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Isolated Massive Star Formation in G28.20-0.05

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Abstract

We report high-resolution 1.3 mm continuum and molecular line observations of the massive protostar G28.20-0.05 with Atacama Large Millimeter/submillimeter Array. The continuum image reveals a ring-like structure with 2000 au radius, similar to morphology seen in archival 1.3 cm Very Large Array observations. Based on its spectral index and associated H30 α emission, this structure mainly traces ionized gas. However, there is evidence for $\sim 30 M_{\odot}$ of dusty gas near the main millimeter continuum peak on one side of the ring, as well as in adjacent regions within 3000 au. A virial analysis on scales of ~ 2000 au from hot core line emission yields a dynamical mass of $\sim 80 M_{\odot}$. A strong velocity gradient in the H30 α emission is evidence for a rotating, ionized disk wind, which drives a larger-scale molecular outflow. An infrared spectral energy distribution (SED) analysis indicates a current protostellar mass of $m_{*} \sim 40 M_{\odot}$ forming from a core with initial mass $M_c \sim 300 M_{\odot}$ in a clump with mass surface density of $\Sigma_{\text{cl}} \sim 0.8 \text{ g cm}^{-2}$. Thus the SED and other properties of the system can be understood in the context of core accretion models. A structure-finding analysis on the larger-scale continuum image indicates G28.20-0.05 is forming in a relatively isolated environment, with no other concentrated sources, i.e., protostellar cores, above $\sim 1 M_{\odot}$ found from ~ 0.1 to 0.4 pc around the source. This implies that a massive star can form in relative isolation, and the dearth of other protostellar companions within the ~ 1 pc environs is a strong constraint on massive star formation theories that predict the presence of a surrounding protocluster.

Unified Astronomy Thesaurus concepts: Massive stars (732); Interstellar medium (847); Star formation (1569); Star forming regions (1565)

1. Introduction

Massive ($> 8 M_{\odot}$) stars impact many areas of astrophysics. However, the mechanism of their formation is still under debate. Two main scenarios are (i) core accretion (e.g., the Turbulent Core Accretion model of McKee & Tan 2003) and (ii) competitive accretion (e.g., Bonnell et al. 2001; Wang et al. 2010; see, e.g., Tan et al. 2014 for a review). The former is a scaled-up version of the standard model of low-mass star formation (Shu et al. 1987), although with the internal pressure of the massive prestellar core being dominated by turbulence and/or magnetic fields, rather than thermal pressure. Such conditions make it likely that the collapse will be more disordered than in the low-mass case, perhaps including significant accretion via overdense filaments and other substructures, e.g., as seen in magnetohydrodynamical simulations of such structures (e.g., Seifried et al. 2012; Myers et al. 2013; Hsu et al. 2021). A characteristic feature of core accretion models is a more direct linkage of the prestellar core mass function (CMF) and the stellar initial mass function (IMF), although perhaps mediated by effects of a varying core-

to-star formation efficiency and binary or small- N multiple formation by disk fragmentation within a core.

In competitive accretion, stars chaotically gain their mass via the global collapse of a cluster-forming clump without passing through the massive prestellar core phase. In the context of the competitive accretion model, there is no correlation between the CMF and the IMF as the accretion involves ambient gas materials of the cloud.

Identifying relatively isolated massive protostars provides a direct way to constrain massive star formation. These types of sources, i.e., with limited surrounding fragmentation and star formation, indicate that collapse from a massive core has occurred in a relatively monolithic manner. For instance, Csengeri et al. (2017) studied 35 sources with Atacama Large Millimeter/submillimeter Array (ALMA) and found that most of them show limited fragmentation, with at most 3 cores per clump. Louvet et al. (2019) also found low levels of fragmentation in the massive cores of the NGC-6334 region. On the other hand, Cyganowski et al. (2017) studied the massive star-forming region G11.92-0.61 finding that the three massive protostars in the region are surrounded by at least 16 lower-mass protostellar sources within a region about 0.3 pc in radius.

Protostars forming via core accretion, especially in relatively uncrowded environments, are more likely to involve an ordered transition from the infall envelope to a Keplerian disk, as has



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Table 1
Summary of ALMA Observations of G28.20-0.05

Start of Observation (epoch)	Observation Time (min.)	Config.	Antennas Used	Baseline Lengths (m)	Averaged Interval (sec.)	MRS (")	Beam Size (")
2016/4/24 07:36:04.4 (2016.31)	12.3	C (C36-2)	41	14.7–377	6.05	11.0	0.735×0.813
2016/9/11 02:45:50.2 (2016.78)	16.2	I (C36-5)	37	47.9–1400.0	6.05	3.40	0.201×0.217
2017/9/30 01:42:37.4 (2017.75)	45.3		41				
2017/11/1 00:08:10.3 (2017.84)	41.8	E (C40-9)	49	347–9740	2.02	0.340	0.0260×0.0480
2017/11/5 22:49:20.9 (2017.85)	72.5		47				

been claimed in G339.88-1.26 by Zhang et al. (2019a). They are also more likely to exhibit relatively ordered outflows, i.e., launched orthogonally to the accretion disk and maintaining their orientation for relatively long periods.

Additional observational studies of isolated massive protostars are important to test theoretical models, as they are relatively simple systems that can have high discriminatory power between the different formation scenarios. In this work, we analyze 1.3 mm (band 6) continuum and line data obtained by ALMA observations of the massive protostar G28.20-0.05. This source has been characterized as being a high luminosity ($\sim 1.4\text{--}1.6 \times 10^5 L_\odot$) (Hernández-Hernández et al. 2014; Maud et al. 2015) shell-like hypercompact H II region and a hot molecular core (Walsh et al. 2003; Sewilo et al. 2004; Qin et al. 2008) at a near kinematic distance of $d = 5.7^{+0.5}_{-0.8}$ kpc (Fish et al. 2003), based on a systemic velocity of $v_{\text{sys}} = 95.6 \pm 0.5 \text{ km s}^{-1}$ (Qin et al. 2008), which is consistent with our observations of hot core line tracers in the source (see Section 3.3). We note that some previous studies adopted the far kinematic distance of 9.1 kpc (Kurtz et al. 1994; de la Fuente et al. 2020); however, as discussed later in Section 3.2, we are able to make a new astrometric confirmation of the near distance and so adopt $d = 5.7$ kpc throughout this work.

Based on Submillimeter Array (SMA) 1.3 mm continuum emission that is assumed to trace dusty gas within a radius of 0.48 pc, G28.20-0.05 has been estimated to have a gas mass within this region of $33 M_\odot$, and thus a mass surface density of $9.52 \times 10^{-3} \text{ g cm}^{-2}$ (Hernández-Hernández et al. 2014). However, such an estimate is quite uncertain due to assumptions about dust temperatures and may also be subject to missing flux. Previous studies (e.g., Sollins et al. 2005) have suggested the presence of two components: (i) an infalling equatorial torus of molecular gas containing a central ionized region; and (ii) an extended molecular shell, which is associated wide-angle outflow or wind. Furthermore, Klaassen et al. (2009) presented SMA observations and inferred from a velocity gradient perpendicular to the outflow direction that warm molecular gas (e.g., as traced by SO_2) is undergoing bulk rotation. Klaassen et al. (2011) detected a large and wide-angle $^{12}\text{CO}(2-1)$ outflow based on the JCMT observations. Qin et al. (2008) presented a chemical study of the source with the

SMA to measure the kinetic temperature and column density of the source. Based on multiple K-components of CH_3CN , the authors measured a rotational temperature of about 300 K.

This paper is organized as follows. In Section 2, we summarize the ALMA observations and the reduction procedures. Here we also summarize Hubble Telescope (HST) near-infrared (NIR) observations of the source. We study the continuum and molecular line properties of the protostar in Section 3, including a discussion of overall morphology, kinematics, and dynamics. We measure and model the spectral energy distribution (SED) of the protostar with multiwavelength data in Section 4. In Section 5, we discuss the fragmentation and multiplicity properties of the source. Finally, a summary is presented in Section 6.

2. Observations

2.1. ALMA Observations and Data Reduction

G28.20-0.05 was observed with ALMA in band 6 via a Cycle 3 project (PI: Y. Zhang; 2015.1.01454.S) with compact (C36-2, C) and intermediate (C36-5, I)¹⁰ array configurations and via a Cycle 4 project (PI: J. Tan; 2016.1.00125.S) with an extended (C40-9, E)¹¹ configuration (see Table 1). In each case, a single pointing observation was made with a primary beam size (half power beamwidth) of $26''.9$.¹² All the observations have the same spectral setup, covering frequencies from 216.7 to 234.9 GHz. Information about the spectral bandpass used in the observations is given in Table 2. The total integration times were 12.3, 16.2, and 160 minutes in these configurations, respectively. For the Cycle 3 compact configuration observation, J1751+0939 (1.74 Jy) was used for bandpass and flux calibration, and J1830+0619 (0.31 Jy) was used for phase calibration. For the intermediate configuration observation, J1924-2914 (4.03 Jy) was used for bandpass and flux calibration, and J1851+0035 (0.24 Jy) was used for phase calibration. For the Cycle 4 observations, J1924-2914 (8.48 Jy)

¹⁰ <https://almascience.eso.org/documents-and-tools/cycle3/alma-technical-handbook> (Table 7.1).

¹¹ <https://arc.iram.fr/documents/cycle4/ALMACycle4TechnicalHandbook-Final.pdf> (Table 7.1).

¹² <https://almascience.nrao.edu/about-alma/alma-basics>

Table 2
Summary of Setups of ALMA Spectral Windows

Spectral Window	Molecular Line	Frequency Range (MHz)	Channel Spacing (kHz)	rms (mJy beam ⁻¹)
Spw0	CH ₃ OH(4 _{2,3} – 5 _{1,4})	232928.10–234928.10	15625.00	0.600
Spw1	H(30) α	231587.86–232056.61	488.28	0.810
Spw2	¹² CO(2–1)	230297.25–230765.99	488.28	0.310
Spw3	H ₂ CO (9 _{1,8} – 9 _{1,9})	218714.73–218656.14	122.07	0.717
Spw4	CH ₃ OH(4 _{-2,3} – 3 _{-1,2})	21839.79–218365.55	122.07	0.589
Spw5	C ¹⁸ O(2–1)	219514.81–219485.58	122.07	0.460
Spw6	CH ₃ CN	220278.12–220248.88	122.07	0.890
Spw7	SiO(5–4)	217147.97–217031.03	488.28	0.279
Spw8	SO ₂	216685.46–216451.08	488.28	0.710

was used for bandpass and flux calibration, and J1834-0301 (0.26 Jy) was used for phase calibration.

All data reduction and imaging was performed with CASA software (McMullin et al. 2007) using version 5.6.0–60. After pipeline calibration, we performed the continuum and line imaging separately. For the continuum imaging, we image and self-calibrate each spectral window based on line-free channels with task *tclean* using *briggs* weighting with a robust parameter of 0.5. We identify line-free channels as follows. We first inspect the raw spectrum of each spectral window and identify representative line-free channels. We then define a threshold by the root mean square ($\sigma_{\text{rms}} = \sqrt{\sum I_i^2 / N}$) of those channels. Any channels that are within four times the rms are counted as line-free channels. We then use these line-free channels to form the individual spectral window continuum images, as well as the total continuum image. We perform four iterations of phase-only calibrations with solution intervals of 30, 10 s, and *int* with manual masking. Finally, there is an iteration of amplitude calibration with manual masking, which we also apply to the line data.

The resulting synthesized beams of each configuration are summarized in Table 1. We self-calibrated each configuration before combining them using the CASA function *concat* to obtain the final combined continuum images. During combination, we weight each configuration based on their average time intervals (Table 1, column (6)).¹³ The final weightings are 1, 1, and 0.33 for the C, I, and E configurations, respectively. The combined (C+I+E) continuum was then cleaned interactively with manual masking with multiscale deconvolver at scales of 0, 10, 50, 150 pixels. The resulting synthesized beam of the final combined continuum image is $0''.060 \times 0''.036$. A summary of the observations and configuration setup is given in Table 1. We note that the maximum recoverable scale (MRS) ranges from $0''.34$ to $11''$, while the resolution ranges from about $0''.026$ to $0''.813$.

The rms noise level in a given continuum image (before primary beam correction) is done by sampling 5000 regions, each with an area equal to that of the synthesized beam. We then fit a Gaussian to this distribution of fluxes and estimate the 1σ noise level from the standard deviation of this Gaussian. The 1σ noise levels in the C, C+I and C+I+E configuration images are 1.29, 0.358, and 0.170 mJy beam⁻¹, respectively.

For molecular line imaging, we first subtract the baseline in the UV plane using *uvcontsub* in CASA. The continuum emission is obtained by subtracting line-free channels using the *uvcontsub* function. We apply manual masking during line imaging. We use *tclean* to image the emission lines, again with

briggs weighting and a robust factor of 0.5 and multiscale deconvolver at scales of 0, 10, 50, 150 pixels. Unless otherwise mentioned specifically, science measurements are performed on the primary beam-corrected images.

2.2. HST Observations

G28.20-0.05 was observed with the Hubble Space Telescope (HST) on 2016 August 8 (epoch 2016.61) with the Wide Field Camera 3 in the NIR channel (Project ID: 14494, PI: J. C. Tan). Two broad filters, F110W and F160W, covering the J (1.1 μm) and H (1.6 μm) bands and two narrowband filters, F128N and F164N, targeting the Pa β (1.28 μm) and [Fe II] (1.64 μm) lines were used. The diffraction-limited spatial resolutions for the images are $0''.13$, and $0''.16$ for the J and H bands, respectively. The pixel scale is $0''.13$. The field of view (FOV) is $2' \times 2'$, and the image is centered at (R.A., decl.) = (18:42:58.48, $-04:13:57.8$). The integration times were 202.93 s in the F110W and F160W filters and 399.23 s in the F128N and F164N filters. The reduced images were downloaded from the Hubble Legacy Archive,¹⁴ and custom Python scripts were used to analyze them. We used the Python package *DrizzlePac*¹⁵ to align the HST image to Gaia-determined astrometry (Gaia Collaboration et al. 2021). This resulted in an astrometric accuracy of 9 mas, which is consistent with their documentation.¹⁶

2.3. SOFIA Observations

G28.20-0.05 was observed with the Stratospheric Observatory for Infrared Astronomy (SOFIA) Faint Object infraRed CAmera (FORCAST; Herter et al. 2018) at 7.7, 19.7, 31.5, and 37.1 μm in 2022 February (epoch 2022.05) as part of the SOFIA Massive (SOMA) star formation survey (Project ID 09_0085, PI: J. C. Tan). The corresponding beam FWHM for all four bands is $3''.8$. The full imaging data from this observation will be presented elsewhere as part of the SOMA survey. Here we use these data to measure background-subtracted fluxes of the protostar to help constrain SED. To obtain these fluxes, standard analysis methods following those of the SOMA survey papers (De Buizer et al. 2017; Liu et al. 2019, 2020; Fedriani et al. 2022) were used. The photometric calibration error is estimated to be in the range of $\sim 3\%$ – 7% . The astrometric precision is about $0''.1$ for the SOFIA 7 μm

¹³ <https://casaguides.nrao.edu/index.php/DataWeightsAndCombination>

¹⁴ https://hla.stsci.edu/hla_welcome.html

¹⁵ <https://drizzlepac.readthedocs.io/en/deployment/index.html>

¹⁶ <https://hst-docs.stsci.edu/drizzpac/chapter-4-astrometric-information-in-the-header/4-5-absolute-astrometry>

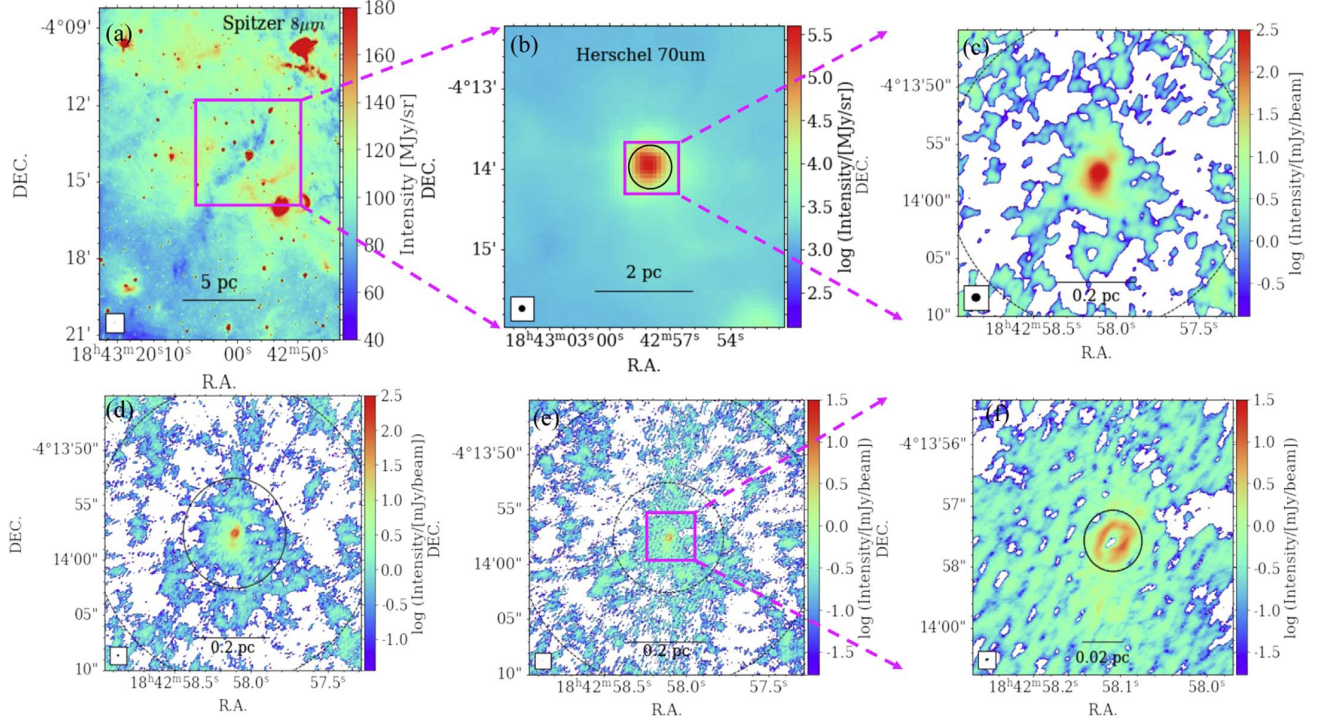


Figure 1. Multiscale views of the G28.20-0.05 massive protostar. (a) Top left: Spitzer $8\ \mu\text{m}$ image ($1''.9$ resolution) of the ~ 20 pc scale region. A filamentary IRDC, from which the massive protostar appears to have formed, is visible as a dark shadow. (b) Top middle: Herschel $70\ \mu\text{m}$ image ($5''.2$ resolution) of the ~ 5 -pc scale region. The circle shows the aperture used for MIR to FIR SED photometry of the source (see text). (c) Top right: ALMA 1.3 mm compact (C) configuration only continuum map (beam size of $0''.813 \times 0''.735$), with intensity scale from 0.1σ ($\sigma = 1.29\ \text{mJy beam}^{-1}$) to $316\ \text{mJy beam}^{-1}$. The dashed circle represents the ALMA 12 m primary beam. (d) Bottom left: the same field of view as (c), but now showing the ALMA 1.3 mm compact + intermediate (C+I) configurations continuum map (beam size of $0''.201 \times 0''.217$), with intensity scale from 0.1σ ($\sigma = 0.358\ \text{mJy beam}^{-1}$) to $316\ \text{mJy beam}^{-1}$. The solid circle shows a $5''$ radius aperture enclosing the main continuum structures, which is one scale used for flux measurements. (e) Bottom middle: as (c), but now showing the ALMA 1.3 mm compact + intermediate + extended (C+I+E) configurations continuum map (beam size of $0''.060 \times 0''.036$), with intensity scale from 0.1σ ($\sigma \sim 0.170\ \text{mJy beam}^{-1}$). (f) Bottom right: a zoom-in of panel (e) to the inner region of G28.20-0.05, which shows a ring-like structure. The solid circle shows a $0''.5$ radius aperture, which is used to measure the flux of the ring.

image and $0''.4$ at the longer wavelengths (see De Buizer et al. 2017, for further details).

G28.20-0.05 was also observed with SOFIA’s High-resolution Airborne Wideband Camera Plus (HAWC+; Dowell et al. 2010; Harper et al. 2018) at $53\ \mu\text{m}$ (band A) and $214\ \mu\text{m}$ (band E) on 2021 September 7 (epoch 2021.68; Project ID 09_0164, PI: C.-Y. Law). The full imaging data, including polarimetric imaging properties, of the source will be presented elsewhere (C.-Y. Law et al., 2022 in preparation). In this paper, we use the fluxes derived from these images to further constrain the SED of the source. The full width at half maximum at the band A and band E center wavelengths are $4''.85$ and $18''.2$. The observations were performed using the Nod-Match chop mode with a Lissajous scan pattern. The raw data were processed by the SOFIA/HAWC+ instrument team using the data reduction pipeline version 3.0.0. This pipeline includes different data processing steps, including corrections for dead pixels and the intrinsic polarization of the instrument and telescope (Harper et al. 2018; Santos et al. 2019).

2.4. Other Ancillary Data

The following archival imaging data for G28.20-0.05 were also retrieved and analyzed. Spitzer IRAC (Werner et al. 2004; Fazio et al. 2004) data at 3.6 , 4.5 , 5.8 , and $8.0\ \mu\text{m}$ from the Galactic Legacy Infrared Midplane Survey Extraordinaire Spitzer legacy survey (Benjamin et al. 2003; Churchwell et al. 2009) were

obtained and analyzed. The mean spatial resolutions are $1''.6$, $1''.7$, $1''.8$, and $1''.9$, respectively (Fazio et al. 2004).

Herschel PACS (Poglitsch et al. 2010) and SPIRE (Griffin et al. 2010) images at 70 , 160 , 250 , 350 , and $500\ \mu\text{m}$ were obtained from the Herschel High-Level Images¹⁷ in the Herschel Science Archive (HSA 2020). These Herschel images are processed to the highest level available through the Standard Product Generation pipeline (version 14.0). The image product level of the PACS and SPIRE data used is 2.5 or 3.0. The angular resolutions are $5''.2$, $12''$, $18''$, $25''$, and $36''$.

Archival Very Large Array (VLA) 1.3 cm data for G28.20-0.05 (Sewilo et al. 2011) were retrieved and analyzed. The VLA observations were carried out on 2006 March 14 (epoch 2006.2) with the K-band A-array (Program ID AZ168). The angular resolution of the continuum image is $0''.09$.

3. Characterizing the Protostar

3.1. Morphology

Figure 1 presents an overview of the G28.20-0.05 protostellar system and its surroundings. Panel (a) shows the Spitzer-IRAC $8\ \mu\text{m}$ image of the large-scale environment around the source, spanning 7 by 10 pc. The protostar is visible as a mid-infrared (MIR) -bright compact source that is embedded in

¹⁷ <https://irsa.ipac.caltech.edu/data/Herschel/HHLI/index.html>

an MIR-dark filament, i.e., an IRDC. We note that the G28.20-0.05 source is close in the sky position (about $6'$) to the well-studied, massive IRDC G028.37+00.07 (also known as Cloud C in the sample of Butler & Tan 2009, 2012), which has a similar estimated kinematic distance of 5 kpc. The uncertainties in kinematic distances are such that it is possible these sources could be in close proximity, with the projected separation being about 10 pc.

Figure 1(b) presents the Herschel 70 μm continuum map of G28.20-0.05. The dynamic range in intensity of this image spans more than a factor of 1000. In this image, the central source appears to be relatively isolated with no detection greater than 1% of peak emission within a $2'$ (3.3 pc) radius around it.

Figure 1(c) shows the primary beam-corrected ALMA 1.3 mm continuum image derived from the compact (C) configuration observation. As described above, the 1σ noise level in the central part of this image is $1.29 \text{ mJy beam}^{-1}$. The intensity scale is set to have a minimum value of 0.1σ . The image shows a central, compact source surrounded by a halo of fainter emission, but with an absence of other bright sources.

Figure 1(d) shows the primary beam-corrected 1.3 mm continuum image derived from the compact and intermediate combined (C+I) data, with 1σ noise level of $0.358 \text{ mJy beam}^{-1}$ in the central regions. Again, the intensity scale is set to have a minimum value of 0.1σ . This image reveals finer details and substructure of the central source. However, again, there is no clear evidence of strong, compact secondary sources in the wider FOV. We return to this topic with a quantitative analysis of this image for the presence of secondary sources in Section 5.

Figure 1(e) shows the primary beam-corrected 1.3 mm continuum image derived from all the configurations combined (C+I+E), while Figure 1(f) presents a zoom-in view of the central source. The range of intensities shown extends down to 0.1σ , with the 1σ noise level being $0.170 \text{ mJy beam}^{-1}$ in the central regions. The image reveals a ring-like structure with a radius from its central minimum to its bright rim of $\sim 0.01 \text{ pc}$ (2000 au). Three peaks have been identified within the ring. The main peak of the continuum emission is on the SW side at R.A. = 18:42:58.0997, decl. = $-4:13:57.636$. A secondary peak is found on the NE side, and a third relatively faint peak is toward the northern part of the ring. Outside of the ring, more extended, fainter structures are visible, especially on each side that is aligned to the apparent long axis of the ring, i.e., NW to SE.

3.2. Radio to Millimeter Spectral Index to Probe Ionized and Dusty Gas

Sewilo et al. (2008, 2011) carried out high-resolution VLA observations toward G28.20-0.05 at 1.3 cm (22.4 GHz, i.e., radio K band) and detected a ring-like structure that is similar to the one we see in the 1.3 mm continuum. Figures 2(a) and (b) present the archival VLA 1.3 cm continuum image overlaid with the E-only and C+I+E combined 1.3 mm continuum images, respectively. On first inspection, the images show very similar morphology at these two wavelengths. Assuming the 1.3 cm continuum traces free-free emission from ionized gas, this suggests that a significant portion of the 1.3 mm continuum is also contributed by such emission.

On closer examination of the VLA and ALMA images, we notice an apparent offset in peak positions and overall ring

structure. Based on 2D Gaussian fits to the 3 peak positions, this offset has a magnitude of $(34.8 \pm 9.5) \text{ mas}$ in a direction of P.A. = 216° (from VLA to ALMA). This offset is larger than the astrometric uncertainties of the VLA (9 mas)¹⁸ and ALMA (3 mas)¹⁹ observations. The VLA observations were carried out in 2006 March 14, i.e., 11.65 yr before our ALMA observations (using the 2017 November 5 date of the longest E configuration observation). Thus, the observed offset corresponds to a proper motion of $(2.99 \pm 0.82) \text{ mas yr}^{-1}$, i.e., $(81 \pm 22) \text{ km s}^{-1}$ at the 5.7 kpc distance of source. The expected proper motion due to Galactic orbital motion (assuming, for simplicity, a constant rotation curve of amplitude 200 km s^{-1} , a solar galactocentric distance of 8.0 kpc, and a kinematic distance to the source of 5.7 kpc in the direction of $l = 28^\circ.2$) is 109 km s^{-1} in the direction of decreasing l , i.e., W in Galactic coordinates. The P.A. of this direction along the Galactic plane in R.A.–decl. projection is 207° . Thus, the observed proper motion, given the uncertainties, is consistent with being entirely in this direction along the Galactic plane. Additional velocity components of $\sim 10 \text{ km s}^{-1}$ due to noncircular motions in the Galaxy, e.g., due to spiral arms or local turbulence, are also likely to be present, which can also help explain the difference between the observed motion and that predicted by the simple Galactic orbital model. We note that if the source was at the far kinematic distance (9.1 kpc), then the expected motion would be larger, i.e., 243 km s^{-1} , corresponding to 5.63 mas yr^{-1} . Thus, overall, we conclude that the observed proper motion is consistent with that expected due to Galactic orbital motion and with the magnitude strongly favoring the source being at the near kinematic distance of 5.7 kpc.

We proceed by correcting for the apparent offset, i.e., by shifting the VLA image so that it best aligns with the ALMA image. These overlaid images, the equivalent of Figures 2(a) and (b), are shown in Figures 2(c) and (d). Apart from the general close agreement between the images, we also note the presence of an extended spur of emission in the 1.3 cm image extending from the N of the ring.

We next evaluate the spectral index, α_ν , map of the source based on the ratio of the intensities at 1.3 cm and 1.3 mm. The spectral index can help diagnose the physical processes responsible for the emission. In particular, the regions where dust starts to make a dominant contribution to the 1.3 mm flux would have larger values of α_ν . We first regrid the ALMA image to the VLA resolution (i.e., a pixel scale of $0''.03$) using the imgrid function in CASA. The spectral index is defined via

$$\alpha_\nu = \log(I_{\nu_1}/I_{\nu_2}) / \log(\nu_1/\nu_2), \quad (1)$$

where $\nu_1 = 22.4 \text{ GHz}$, and $\nu_2 = 234 \text{ GHz}$. When making the spectral index map, we only consider the pixels that are 4 times the corresponding measured rms noise levels in both the ALMA and VLA images, i.e., 0.0689 and $0.00679 \text{ Jy arcsec}^{-2}$, respectively. Figures 2(e) and (f) present the maps of α_ν using the E and C+I+E ALMA images, respectively.

We see that α_ν takes values of about 0.1 in the main ring structure, as based on VLA to ALMA C+I+E data. As

¹⁸ <https://science.nrao.edu/facilities/vla/docs/manuals/oss/performance/positional-accuracy>

¹⁹ <https://help.almascience.org/kb/articles/what-is-the-astrometric-accuracy-of-alma>

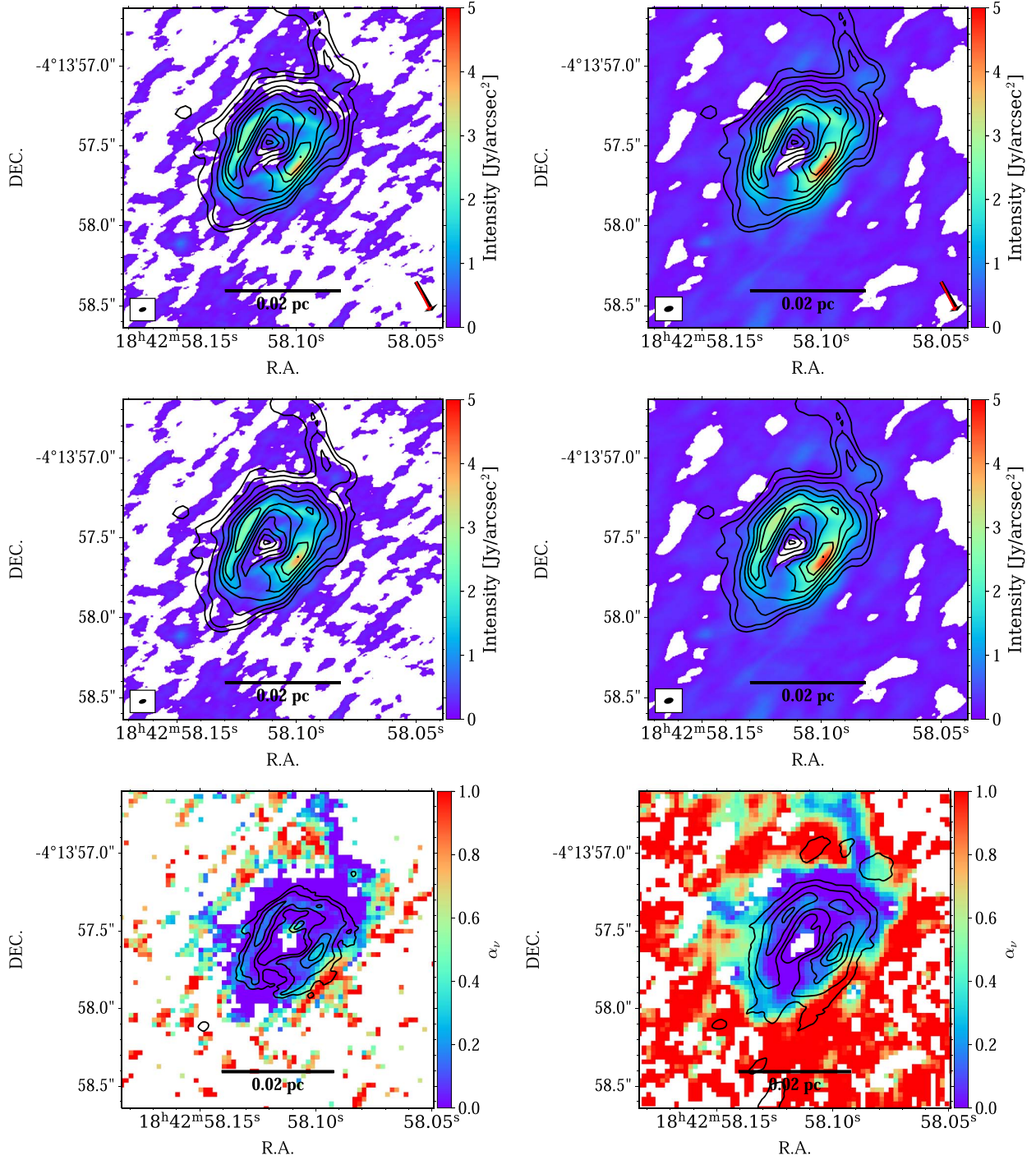


Figure 2. (a) Top left: ALMA 1.3 mm continuum image (E configuration only) of inner region of G28.20-0.05. The VLA 1.3 cm continuum is shown with contours [0.1, 0.2, 0.4, 0.8, 1.2, 1.6, 2.0, 2.4 Jy arcsec⁻²]. The beam size of 1.3 mm image is $0''.048 \times 0''.026$ (see inset), while that at 1.3 cm is $0''.09$. An offset of 34.8 mas in the direction of P.A. = 216° is seen between the peaks of the VLA and ALMA images (see red arrow in lower right, while black arrow, almost overlapping, shows the direction to Galactic center). (b) Top right: as (a), but now showing 1.3 mm C+I+E combined image, with beam size $0''.060 \times 0''.036$ (see inset). (c) Middle left: as (a), but now with the 1.3 cm image translated to align with the 1.3 mm image. (d) Middle right: as (b), but now with the 1.3 cm image translated to align with the 1.3 mm image. (e) Bottom left: map of spectral index, $\alpha_\nu = \log(I_{\nu_1}/I_{\nu_2})/\log(\nu_1/\nu_2)$, where $\nu_1 = 22.4$ GHz and $\nu_2 = 234$ GHz, i.e., evaluated using 1.3 cm and 1.3 mm (E configuration) fluxes. The contours are the 1.3 mm continuum E configuration image, with levels at [0.5, 1, 1.5, 2, 2.5, 3, 3.5, 4, 4.5, 5, 5.5, 6] Jy/arcsec². (f) Bottom right: as (e), but now using the 1.3 mm C+I+E configuration image. The overlaid contours levels are the same as in (e).

expected, when only ALMA E configuration is used, smaller values of α_ν are generally seen, which is likely due to missing flux at 1.3 mm in this case. We notice that toward the main

1.3 mm continuum peak there is a local enhancement of α_ν to values of about 0.5. There are also larger values of α_ν , i.e., $\gtrsim 1$, seen immediately surrounding the ring.

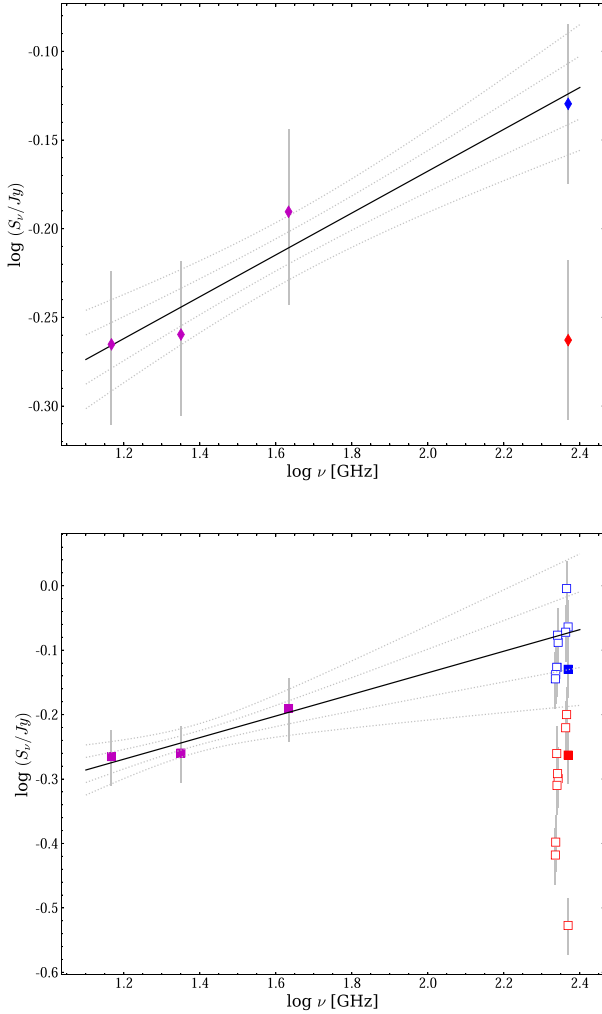


Figure 3. (a) Top: radio to millimeter SED of the inner $0''.5$ (2850 au) radius region of G28.20-0.05. Two data points are shown at 230 GHz (1.3 mm) from the ALMA observations. The higher (blue) point is based on the C+I+E combined image, while the lower (red) one is based on the E-only configuration image. Information about the other data points, which are from VLA observations, is given in Table 3. A power-law fit to the SED (including ALMA C+I+E measurement) is shown, along with 1σ and 2σ confidence intervals, with derived spectral index of $\alpha_\nu = 0.118 \pm 0.020$. (b) Bottom: as (a), but now showing separate in-band ALMA measurements from 43 to 14.7 GHz (C+I+E—blue open squares; E—red open squares). Solid squares show the equivalent average ALMA fluxes with these configurations from (a). Now the power-law fit is only done to the VLA data points and then extrapolated to the ALMA frequencies. This power law has $\alpha_\nu = 0.168 \pm 0.058$.

To obtain an average value of the spectral index of the inner region, we integrate the flux inside a radius of $0''.5$. At 1.3 cm this flux is 0.550 Jy. At 1.3 mm the flux is 0.546 Jy in the E configuration image and 0.742 Jy in the C+I+E image. Thus the average values of α_ν of the inner region are -0.00311 ± 0.06140 and 0.128 ± 0.061 , respectively. These data are shown in Figure 3(a), along with previous reported flux measurements from the VLA at 14.7 GHz (0.543 Jy with source size of $1''.0 \times 0''.7$) and 43 GHz (0.645 Jy with source size of $0''.9$; Sewilo et al. 2011; see also Table 3). We see that the three VLA data points and the ALMA C+I+E data point

Table 3
Integrated Flux Densities from 3.6 μm to 6 cm

Facility	Wavelength (μm)	Integrated Intensity (Jy)	Aperture Radius ($''$)
Spitzer	3.6	0.424 ± 0.598	15.5
Spitzer	4.5	1.34 ± 0.15	15.5
Spitzer	5.8	5.07 ± 0.57	15.5
SOFIA	7.7	9.20 ± 0.92	15.5
Spitzer	8.0	7.74 ± 1.65	15.5
SOFIA	19.5	48.5 ± 4.9	15.5
SOFIA	31.5	478 ± 48	15.5
SOFIA	37.1	696 ± 70	15.5
SOFIA	53	1449 ± 144	15.5
Herschel	70	1561 ± 77	15.5
Herschel	160	1222 ± 282	15.5
SOFIA	214	746 ± 187	15.5
Herschel	350	167 ± 82	15.5
Herschel	500	26.9 ± 33.5	15.5
ALMA	1282	$0.827(2.63)^a$	0.5(5.0)
ALMA	1294	$0.842(2.28)^a$	0.5(5.0)
ALMA	1301	$0.778(2.41)^a$	0.5(5.0)
ALMA	1362	$0.756(2.31)^a$	0.5(5.0)
ALMA	1367	$0.766(2.32)^a$	0.5(5.0)
ALMA	1372	$0.737(2.58)^a$	0.5(5.0)
ALMA	1374	$0.739(2.52)^a$	0.5(5.0)
ALMA	1382	$0.720(2.05)^a$	0.5(5.0)
ALMA	1385	$0.726(2.13)^a$	0.5(5.0)
VLA	6900	0.645 ± 0.065^b	0.9
VLA	13000	0.548 ± 0.055^c	0.5
VLA	20000	0.494 ± 0.050^c	0.8×0.6^g
VLA	36000	0.297 ± 0.045^d	3.6×3.8^g
ATCA	45000	0.326 ± 0.033^e	1.9
VLA	60000	0.150 ± 0.015^f	2.19

Notes. Information on the derivation of the MIR to FIR ($\leq 500 \mu\text{m}$) is given in the main text. The following notes relate to the millimeter to centimeter flux measurements.

^a The first number is the flux within the $0''.5$ radius aperture. The second number, in parentheses, is the C+I combined flux within the $5''$ radius aperture. The flux uncertainties are assumed to be 10%.

^b Sewilo et al. (2008).

^c Sewilo et al. (2011), but the 1.3 cm flux has been rederived here.

^d Kurtz et al. (1994).

^e Walsh et al. (1998); uncertainties were not provided, so we adopt a fiducial value of 10%.

^f Purcell et al. (2008).

^g The integrated flux was measured within the source size defined by the long and short axes, which are listed here.

can be well fit by a single power law, i.e., $\alpha_\nu = 0.118 \pm 0.020$. This suggests that a significant fraction of the 1.3 mm continuum flux on these scales is contributed by free-free emission from ionized gas, since approximately power-law behavior is often seen in the radio SEDs of ionized structures in the frequency range where they are transitioning from being partially optically thick to optically thin. However, it remains possible that the free-free emission spectrum deviates from this single power-law description, i.e., if it reaches the fully optically thin limiting value of $\alpha_\nu = -0.1$ by ~ 100 GHz. In this case, a greater proportion of the 1.3 mm flux would be expected to be contributed by dust.

As a further examination on the presence of dust in this region, we evaluate the ALMA in-band SED, i.e., within band

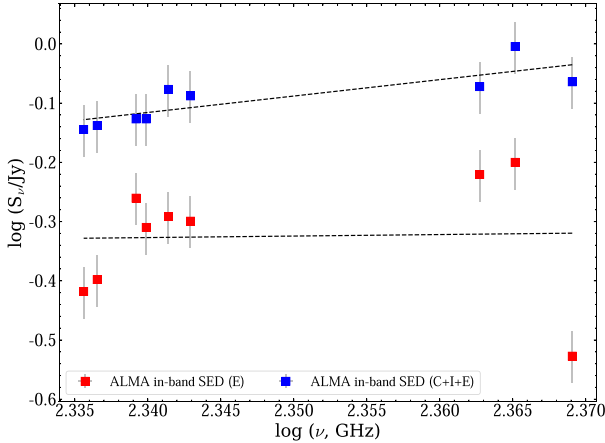


Figure 4. In-band 1.3 mm SED of the inner $0''.5$ (2,850 au) radius region of G28.20-0.05 based on E-only (red points) and C+I+E (blue points) configuration images. Power-law fits to these SEDs are shown with derived spectral indices of $\alpha_\nu = 0.255 \pm 0.930$ for E-only and $\alpha_\nu = 2.77 \pm 0.71$ for C+I+E.

6 from 216.575 to 233.926 GHz. We make the same continuum measurements of the inner $0''.5$ region, but now carried out separately in the line-free regions of each of the spectral windows of the observation. These data are shown in Figure 3(b), as well as in more detail in Figure 4. Here we assume a 10% calibration uncertainty in the integrated fluxes of each measurement, which dominates over other errors. From the in-band data, we measure the following spectral indices (α_ν), depending on which image is used: 0.255 ± 0.930 (E); 2.77 ± 0.71 (C+I+E). We note the following results for other combinations: 1.65 ± 0.26 (I); 1.43 ± 0.35 (C+I). These results suggest the potential presence of dust in the inner region, which leads to a steepening of the spectral index compared to the values seen at longer wavelengths.

If dust is contributing significantly, then we may expect local spatial variations to its contributions. In Figure 5, we present the continuum images (both for E-only and C+I+E) at 216.575 and 233.926 GHz, i.e., from spectral windows 8 and 0, respectively. We also use these data to present in-band spectral index maps.

While the morphologies are generally quite similar, we notice a modest enhancement of the higher-frequency emission in the vicinity of the main continuum peak. In this region of the ring, the in-band value of α_ν has values ~ 1 to 2. Furthermore, the region just outside the ring also shows even larger values. These two features are consistent with those seen in the VLA to ALMA spectral index map.

We thus draw the tentative conclusion that there is dust present in the inner region around the G28.20-0.05 protostar, especially around the main continuum peak and in surrounding regions outside of the ring. We will see below that there is evidence that the protostar is located at the position of the main continuum peak. However, these results motivate the need for high-resolution imaging at other frequencies, especially around ~ 300 GHz and higher to better confirm the presence of dust on these scales.

We proceed by making an approximate estimate of the 1.3 mm continuum flux from dust inside $0''.5$. For this we simply take the difference in fluxes between 234 and 217 GHz based on the in-band power-law fit, i.e., $0.94 - 0.75 = 0.19$ Jy.

For optically thin dust emission, 1.3 mm flux density corresponds to a total (gas + dust) mass surface density of

$$\begin{aligned} \Sigma_{\text{mm}} &= 369 \frac{F_\nu (1'')^2}{\text{Jy } \Omega} \frac{\lambda_{1.3}^3}{\kappa_{\nu, 0.00638}} \\ &\times [\exp(0.111 T_{d,100}^{-1} \lambda_{1.3}^{-1}) - 1] \text{ g cm}^{-2} \\ &\rightarrow 43.2 \frac{F_\nu (1'')^2}{\text{Jy } \Omega} \text{ g cm}^{-2}, \end{aligned} \quad (2)$$

where F_ν is the total integrated flux over solid angle Ω , $\kappa_{0.00638}$ is the dust absorption coefficient normalized to $0.00638 \text{ cm}^2 \text{ g}^{-1}$ (e.g., Cheng et al. 2018). This fiducial value has been derived assuming an opacity per unit dust mass of $0.899 \text{ cm}^2 \text{ g}^{-1}$ (i.e., from the moderately coagulated thin ice mantle model of Ossenkopf & Henning 1994) and a gas-to-refractory-component-dust ratio of 141 (Draine 2011). We note that the mass surface density is sensitive to the temperature of the dust, with the above value normalized to a typical hot core temperature of 100 K (see Section 3.3), i.e., $T_{d,100} \equiv T_d/100 \text{ K} = 1$. The appropriate value of T_d that should be used for this estimate is quite uncertain. Indeed, in reality there will be a range of temperatures along a given line of sight. From the types of hot core lines detected in the system (see Section 3.3) and anticipating that gas and dust temperatures are well coupled in these high-density conditions, we consider that a factor of two uncertainty in average line-of-sight temperature is reasonable. In this case, values of $T_d = 50$ and 200 K would change the coefficient in Equation (2) by factors of 2.12 and 0.486, respectively.

Applying Equation (2) to the inner $0''.5$ circular aperture of G28.20-0.05, i.e., with $\Omega = 0.785 \text{ arcsec}^2$, we estimate $\Sigma_{\text{mm}} = 10.5 \text{ g cm}^{-2}$ (averaged over this region). This corresponds to a total (gas + dust) mass of $30.3 M_\odot$. If T_d is in the range from 50 to 200 K, the mass would thus be in the range from about 60 to $15 M_\odot$.

It is possible that the optically thin assumption used for these mass estimates is not valid. To examine this possibility, we evaluate the dust optical depth $\tau = \kappa_\nu \Sigma_{\text{mm}}$. For our fiducial estimate of $\Sigma_{\text{mm}} = 10.5 \text{ g cm}^{-2}$, we have $\tau = 0.067$, which implies the optically thin approximation is valid. As shown in Figure 6, only if the dust temperature is as low as ~ 20 K does τ start to become significant. However, as discussed below, such low temperatures are not expected to be realistic for this region that is so close to a massive protostar. On the other hand, these estimates assume the dust is spread out uniformly over the $0''.5$ scale region. The actual distribution is likely to show some spatial concentration and thus involves higher values of Σ_{mm} . If the 0.19 Jy emission from dusty gas is concentrated in a region of 10 times smaller area, then Figure 6 shows that $\tau \sim 1$ for $T \lesssim 100$ K, and the method using the optically thin assumption would underestimate the mass by a significant factor. We will return to this mass estimate in Section 3.3 in the context of a dynamical mass estimate of the region.

3.3. Hot Core Environment

The ALMA spectral setup includes various molecular lines (see Table 2). The structure and kinematics of hot molecular core emission lines that trace dense and warm molecular gas can also be used to characterize the protostar. Figure 7 shows moment 0 maps of a lower excitation line of $\text{H}_2\text{CO}(3_{2,1} - 2_{2,0}$, $E_{\text{up}} = 68.1 \text{ K}$), a higher excitation line of $\text{H}_2\text{CO}(9_{1,8} - 9_{1,9}$,

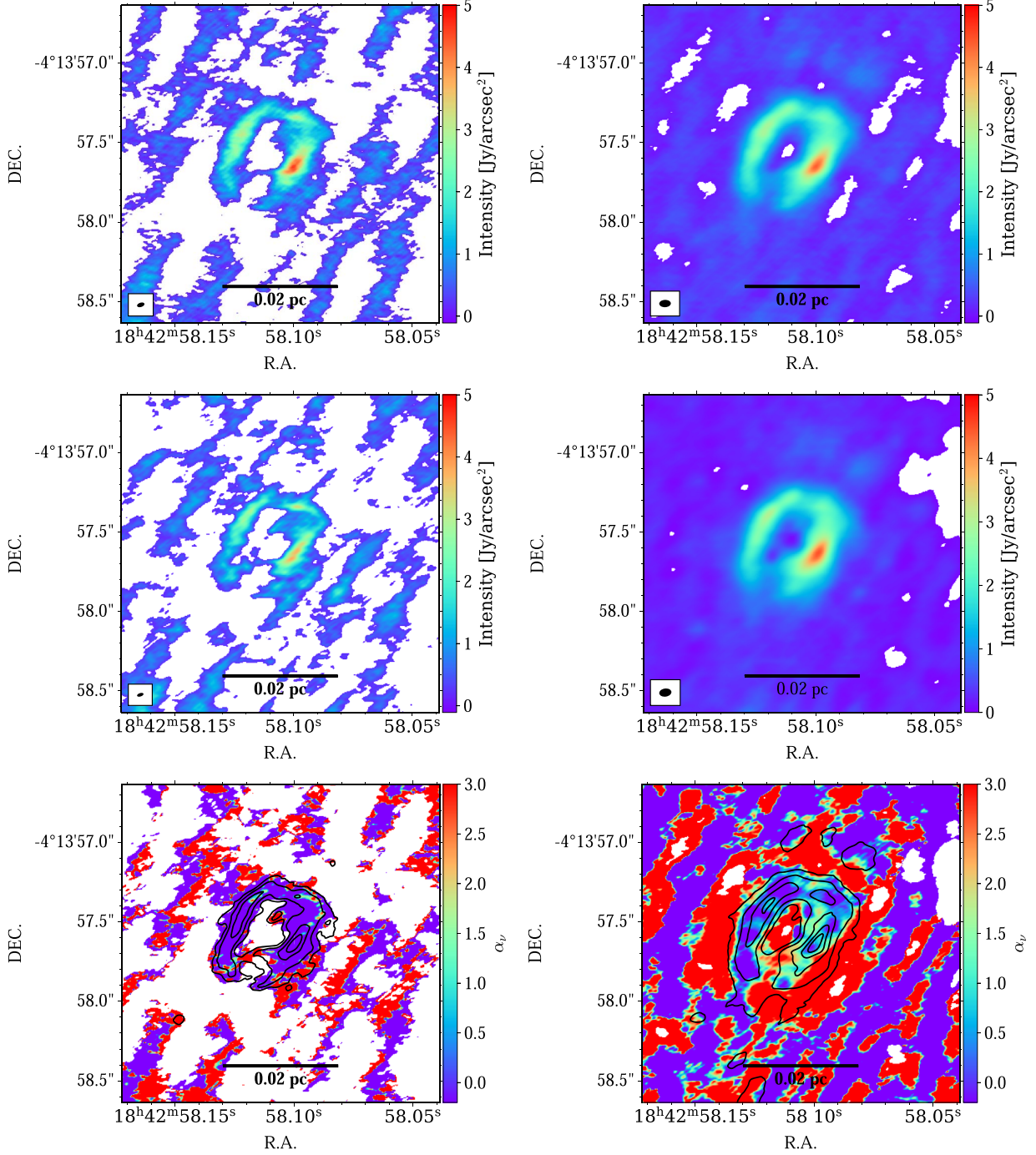


Figure 5. (a) Top left: 216.575 GHz (spw8) continuum image in E configuration only of the inner region of G28.20-0.05. The beam size is $0''.042 \times 0''.021$. (b) Top right: as (a), but now for C+I+E combined data. The beam size is $0''.073 \times 0''.046$. (c) Middle left: as (a), but for 233.926 GHz (spw0). The beam size is $0''.048 \times 0''.027$. (d) Middle right: as (b), but for 233.926 GHz (spw0). The beam size is $0''.077 \times 0''.051$. (e) Bottom left: E-only spectral index map between spw8 and spw0, i.e., $\alpha_\nu = \log(I_{\nu_1}/I_{\nu_2})/\log(\nu_1/\nu_2)$, where $\nu_1 = 216.575$ GHz, and $\nu_2 = 233.926$ GHz. The overlaid continuum contours levels are the same as in Figure 2. (f) Bottom right: as (e), but for C+I+E combined data.

$E_{\text{up}} = 174$ K), CH_3OH ($4_{-2,3} - 3_{-1,2}$, $E_{\text{up}} = 45.46$ K), and CH_3OCH_3 ($22_{4,19} - 22_{3,20}$, $E_{\text{up}} = 253.41$ K). We see that the two higher excitation species are concentrated in a region that is close to and overlapping with the main millimeter continuum peak, but with a slight offset of about $0''.2$ (i.e., ~ 1000 au). Some emission from these species is also seen extending

around and just exterior to the millimeter continuum ring. The two lower excitation species have a more extended distribution with their strongest emission just exterior to the millimeter continuum ring. These results indicate that there is dense, warm molecular gas present just outside the ring, but also even hotter gas near the main millimeter continuum peak and likely to be

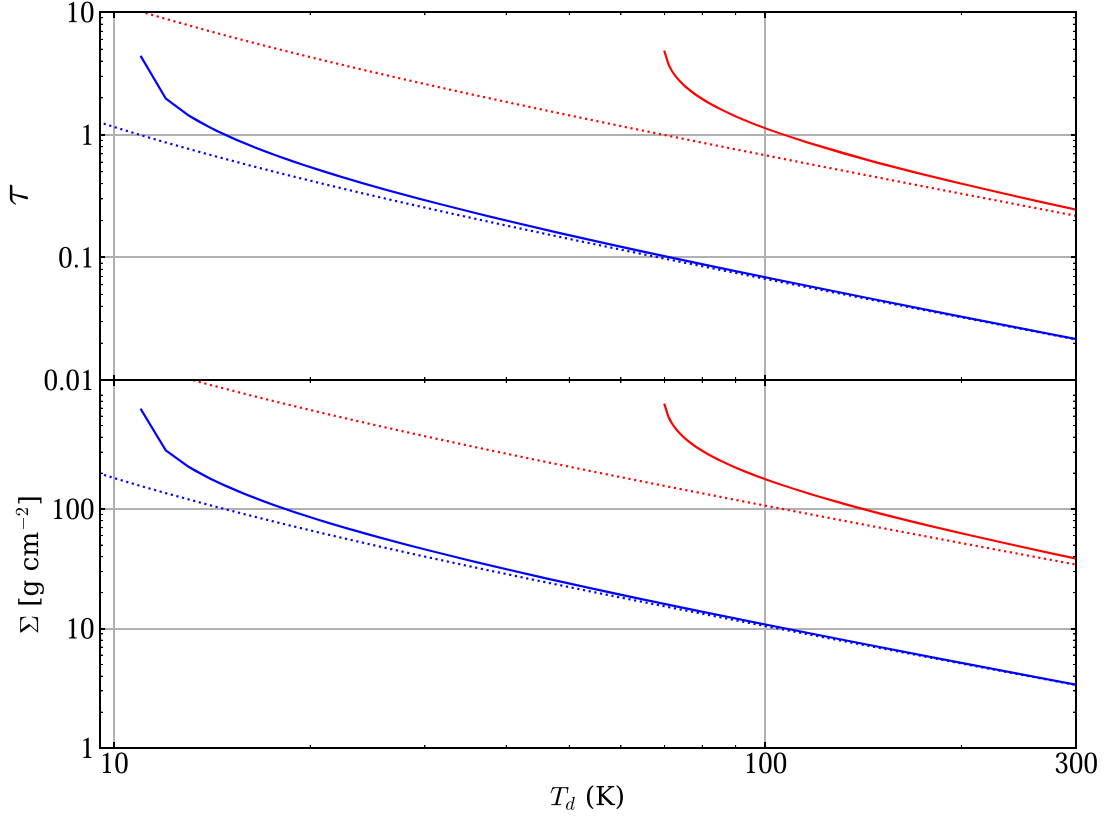


Figure 6. (a) Top: optical depth (τ) of inner-scale dust around G28.20-0.05, i.e., within a projected radius of $0''.5$ (2850 au), vs. assumed dust temperature. The blue dotted line shows τ evaluated from a uniform face-on slab of material in this region, with its mass surface density estimated assuming 0.19 Jy is due to optically thin dust emission at 1.3 mm. The solid blue line shows the equivalent τ , but allowing for optical depth in the slab. The red dotted and solid lines show the equivalent cases when assuming this 1.3 mm continuum emission is concentrated in a 10 times smaller area, e.g., a uniform slab of radius 900 au. (b) Bottom: as (a), but now showing the implied mass surface densities, Σ_{mm} , of the slabs.

heated by a source at this location. The upper state energies of these transitions indicate that a typical value of ~ 100 K for the dust temperature used above is a reasonable choice in the inner $0''.5$ scale region.

In addition, other molecules are detected in the spectra of G28.20-0.05, both relatively simple, such as SO_2 ($22_{2,20} - 22_{2,21}$, $E_{\text{up}} = 248.44$ K), H_2S ($2_{2,0} - 2_{1,1}$, $E_{\text{up}} = 83.98$ K), to more complex hot core lines, such as $\text{C}_2\text{H}_5\text{CN}$ ($27_{1,27} - 26_{1,26}$, $E_{\text{up}} = 157.73$ K). Thus, G28.20-0.05 appears to be a relatively chemically rich massive protostar, e.g., compared to other sources studied with the same spectra setup, such as IRAS 07299-1651 (Zhang et al. 2019b), G339.88-1.26 (Zhang et al. 2019a), or G35.20-0.74N (Zhang et al. 2022). A detailed chemical characterization of G28.20-0.05 will be presented in a companion paper to this one (P. Gorai et al. 2022, in preparation).

In Figure 8, we show the average spectra of the four lines shown in Figure 7. These spectra exhibit a central main Gaussian peak, but with evidence of high-velocity line wings, especially to more redshifted velocities. The lines peak at velocities close to the reported literature source systemic velocity of 95.6 km s^{-1} (see Section 1). Hence, we adopt this value as the systemic velocity of the source throughout this work.

In Figure 9, we present the moment 0, 1, and 2 maps of the CH_3OCH_3 line. The velocities near the main millimeter continuum peak are seen to be close to the systemic velocity of $+95.6 \text{ km s}^{-1}$, but become blueshifted by several kilometers

per second as one moves around the ring. The moment 2 map, which shows the estimate of the 1D line-of-sight velocity dispersion, σ , exhibits values as high as 2.5 km s^{-1} near the main millimeter continuum peak.

We now use the velocity dispersion of the highest excitation species, i.e., the CH_3OCH_3 line, to estimate a dynamical mass of the protostar assuming it traces virialized motions of a region extending out to radius, $R = 1,700 \text{ au}$ ($0''.3$). This radius is justified as being the approximate extent of the emission from the main millimeter continuum peak. The measured 1D velocity dispersion in this region is $\sigma = 2.95 \text{ km s}^{-1}$. Thus the dynamical mass assuming simple virial equilibrium ignoring magnetic fields and surface pressure terms (see, e.g., Bertoldi & McKee 1992) is

$$M_{\text{dyn}} = 5\sigma^2 R / G \simeq 84 M_{\odot}, \quad (3)$$

with this evaluation further assuming that the gas is distributed as an uniform sphere. We consider that the uncertainty in this mass estimate is at least $\sim 20\%$ due to a combination of kinematic distance uncertainty to the source, choice of radius of region traced by CH_3OCH_3 emission, and simplifying assumptions in the application of the virial theorem to the region, such as density structure, surface pressure terms, and effects of magnetic fields. Nevertheless, we see that the dynamical mass estimate is comparable to the previous estimate of dusty gas mass (see Section 3.2), but is about a factor of two

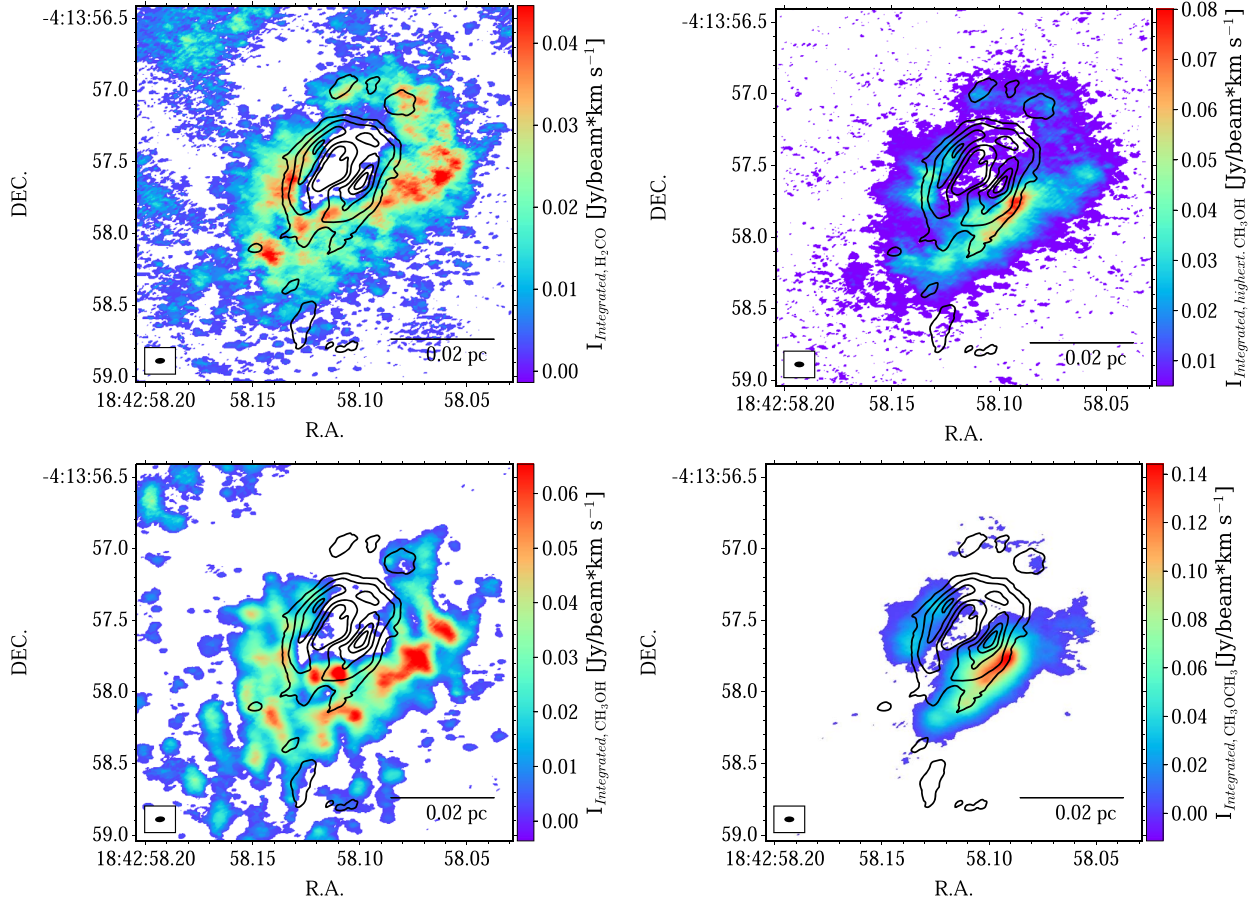


Figure 7. (a) Top left: moment 0 map of H_2CO ($3_{2,1} - 2_{2,0}$, $E_{\text{up}} = 68.1$ K) only considering cells above 1σ of the spectral rms measured over representative emission free channels. The black contours show the 1.3 mm continuum emission ($0.5, 1, 1.5, 2, 2.5, 3, 3.5, 4, 4.5, 5, 5.5, 6$ Jy arcsec $^{-2}$). (b) Top right: as (a), but for H_2CO ($9_{1,8} - 9_{1,9}$, $E_{\text{up}} = 174$ K). (c) Bottom left: as (a), but for CH_3OH ($4_{-2,3} - 3_{-1,2}$, $E_{\text{up}} = 45.46$ K). (d) Bottom right: as (a), but for CH_3OCH_3 ($22_{4,19} - 22_{3,20}$, $E_{\text{up}} = 253.41$ K).

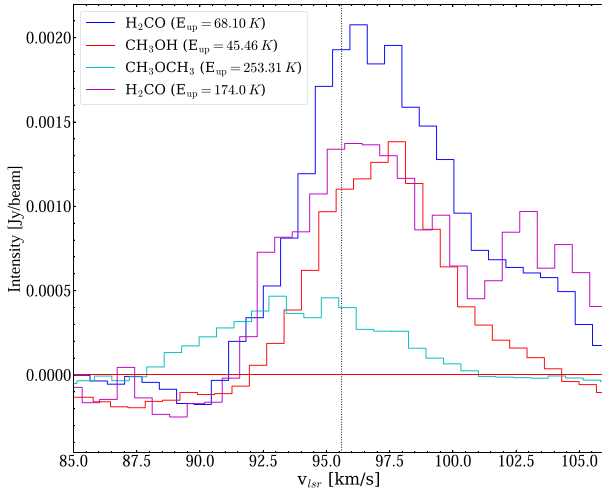


Figure 8. Averaged H_2CO ($3_{2,1} - 2_{2,0}$, $E_{\text{up}} = 68.1$ K) (blue), CH_3OH ($4_{-2,3} - 3_{-1,2}$, $E_{\text{up}} = 45.46$ K) (red), H_2CO ($9_{1,8} - 9_{1,9}$, $E_{\text{up}} = 174$ K) (magenta), and CH_3OCH_3 ($22_{4,19} - 22_{3,20}$, $E_{\text{up}} = 253.41$ K) (cyan) spectrum of G28.20-0.05 over an aperture $3''$ radius. We notice the lines both show good Gaussian-like single peak shape. The black dotted line shows the systemic velocity from the literature at 95.6 km s $^{-1}$, which is consistent to the peaks of both lines. Hence, we adopt the literature value as the systemic velocity of the source and used through this work.

larger. The dynamical mass is expected to be larger since it probes the potential of the total mass enclosed in the region, i.e., of the gas and the protostar.

3.4. $\text{H}30\alpha$ Emission

The ALMA spectral setup also includes the $\text{H}30\alpha$ recombination line that traces ionized gas. Figure 10 presents the moment 0, 1, and 2 maps of $\text{H}30\alpha$, only including pixels with values five times larger than the root mean square noise measured from line-free channels in the neighborhood of the spectral line. Figure 11 presents channel maps of the $\text{H}30\alpha$ emission. The moment 0 map shows a structure, including ring and extended NW–SE emission, that has close correspondence to the 1.3 mm continuum emission. This is additional evidence that a large fraction of the 1.3 mm continuum emission is tracing ionized gas, as already concluded from the 1.3 mm spectral index analysis.

The moment 1 map reveals a very strong velocity gradient toward the main millimeter continuum peak, which is also clearly seen in the channels maps (Figure 11). At the location of the peak, the velocity is close to the $+95.6$ km s $^{-1}$ systemic velocity of the protostar inferred from molecular lines (see above). Then, in the direction of elongation of the millimeter continuum source, there is an ordered, relatively smooth