



Publication Year	2019
Acceptance in OA	2021-01-14T15:08:39Z
Title	Search instructions for globular clusters in formation at high redshifts
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Publisher's version (DOI)	10.1093/mnras/stz785
Handle	http://hdl.handle.net/20.500.12386/29776
Journal	MONTHLY NOTICES OF THE ROYAL ASTRONOMICAL SOCIETY
Volume	485

Search instructions for globular clusters in formation at high redshifts

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Accepted 2019 March 7. Received 2019 February 8; in original form 2018 November 8

ABSTRACT

The formation of globular clusters (GC), with their multiple stellar generations, is still an unsolved puzzle. Thus, interest is rising on the possibility to detect their precursors at high redshift, hence directly witnessing their formation. A simple set of assumptions are empirically justified and then used to predict how many such precursors formed between redshift 3 and 10 could actually be detected by the Near Infrared Camera (NIRCam) instrument on board of *James Webb Space Telescope*. It is shown that the near power-law shape of the rest-frame UV continuum of young globular cluster precursors (GCPs) implies that both colours and luminosities in NIRCam long-wavelength passbands depend remarkably weakly on formation redshift. Thus, the predicted number counts depend only little on the actual formation redshifts in the mentioned range, with the exception of the bluest passbands for which counts can be strongly suppressed by intergalactic absorption along the line of sight. Instead, counts depend strongly on the actual mass of GCPs, in such a way that one NIRCam pointing should detect of the order of 10 GCPs to mag ~ 30 if their mass distribution was the same of today GCs, or over 1000 if their mass was 10 times higher. Therefore, GCP number counts will set fairly tight constraints on the initial mass of GCs. An encouraging agreement with the number density of candidate GCPs at $z = 6-8$, revealed by the Hubble Frontier Fields programme, suggests that their initial mass could be at least four times higher than that of their local descendants if all were to end up as GCs.

Key words: globular clusters: general – galaxies: evolution – galaxies: formation – galaxies: high redshift.

1 INTRODUCTION

The formation of globular clusters (GC) along with their multiple stellar populations remains a major unsolved issue in astrophysics. A new opportunity to attack the problem has recently emerged in view of the *James Webb Space Telescope* (JWST) operations, i.e. the direct observation of forming GCs at high redshifts. Actually, the first suggestion of its possible feasibility is quite old, with Carlberg (2002) having made early predictions on the expected luminosity function and clustering of high-redshift GCs (up to $z \lesssim 10$). More recently, the observability of GCs in formation at high redshift has been addressed by Katz & Ricotti (2013), Trenti, Padoan & Jimenez (2015), Renzini (2017), and Zick, Weisz & Boylan-Kolchin (2018), whereas there are hints that some GC precursors (GCPs) may have been already detected (Vanzella et al. 2016, 2017a,b; Bouwens et al. 2018). As emphasized in these papers, the search for *first*

galaxies, GCPs, and the agents of cosmic reionization are tightly interconnected from an observational point of view, and also include a possible direct role of GCPs in the reionization (see also Ricotti 2002; Schraerer & Charbonnel 2011; Boylan-Kolchin 2018).

In this paper we present some of the expected properties of GCPs at high redshifts, such as luminosities, colours, and luminosity functions, specifically for the passbands of the Near Infrared Camera (NIRCam) on board of JWST. In doing so we capitalize on the most salient properties of GCs in our Galaxy and in other galaxies in the local Universe, including their old age, broad metallicity distribution, mass function, compactness and puzzling multiple populations. These properties are succinctly summarized here.

The bulk of GCs are assigned ages of 12.5 ± 1 Gyr, with a possible trend of metal rich ones being slightly (~ 1 Gyr) younger (Marin-Franch, Aparicio & Piotto 2009; VandenBerg et al. 2013; Brown et al. 2014). A few clusters some Gyrs younger than this also exist, though they tend to have masses lower than typical GCs (Marin-Franch et al. 2009). Here we ignore this minor component and consider the mentioned age range as encompassing virtually all GCs in the Milky Way (MW). We also assume that this age range

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applies to the bulk of GCs, not only in the MW but also in the local Universe as a whole (see e.g. Puzia et al. 2005). A lookback time of 12.5 Gyr corresponds to $z \simeq 5$ and when accounting for the ± 1 Gyr uncertainty the redshift range we consider for the GC formation epoch becomes $3 \lesssim z \lesssim 10$. It has been argued, and there is some evidence in support, that massive star clusters form in gas-rich mergers even today (Ashman & Zepf 1992), though the old age of GCs in the MW suggests that our Galaxy did not experience much merger-driven GC formation in the last 10 Gyr, or more. We also note that about 20 per cent Galactic GCs have ages younger than ~ 11 Gyr according to Marin-Franch et al. (2009) and VandenBerg et al. (2013), hence their progenitors would be found at redshifts lower than ~ 3 , near the peak of cosmic star formation density at $z \sim 2$ (Madau & Dickinson 2014). GCs forming at this epoch ($z \sim 2$) were likely embedded in a metal-rich, high-extinction environment, hence more difficult to detect. Moreover, many of these younger GCs are very sparsely populated, hence are irrelevant in the present contest, unless they were orders of magnitude more massive at formation. For these reasons, we focus on the redshift range beyond 3, as that offering the best chances for GCP detection.

The present mass function of GCs, both in the MW and in other studied galaxies, is well represented by a lognormal distribution. From Harris et al. (2013) we adopt 1.5 Mpc^{-3} for the local number density of GCs, with their mass distribution peaking at $\sim 2 \times 10^5 \text{ M}_\odot$. For the distribution itself, we adopt the Gaussian distribution in $\log(M)$ as from Harris et al. (2014), their equation (1), and use it to describe the mass distribution:

$$\frac{dN}{d\log M} = N_0 \exp \left[-\frac{(\log M - \log M^*)^2}{2 \times 0.52^2} \right], \quad (1)$$

where M^* is the mass at the peak of the Gaussian, with $\log M^*/\text{M}_\odot = 5.3$ and 0.52 is the σ of the Gaussian, as from fig. 4 in Harris et al. (2014). Integrating this distribution from $-\infty$ to $+\infty$ and setting it to 1.5 Mpc^{-3} one gets the normalization $N_0 = 1.15 \text{ Mpc}^{-3}$. Adopting a GCP mass function with the same Gaussian shape as that of local GCs is a conservative assumption, in principle giving a lower limit to the expected number counts of GCPs. Indeed, it has been argued that the mass function at formation may have been much different from that of the GCs surviving today, such as a power law (e.g. Fall & Rees 1977; Gnedin & Ostriker 1997; Vesperini 1998; Fall & Zhang 2001). However, disruption is expected to affect predominantly lower mass clusters, hence fainter GCPs that may well be below detectability even with *JWST*. In practice, the precise shape of the CGP mass function below the peak is completely irrelevant in the present context.

For our reference case, we assume that the mass function of GCPs has exactly this shape, however with the peak mass M^* being 10 times higher than in the local Universe, i.e. $\sim 2 \times 10^6 \text{ M}_\odot$. This hypothetically higher value of M^* at GC formation is meant to comply with the widely invoked necessity of GCs being substantially more massive at birth in order to account for the multiple populations that are ubiquitous among MW GCs (more below). One generally refers to it as the ‘mass budget problem’ (e.g. Renzini et al. 2015, and references therein). We emphasize that we are not arguing for the mass budget factor to be 10, as this value is used only for illustrative purposes, with the understanding that its actual value can only be established by future observations, in particular with NIRCam on board *JWST*. Our paper is meant to provide easily scalable predictions that future observations can test, hence setting direct observational constraints on the actual value of the mass budget factor. The assumption for the factor of the order of 10 upscale of the mass of GCPs can be supported by

the following arguments. The mere mass-loss from individual stars (stellar winds and supernovae) accounts for a ~ 40 per cent mass reduction from formation to the present. On top of it, star losses via evaporation, tidal interactions and the like would account for a further mass reduction that is difficult to quantify and that depends on the structure of the GCPs that may have been different from that of the surviving GCs. For example, it has been argued that all GCs, or at least the metal poor ones, may have formed as nuclei of dwarf galaxies with most such hosts having later dissolved with their bare nuclei becoming the GCs of today (Searle & Zinn 1978). In this respect, the Fornax DSph galaxy and the Sagittarius galaxy, with their exceptionally high GC frequency (e.g. Brodie & Stradler 2006, see also Georgiev et al. 2010) lend some support to this notion. Highly reminiscent of this scenario is the recent finding at $z \simeq 6$ of a star-forming dwarf with a size of $\sim 400 \text{ pc}$ and a stellar mass of $\sim 2 \times 10^7 \text{ M}_\odot$ hosting a compact, unresolved nucleus with $R_e \lesssim 13 \text{ pc}$ and a mass of $\sim 10^6 \text{ M}_\odot$, perhaps the best example so far of a GCP (Vanzella et al. 2019). Moreover, a substantially higher mass for GCPs has been invoked by virtually all scenarios for the formation of the multiple population phenomenon, though none of such scenarios is able to account for all the complexities of the observational evidence.

An upper limit to the mass budget factor is set by considering Galactic GCs in the context of the Galactic stellar halo. The total mass in halo GCs is $\sim 3 \times 10^7 \text{ M}_\odot$ and the mass of the halo is about 30 times higher. So, even if the whole halo was formed by stripped GCPs, the GCPs could not have been more than ~ 30 times more massive than the combined present mass of halo GCs. In any event, the size of this adopted mass upscale is perhaps the most important unknown quantity that high redshift observations may allow us to measure, with the understanding that only GCPs in the high-mass portion of the distribution will have a chance to be detected, as we shall see in the sequel.

In the MW the metallicity distribution of GCs is very broad, from less than 1/100 solar to nearly solar, with a hint of bimodality that is evident in the GC families of massive external galaxies (e.g. Brodie & Stradler 2006; Harris 2010). In the MW the metal rich GCs, with, say $[\text{Fe}/\text{H}] \gtrsim -0.5$, are confined within the bulge (e.g. Barbuy, Bica & Ortolani 1998) and even the most metal rich ones appear to be coeval with the bulge itself (Ortolani et al. 1995). A few GCs with supersolar metallicity may well exist in other massive galaxies. Now, the metal poor GCs must pre-date the formation of the major part of the body of today’s host galaxies, which will certainly help their detection also for being virtually unobscured by dust not unlike faint very high redshift galaxies with their steep UV continuum (Bouwens et al. 2014, 2015; Vanzella et al. 2019). Conversely, metal rich GCPs must have formed only after a substantial galaxy was already in place, having reached high metallicities and therefore the young GCPs are likely to have been deeply embedded in dust, hence substantially extincted in the UV, and therefore much more difficult to detect. In summary, our best chances to detect GCPs at high redshift are offered by the metal poor ~ 50 per cent fraction of the total population and of it by those in the high-mass side of the distribution.

All studied GCs harbour multiple stellar populations of various complexity, as most vividly illustrated by *Hubble Space Telescope* (*HST*) multiband photometry (e.g. Piotto et al. 2015; Milone et al. 2017). The natural interpretation of the multiple population phenomenon is in terms of successive stellar generation, i.e. as a series of two or more burst of star formation, with second-generation bursts being even stronger than the first one in most massive GCs (Milone et al. 2017). We will not try to model such

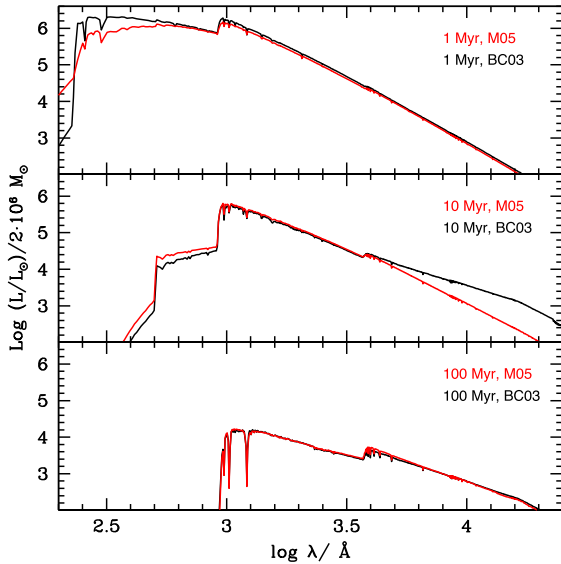


Figure 1. The spectral energy distributions of SSP models for three ages representative of very young star clusters, namely 1, 10, and 100 Myr (from top to bottom). Red and black curves refer to models by M05 and BC03, respectively, with the same IMF, and similar metallicity ($Z = 10^{-3}$ and $Z = 4 \times 10^{-3}$), respectively.

multiple bursts, but following Elmegreen (2017) we assume that at least the main burst is completed in less than ~ 1 Myr, a time shorter than the evolutionary times of massive stars (~ 3 to ~ 30 Myr, depending on mass). Such a short time is a direct consequence of the compactness of GCs, for which Elmegreen estimates a free-fall time of ~ 0.03 Myr with star formation being completed in ~ 0.3 Myr. Thus, for the purposes of this paper, this argument allows us to approximate GCPs with simple stellar populations (SSPs), i.e. a set of coeval, chemically homogeneous stars, though we know that they are not. We shall return on this point in Section 6.

In summary, the number density of GCPs we are going to estimate refer to precisely the precursors of those objects that we recognize as GCs in the local Universe and whose mass function is given by equation (1), as adopted from Harris et al. (2014). It may well be that in the early Universe objects existed similar to such CGPs but which have disappeared in the meantime. We are not trying to include such objects (see Carlberg 2002 for an attempt to do so) and therefore the estimates presented in this paper can be regarded as lower limits. Even so, the local volume density of GCs (1.5 Mpc^{-3}) implies that one single frame of NIRCcam will include over 200 000 GCPs in the redshift range 3–10 (Renzini 2017), caught in whatever stage of their formation and evolution, from being still a gas cloud before forming stars, to be at the peak of its star formation rate, to possibly having already dimmed below detectability. The question is, how many of them could be caught as bright enough to be detected?

The standard cosmology ($H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_m = 0.3$, $\Omega_\Lambda = 0.7$) and AB magnitudes are adopted.

2 THE MODEL SSPs

We use the stellar population models of Maraston (2005), hereafter M05,¹ to describe the early spectroscopic and photometric evolution of the progenitors of present-day globular clusters, assuming they

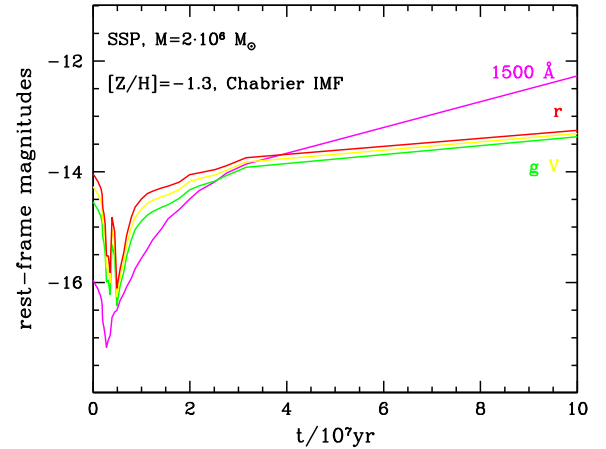


Figure 2. Rest-frame absolute AB magnitudes at $1500 \text{ \AA}$ and for the r , g , and V filter passbands for the SSP models considered here, shown for $t < 100$ Myr and our typical total stellar mass M_* of $2 \times 10^6 M_\odot$.

formed in the redshift range $3 < z < 10$. Though these models are available for ages from 0 to 15 Gyr, in this work we shall focus at most on the first billion year of evolution as older models fade below *JWST* detection limits (as we shall show later). For simplicity we consider only one chemical composition, namely a fractional abundance of heavy elements Z as $[Z/H] = -1.35$, which lies near the middle of the metallicity distribution of present-day Milky Way GCs (Harris 2010). In any event, at the stellar ages of interest here the opacity in the envelope of massive young stars is dominated by electron scattering, hence the spectral energy distribution (SED) of young SSPs is fairly insensitive to metallicity. Finally, for all models we adopt the initial mass function (IMF) of Chabrier (2003), from 0.1 to $100 M_\odot$.

2.1 Rest-frame spectral and photometric evolution

In Fig. 1 we show the rest-frame spectra of three selected models for ages that are relevant to this work (1, 10, and 100 Myr, from top to bottom) and compare them to analogue SSP models by Bruzual & Charlot (2003), hereafter BC03, which are based on different stellar evolutionary tracks and on the same library of stellar spectra as M05. For completeness we display the models over a wide wavelength range ($90\text{--}25\,000 \text{ \AA}$), but note that differences between them at the shortest wavelengths or in the rest-frame near-IR are not relevant in the present context. Below $\text{Ly}\alpha$ the flux is absorbed by the intergalactic medium (IGM) at high- z , while the rest-frame near-IR lies outside the NIRCcam range. In the range $\sim 1000\text{--}9000 \text{ \AA}$ rest frame, which is the one sampled by NIRCcam passbands at the redshifts of interest, the models are very similar, hence our predictions would have been the same if using the BC03 stellar population models.

Fig. 2 displays the rest-frame magnitudes at $1500 \text{ \AA}$ and for the r , g , and V filter passbands for the M05 SSP models, shown for $t < 100$ Myr and for a total stellar mass of 2×10^6 . In the UV the models are brightest at an age of 3 Myr, when the most massive stars start to die, whereas in optical bands they are brightest at slightly later times due to the appearance of red supergiants. After reaching the brightest luminosity $\sim 3\text{--}10$ Myr since formation (depending on wavelength), all models fade monotonically in all considered bands. Hence, catching them within the first ~ 10 Myr since formation gives the best chance to observe the progenitors of present-day globular clusters. For instance, comparable UV

¹www.maraston.eu

Table 1. Example of observer-frame magnitudes in *JWST* filters for a GCPs of mass $\log(M/M_\odot) = 6.3$. The first column gives the redshift at which the object is observed, having already aged about 3 Myr since its formation, hence formed at a slightly higher redshift. The full table can be found at: <https://sites.google.com/inaf.it/pozzetti-gcps/home>.

Redshift	log(age/yr)	m_{F070W}	m_{F090W}	m_{F115W}	m_{F150W}	m_{F200W}	m_{F277W}	m_{F356W}	m_{F444W}
10.00	6.50	32.02	31.69	30.22	30.29	30.54	30.89	31.23	31.37
9.00	6.50	31.72	30.78	30.04	30.22	30.48	30.85	31.16	31.27
8.00	6.50	31.42	30.03	29.92	30.14	30.41	30.82	31.02	31.24
7.00	6.50	30.58	29.65	29.80	30.04	30.34	30.75	30.91	31.22
6.00	6.50	29.57	29.47	29.67	29.93	30.27	30.58	30.86	31.20
5.00	6.50	29.12	29.30	29.53	29.81	30.18	30.41	30.81	31.15
4.00	6.50	28.86	29.09	29.34	29.67	29.91	30.31	30.73	31.05
3.00	6.50	28.57	28.81	29.11	29.38	29.65	30.17	30.57	30.89
2.00	6.50	28.10	28.41	28.63	28.90	29.35	29.86	30.28	30.65

luminosities (between $M_{1500} = -17$ and -15) have been derived for candidate GCPs with similar masses ($1-20 \times 10^6 M_\odot$) found by Vanzella et al. (2017a,b, 2019), and for tiny lensed sources identified by Bouwens et al. (2017) at $z \sim 6-8$.

2.2 Redshift evolution of population models

The observer-frame properties of the models in the NIRCcam filters are calculated by red-shifting the rest-frame model SEDs to a family of redshifts (from $z = 10$ down to $z = 3$, in steps of $\Delta z = 0.1$) and for a series of times since formation, hence redshift at which they are observed. The cosmological dimming is calculated using the Flake code (*Flexible-k-and-evolutionary-correction*, Maraston, in preparation). For an assumed cosmology, the procedure, calculates the observed-frame magnitudes in arbitrary photometric filters for all model ages and redshifts, including $z = 0$. A (large) table of possible evolutionary paths is output, which besides providing the observed-frame and absolute magnitudes of objects with arbitrary ages and star formation histories, it also allows a quick evaluation of the *K*-correction in various filters without the need to approximate. Moreover, as all model ages are considered at all redshifts, there is no need to assume one specific formation redshift in order to follow the evolution, as any choice for this parameter is possible. In this work we shall experiment with a set of formation redshifts and other parameters, as described in the following sections. Table 1 provides an example of the observer-frame magnitudes in all *JWST* filters at different redshift and ages for our typical GCPs of mass $2 \times 10^6 M_\odot$ based on the adopted SSP models. In particular we list observer-frame magnitudes near the brightest phase, i.e. at age of 10 Myr since its formation, and hence formed at a slightly higher redshift. The full table can be found here.² Note that these magnitudes do not yet include the high- z absorption by the intervening IGM.

3 COLOURS AND LUMINOSITIES OF YOUNG SSPS AS SEEN BY JWST

In this section we make predictions on the detectability of GCPs by *JWST*, using NIRCcam imaging under the assumptions mentioned above. NIRCcam offers high sensitivity imaging³ from 0.6 to 5.0 μm in eight broad-band filters (*F070W*, *F090W*, *F115W*, *F150W*, *F200W*, *F277W*, *F356W*, *F444W*) and consists of two modules pointing to adjacent fields of view, separated by 4.4 arcsec. Each

module observes simultaneously in a short-wavelength channel (0.6–2.3 μm) and in a long-wavelength channel (2.4–5.0 μm). The total field of view (FoV) of each NIRCcam pointing is 9.7 arcmin².

Using the stellar population models presented in the previous section, we show in Fig. 3 the observed spectra for a GCP/SSP of mass $2 \times 10^6 M_\odot$, at different ages and redshifts from $z = 3$ to $z = 10$, from young (10^6 yr) to old (10^9 yr) ages. We focus to those on the youngest ages ($\sim 10^{6.5}$ yr) as they correspond to the brightest phase of GCPs, hence with the highest chance of being detected. Indeed, after the brightest phase, the flux drops quite rapidly by at least 1 mag in a time-scale of few Myr at all wavelengths and redshifts. Notice that for such young ages the spectrum longward of the Lyman break is well represented by a power law with $F_\lambda \sim \lambda^\beta$ with only a mild evolution during the first ~ 10 Myr from $\beta \simeq -3$ to -2.5 . Steep UV spectral slopes are indeed ubiquitous among very high redshift galaxies, getting steeper with decreasing luminosity and approaching $\beta \sim -3$ in the luminosity range expected for GCPs (see fig. 1 in Bouwens et al. 2014). Such steep spectral slopes are also very similar to those observed in candidate GCPs (Vanzella et al. 2016, 2019). This steep UV slope of young SSPs plays a critical role in determining the predicted luminosities and colours of detectable GCPs as a function of their formation redshift. For this reason, in the reddest channels the flux in the brightest phase is not dramatically lower at $z = 10$ compared to $z = 3$.

Fig. 3 illustrates how NIRCcam filters sample different spectral ranges depending on redshift, from UV to optical going from bluest to reddest filters and from lower to higher redshifts. At $z > 5$ the bluest filters cover a range of wavelengths shorter than Lyman break and Ly α , which are affected by the absorption by the high- z hydrogen in the IGM intervening along the line of sight. At $z = 7$ this effect is important for the *F070W*, and *F090W* passbands and at $z = 10$ also for the *F115W* passband, hence GCPs at such redshift will appear as drop-out objects in those bands. As a consequence, at $z \geq 7$ the GCPs are detectable only in the complementary longer wavelength channels.

To guide the eye, we show in Fig. 4 the rest-frame wavelengths as sampled by the various NIRCcam filters at redshifts in the range $2 < z < 10$, and the corresponding rest-frame spectral ranges. In particular, the short-wavelength channel samples the UV rest frame at $z > 5$. For example, in the range $3 < z < 10$ the filter *F200W* observes the rest frame from $\lambda_{\text{REST}} = 1800$ to 5000 Å. Conversely, in the same range the rest frame at 1500 Å will be observed only by filters bluer than *F150W*. The long-wavelength camera instead would cover almost exclusively the rest-frame optical range ($2500 \text{ Å} < \lambda_{\text{REST}} < 10\,000 \text{ Å}$), even at $z = 10$. Furthermore, as already pointed out, the effect of the IGM absorption,

²<https://sites.google.com/inaf.it/pozzetti-gcps/home>

³<https://jwst-docs.stsci.edu/display/JTI/NIRCcam+Sensitivity>

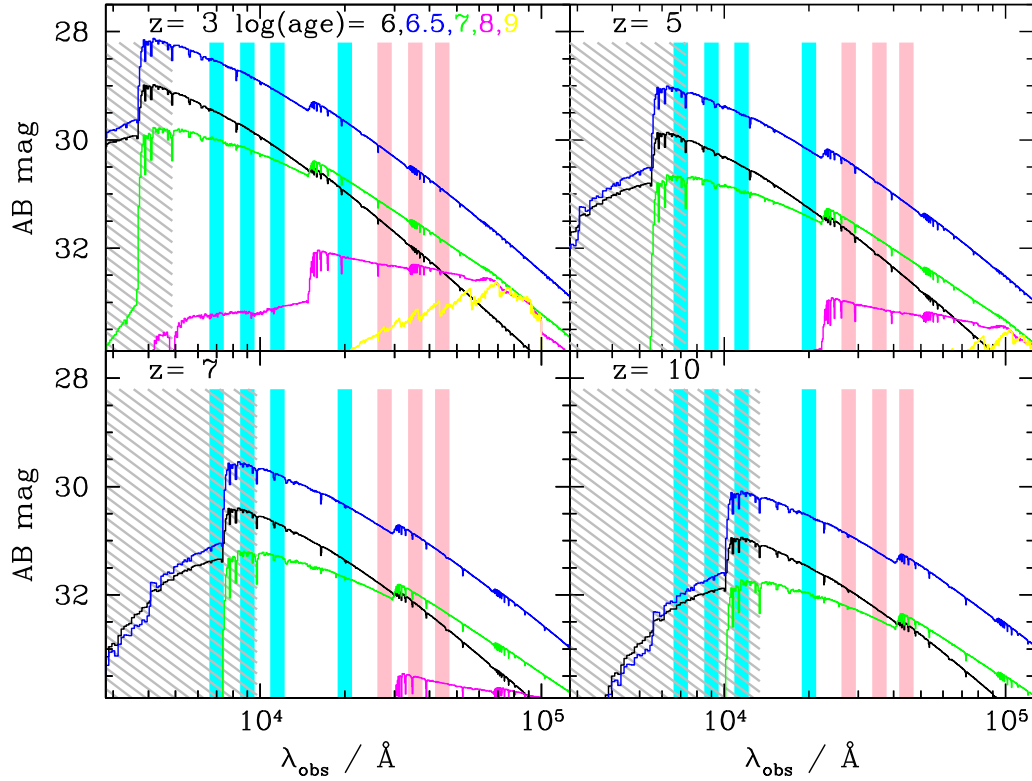


Figure 3. Expected observer-frame spectra for our typical GCP of mass $2 \times 10^6 M_{\odot}$, at different ages (from 10^6 to 10^9 yr) and at different redshifts (at $z = 3, 5, 7, 10$ in the four panels). Fluxes are expressed in AB magnitudes. Vertical bands reproduce the *JWST* filters from the two short and long-wavelength channels (in blue and pinkish, respectively). In the grey shaded region spectra are affected by high- z IGM absorption.

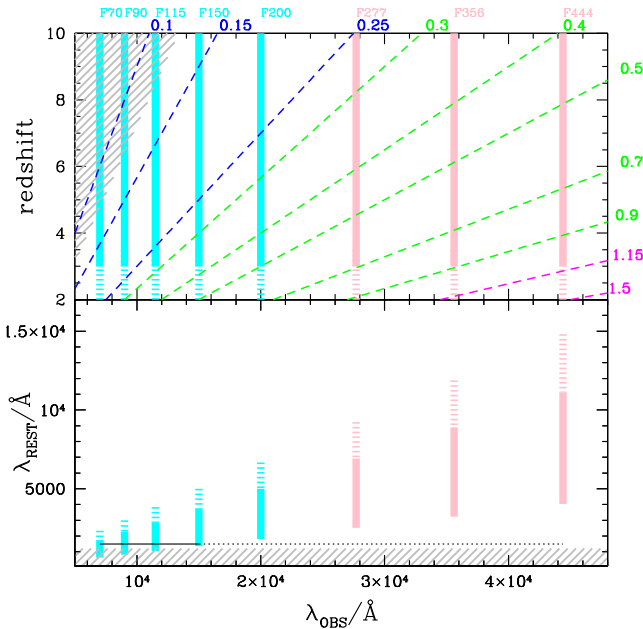


Figure 4. The rest-frame wavelength sampled by Nircam filters for objects at different redshifts. *Top panel:* The dashed lines labelled by the rest-frame wavelength (in μm) show the corresponding observed wavelength as a function of redshift. The vertical bands give the central wavelength of the Nircam filters, as in Fig. 3. Thus, the intercepts of the dashed lines with the vertical bars give the redshifts at which a given rest-frame wavelength will be sampled by the various Nircam passbands. *Bottom panel:* The vertical bars give the full rest-frame wavelength range sampled by the Nircam filters for objects in the redshift range $2 < z < 10$ (dashed bands for $2 < z < 3$, continuous bands in the redshift range $3 < z < 10$).

dropping-out the object from bluer passbands, is important only for F070W, F090W, and F115W and for $z > 5, 6.5$, and 8.5 , respectively.

Using the table of redshifted magnitudes described in Section 2.2, we derive the expected fluxes/magnitudes in the various Nircam passbands as a function of age and redshift/epoch of formation (z_f and t_f , hereafter), to which we add the IGM absorption effect by attenuating the resulting flux/magnitudes using the prescriptions of Madau (1995). Fig. 5 shows the expected magnitudes, as a function of time since formation and as a function of redshift, for representative formation redshifts in our interval ($z_f = 3, 5, 7, 10$). Adopting a typical GCP with mass $\log(M^*/M_{\odot}) = 6.3$ we find that the maximum fluxes range between 29 and 31 in magnitude, depending on the filter. As already mentioned, the brightest phase is very fast (few Myr) and peaks at very young ages ($\sim 10^{6.5}$ yr), hence it lies at redshifts very close to z_f . For ages older than 10–100 Myr, GCPs fade by several magnitudes (from 2 to 3, i.e. a factor 5 to 10 in flux). Note that the bright phase is always short, independently on the formation redshift and on the filter, lasting few Myrs before dropping by 1 mag, or at most up to 100 Myr to fade by 2 mag for low-formation redshift ($z_f = 3$) and reddest filters, but always covering a broader redshift range for higher formation redshift. We also note that the brightest fluxes are expected at the lowest redshift and in the bluest filters, given that the maximum flux is in the UV rest-frame (see previous section). At higher redshift and redder filters we expect GCPs to be fainter due to, respectively, higher distance and filters sampling an intrinsically fainter part of the SED. However, for filters redder than F150W the difference in the maximum fluxes between different z_f is always less than ~ 1 mag. Actually, as evident from Fig. 5, during the first 10 Myr since formation the observed magnitude range spanned by GCPs only slightly change with the formation

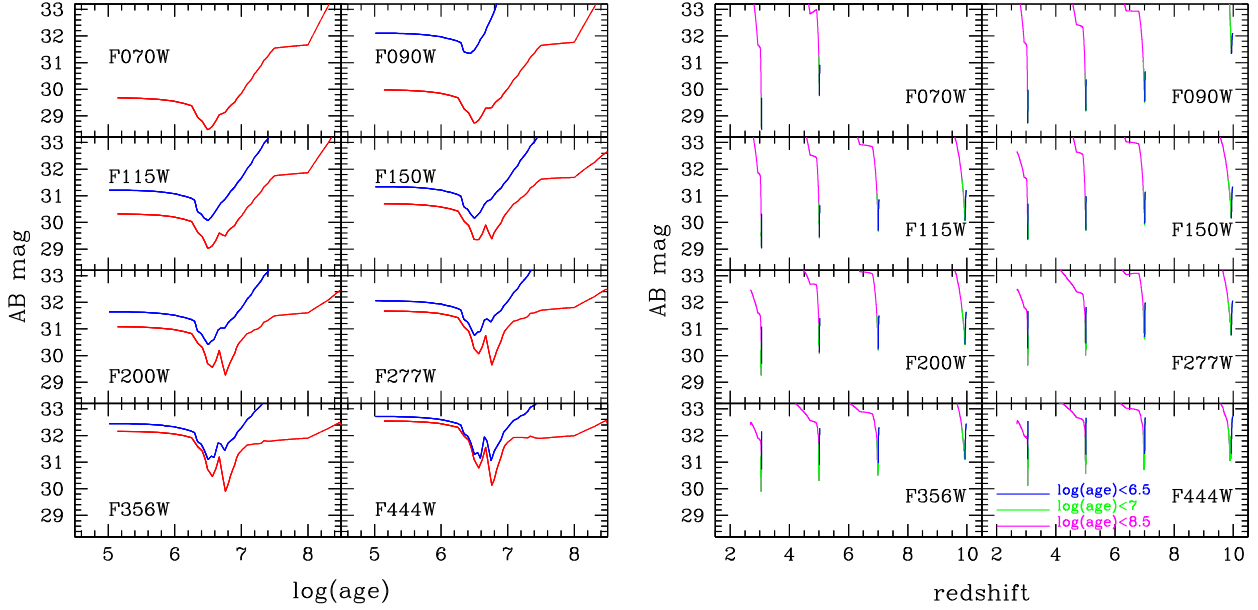


Figure 5. Observed NIRCcam magnitudes expected for a typical GCP of mass $2 \times 10^6 M_{\odot}$, adopting different formation redshifts. Different panels show different NIRCcam bands. *Left-hand panels* show magnitudes as a function of time since formation, for a GCP formed at $z_f = 3$ (in red) and $z_f = 10$ (in blue). *Right-hand panels* show magnitudes as a function of redshift for $z_f = 3, 5, 7, 10$. Different colours refer to different ages after the formation, as encoded in the bottom-right panel.

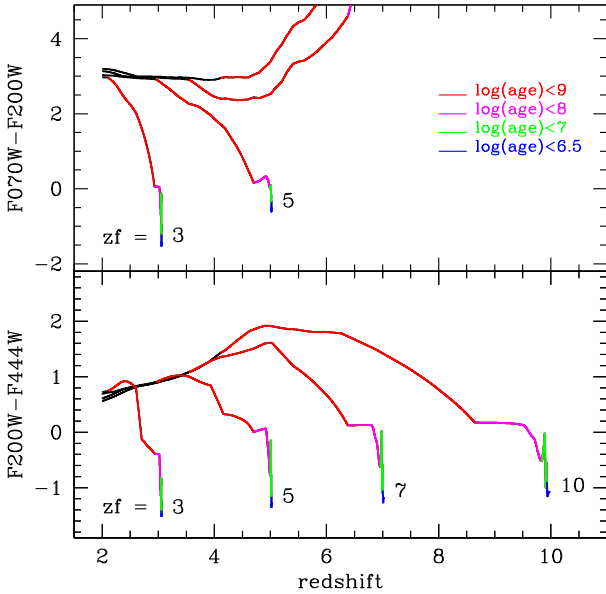


Figure 6. Observed $F070W-F200W$ and $F200W-F444W$ colour evolution of GCPs forming at redshifts $z_f = 3, 5, 7, 10$. Different colours refer to age ranges as indicated in the insert. For $z_f = 7$, and 10 GCPs would appear as $F070W$ dropouts.

redshifts. This is a consequence of the near power-law shape of the spectrum of young SSPs (see again Fig. 3), hence a negative k correction largely compensates for the increasing distance with redshift.

As a further example, Fig. 6 shows the expected $F070W-F200W$ and $F200W-F444W$ colours as a function of redshift for the same z_f values. In both cases the bluest colours are found at $z \sim z_f$, i.e. at formation and shortly thereafter. The $F070W-F200W$ colours become very red at $z > 5$, due the effect of IGM absorption in the

bluest of the two filters. Also for the colours, as for magnitudes, there are no big differences for different z_f . In particular, during the first 10 Myr the $F200W-F444W$ colours are quite insensitive to z_f , which is again due to the power-law shape of the SED. In conclusion, the shape of the GCP spectrum during the first ~ 10 Myr since formation has the interesting effect that both luminosities and colours are quite insensitive to the formation redshift, unless photons that would be detected in a given passband suffer from IGM absorption. Actually, the only way of measuring a photometric redshift for GCP candidates will be through the dropout technique, as is currently the case for very high redshift galaxies.

Finally, we illustrate the relation between magnitude and mass in the different filters and for various formation redshifts. As luminosity scales linearly with mass, their relation can be written as

$$\log(M/M_{\odot}) = 0.4 \times (F_0 - F) + \log(M_0/M_{\odot}), \quad (2)$$

where F_0 is the magnitude in a generic filter at a given reference mass M_0 . In Fig. 7 we show that this relation has a minimum in the bright phase at young age (< 10 Myr) and how it depends on age and z_f . For the bluest filter ($F70W$) there is a strong dependence on z_f due to the effect of high- z IGM on the UV flux. Instead, for the reddest filter ($F444W$) this relation depends only mildly on z_f (less than 1 mag, or less than a factor 2 in mass) and on age. In each set of panels the yellow line coincides with the left edge of the band in the lowest redshift panel and for reference it is then reproduced identically in the other three panels. This helps to appreciate how little these magnitude-mass relations depend on formation redshift. Similar, self-explanatory figures for other passbands can be found in the appendix. From these relations we can define the minimum mass for a GCP to be observed at a given magnitude in a given filter. For the representative mass $\log(M_0/M_{\odot}) = 6.3$ the corresponding magnitudes in the equation above are, for $z_f = 3$, $F_0 = 28.55, 28.75, 29.05, 29.35, 29.25, 29.65, 29.85$, and 30.15 for the various NIRCcam filters, respectively from the blue to the red. Similarly, for a given observed magnitude $F = 30$, the minimum

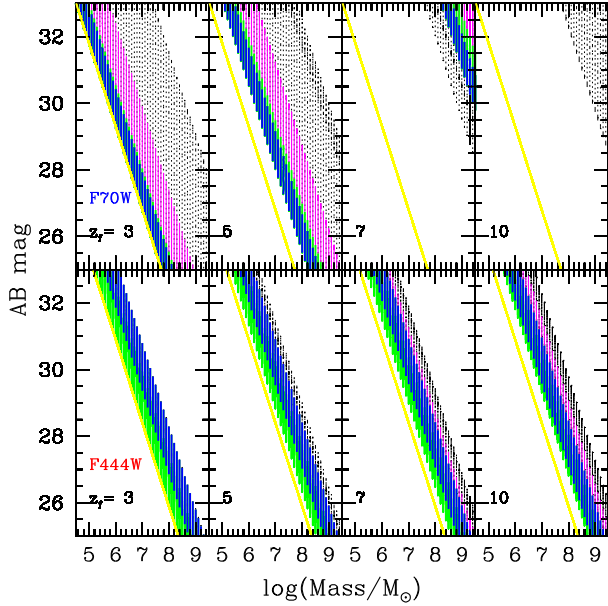


Figure 7. Observed magnitude–mass relation for a GCP, adopting different formation redshifts ($z_f = 3, 5, 7, 10$). Different colours refer to different ages after the formation, as in Fig. 5, and in grey for ages greater than 1 Gyr. Top panels for the *F70W* filter, bottom panels for *F444W*. The yellow line is at a fixed position in all panels of a given filter.

corresponding masses, for $z_f = 3$, can be derived from the previous equation and are about $\log M/M_\odot = 5.7, 5.8, 5.9, 6.0, 6.0, 6.1, 6.2$ and 6.3 for the various NIRCcam filters, from the bluest to the reddest one, respectively. These minimum masses increase with increasing z_f , but only slightly for long-wavelength passbands (see Fig. 7). Thus, reaching magnitude 30 in all bands the minimum detectable GCP mass increases with wavelength, but only by a factor ~ 4 from the bluest to the reddest NIRCcam filter.

For comparison candidate GCPs at $3 \lesssim z \lesssim 6$ in Vanzella et al. (2017a, 2019) are as bright as $-17 \lesssim M_{1500} \lesssim -15$, or $29 \lesssim m_{F105W} \lesssim 32$, and derived masses in the range $\sim 5 \times 10^5$ and $\sim 10^7 M_\odot$ and ages less than 10 Myr.

4 PREDICTED NUMBER COUNTS OF GLOBULAR CLUSTERS PRECURSORS

Forecasts for any future survey require as input the luminosity/mass function in order to determine the number of objects above a given sensitivity. Indeed, in a homogeneous and isotropic universe, the number of objects of each type (GCPs in our case) brighter than a given magnitude m_λ can be calculated from the integral:

$$N_{z_f}(< m_\lambda) = \int_{z_{\min}}^{z_f} \int_{M_{\min}(m_\lambda, z, z_f)}^{\infty} \frac{dN}{d\log M} \frac{dV}{dz} d\log M dz, \quad (3)$$

where z_f is the adopted redshift of formation for our GCPs and $\frac{dN}{d\log M}$ (function of M) is the GCPs mass distribution from equation (1), with $\log(M^*) = 6.3$, normalized to include all GCPs, i.e. $N = 1.15 \text{ Mpc}^{-3}$. We assume that our GCPs have concluded their primordial phase at $z_{\min} = 2$. This assumption does not affect dramatically our computation, since their luminous and mass loss phases are even shorter than 1.5 Gyr (the maximum time elapses between $z = 2$ and the maximum redshift of formation assumed, $z_f = 10$). The integration over $d\log M$ extends from $M_{\min}$ which depends on redshift, age, z_f and magnitude limit (m_λ) in a given band and it can

be derived from our model SSPs and take into account the aging of the GCP population. In practice, we invert the relation between magnitude and mass using Table 1. The minimum mass $M_{\min}$ in equation (1) is then derived from the equation:

$$\log M_{\min} = -0.4 (m_\lambda - m_\lambda^{6.3}(z, z_f)) + 6.3, \quad (4)$$

where $m_\lambda^{6.3}$ is the magnitude observed in a given filter at λ for a GCP of mass $\log(M/M_\odot) = 6.3$, at a given redshift (z) taking into account the evolution in time since formation (t) for any given formation redshift (z_f). This has been derived from the intrinsic evolving spectra as described in Section 3 and further attenuated by the IGM.

Finally, assuming that the bulk of GCs are formed in the range $3 < z_f < 10$, corresponding to a lookback time in the range of $t_{\text{lb}} = 11.5\text{--}13$ Gyr, then the total number densities of GCPs ‘brighter than m_λ is given by:

$$N_{\text{tot}}(< m_\lambda) = \int_{t_{\text{lb}}^{\text{inf}}}^{t_{\text{lb}}^{\text{sup}}} \int_{z_{\min}}^{z_f(t_{\text{lb}})} \int_{M_{\min}(m_\lambda, z, z_f)}^{\infty} \mathcal{F}(t_{\text{lb}}) \frac{dN}{d\log M} \frac{dV}{dz} d\log M dz dt_{\text{lb}}, \quad (5)$$

where $\mathcal{F}(t_{\text{lb}})$ is the fraction of globular clusters produced per unit time. Here we assume it constant in time (not in redshift) and therefore $\mathcal{F}(t_{\text{lb}}) = 1/(t_{\text{lb}}^{\text{sup}} - t_{\text{lb}}^{\text{inf}})$.

From equations (3) and (5) we obtain the redshift distribution per unit redshift ($\frac{dN}{dz}(< m_\lambda; z, z + dz)$) by integrating over the specific redshift range ($z, z + dz$). We stress here that at all observed magnitudes, all masses, formation redshifts and times since formation can contribute to the number counts.

The cumulative number densities per arcmin² predicted by our model are shown in Fig. 8 in the various NIRCcam filters, from the bluest (*F070W*) to the reddest (*F444W*) and for our adopted mass function with $\log(M^*/M_\odot) = 6.3$, assuming a birth rate constant in time in the redshift range $z_f = 3\text{--}10$. We also show the counts adopting a single z_f , i.e. all GCPs form at the same z_f . The corresponding tables with cumulative and differential counts can be found at: <https://sites.google.com/inaf.it/pozzetti-gcps/home>. We note that the predicted counts are fairly insensitive to the formation redshift, with the exception of the three bluest passbands, because of the dropout effect. The similarity of the predicted counts per unit area (not per unit volume) for different formation redshifts is due to a combination of effects. First, as already discussed, the magnitudes/fluxes during the brightest phase are fairly insensitive to the formation redshift, being at most 1 mag fainter for $z_f = 10$ compared to $z_f = 3$. Furthermore, even if the duration of the bright phase is similar for different z_f in term of time, it is always broader in term of redshift range for high z_f . This will end in a larger volume per unit area for high z_f GCPs, which compensates for the slightly fainter fluxes, hence determining similar effective counts per unit area.

5 PERSPECTIVES AT DETECTING GCS IN FORMATION, WITH CAVEATS

We show in Figs 9 and 10, in particular, the same for the *F200W* and *F444W* passbands. Notice that the upper scale gives the minimum mass for a GCP (which form at $z_f = 3$) being as bright as its peak luminosity as indicated by the lower scale. For example, only GCPs more massive than $\sim 10^7 M_\odot$ could be brighter than mag = 28 and only those more massive than $10^6 M_\odot$ could shine brighter than mag = 30. At these masses, as shown in Fig. 2, we expect, for instance for a GCP of $\sim 2 \times 10^6 M_\odot$ to be as bright as $M_{1500} \sim -17$

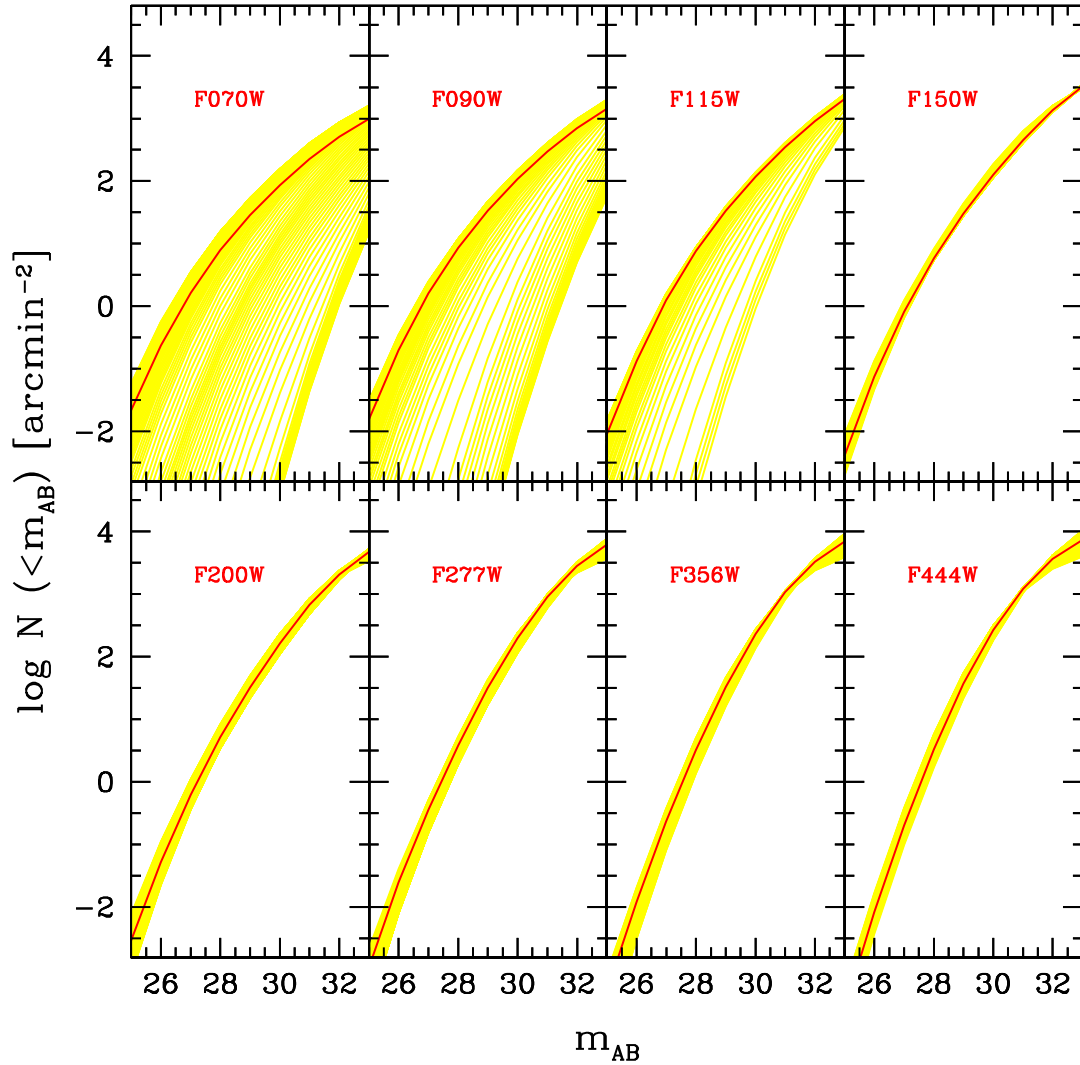


Figure 8. Expected GCPs cumulative counts in all NIRCcam bands as from equation (5) for $\log(M^*) = 6.3$. Each panel refers to a different NIRCcam band. In yellow the number counts for various formation redshifts in the range $z_f = 3-10$, assuming that all GCPs form at a given z_f . The red lines show the total counts adopting a continuous formation redshifts (constant in time) within the same redshift interval.

at its peak. The upper and lower borders in the counts, with single z_f , correspond to all GCPs forming at $z = 3$ or at $z = 10$, respectively, with the exception of the very faint magnitudes case, in particular for *F444W* passband, where there is an inversion with those formed at the highest redshift starting to dominate the counts. This effect is mainly due to the fact that GCPs formed at $z_f = 3$ are brighter (by 1 mag) and therefore they reach the maximum density at the peak of the Gaussian Mass Function (at $\log M/M_\odot = 6.3$ and magnitude ~ 30.5) and thereafter, at magnitudes brighter than those of GCPs formed at $z_f = 10$, start to diminish in density relative to them. The figures also include the expected number counts for $\log(M^*/M_\odot) = 5.3$, i.e. assuming GCPs formed with the same mass function of present-day GCs, as if they had suffered no mass loss at all. This is clearly a strict lower limit to the expected number counts. To the extent to which an M^* 10 times higher than that can be regarded as an upper limit, then we expect that the real counts will fall somewhere in between the dotted line and the yellow band. Note that the vertical scale in these two figures gives the number counts per NIRCcam FoV, hence for $\log M^*/M_\odot = 6.3$ one expects NIRCcam to detect of the order of ~ 1000 GCPs down to $\text{mag} = 30$ in either the *F200W* or the

F444W passbands. In the most conservative case, this number falls down to ~ 10 detections per NIRCcam pointing. So, in conclusion, one expects from ~ 10 to ~ 1000 GCP detections, depending on the actual value of the ‘mass budget factor’ (*mbf*) in the range 1–10, that future NIRCcam observations will actually allow us to estimate. The first opportunity to check these numbers will be offered by the *JWST* Early Release Science (ERS) observations that will include the coverage of $\sim 100 \text{ arcmin}^2$ with NIRCcam⁴ in the five reddest band down to $\text{mag} \sim 29$ in the *F200W* band (28.6 in the *F444W* band). We estimate that ~ 3700 (1400) candidate GCPs should be detected in the *F200W* (*F444W*) band during ERS for $mbf = 10$, which are drastically reduced to less than 20 for $mbf = 1$. Then ERS observations will be followed by the NIRCcam guaranteed time observations (GTO)⁵ planned to reach $\text{mag} = 29.8$ (at 10σ for point sources) over an area of 46 arcmin^2 . At this limiting

⁴<https://jwst.stsci.edu/observing-programs/program-information?id=1345>

⁵<https://jwst-docs.stsci.edu/display/JSP/JWST+GTO+Observation+Specifications>

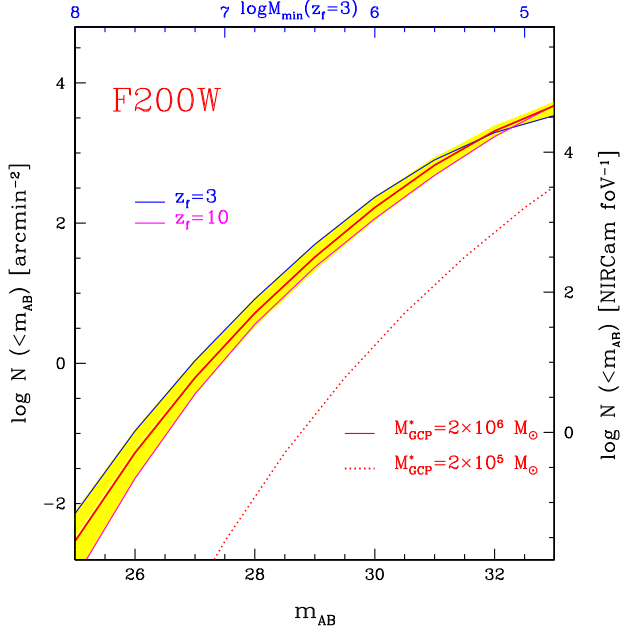


Figure 9. Expected GCPs cumulative counts per arcmin² (also reported per NIRCam FoV on right y-axis) for the *F200W* band, replicated from Fig. 8. The dotted red line shows the cumulative counts assuming that GCPs formed with the same mass function of present-day GC (i.e. $\log M^*/M_\odot = 5.3$). In yellow the number densities for various formation redshift, of which in blue for $z_f = 3$ and in magenta for $z_f = 10$.

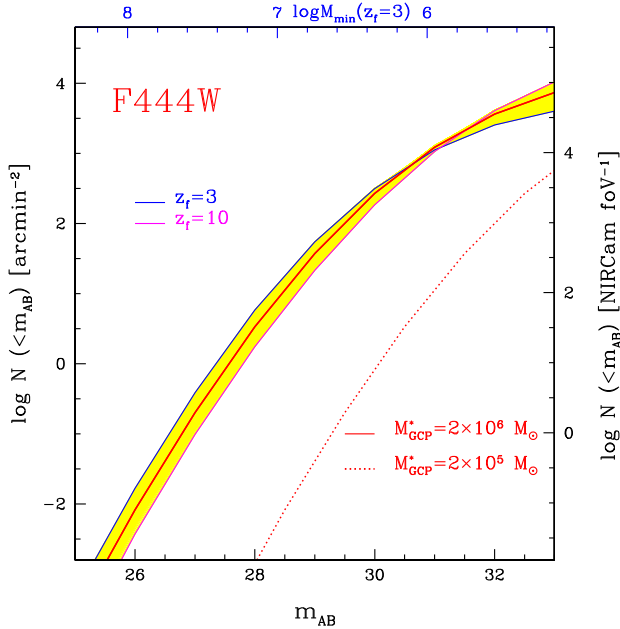


Figure 10. The same as in Fig. 9, for the *F444W* band.

magnitude, objects more massive than $\sim 10^6 M_\odot$ should be detected while near maximum light, providing up to ~ 5500 (8300) objects in the *F200W* (*F444W*) band for $mbf = 10$, reduced to just ~ 55 (20) detections for $mbf = 1$. Moreover, the GTO team plans also to reach $\text{mag} = 28.8$ over an area of 190 arcmin², which from Fig. 8 corresponds to detecting ~ 20 or ~ 4000 GCPs, for a mass budget factor 1 or 10, respectively. Combining together the ERS data and

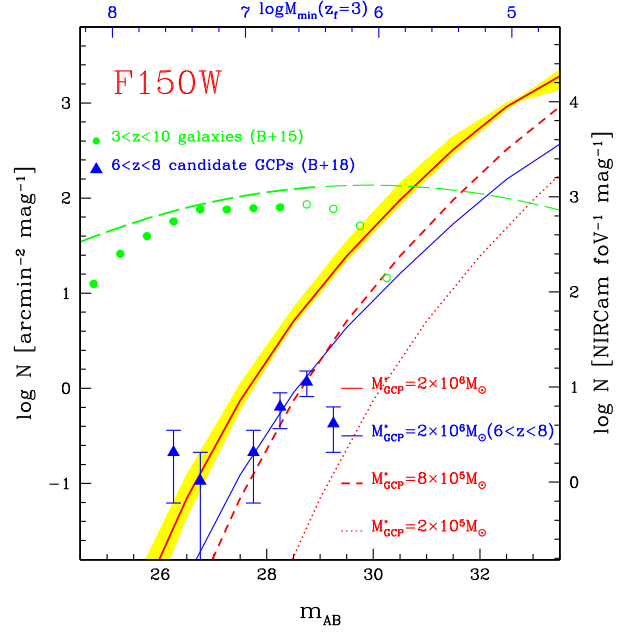


Figure 11. The differential number counts for the *F150W* band. In red the total GCPs counts adopting a continuous formation redshifts ($z_f = 3, 10$). The green dots show the counts for the $3 < z < 10$ galaxies from Bouwens et al. (2015) in the *HST F160W* passband. The green long-dashed line shows the total number counts for galaxies at all redshifts (Madau & Pozzetti 2000). In addition, we show the counts of candidate GCPs revealed in the HFF by Bouwens et al. (2018) in the redshift range $6 < z < 8$ (blue triangles with error bars). For comparison we show also our model (with $M^* = 2 \times 10^6 M_\odot$), limited to the same redshift range (blue solid line), and the models adopting $M^* = 8 \times 10^5 M_\odot$ and $2 \times 10^5 M_\odot$ (dashed and dotted red line, respectively).

the GTO deep and broad observations, NIRCam should detect from ~ 100 to ~ 14000 GCPs, respectively for $mbf = 1$ and 10. But see caveats in the following section.

Moreover, Fig. 11 shows the *differential* GCP number counts for the *F150W* passband. We compare our GCP predictions to the current number counts in the *HST F160W* passband for galaxies in the redshift range 3–10 (Bouwens et al. 2015) along with the number counts for galaxies at all redshifts, from a compilation of several extragalactic surveys, fitted and extrapolated to unobserved fainter magnitudes (Madau & Pozzetti 2000). Apparently, for a mass budget factor of 10, GCPs start dominating the high redshift counts just beyond the Bouwens et al. (2015) limit, and beyond $m_{AB} \simeq 30.5$, corresponding to a peak luminosity of $M_{1500} \sim -15.5$, could dominate over the total galaxy counts. This transition from galaxy-dominated to GCP-dominated number counts should produce a sharp inflection in the overall luminosity function from whose location it should be possible to determine the mass budget factor.

Finally, the comparison between our number counts predictions and candidates GCPs from Bouwens et al. (2018), revealed as compact objects (with size < 40 pc, thanks to lensing effect) by the Hubble Frontier Fields (HFF) programme, seems to suggest an encouraging agreement, but better be cautious about its interpretation. Formally, the observed counts of these compact objects at $z = 6-8$ suggest that M^* could be at least $8 \times 10^5 M_\odot$, i.e. 4 times more massive than local GCs, or even 10 times ($M^* = 2 \times 10^6 M_\odot$) if GCs were to form at a constant rate in the range $z = 3-10$, as assumed in our model. In such case, Bouwens et al. (2018) would have seen only the fraction (20 per cent) formed in that redshift

range $6 < z < 8$, most likely being all metal poor. However, not all such objects may eventually end up as GCs ~ 13 Gyr later, especially those with masses well above $\sim 10^7 M_\odot$, so we cannot claim to have already measured the mass budget factor. Clearly, reaching down to $\text{mag} = 30\text{--}31$ will be critical in this respect.

All number counts presented in this section can be found on the same site along with the full Table 1 (<https://sites.google.com/inaf.it/pozzetti-gcps/home>).

6 PERSPECTIVES AT DETECTING GCS IN FORMATION, WITH CAVEATS

The number counts presented in the previous section follow from a series of assumptions that may or may not be verified in nature, either in one direction or the other. One assumption is that GCPs can be described as SSPs. This is obviously at variance with the multiple generation phenomenon, so it appears that real GCs have formed in a more or less extended series of individual bursts, interleaved by inactive periods of unknown duration (see Calura et al. in preparation, for a model with multiple populations). If all star formation activity was confined within less than ~ 10 Myr, then our single SSP assumption should not be grossly in error. For longer separations between bursts, the above number counts should be considered as referring to the brightest event, hence the number counts should be reduced by a factor of ~ 3 if such brightest event produced just 1/3 of the final mass of the GCP. In any event, at least in some scenario, the *supernova avoidance* requirement (Renzini et al. 2015) dictates that all bursts should take place within few Myr, or being separated by more than ~ 30 Myr having allowed supernova ejecta do leave the system.

Our modelling may suffer by another limitation in that it considers GCPs consisting only of single, non-rotating stars. Direct evidence indicates that the majority of young massive (O-type) stars are members of binary systems (Sana et al. 2012). Thus, binary members interactions can affect the resulting SED compared to the SSP approximations, in particular during the first ~ 10 Myr since formation, though synthetic stellar populations including massive binaries indicate that the effects on the SED are relatively modest (Eldridge et al. 2017) as is the effect of rotation (Leitherer et al. 2014).

Another assumption is that GC stars formed following the IMF proposed by Chabrier (2003), which may or may not be the case. Here we can only say that massive stars must have formed, given the presence of many pulsars in today's GCs (e.g. Manchester et al. 1991), hence the IMF cannot have been too steep. It cannot have been too flat either, otherwise clusters would have dissolved in response to stellar mass loss (Chernoff & Weinberg 1990). Our exercise assumes a slope near Salpeter (-2.35) between ~ 1 and $\sim 100 M_\odot$. If the IMF was (slightly) flatter (steeper) than this, then the number counts presented in the previous section would have been underestimated (overestimated). The effect of varying the IMF is explored in Jeřábková et al. (2017).

Then we have assumed that GCPs at high redshift suffered negligible reddening in the rest-frame UV. As already mentioned, this is likely the case for metal poor GCPs, say those with $[\text{Fe}/\text{H}] \lesssim -1.0$, which account for roughly half of the GCs in the local Universe. However, the metal-rich half of the local GC populations must have formed when the build-up of the today hosting galaxy was already quite advanced, given the mass-metallicity relation of high-redshift galaxies (Erb et al. 2006; Kashino et al. 2017). Within the MW, most metal rich GCs belong to the Galactic bulge and must have formed along with bulge itself, over ~ 10 Gyr ago (Ortolani et al. 1995; Renzini et al. 2018), hence in a chemically enriched and

dusty environment, such that GCPs must have suffered substantial UV extinction making unlikely they could be detected. If so, all predictions made in the previous section should be cut by a factor ~ 2 .

The physical nature itself of GCPs remains basically unknown. They might have been just compact, somewhat more massive GCs or they may have been the nuclei of dwarf galaxies, as was suggested by Searle & Zinn (1978), most of which later dissolved. Indeed, the Fornax and Sagittarius dwarfs contain an unusually large number of GCs for their mass. In this respect, one question is whether one should consider the whole mass of the dwarf as the mass of the GCP, or just that of the compact object hosted by the dwarf. Following the argument of Elmegreen (2017), given the likely $\sim \text{kpc}$ size of dwarfs the typical time-scale of star formation was of the order of $\sim 10^8$ yr rather 10^6 yr as for GCs, and Zick et al. (2018) have shown that when forming individual GCs in a Fornax-like precursor would have overshadowed the underlying galaxy. Thus, dwarfs hosting forming GCs could not be adequately described by our SSP approximation for GCPs, hence dwarfs parent to GCs are unlikely to be included in the number counts presented in the previous section, as they would be substantially fainter than massive GCPs younger than ~ 10 Myr. In this respect, once a suitable number of candidate GCPs will be found, then stacking them could actually reveal the presence of host dwarfs, and a concrete example has been documented by Vanzella et al. (2019). Still, even if not qualifying as GCPs from the observational point of view, dwarfs might have provided nucleary processed material for the formation of GC multiple stellar generations. Hence, detecting and characterize the immediate environment of GCPs should provide critical insight on the process of GC formation. Distinguishing GCPs from their dwarf hosts (or in general from high redshift galaxies) will not be trivial. GC-size objects will appear as point-like in NIRCам images, given their ~ 200 pc resolution at these redshifts (Renzini 2017), and hosting dwarfs of few 100 pc diameter will be only marginally resolved, unless lensed as in the object of Vanzella et al. (2019).

For a better chance to distinguish true GCPs from their host, or from other high-redshift dwarfs or close multiple GCPs or multiple knots of star formation, we will have to take advantage of the higher spatial resolution provided by lensing (such as in Vanzella et al. 2019) or of the next generation of extremely large telescopes (ELT) assisted by advanced adaptive optics. For example, the 39m European ELT will provide a ~ 6 times better spatial resolution compared to *JWST*, corresponding to ~ 30 pc at these redshifts.

On the side of the mass budget factor, the mere stellar mass loss via stellar winds would account for a factor ~ 1.7 in GC mass reduction from formation to the present. On top of this, star evaporation and stripping via disc shocking and tidal interactions would further reduce the GC masses which according to the *N*-body simulations of Webb & Leigh (2015) could be as high as a factor of ~ 10 with an average of a factor ~ 4.5 . Thus, a mass budget factor of the order of 10 does not appear to be unconceivable. Again, direct counts will be the only way to estimate this critical factor.

With each of its pointings, NIRCам will sample a comoving volume between $z = 3$ and 10 of over $160\,000 \text{ Mpc}^3$. Brightest cluster galaxies (BCG) as massive as M87, likely with a similar share of $\sim 10\,000$ GCs, come with a space density of $\sim 10^{-5} \text{ Mpc}^{-3}$ (Bernardi et al. 2013), hence there is a fair chance that each NIRCам pointing will include one BCG precursor along with the precursor of the galaxy cluster hosting it. If Nature has been benign enough to make bright GCPs, we will have the opportunity to learn much about the star formation and its clustering preceding the appearance of massive galaxies, with clustered GCPs working as signposts of incipient massive galaxy formation.

7 SUMMARY AND CONCLUSIONS

Having assumed and justified that the spectrum of young GCPs, or at least of their ~ 50 per cent metal poor fraction, could reasonably be described by simple stellar population models, we have shown that:

(i) Given the power-law shape of the spectrum of young GCPs, both colours and fluxes/magnitudes in NIRCcam passbands are fairly insensitive to the actual formation redshift. Only in the bluest passbands (namely, $F70W$, $F90W$, and $F115W$) colours, luminosities and counts are dramatically affected by hydrogen absorption in the intervening IGM along the line of sight.

(ii) As a consequence, we show that for $F150W$ and redder filters the expected number counts of GCPs is fairly insensitive to the actual distribution of formation redshifts, in the range $3 < z < 10$.

(iii) Number counts depend instead critically on the actual mass distribution of GCPs, i.e. on how much more massive they were compared to their GC progeny. The ratio of the initial to present GC mass, commonly referred to as the *mass budget factor*, is the primary controller of the GCP number counts. For such factor being 1 (GCPs as massive as today GCs, i.e. no mass loss) NIRCcam should detect of the order of 10 GCPs per pointing down to $\text{mag} \simeq 30$, a number that jumps to ~ 1000 if such factor is instead 10. The recently observed number density of candidate GCPs at $z = 6-8$, revealed by the HFF programme, suggests that the *mass budget factor* could be at least 4, i.e. GCPs being at least four times more massive than their local descendants, if all were to end up as GCs.

(iv) For a mass budget factor of 10, GCPs should start to dominate the number counts of high- z galaxies just beyond the limits currently achieved so far, i.e. $m_{F160W} \simeq 29$ (Bouwens et al. 2015).

(v) Thus, actual number counts will set stringent constraints on the mass budget factor, providing crucial information on the formation and early evolutionary stages of GCs and helping to decipher their multiple generation phenomenon.

(vi) Like some GCs today, GCPs may have been hosted by dwarf galaxies which could be detected either directly (for a recent tantalizing finding see Vanzella et al. 2019) or by stacking many detected GCPs, thus characterizing the environment having nursed GCs at their formation epoch.

(vii) As metal poor GCs formed well before the bulk of the galaxy hosting them today, clustering of GCPs on ~ 100 kpc scale would mark the signpost of incipient massive galaxy formation. In this respect, reaching ~ 1 magnitude deeper, i.e. down to $\text{mag} = 31$ as in a future *JWST Ultra Deep Field*, should boost number counts by a factor ~ 4 , thus greatly facilitating the GCP clustering analysis.

(viii) The only way of measuring GCP photometric redshifts appears to be via the drop-out technique, as bluest photons are absorbed by intervening hydrogen. Using only NIRCcam, drop-outs in the filters $F070W$, $F090W$, and $F115W$ will correspond to GCPs at $z > 5$, > 6.5 and > 8.5 , respectively. Photometric redshifts for objects at $z < 5$ will have to rely on bluer data from either *HST* or the ground.

We also qualitatively discuss a series of caveats concerning all these predictions, including extinction (likely affecting the metal rich fraction of GCPs), the possibility of the IMF of GCPs being different from the assumed Chabrier (2003) IMF, the effect of the GCP mass being built up with multiple episodes of star formation, and finally the possible role of massive binaries. We do not explore the possibility of young GCPs hosting a supermassive star of $\sim 10^4 - 10^5 M_\odot$, that could significantly contribute to their overall luminosity (Denissenkov & Hartwick 2014).

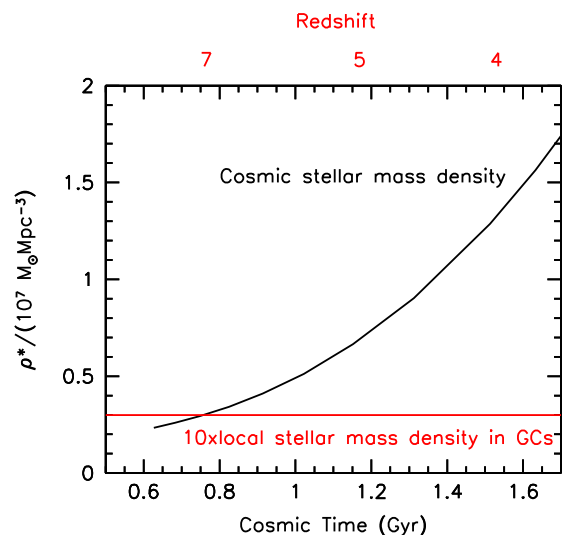


Figure 12. The early evolution of the comoving cosmic star mass density as from Madau & Dickinson (2014) (black line), adjusted to our adopted Chabrier IMF. The red horizontal line is drawn at the level of $3 \times 10^6 M_\odot \text{Mpc}^{-3}$, corresponding to 10 times the local stellar mass density in globular clusters.

Finally, we briefly comment on the possible role of GCPs on cosmic reionization, an issue already mentioned in Section 1. Fig. 12 shows the evolution of the comoving stellar mass density as a function of cosmic time as from Madau & Dickinson (2014), adjusted to our choice of the Chabrier IMF. The red horizontal line is drawn at 10 times the local mass density in GCs, which is $\sim 3 \times 10^5 M_\odot \text{Mpc}^{-3}$. From the figure, we see that GCs, together with their possible dwarf galaxy hosts, may have dominated the cosmic mass density and, therefore, the reionization if the bulk of them formed at $z \gtrsim 7$, whereas their contribution would have been marginal if formed predominantly at $z \lesssim 4$. This is consistent with the early estimate by Ricotti (2002), thus leaving open the connection with reionization, especially if the escape fraction of ionizing photons from GCPs were close to unity as argued by Ricotti (2002) and Katz & Ricotti (2014).

ACKNOWLEDGEMENTS

We would like to thank Eros Vanzella, Francesco Calura, Enrico Vesperini, and Gianni Zamorani for stimulating discussions and useful input on the evolution of GCs. We also thank Rychard Bouwens for providing his data in electronic form and for useful input. We thank the referee Massimo Ricotti for his constructive comments and suggestions. AR and LP acknowledge support from an INAF/PRIN-SKA 2017 grant and grants PRIN MIUR 2015, ASI n.1/023/12/0 and ASI n.2018-23-HH.0.

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SUPPORTING INFORMATION

Supplementary data are available at *MNRAS* online.

Table 1. Example of observer-frame magnitudes in *JWST* filters for a GCPs of mass $\log(M/M_{\odot}) = 6.3$. The first column gives the redshift at which the object is observed, having already aged about 3 Myr since its formation, hence formed at a slightly higher redshift.

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APPENDIX A: EVOLUTION WITH TIME AND MAG-MASS RELATIONS IN ALL OTHER FILTERS

For completeness, Figs A1, A2, and A3 are analogue to Fig. 7, but for the other NIRC2 filters.

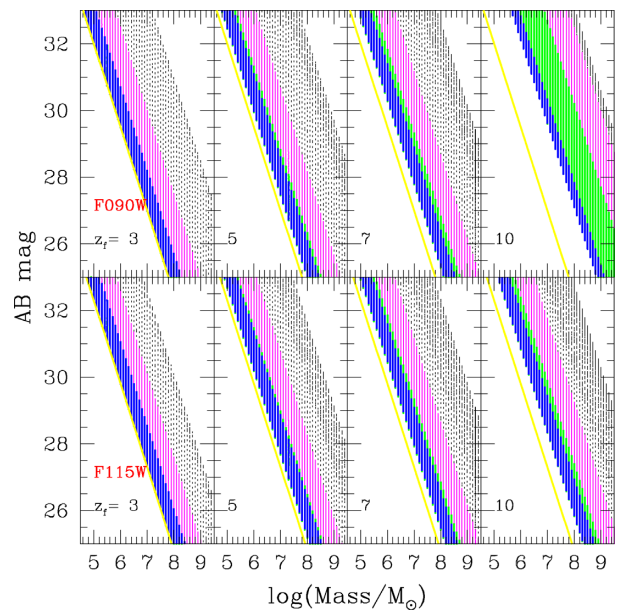


Figure A1. The same as Fig. 7 for the F090W and F115W bands, as indicated.

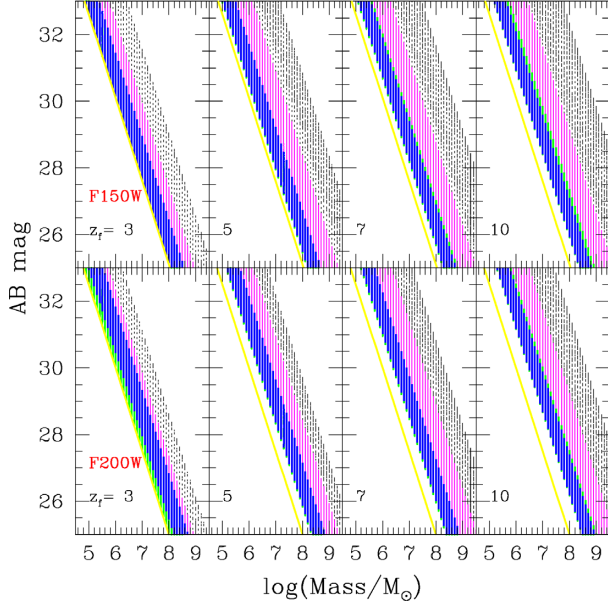


Figure A2. The same as Fig. 7 for the F150W and the F200W bands, as indicated.

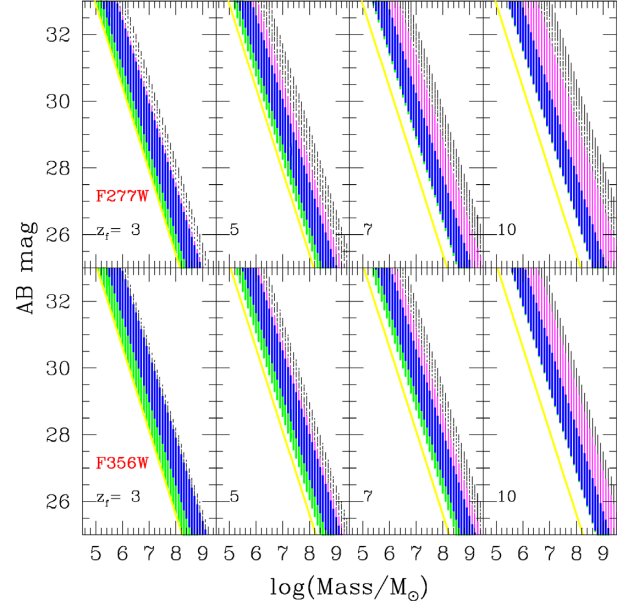


Figure A3. The same as Fig. 7 for the F277W and the F356W bands, as indicated.

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