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Authors	Tan, Jonathan C., Kong, Shuo, FONTANI, FRANCESCO, Caselli, Paola, Butler, Michael J.
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AN ORDERED BIPOLAR OUTFLOW FROM A MASSIVE EARLY-STAGE CORE

JONATHAN C. TAN¹, SHUO KONG², YICHEN ZHANG³, FRANCESCO FONTANI⁴, PAOLA CASELLI⁵, AND MICHAEL J. BUTLER⁶¹Departments of Astronomy & Physics, University of Florida, Gainesville, FL 32611, USA; jctan.astro@gmail.com²Department of Astronomy, University of Florida, Gainesville, FL 32611, USA³Departamento de Astronomía, Universidad de Chile, Casilla 36-D, Santiago, Chile⁴INAF—Osservatorio Astrofisico di Arcetri, I-50125, Florence, Italy⁵Max Planck Institute for Extraterrestrial Physics (MPE), Giessenbachstr. 1, D-85748 Garching, Germany⁶Max Planck Institute for Astronomy, Königstuhl 17, D-69117 Heidelberg, Germany

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ABSTRACT

We present ALMA follow-up observations of two massive, early-stage core candidates, C1-N and C1-S, in IRDC G028.37+00.07, that were previously identified by their $\text{N}_2\text{D}^+(3-2)$ emission, and show high levels of deuteration of this species. The cores are also dark at far-infrared wavelengths up to $\sim 100\ \mu\text{m}$. We detect $^{12}\text{CO}(2-1)$ from a narrow, highly collimated bipolar outflow that is being launched from near the center of the C1-S core, which is also the location of the peak 1.3 mm dust continuum emission. This protostar, C1-Sa, has associated dense gas traced by $\text{C}^{18}\text{O}(2-1)$ and $\text{DCN}(3-2)$, from which we estimate that it has a radial velocity that is near the center of the range exhibited by the C1-S massive core. A second outflow-driving source is also detected within the projected boundary of C1-S, but it appears to be at a different radial velocity. After considering the properties of the outflows, we conclude that C1-Sa is a promising candidate for an early-stage massive protostar and as such it shows that these early phases of massive star formation can involve highly ordered outflow, and thus accretion, processes, similar to models developed to explain low-mass protostars.

Key words: ISM: clouds – ISM: jets and outflows – stars: formation

1. INTRODUCTION

Understanding how massive stars form is an important goal, since their radiative, mechanical, and chemical feedback play leading roles in regulating the interstellar medium, star formation activity, and the overall evolution of galaxies. Core accretion (e.g., McKee & Tan 2003, hereafter MT03) is one class of models for massive star formation, which involves the initial conditions of a high-mass, self-gravitating starless core, followed by relatively ordered collapse to a central disk and protostar (see Tan et al. 2014, p. 149 for a review). These models are scaled-up from those developed for low-mass star formation (e.g., Shu et al. 1987), but in the case of the MT03 Turbulent Core Model, they involve non-thermal forms of pressure support, i.e., turbulence and magnetic fields, for the initial core to be in approximate pressure and virial equilibrium.

Alternatively, competitive accretion models (e.g., Bonnell et al. 2001; Wang et al. 2010) involve a massive star gaining most of its mass by competitive, chaotic Bondi–Hoyle accretion in the center of a crowded protocluster of mostly low-mass stars. In these models, the initial conditions of massive star formation, i.e., the gas immediately surrounding the protostar that is destined to become a high-mass star, do not involve massive starless, self-gravitating cores, but rather low-mass cores, with most of the mass reservoir accumulating later from the protocluster clump.

To try and distinguish between these theories we have developed a method for searching for massive starless and early-stage core candidates. We first identify target regions in infrared dark clouds (IRDCs) using mid-infrared, i.e., $8\ \mu\text{m}$ *Spitzer*-IRAC, extinction (MIREX) mapping (Butler & Tan 2009, 2012). We select regions that are peaks in the resulting mass surface density, Σ , map. We further check that these regions are dark in $24\ \mu\text{m}$ (*Spitzer*-MIPS) and $70\ \mu\text{m}$ (*Herschel*-PACS) images. We then search for $\text{N}_2\text{D}^+(3-2)$

emission with ALMA, since the abundance of this species, i.e., the deuteration fraction $D_{\text{frac}}^{\text{N}_2\text{H}^+} \equiv [\text{N}_2\text{D}^+]/[\text{N}_2\text{H}^+]$, is known to rise in cold ($T < 20\ \text{K}$), dense ($n_{\text{H}} \gtrsim 10^5\ \text{cm}^{-3}$) conditions, especially when CO molecules are largely frozen-out onto dust grain ice mantles and the ortho-to-para ratio of H_2 drops to low values (e.g., Kong et al. 2015). N_2D^+ is known to be a good tracer of low-mass starless cores that are on the verge of collapse, i.e., pre-stellar cores (Caselli & Ceccarelli 2012), as well as early-stage low-mass Class 0 sources (Emprechtinger et al. 2009).

We carried out a pilot search of 4 IRDC regions (C1, F1, F2, G2) with ALMA in Cycle 0 (compact configuration, $2''/3$ resolution), identifying 6 $\text{N}_2\text{D}^+(3-2)$ cores (Tan et al. 2013, hereafter T13) by projection of their 3σ $l-b-v$ space $\text{N}_2\text{D}^+(3-2)$ contours. The two most massive cores were in the IRDC G028.37+00.07 (hereafter Cloud C) C1 region: we refer to these as C1 north and south (C1-N, C1-S). We estimated masses in two ways: (1) from the MIREX map, finding that C1-N has $61 \pm 30\ M_{\odot}$ and C1-S has $59 \pm 30\ M_{\odot}$, with $\sim 50\%$ systematic uncertainty due to distance ($5 \pm 1\ \text{kpc}$) and dust opacity ($\sim 30\%$) uncertainties; (2) from millimeter dust continuum emission, finding that C1-N has $16^{+33}_{-27}\ M_{\odot}$ and C1-S has $63^{+129}_{-27}\ M_{\odot}$, with uncertainties mostly due to the adopted dust temperature of $T = 10 \pm 3\ \text{K}$, together with distance and dust emissivity uncertainties.

These ALMA observations resolve the cores with about three beam diameters. C1-S appears as quite round, centrally concentrated, and monolithic, while C1-N shows evidence of multiple fragments. Given their high levels of deuteration (Kong et al. 2016) and their dark appearances in *Herschel*-PACS images, even at wavelengths as long as $100\ \mu\text{m}$, C1-S, and perhaps also C1-N, are among the best known candidates of massive starless or early-stage cores.

However, we note that Wang et al. (2006) reported a water maser detection in this area (just outside C1-S's lowest

$\text{N}_2\text{D}^+(3-2)$ contour), though at a different velocity (59.5 km s^{-1}) and in a single channel (0.66 km s^{-1} wide). This water maser was not detected in the more sensitive observations of Chambers et al. (2009) and Wang et al. (2012). We also note that Pon et al. (2015) have detected CO(8-7) and (9-8) emission toward C1-N and S with *Herschel*-HIFI ($\sim 20''$ resolution) and argue that this emission results from turbulence dissipating in low-velocity shocks, which could be either driven by large-scale turbulent motions from the surrounding cloud or from protostellar outflow activity.

Here we search for potential protostars and outflow activity via $^{12}\text{CO}(2-1)$ and other tracers using an ALMA Cycle 2 observation. Below we describe the observations (Section 2), present our results (Section 3), and discuss their implications (Section 4).

2. OBSERVATIONS

We use data from our ALMA Cycle 2 project (2013.1.00248.S, PI: Tan), which observed the C1 region in a compact configuration on 2015 April 05, yielding sensitivity to scales from $\sim 10''$ to $\sim 1''$. The position of the field center was R.A. = 18:42:46.5856, decl. = -04:04:12.361 (FK5 J2000 system) ($l = 28.3230$, $b = +0.06750$). It was chosen to be between C1-N and C1-S, slightly closer to C1-S. Thus both cores are within the $27''$ field of view.

The spectral set-up included a continuum band centered at 231 GHz with a width 1.875 GHz, i.e., from 230.0625 to 231.9375 GHz. The achieved sensitivity was 0.045 mJy per $1''.51 \times 0''.84$ beam. In this continuum band, each channel has a width 1.129 MHz, i.e., a velocity resolution 1.465 km s^{-1} . The $^{12}\text{CO}(2-1)$ line frequency is 230.538 GHz. C1-S's radial velocity from its $\text{N}_2\text{D}^+(3-2)$ emission is $+79.40 \pm 0.01 \text{ km s}^{-1}$, with a one-dimensional (1D) dispersion of 0.365 km s^{-1} (i.e., FWHM = 0.860 km s^{-1}), so the sky frequency of $^{12}\text{CO}(2-1)$ from this source is 230.477 GHz. Thus we are sensitive to the presence of $^{12}\text{CO}(2-1)$ emission with moderate velocity resolution. Note that ambient CO molecules in the core and even the wider scale IRDC are expected to be largely frozen-out onto dust grain ice mantles (Hernandez et al. 2011). However, $^{12}\text{CO}(2-1)$ emission near ambient velocities is still likely to be very optically thick from this region. In our data the nine channels closest to the expected ambient velocity of C1-S have velocities centered at: 74.1, 75.4, 76.7, 78.0, 79.3, 80.6, 81.9, 83.2, 84.5 km s^{-1} .

The other basebands were tuned to observe $\text{N}_2\text{D}^+(3-2)$, $\text{C}^{18}\text{O}(2-1)$, DCN(3-2), $\text{DCO}^+(3-2)$, $\text{SiO}(\nu = 1)(5-4)$ and $\text{CH}_3\text{OH } \nu_t = 0 \ 5(1, 4) - 4(2, 2)$. These data will be presented and analyzed in full in a future paper, while in this Letter we focus mostly on the results of the broad continuum band and its $^{12}\text{CO}(2-1)$ line, along with some results from $\text{C}^{18}\text{O}(2-1)$ and DCN(3-2), that help probe denser gas of protostellar cores.

3. RESULTS

Figure 1(a) presents the MIREX Σ map (Butler et al. 2014) of the C1 region, together with contours of $\text{N}_2\text{D}^+(3-2)$ integrated intensity (T13), which define C1-S and C1-N. Also shown are the locations of potential protostars defined by 1.3 mm dust continuum peaks detected in the ALMA Cycle 2 image (Figure 1(b)), which has been cleaned with natural weighting and had a primary beam correction applied.

Figure 1(b) also shows integrated intensities of (continuum subtracted) $^{12}\text{CO}(2-1)$ with red contours tracing $v_{\text{LSR}} = 85.8$ to 124.8 km s^{-1} , i.e., redshifted velocities up to 45.4 km s^{-1} with respect to C1-S's central velocity, and blue contours tracing from 33.8 to 72.8 km s^{-1} , i.e., blueshifted velocities up to 45.6 km s^{-1} from C1-S. Figure 1(c) shows just high-velocity outflow gas that is $> 20 \text{ km s}^{-1}$ away from C1-S's ambient velocity. Figure 1(d) shows the integrated intensity of "ambient" $^{12}\text{CO}(2-1)$, i.e., from 75.4 to 83.2 km s^{-1} . Figures 1(e) and (f) show integrated intensities of $\text{C}^{18}\text{O}(2-1)$ and DCN(3-2), respectively, both potentially helpful for identifying dense gas associated with protostellar cores (e.g., Parise et al. 2009), and thus their radial velocities.

From these images we clearly identify two protostellar sources that are spatially overlapped with the C1-S core (Table 1). We refer to the more central source as C1-Sa and define its spatial position as the location of the 1.3 mm continuum peak. This peak, with a flux density of 11 mJy/beam, is $1''.31$ from C1-S's center as defined by $\text{N}_2\text{D}^+(3-2)$ (T13). Recall that for an estimated (kinematic) source distance of $5 \pm 1 \text{ kpc}$, $1''$ corresponds to 5000 au, i.e., 0.024 pc, with $\sim 20\%$ uncertainties. So the spatial location of C1-Sa is quite close to center of the C1-S core, which has a radius of $3''.61$ (i.e., 0.0875 pc; 18,000 au).

Figures 1(g), (h), and (i) show the spectra of $^{12}\text{CO}(2-1)$, $\text{C}^{18}\text{O}(2-1)$ and DCN(3-2), also in comparison with the T13 observation of $\text{N}_2\text{D}^+(3-2)$, toward the protostars. We estimate the radial velocity of the C1-Sa protostar from the $\text{C}^{18}\text{O}(2-1)$ and DCN(3-2) spectra toward the continuum peak. The $\text{C}^{18}\text{O}(2-1)$ spectrum shows a main peak at $+79.01 \pm 0.12 \text{ km s}^{-1}$, while DCN(3-2) shows a single peak at $+79.8 \pm 0.2 \text{ km s}^{-1}$. Thus it seems very likely that C1-Sa is forming inside the C1-S $\text{N}_2\text{D}^+(3-2)$ core, which has a mean velocity of $+79.4 \text{ km s}^{-1}$ and an FWHM of 0.86 km s^{-1} .

The Galactic coordinate frame position angle of C1-Sa's $^{12}\text{CO}(2-1)$ outflow axis, which we define to the blueshifted axis, is $\simeq 155^\circ$. The outflow is highly collimated and is seen to extend $\sim 12''$ (60,000 au, 0.3 pc), and is quite symmetric, i.e., the P.A. of the redshifted lobe is almost 180° greater than that of the blueshifted lobe. Also, the observed extent of the outflow is similar in each direction, although the highest velocity flow is more extended on the redshifted side. The outflow can be traced down to 3σ above the noise level, without bunching of the contours, so we expect the observed extent to simply be due to observational sensitivity and the actual extent could be much larger.

We re-checked our ALMA Cycle 0 data, which included a requested bandpass set to an intermediate frequency between DCN(3-2) and $\text{SiO}(\nu = 0)(5-4)$. However, this frequency was later mistakenly shifted to be closer to DCN(3-2), causing the SiO line center to be unobserved: only the potential blue wing up to $v_{\text{LSR}} = +69 \text{ km s}^{-1}$ was observed. For this reason, T13 did not report the detection of any SiO emission toward C1-S. However, now we do see indications of the blue wing of SiO ($\nu = 0)(5-4)$ overlapping with the central part of the blue lobe of the $^{12}\text{CO}(2-1)$ outflow and extending to $v_{\text{LSR}} \simeq +50 \text{ km s}^{-1}$. We conclude that it is likely that the outflow from C1-Sa also emits strongly in $\text{SiO}(\nu = 0)(5-4)$ across its full velocity range.

The second source, C1-Sb, has a much weaker 1.3 mm continuum flux of 1.3 mJy/beam and is located $2''.0$ from the center of the C1-S $\text{N}_2\text{D}^+(3-2)$ core and $3''.3$ from C1-Sa. The

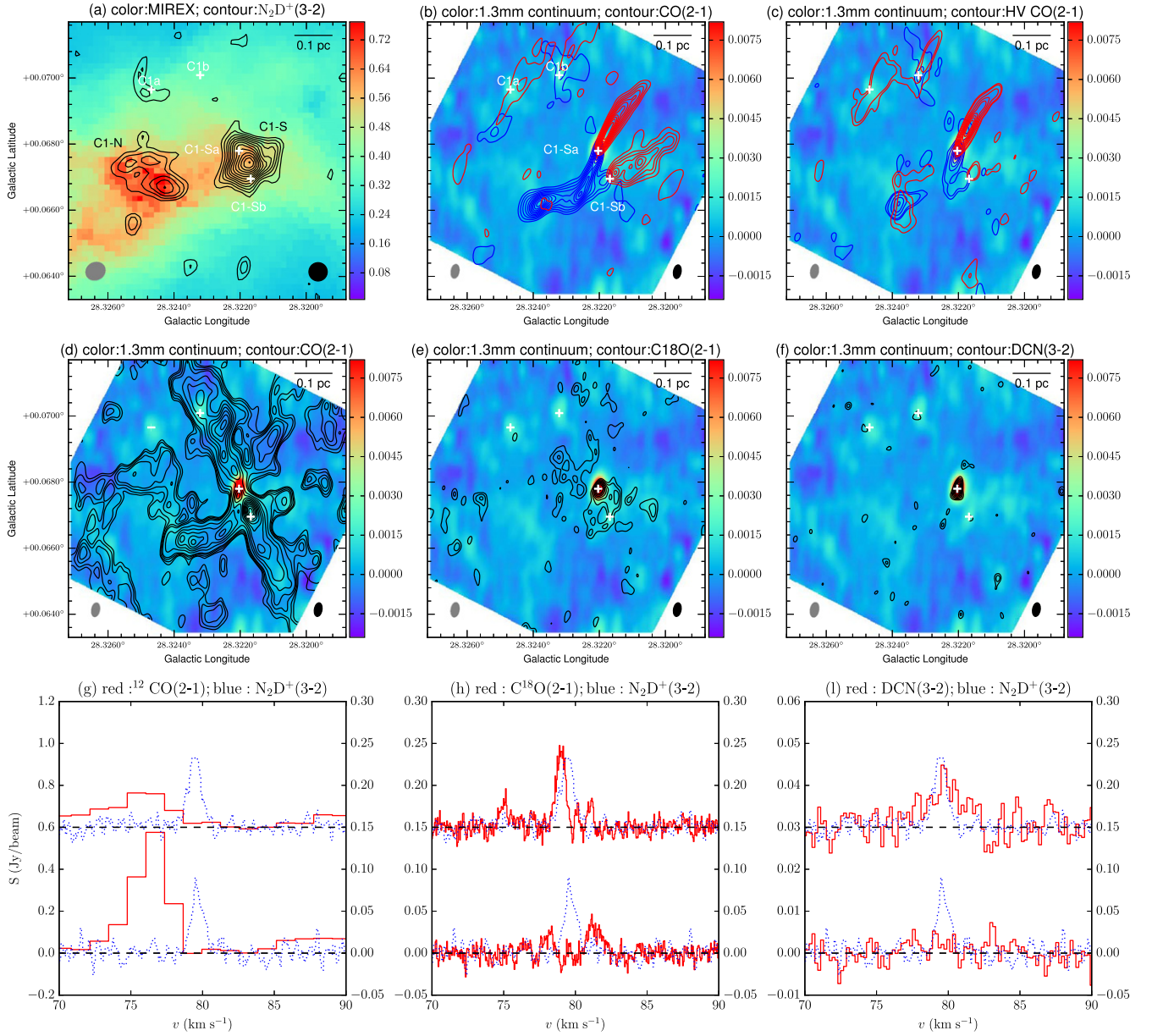


Figure 1. (a) Top left: MIREX Σ map (in g cm^{-2}) of the C1 region, also showing $\text{N}_2\text{D}^+(3-2)$ emission (black contours) observed with ALMA (Cycle 0) (T13) (the ALMA beam is the gray ellipse in the lower left panel; the *Spitzer* beam that sets the MIREX map resolution is in the lower right panel). $\text{N}_2\text{D}^+(3-2)$ contours are shown from 2, 3, 4 σ ..., with $1\sigma \simeq 10 \text{ mJy beam}^{-1} \text{ km s}^{-1}$. The C1-S core is prominent to the center-right, while the more fragmented C1-N core is to the center-left. Protostar candidates (“+” symbols) are based on millimeter-continuum peaks (panel b). (b) Top middle: 231 GHz continuum emission (color scale in Jy/beam) from ALMA Cycle 2 observation with $1''/2$ beam shown in the lower left panel. Red contours show the integrated intensity of $^{12}\text{CO}(2-1)$ from $v_{\text{LSR}} = 85.8$ to 124.8 km s^{-1} ; blue contours show emission from 33.8 to 72.8 km s^{-1} . Contour levels start from 30, 60, 90 σ ..., where $\sigma = 11.6 \text{ mJy beam}^{-1} \text{ km s}^{-1}$. The beam at the line frequency is to the lower right. (c) Top right: same as (b), but now only showing high-velocity (HV) gas that is $+20 \text{ km s}^{-1}$ or greater (red contours) and -20 km s^{-1} or less (blue contours) than v_{LSR} of C1-S. The contour levels are shown from 20, 30, 60, 90 σ ..., where $\sigma = 9.3 \text{ mJy beam}^{-1} \text{ km s}^{-1}$. (d) Center left: same as (b), but now contours show “ambient” $^{12}\text{CO}(2-1)$ integrated intensity over velocities 75.4 – 83.2 km s^{-1} , ranging from 3, 10, 20, 30, 60, 90 σ ..., with $1\sigma = 5.9 \text{ mJy beam}^{-1} \text{ km s}^{-1}$. (e) Center: same as (b), but now the contours show the $\text{C}^{18}\text{O}(2-1)$ integrated intensity over velocities 74 to 84 km s^{-1} , ranging from 3, 5, 7, 9, 11, 13, 15, 18 σ , with $1\sigma = 8.5 \text{ mJy beam}^{-1} \text{ km s}^{-1}$. (f) Center right: same as (b), but now the contours show the DCN(3-2) integrated intensity over velocities of 77 to 81 km s^{-1} , ranging from 3, 4, 5, 6, 7, 8, 9 σ , with $1\sigma = 3.8 \text{ mJy beam}^{-1} \text{ km s}^{-1}$. (g) Bottom left: spectra of $^{12}\text{CO}(2-1)$ (red solid lines) and $\text{N}_2\text{D}^+(3-2)$ (dotted blue lines) extracted over 1 beam area toward C1-Sa (top, offset up) and C1-Sb (bottom). (h) Bottom middle: same as (g), but showing $\text{C}^{18}\text{O}(2-1)$ and $\text{N}_2\text{D}^+(3-2)$. (i) Bottom right: same as (g), but showing DCN(3-2) and $\text{N}_2\text{D}^+(3-2)$.

$\text{C}^{18}\text{O}(2-1)$ spectrum toward C1-Sb shows a main peak at $+81.36 \pm 0.42 \text{ km s}^{-1}$ and a secondary peak (with about half the equivalent width) at $+78.16 \pm 0.45 \text{ km s}^{-1}$. The DCN(3-2) spectrum shows no particularly strong features, although a 3σ peak is seen in the integrated intensity map (Figure 1(f)). We tentatively assign C1-Sb’s radial velocity to be that of the main $\text{C}^{18}\text{O}(2-1)$ spectral feature. We discuss below that this

assignment is potentially supported by examination of the channel maps of the $^{12}\text{CO}(2-1)$ “ambient” gas. If this radial velocity is correct, then it would suggest that C1-Sb is not physically associated with the C1-S $\text{N}_2\text{D}^+(3-2)$ core, and in fact may be part of a gas structure that is linked to the C1-N core. However, we cannot exclude the possibility that C1-Sb is also forming from the C1-S core.

Table 1
Cores and Protostars in the C1 Region

Name	$l(^{\circ})$	$b(^{\circ})$	$S_{1.3\text{ mm}}$ (mJy)	v_{LSR} (km s $^{-1}$)	P.A.outflow ($^{\circ}$)
C1-N	28.32503	0.06724	6.94 ± 0.72	81.18 ± 0.03^a	...
C1-S	28.32190	0.06745	26.7 ± 0.77	79.40 ± 0.01^a	...
C1-Sa	28.322093	0.067698664	19.5 ± 0.1	79.01 ± 0.12^b	155
C1-Sb	28.321752	0.066847223	2.7 ± 0.1	81.36 ± 0.42^c	113
C1a	28.324765	0.069543149	3.6 ± 0.1	80.2^d	150
C1b	28.323272	0.069987301	3.5 ± 0.1	...	150

Notes.

^a v_{LSR} of C1-N and C1-S estimated from $\text{N}_2\text{D}^+(3-2)$ (T13). v_{LSR} of protostars estimated from strongest $\text{C}^{18}\text{O}(2-1)$ peak, but see the notes below for individual sources.

^b C1-Sa: v_{LSR} estimated from the strongest $\text{C}^{18}\text{O}(2-1)$ peak. Secondary peaks at 75.08 ± 0.05 km s $^{-1}$, 81.08 ± 0.03 km s $^{-1}$, while DCN(3-2) has a single peak at 79.8 ± 0.2 km s $^{-1}$.

^c C1-Sb: v_{LSR} estimated from the strongest $\text{C}^{18}\text{O}(2-1)$ peak. Secondary peak at 78.16 ± 0.05 km s $^{-1}$. DCN(3-2) is too weak to measure v_{LSR} .

^d C1a: v_{LSR} tentatively estimated from weak $\text{N}_2\text{D}^+(3-2)$ emission (T13).

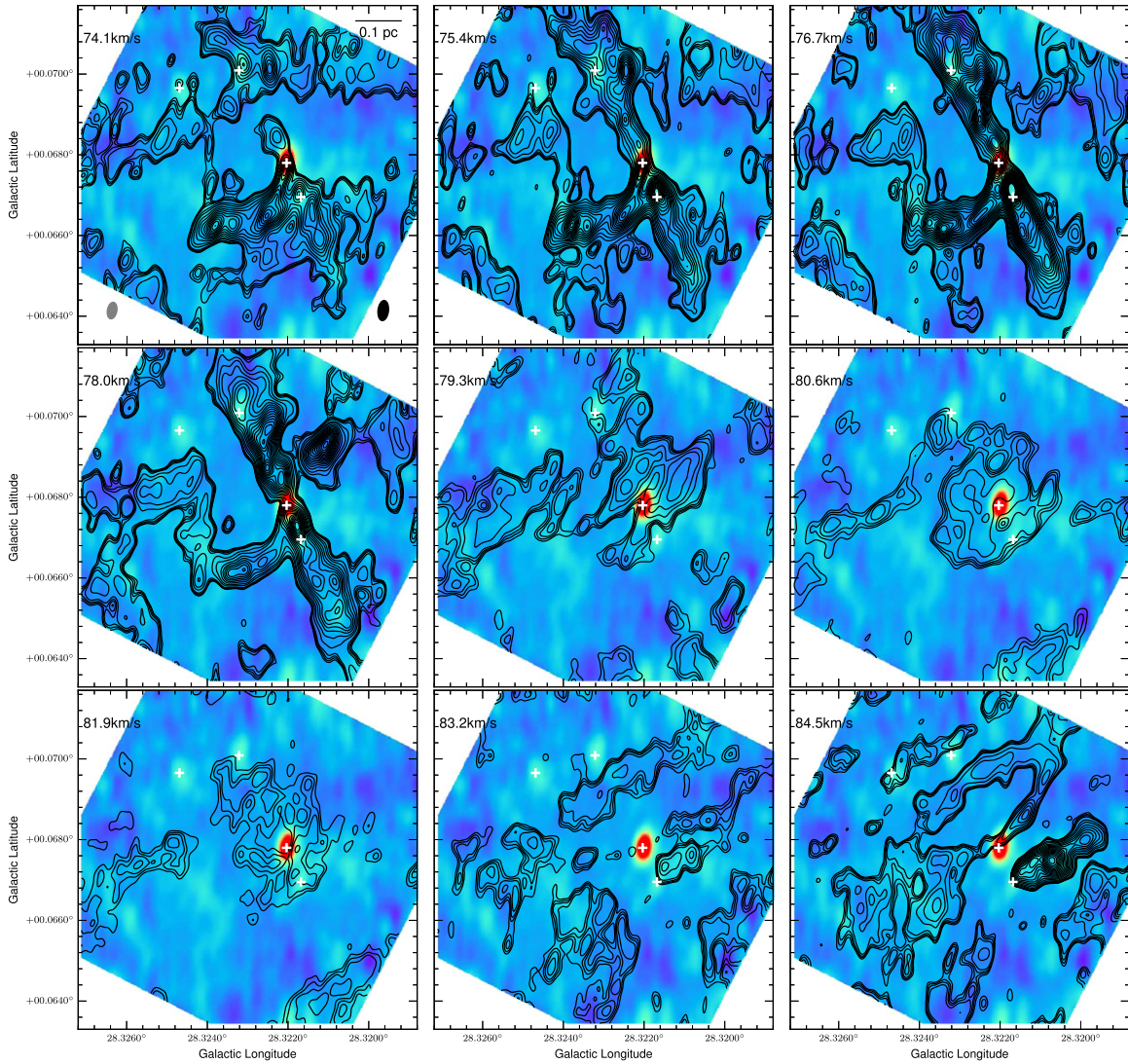


Figure 2. Channel maps of $^{12}\text{CO}(2-1)$ emission, with contours starting from 3, 5, 7, 10, 20, 30, 60, 90 σ ..., where $\sigma = 2.1$ mJy beam $^{-1}$ km s $^{-1}$. The background image shows 231 GHz continuum (Figure 1(b)).

The outflow from C1-Sb has a similar extent as that from C1-Sa and appears to have a wider opening angle. It has a P.A. = 113 $^{\circ}$, and on its blueshifted side the outflow spatially overlaps with that from C1-Sa. Analysis of lower

intensity contours (down to 3 σ) indicates a bunching close to the higher-intensity contours. This may indicate that, unlike for C1-Sa, we are seeing the full extent of C1-Sb's outflow.

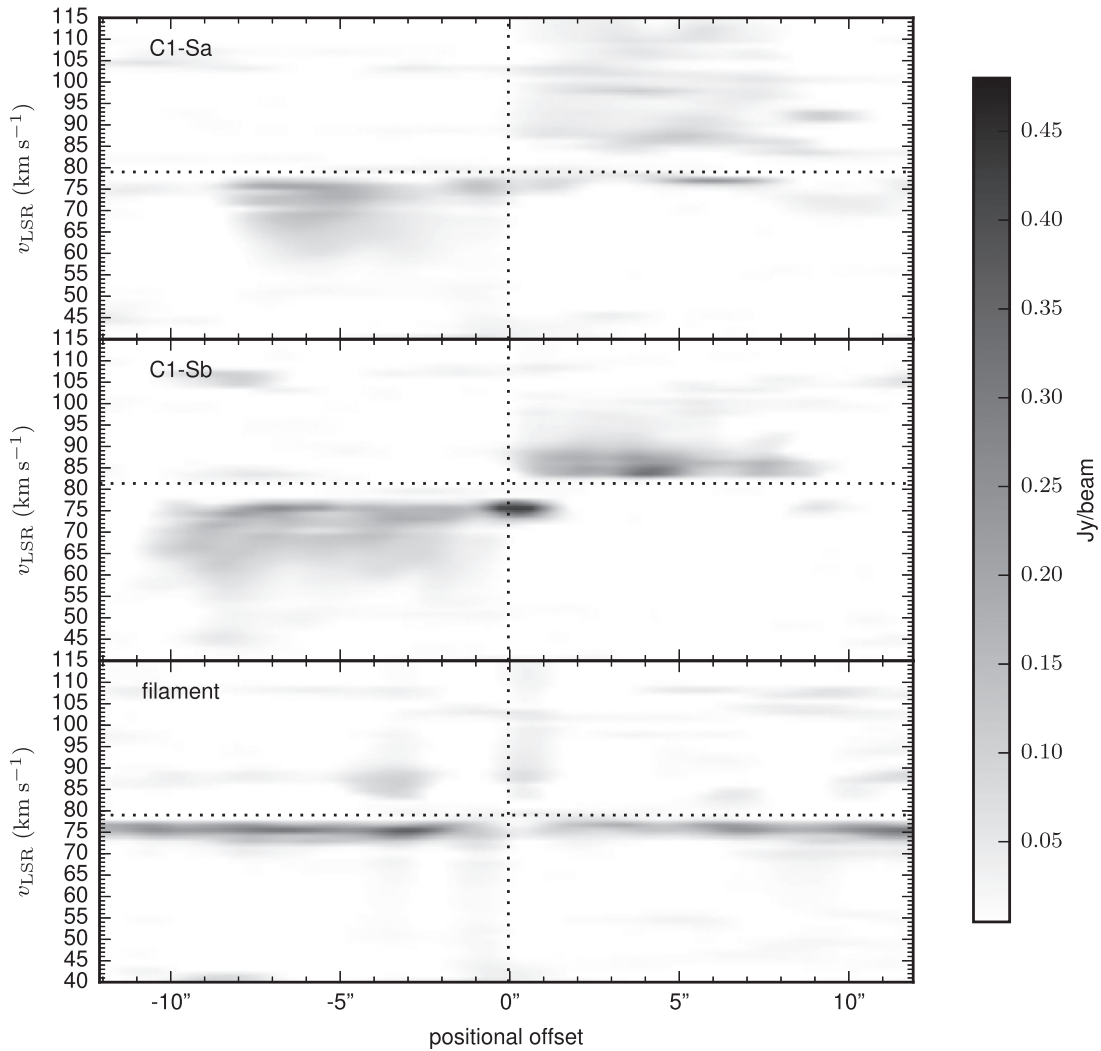


Figure 3. Position–velocity diagram of $^{12}\text{CO}(2-1)$ emission along the axis of the C1-Sa outflow (top), C1-Sb outflow (middle), and the ambient “filament” (bottom). Horizontal lines are included for reference at 79.0 km s^{-1} for C1-Sa/filament and at 81.4 km s^{-1} for C1-Sb.

In Figure 1(b) we identify two more candidate protostars that are away from the C1-S and C1-N cores. C1a is located in a region with faint $\text{N}_2\text{D}^+(3-2)$ emission with a radial velocity of 80.6 km s^{-1} . There is relatively faint $^{12}\text{CO}(2-1)$ emission, which may be driven from this source with P.A. $\sim 130^\circ$. There is no significant $\text{C}^{18}\text{O}(2-1)$ emission associated with this source, and only a weak 3σ feature in the $\text{DCN}(3-2)$ integrated intensity map. C1b is close to C1a and has a similar P.A. as its $^{12}\text{CO}(2-1)$ emission of about 130° . There is no significant $\text{C}^{18}\text{O}(2-1)$ or $\text{N}_2\text{D}^+(3-2)$ emission at C1b, and only a very weak feature in $\text{DCN}(3-2)$. Note that the protostellar nature of C1a and C1b is uncertain, especially since some of the observed $^{12}\text{CO}(2-1)$ features may be affected by side-lobe contamination from C1-Sa and C1-Sb.

Figure 2 shows nine channel maps of “ambient” $^{12}\text{CO}(2-1)$ emission from the C1 region, with high-velocity and low-velocity ends connecting with the “outflow” velocities plotted in Figure 1(b). The C1-Sa outflow lobes are visible in the low-velocity and high-velocity channels. Figure 2 also suggests that the C1-Sb outflow has a driving source with a $v_{\text{LSR}} \simeq +81 \text{ km s}^{-1}$, since the blueshifted lobe is already apparent in the 79.3 km s^{-1} channel, while the redshifted lobe appears to vanish by the 81.9 km s^{-1} channel.

There are several other striking features seen in Figure 2. First, there is a very elongated “filament” that peaks in the 76.7 and 78.0 km s^{-1} channels, but is visible from 74.1 to 79.3 km s^{-1} . The filament overlaps with the C1-Sa, C1-Sb, and C1b protostars and its orientation is almost perpendicular to their outflows. An interpretation of this filament as an ambient gas feature, rather than as a collimated bipolar outflow, is discussed below, considering its position–velocity diagram. Second, there is a relatively weak, but still highly significant, quasi-spherical “core” of gas seen in the 80.6 and 81.9 km s^{-1} channels. Third, there is $^{12}\text{CO}(2-1)$ emission in the vicinity of the C1-N $\text{N}_2\text{D}^+(3-2)$ core. Fourth, there are additional emission features on the periphery of the image.

Figure 3 shows position–velocity diagrams of $^{12}\text{CO}(2-1)$ emission along the outflow axes of C1-Sa and C1-Sb and along the axis of the “ambient filament.” These are defined by rectangular regions of $3''$ wide running length-wise along the P.A. of each outflow or the filament.

The ambient nature of the filament is readily apparent in these figures. It has a very narrow velocity dispersion, which is comparable to other ambient gas tracers, like $\text{C}^{18}\text{O}(2-1)$ and $\text{N}_2\text{D}^+(3-2)$, and does not show a significant gradient in radial velocity. The lack of such a gradient could be explained by an

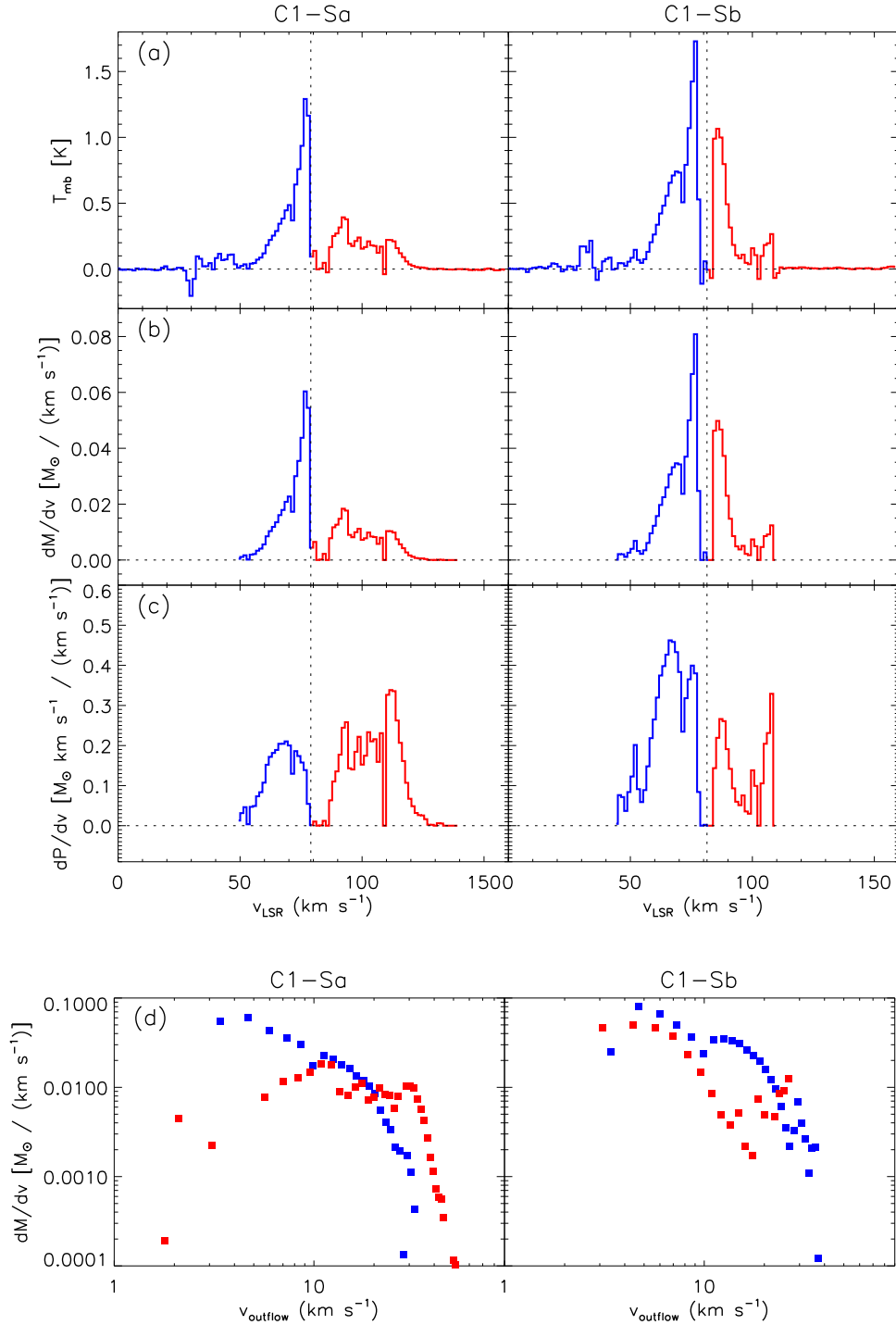


Figure 4. (a) Top row: $^{12}\text{CO}(2-1)$ spectra of C1-Sa (left) and C1-Sb (right), each extracted from a $24'' \times 3''$ aperture centered on the protostars and aligned along the outflow axes. The estimated ambient velocities of the protostars are shown with dotted lines. (b) Second row: distribution of outflow mass (see the text). (c) Third row: distribution of outflow momentum (see the text). (d) Bottom row: comparison of mass distributions of blueshifted and redshifted outflows vs. outflow velocity, v_{outflow} , i.e., relative to the ambient velocity of each protostar.

outflow that was precisely aligned in the plane of the sky, but this would still be expected to have a relatively broad velocity dispersion, which the filament feature does not exhibit.

Figure 4 shows $^{12}\text{CO}(2-1)$ spectra and derived mass and momentum distributions of the C1-Sa and C1-Sb outflows, extracted from the same regions used for Figure 3. To derive the mass, one must assume an excitation temperature, with values of $T_{\text{ex}} \simeq 10\text{--}50\text{ K}$ typically being used. $T_{\text{ex}} = 17\text{ K}$

minimizes the mass estimate, while 50 K increases this by a factor of 1.5.

For C1-Sa, integrating from $50 < (v_{\text{LSR}}/\text{km s}^{-1}) < 140$ and assuming $T_{\text{ex}} = 17\text{ K}$, the blue/red lobes have masses $m_w = 0.50/0.32 M_{\odot}$ and momenta $p_w = 3.6/6.6 M_{\odot} \text{ km s}^{-1}$, respectively. Similarly, for C1-Sb, integrating from $45 < (v_{\text{LSR}}/\text{km s}^{-1}) < 110$ the blue/red lobes have masses $0.73/0.38 M_{\odot}$ and momenta $8.5/3.5 M_{\odot} \text{ km s}^{-1}$, respectively.

Note that the estimates for each blueshifted lobe are affected by the overlap of the sources, leading to a modest overestimation of their properties, especially relative to the redshifted lobes. However, in general the absolute values of the above estimates should be viewed as lower limits, not only because of the choice of $T_{\text{ex}} = 17$ K, but also because of inclination effects (which boost momentum estimates by $1/\cos(i)$, where i is the inclination of the outflow axis to the line of sight, with a random expectation value of $i = 60^\circ$), and because of optical depth effects (both within the outflowing gas, which may boost momentum by factors of ~ 6 (Zhang et al. 2016), and due to foreground absorption).

For C1-Sa, the mass-weighted mean velocities ($v_w = p_w/m_w$) of the blue/red lobes are $v_w = 7.3/21 \text{ km s}^{-1}$, while for C1-Sb they are $12/9.2 \text{ km s}^{-1}$. Assuming that the lengths of all the outflows lobes are $\simeq 12''$ (60,000 au), the dynamical times for the blue/red lobes of C1-Sa are $t_w = (3.9/1.3) \times 10^4 \text{ year}$ and for C1-Sb are $(2.4/3.1) \times 10^4 \text{ year}$. The correction factor for inclination is $\cos(i)/\sin(i)$, i.e., 0.58 for $i = 60^\circ$. The mass outflow rates for the C1-Sa blue/red lobes are $(1.3/2.5) \times 10^{-5} M_\odot \text{ yr}^{-1}$, and for C1-Sb are $(3.0/1.2) \times 10^{-5} M_\odot \text{ yr}^{-1}$. The correction factor for inclination is $\sin(i)/\cos(i)$, i.e., 1.73 for $i = 60^\circ$. Finally, the momentum injection rates for the C1-Sa blue/red lobes are $\dot{p}_w = (0.9/5.0) \times 10^{-4} M_\odot \text{ km s}^{-1} \text{ yr}^{-1}$ and for the C1-Sb blue/red lobes are $(3.5/1.1) \times 10^{-4} M_\odot \text{ km s}^{-1} \text{ yr}^{-1}$. The inclination correction factor is $\sin(i)/\cos^2 i$, i.e., 3.46 for $i = 60^\circ$. Allowing for a mass underestimation factor of 3 and assuming $i = 60^\circ$, the overall estimates of the total momentum fluxes are boosted by a factor of 10, i.e., totals for C1-Sa of $\dot{p}_w \sim 5.9 \times 10^{-3} M_\odot \text{ km s}^{-1} \text{ yr}^{-1}$ and for C1-Sb of $\sim 4.6 \times 10^{-3} M_\odot \text{ km s}^{-1} \text{ yr}^{-1}$.

4. DISCUSSION

Outflow momentum flux is expected to be the most reliable direct tracer of protostellar properties, since it should be independent of the effects of ambient environment (unlike mass flux, which depends mostly on the mass that has been swept-up by the primary outflow). Models of massive protostar formation (Zhang et al. 2014) based on the Turbulent Core Model (MT03) for cores of $60 M_\odot$ in a clump environment with $\Sigma_{\text{cl}} \simeq 0.4 \text{ g cm}^{-2}$ (relevant for C1-S) predict $\dot{p}_w \lesssim 6 \times 10^{-3} M_\odot \text{ km s}^{-1} \text{ yr}^{-1}$, when the protostellar mass is $m_* \lesssim 3 M_\odot$, rising to $\sim 10^{-2} M_\odot \text{ km s}^{-1} \text{ yr}^{-1}$ by the time $m_* \simeq 10 M_\odot$. It is interesting that these estimates are comparable with the observed values of $\dot{p}_w$ from C1-Sa and C1-Sb. This suggests that if C1-Sa is a massive protostar in the process of formation, it is currently at a very early stage, i.e., it has yet to accrete most of its mass. This would be broadly consistent with the protostar having a relatively low luminosity such that it does not appear yet as an MIR source. High angular resolution, high-sensitivity MIR to FIR observations, e.g., with *JWST*, are needed to measure the SED of the protostar, which can then also constrain protostellar models. We conclude that we have detected protostars of relatively low current masses. C1-Sa appears to be embedded within the C1-S massive, cold core, as defined by $\text{N}_2\text{D}^+(3-2)$ emission. It thus has a large mass reservoir from which to continue to grow: we speculate that it is destined to become a massive star.

As traced by $^{12}\text{CO}(2-1)$, C1-Sa's bipolar outflow is highly collimated and has velocities extending to $\sim 50 \text{ km s}^{-1}$. Similar

(blueshifted) velocities are seen in $\text{SiO}(v=0)(5-4)$ emission. Using mass-weighted mean velocities, which are $\sim 10 \text{ km s}^{-1}$, the $i = 60^\circ$ inclination-corrected outflow timescale is $\sim 2 \times 10^4 \text{ year}$. However, since C1-Sa's outflow is likely to extend to larger distances than what we observe, this is probably a lower limit on the duration of protostellar activity. Given the symmetric and linear morphology of the outflow lobes, it appears that C1-Sa (and C1-Sb) have not suffered significant dynamical disturbance from other nearby (proto) stars during the period they have been driving these outflows. This is consistent with the assumptions of core accretion models and is a constraint on competitive accretion models.

For constant instantaneous star formation efficiency from the core, ϵ_c , the fiducial turbulent core model predicts $m_* = \epsilon_c M_c (t/t_{*f})^2$, where the total star formation time is $t_{*f} \rightarrow 1.29 \times 10^5 (M_c/60 M_\odot)^{1/4} (\Sigma_{\text{cl}}/1 \text{ g cm}^{-2})^{-3/4} \text{ year}$. So for the $M_c = 60 M_\odot$ and $\Sigma_{\text{cl}} = 0.4 \text{ g cm}^{-2}$ case, then $t > 2 \times 10^4 \text{ year}$ implies $m_* > 0.36 M_\odot$ (assuming $\epsilon_c \simeq 1$, expected during early stages when outflow cavity opening angles are small) and $\dot{p}_w \gtrsim 1 \times 10^{-4} M_\odot \text{ km s}^{-1} \text{ yr}^{-1}$.

We conclude that C1-Sa is a good candidate for an early-stage massive protostar and as such it shows that these early phases of massive star formation can involve highly ordered outflow, and thus accretion, processes (see also Zhang et al. 2015). The massive C1-S core is potentially also forming a second protostar, C1-Sb: improved estimates of its radial velocity are needed to clarify its association with this core. The results presented complement work that has shown that collimated outflows can be launched from later-stage, more luminous massive protostars (e.g., Beuther et al. 2002). They also illustrate that there are similarities between high-mass protostars and their lower-mass cousins forming from lower-mass cores.

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