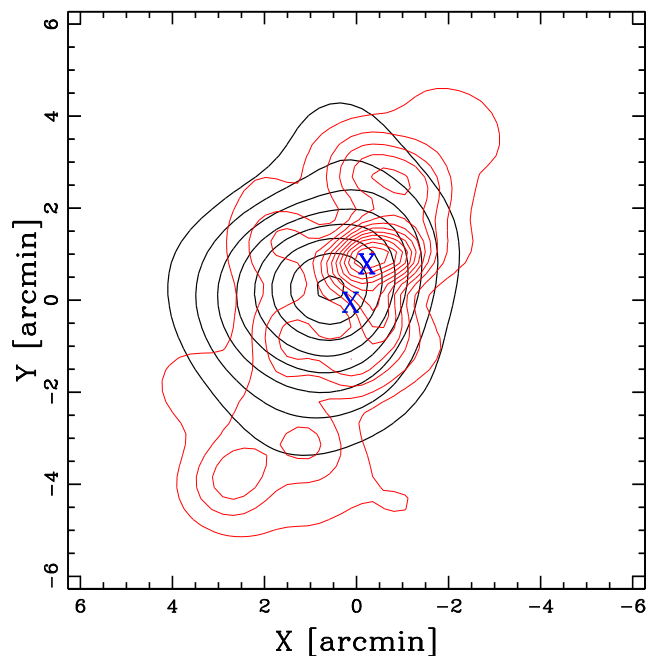


Figure 9. Isodensity contour map of the photometric likely cluster members according to the 2D-DEDICA method. The black thick and red thin contours refer to bright ($r \leq 20$) and faint members, respectively (see the text). The blue crosses indicate the locations of the galaxy peaks ‘S’ and ‘N’ (see Table 2). The plot is centred on the cluster centre (the BCG).

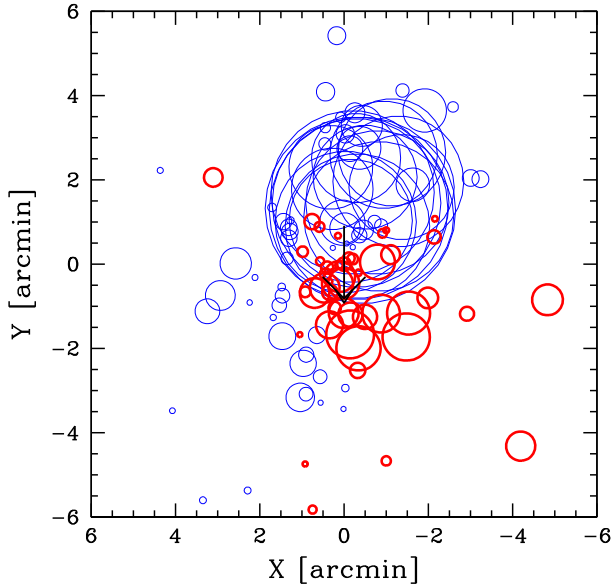


Figure 10. Spatial distribution of the cluster members in a $12 \text{ arcmin} \times 12 \text{ arcmin}$ box, each marked by a circle: the larger the circle, the larger is the deviation of the local mean velocity from the global mean velocity. In particular, the thin/blue and thick/red circles show galaxies whose local value is smaller/larger than the global one according to the modified DS-test. The plot is centred on the cluster centre (the BCG). The black big arrow shows the direction of the velocity gradient (see the text).

$< r \leq 21.5$) we recover the SSE-NNW elongation with a maximum galaxy density close to the ‘N’ peak of Table 2.

This luminosity segregation is not a unique feature of A1703. In Section 6 we discuss this result in more detail.

5.3 3D analysis: combining velocity and position information

As for the 3D analysis, we employed different tools to search for a correlation between velocity and position information, which would be a clear sign of real substructures in the cluster.

First, we searched for an eventual velocity gradient by performing a multiple linear regression fit to the cluster velocity field (den Hartog & Katgert 1996). We find marginal evidence (at the ~ 92 per cent c.l.) of a velocity gradient with $\text{PA} = 180 \pm 29$ degrees in the sample of the 147 spectroscopic cluster members. In particular, the southern region is populated by higher velocity galaxies.

Then, over the same sample we apply the classical Δ -test (Dressler & Schectman 1988, hereafter DS-test), which quantifies substructure searching for subsystems whose mean velocities and/or dispersions deviate from the global cluster values. Very significant substructure (at > 99.9 per cent c.l., checked by running a Monte Carlo shuffling of the galaxy velocities; Dressler & Schectman 1988) is found in A1703 both with the ‘canonical’ DS-test and its modified version, which consider only the local mean velocity as kinematical indicator (see also e.g. Girardi, Boschin & Barrena 2010 for more details). Again, high velocity galaxies tend to populate the southern region of the cluster.

Later, we resorted to the ‘hierarchical tree’ (Htree) algorithm developed by Serna & Gerbal (1996; see also Adami et al. 2018 and Girardi et al. 2019 for recent applications). We apply it to the catalogue of 147 member galaxies. The method computes the relative binding energies of cluster galaxies and performs a hierarchical clustering analysis to detect galaxy subsystems.

The results of the Htree test are convincing: the cluster hosts two main substructures. In the centre of the dendrogram of Fig. 11 we find the group G2. It contains 35 galaxies and, in particular, the BCG. On the right the group G1 is the most prominent structure. It contains 47 galaxies and is itself substructured in two groups,

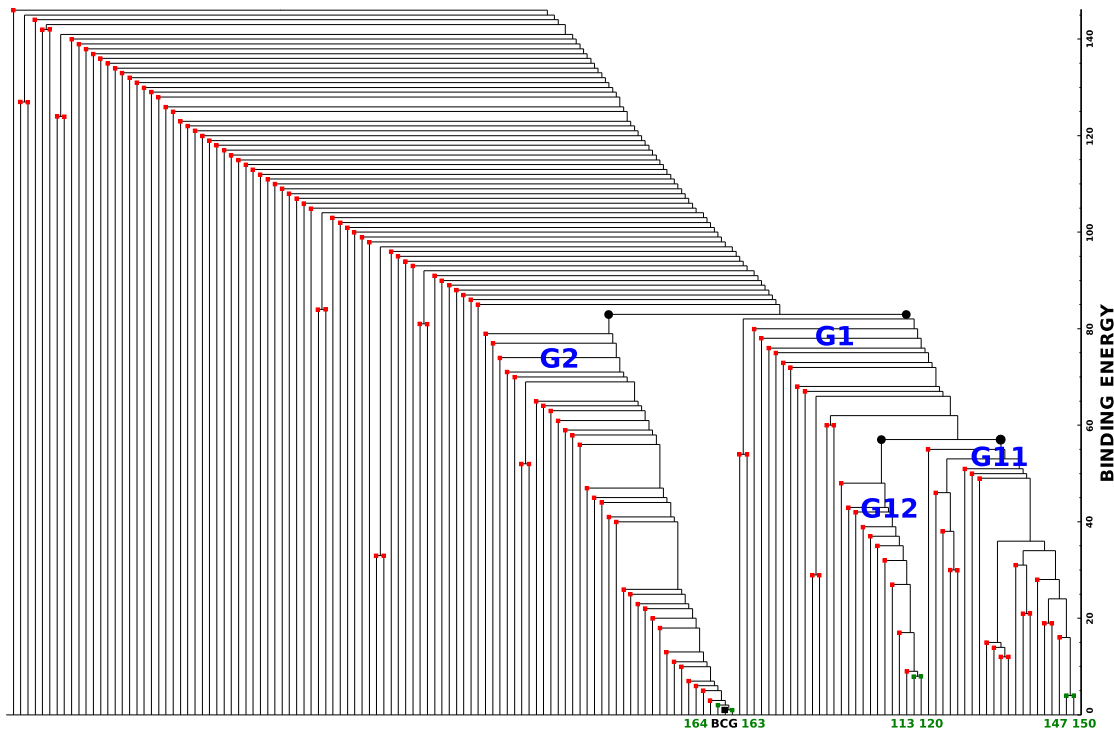


Figure 11. Dendrogram obtained through the Serna & Gerbal algorithm adopting $M/L_r = 150 M_\odot/L_{r,\odot}$. The y-axis is the binding energy, here in arbitrary units with the lowest energy levels on the bottom. Labels indicate prominent galaxies and structures.

Table 3. Substructures detected by the 3D analysis of the spectroscopic members of A1703. For each subclump, the Htree method provides the number of assigned member galaxies N_S , the mean radial velocity with its error, and the radial velocity dispersion with its error.

Subclump	N_S	$\langle V \rangle$ km s $^{-1}$	σ_v km s $^{-1}$
G1	47	$82\,022 \pm 85$	577 ± 48
G2	35	$84\,681 \pm 99$	578 ± 75
G11	21	$82\,466 \pm 83$	370 ± 69
G12	12	$81\,242 \pm 73$	236 ± 72

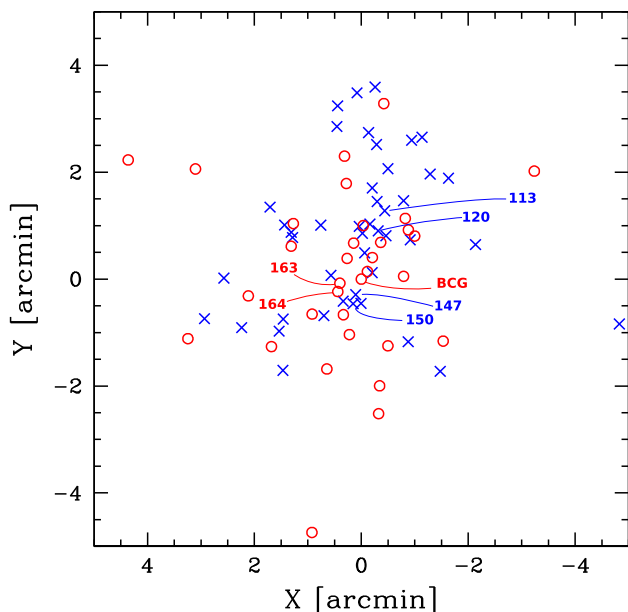


Figure 12. Member galaxies belonging to the groups G1 (blue crosses) and G2 (red circles) in a 10 arcmin $\times$ 10 arcmin box centred on the BCG. Labels highlight galaxies at the lowest binding energy levels of their respective groups (see Fig. 11).

G12 (12 galaxies) and G11 (21 galaxies). We run the Htree test by assuming a constant value of $M/L_r = 150 M_\odot/L_{r,\odot}$ for the mass-to-light ratio of galaxies, but the results are quite robust against the adopted value of M/L_r .

In Table 3 we report the properties of the subclumps found with the Htree method.

In Fig. 12 we show the positions of the galaxies belonging to the groups G1 and G2. G1 galaxies have lower velocities and populate mainly the northern region of the cluster.

Note that the velocity dispersions of the groups of Table 3 are probably underestimated, since they do not include all the galaxies of the cluster and could just be the cores of more massive structures. To overcome this point we used the 3D version of the KMM test (3D-KMM). In particular, we used the galaxy assignments of the groups G1 and G2 as a first guess when fitting two groups. The algorithm fits a two groups partition at the 98.4 percent c.l. The results for the two groups are reported in Table 4. Based on the estimates of the velocity dispersions of the two groups, we find from Munari et al. (2013): $R_{200} = 1.45 \pm 0.1$ Mpc and $M_{200} = (4.6 \pm 1.6) \times 10^{14} M_\odot$ for KMM3D1, $R_{200} = 1.6 \pm 0.2$ Mpc and $M_{200} = (5.8 \pm 2.6) \times 10^{14} M_\odot$ for KMM3D2. Therefore, the total mass would be $M_{200} = (10.4 \pm 3) \times 10^{14} M_\odot$.

Table 4. Substructures detected by the 3D analysis of the spectroscopic members of A1703. For each subclump, the 3D-KMM method provides the number of assigned member galaxies N_S , the mean radial velocity with its error, and the radial velocity dispersion with its error.

Subclump	N_S	$\langle V \rangle$ km s $^{-1}$	σ_v km s $^{-1}$
KMM3D1	86	$82\,119 \pm 85$	782 ± 66
KMM3D2	61	$84\,608 \pm 109$	845 ± 97

Table 5. Substructures detected by the 3D analysis of the spectroscopic members of A1703. For each subclump, the 3D-DEDICA method provides the number of assigned member galaxies N_S , the fitted radial velocity V_{fit} , right ascension, and declination of the density peak, the relative density with respect to the densest sub-clump ρ_S , and the χ^2 value of the galaxy peak.

Subclump	N_S	V_{fit} km s $^{-1}$	$\alpha(\text{J2000}), \delta(\text{J2000})$ h : m : s, $^\circ$: ' : ''	ρ_S	χ^2_S
DED3D1	49	84 346	13 15 04.9, + 51 49 12	0.40	30
DED3D2	43	81 242	13 15 02.8, + 51 50 16	1.00	44
DED3D3	26	82 771	13 15 07.9, + 51 48 36	0.32	20

Finally, the substructure found by the Htree test is confirmed by the analysis performed with the 3D adaptive-kernel method of Pisani (1993, 1996; 3D-DEDICA). The results are reported in Table 5, where the three groups DED3D1, DED3D2, and DED3D3 correspond to the groups G2, G12, and G11 detected by the Htree method.

6 DISCUSSION AND CONCLUSIONS

The value of the global velocity dispersion $\sigma_v = 1324^{+88}_{-71}$ km s $^{-1}$ (in agreement with the estimate of B14) is typical of a massive cluster and is consistent with the X-ray temperature $kT_X = 9.63 \pm 0.75$ keV measured by Ge et al. (2019) under the assumption of energy density equipartition between galaxies and ICM. In fact, we find $\beta_{\text{spec}} = \sigma_v^2 / (kT_X / \mu m_p) = 1.10^{+0.17}_{-0.15}$.

A value of $\beta_{\text{spec}} \sim 1$ is not anomalous for a dynamically relaxed cluster as A1703 was thought to be until very recently. However, we find convincing evidence that this cluster is experiencing a merger of two or more subclumps. The first hint of a disturbed dynamics comes from the 1D analysis of the galaxy velocity distribution. Even if its statistical moments suggest only marginal evidence of deviation from the Gaussianity, the velocity distributions of member galaxies at various cluster-centric distances (see Section 5.1 and Fig. 5) point out the possible existence of two galaxy populations with different mean velocities (separated by ~ 2000 km s $^{-1}$) but well mixed in the central ($R \lesssim 1$ Mpc) region of the cluster. However, the most compelling argument in favour of an ongoing merger in A1703 comes from the very significant peculiar velocity of the BCG (e.g. Martel, Robichaud & Barai 2014). This is in sharp contrast with what is usually found in regular clusters, where the dominant galaxy is well placed at the peak of the global velocity distribution (as in the case e.g. of CL1821 + 643; Boschin et al. 2018).

About the 2D analysis of the galaxy distribution, the 2D-DEDICA method applied on the whole set of 147 member galaxies shows that the cluster is elongated in the SSE-NNW direction (Fig. 6) with two closely spaced peaks (see Table 2), another sign that at least two subclumps are in the process of merging. Indeed, if we consider

only bright members with $r \leq 20$, their spatial distribution is much more circular and exhibits only one peak close to the BCG and the peak ‘S’. We remark that this luminosity segregation is confirmed also by the 2D analysis of the photometric members distribution, thus it does not seem an artefact of the incompleteness of the spectroscopic sample. Moreover, note that the peak ‘N’ and the density peak observed in the distribution of faint member galaxies (Figs 7 and 9) coincide with a secondary mass peak detected by Zitrin et al. (2010, see their fig. 7) from their lensing analysis. Such a spatial segregation between more and less luminous galaxies was found, for instance, in the Coma cluster (Biviano et al. 1996) and, in particular, in the merging cluster Abell 209 (Mercurio et al. 2003a), where bright photometric members are located around the dominant galaxy, while fainter galaxies seem to trace the whole cluster structure in agreement with the X-ray cluster morphology. As in the case of Abell 209 (see also Mercurio et al. 2003b), the segregation observed in A1703 could be interpreted as the result of a merging process caught after the core–core passage, where luminous galaxies trace the remnants of a pre-merging clump hosting the BCG.

Finally, the combined analysis of the velocity and spatial information provides the ultimate evidence of subclustering in A1703. In particular, the classical DS-test detects very significant substructure, with high velocity galaxies mostly located south of the BCG (Fig. 10). The Htree-test makes it possible to identify two main subclusters: G1 and G2 (Fig. 12 and Table 3). G2 hosts the BCG and is populated by high velocity galaxies. G1 exhibits a lower mean velocity and is itself substructured in two subgroups: G11 and G12. G12, in particular, hosts in the bottom of its potential well the galaxy ID 120, placed very close to the ‘N’ peak (Table 2). Comfortingly, the existence of G11, G12, and G2 is independently confirmed also by the 3D version of the adaptive kernel technique (see Table 5). The spatial distribution of G1 and G2 galaxies (see Fig. 12) explains the (marginal) evidence of a velocity gradient (Section 5.3) in the direction North–South and is in agreement with the results of the DS-test.

Using the galaxies assigned to G1 and G2 as a first guess when fitting a two groups partition, the 3D-KMM test divides the spectroscopic sample in two subclumps (Table 4). Based on their velocity dispersions we estimate a total mass $M_{200} \sim 1 \times 10^{15} M_{\odot}$ which should be interpreted as a lower limit for the mass of the cluster, being the upper limit $M_{200} \sim 2 \times 10^{15} M_{\odot}$ derived from the measurement of the global velocity dispersion (Section 4).

Now, we can adopt the above estimates for the mass of A1703 to study in more detail the merger of the two main subclusters G1 and G2. In particular, we use an analytical two-body model (see e.g. Lubin, Postman & Oke 1998 for the details of the method) which is based on the following parameters: M_{sys} , the total mass of the system, D , the projected distance between G1 and G2, and ΔV_{rf} , the relative LOS velocity (in the rest frame). From the biweight estimate of the centres of G1 and G2 we assume $D = 0.46$ Mpc, while their relative velocity is $\Delta V_{\text{rf}} \sim 2000 \text{ km s}^{-1}$. If we also assume that we are seeing the two subclusters at $t = 0.3$ Gyr after the core crossing, which is a typical time suggested by the presence of the suspected radio halo in A1703 (e.g. Brunetti et al. 2009), we can plot a M_{sys} versus α (the projection angle between the plane of the sky and the vector defined by the centres of G1 and G2) graph for three types of solutions: Bound outgoing (expanding) solutions (BO), bound incoming (collapsing) solutions (BI), and unbound outgoing solutions (UO). The results are reported in Fig. 13. The only acceptable solutions in the estimated range of mass of A1703 are BO solutions with $\alpha \sim 65\text{--}70$ deg, i.e. we are seeing a merger

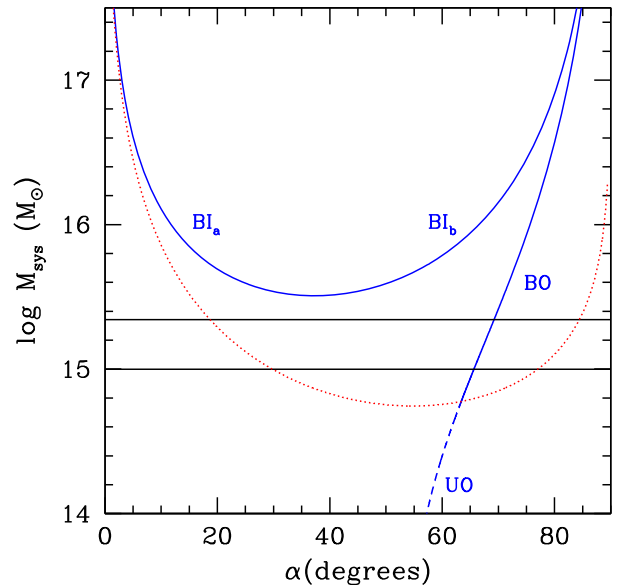


Figure 13. Bound and unbound solutions of the two-body model applied to the subclusters G1 and G2 (thick solid and dashed curves, respectively). The cluster mass is plotted versus the projection angle. Acceptable solutions are the curves intersecting the rectangle defined by the observational values of the cluster mass (bordered by the two horizontal lines). Regions above and below the thin dashed curve are the loci of bound and unbound solutions, respectively, according to the Newtonian criterion of gravitational binding (e.g. Beers, Geller & Huchra 1982).

quite close to the LOS. This is consistent with the hypothesis that the possible extended radio emission observed in the centre of the cluster is a radio halo. In fact, an eventual central radio relic produced by a shock wave propagating along the LOS is a rare event, being a relic preferentially observed in mergers occurring near the plane of the sky and in the outermost regions of a cluster (e.g. Vazza et al. 2012, Golovich et al. 2019).

In conclusion, the observational scenario suggests that A1703 is a massive cluster undergoing strong dynamical evolution, with two or three subclusters involved in a merging process. This evidence arises despite the cluster appears only slightly elongated in the X-rays, but this is not unusual. Indeed, a non-relaxed state is observed in ~ 70 per cent of clusters at $z \sim 0.2$ even in the presence in some cases of a regular shape in the X-ray imaging data (Smith, Kneib & Smail 2005). Thus, our results are in agreement with the disturbed dynamics found by Ge et al. (2019) through the study of several X-ray morphological indicators.

In this merging context, the possible presence of a radio halo in the central regions of the cluster would not be surprising, thus we rule out the hypothesis that A1703 could constitute a new intriguing case of a rare relaxed cluster hosting diffuse radio emission.

Finally, our results are in marked contrast with respect to the picture of an unimodal, dynamically relaxed cluster painted by previous studies based on gravitational lensing (see Section 1). Since cluster mass profiles inferred from gravitational lensing suffer from complicated lens geometry (see Narayan & Bartelmann 1996 for a general review), if A1703 were a rare relaxed cluster at intermediate redshift it would be of great value to test structure formation in the Λ CDM paradigm through the study of its regular mass distribution. On the contrary, the subclustering we find in A1703 seems to rule out this possibility.

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