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J/A+A/634/A34 Complete line list and solar values (Baratella+, 2020)

The Gaia-ESO Survey: a new approach to chemically characterising young open clusters.

Baratella M., D'Orazi V., Carraro G., Desidera S., Randich S., Magrini L., Adibekyan V., Smiljanic R., Spina L., Tsantaki M., Tautvaisiene G., Sousa S.G., Jofre P., Jimenez-Esteban F.M., Delgado-Mena E., Martell S., Van der Swaelmen M., Roccatagliata V., Gilmore G., Alfaro E.J., Bayo A., Bensby T., Bragaglia A., Franciosini E., Gonneau A., Heiter U., Hourihane A., Jeffries R.D., Koposov S.E., Morbidelli L., Prisinzano L., Sacco G., Sbordone L., Worley C., Zaggia S., Lewis J.
<Astron. Astrophys. 634, A34 (2020)>
[=2020A&A...634A..34B](#) (SIMBAD/NED BibCode)

ADC_Keywords: Clusters, open ; Stars, G-type ; Atomic physics

Keywords: stars: abundances - stars: fundamental parameters -
stars: solar-type - open clusters and associations: individual:
IC 2391, IC 2602, IC 4665, NGC 2264, NGC 2516, NGC 2547

Abstract:

Open clusters are recognised as excellent tracers of the Galactic thin disc properties. At variance with intermediate-age and old open clusters, for which a conspicuous number of studies is now available, until a few years ago clusters younger than 150Myr had been mostly overlooked in terms of their chemical composition (with few exceptions). On the other hand, previous investigations seem to indicate an anomalous behaviour of young clusters, which include (but is not limited to) slightly sub-solar iron (Fe) abundances and extreme, unexpected barium (Ba) enhancements. In a series of papers, we plan to expand our comprehension of this topic and investigate whether these chemical peculiarities are instead related to abundance analysis techniques. We present here a new determination of the atmospheric parameters for 23 dwarf stars observed by the Gaia-ESO survey in five young open clusters (age less than 150Myr) and one star forming region (NGC 2264). We exploit a new method based on titanium (Ti) lines to derive spectroscopic surface gravity and, most important, microturbulence parameter. A combination of Ti and Fe lines have been used to obtain effective temperatures. We also infer abundances of FeI, FeII, TiI, TiII, NaI, MgI, AlI, SiI, CaI, CrI and NiI. Our findings are in fair agreement with Gaia-ESO iDR5 results for effective temperatures and surface gravities, but suggest that for very young stars the microturbulence parameter is over-estimated when Fe lines are employed. This impacts the derived chemical composition, causing the metal content of very young clusters to be under-estimated. Our clusters display a metallicity [Fe/H] between +0.04 and +0.12; thus they are not more metal-poor than the Sun. Although based on a relatively small sample size, our explorative study suggests that we may not need to call for ad hoc explanations to reconcile the chemical composition of young open clusters with Galactic chemical evolution models.

Description:

Line list used in this study (table2.dat), complete of the atomic data, references of the log(gf) values adopted and with equivalent width and abundances measured in the Sun.

File Summary:

FileName	Lrecl	Records	Explanations
ReadMe	80	.	This file
table2.dat	75	229	Complete line list

Byte-by-byte Description of file: [table2.dat](#)

Bytes	Format	Units	Label	Explanations
1-	2	A2	---	Element
4-	5	A2	---	Ion
12-	19	F8.3	0.1nm	Wavelength
23-	27	F5.3	eV	Excitation potential
31-	36	F6.3	[-]	loggf
40-	58	A19	---	Adopted log(gf) values
62-	67	F6.2	r_loggf	Reference code for log(gf) values (2)
71-	75	F5.3	0.1pm	Equivalent width measured in the Sun
			---	log(X) sun
				Solar abundances (3)

Note (1): Ionisation stage as follows:

I = neutral
II = singly ionised

Note (2): Reference codes for the log(gf) values correspond to the following references:

GESMCHF = Froese Fischer & Tachiev (2012, Multiconfiguration Hartree-Fock and Multiconfiguration Dirac-Hartree-Fock Collection, Version 2, <http://physics.nist.gov/mchf/>)
[1990JQSRT...43..207C](#) = Chang & Tang ([1990JQSRT...43..207C](#))
[1993JPhB...26.4409B](#) = Butler et al. ([1993JPhB...26.4409B](#))
[1995JPhB...28.3485M](#) = Mendoza et al. ([1995JPhB...28.3485M](#))
GARZ = Garz ([1973A&A...26..471G](#))
BL = O'Brian & Lawler (1991, Phys. rev. A, 44, 7134)
K07 = Kurucz (2007, on-line database of observed and predicted atomic transitions, <http://kurucz.harvard.edu/atoms>)
[1993Phys...48..297N](#) = Nahar (1993, Phys. Src. 48, 297)
SR = Smith & Raggett (1981, J. Phys. B: At. Mol. Phys., 14, 4015)
DIKH = Drozdowski et al. (1997, Zeitschrift fur Physik D Atoms Molecules Clusters, 41, 125)
LGWSC = Lawler et al. ([2013ApJS...205...11L](#))
[2013ApJS...205...11L](#) = Lawler et al. ([2013ApJS...205...11L](#))
MFWi = Martin et al. (1988, Cat. VI/72)
[2013ApJS...208...27W](#) = Wood et al. ([2013ApJS...208...27W](#))
SLS = Sobek et al. (007, Cat. J/ApJ/667/1267)
GESHRL14 = Den Hartog et al. (2014, [2014arXiv1409.8142D](#))
BKK = Bard et al. ([1991A&A...248..315B](#))
GESB79b = Blackwell et al. ([1979MNRAS.186..657B](#))
BWL = O'Brian et al. (1991, Journal of the Optical Society of America B Optical Physics, 8, 1185)
GESB82c = Blackwell et al. ([1982MNRAS.199...21B](#))
BIPS = Blacwell, Ibbetson et al. ([1979MNRAS.186..633B](#))
MRW = May et al. ([1974A&AS...18..405M](#))
[2014MNRAS.441.3127R](#) = Ruffoni et al. ([2014MNRAS.441.3127R](#))
GESB82d = Blackwell, Petford & Simmons ([1982MNRAS.201..595B](#))
GESB86 = Blackwell et al. ([1986MNRAS.220..289B](#))
RU = Raassen & Uylings (1998, Cat. J/A+A/340/300)
2009AA...497..611M = Melendez & Barbuy ([2009A&A...497..611M](#))
[2003ApJ...584L.107J](#) = Johansson et al. ([2003ApJ...584L.107J](#))
[2014ApJS...211...20W](#) = Wood et al. ([2014ApJS...211...20W](#))
K08 = Kurucz (2008, on-line database of observed and predicted atomic transitions, <http://kurucz.harvard.edu/atoms>))
LWST = Lennard et al. ([1975ApJ...197..517L](#))

If the reference code contains +, the adopted log(gf) values are the average from more sources, while if the code contains ", it means that the gf-value is taken from the first source but then re-normalised to an absolute scale using accurate lifetime measurements from the second source.

Note (3): Solar abundances computed with the following atmospheric parameters:
Teff/logg/xi=5790/4.47/1.00

Acknowledgements:

Martina Baratella, [martina.baratella.1\(at\)phd.unipd.it](mailto:martina.baratella.1(at)phd.unipd.it)

(End) Martina Baratella [UniPD], Patricia Vannier [CDS] 10-Jan-2020

The document above follows the rules of the [Standard Description for Astronomical Catalogues](#); from this documentation it is possible to generate *f77* program to load files [into arrays](#) or [line by line](#)

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